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# On the distance to the North Polar Spur and the local CO-H<sub>2</sub> factor<sup>★,★★</sup>

R. Lallement<sup>1</sup>, S. Snowden<sup>2</sup>, K. D. Kuntz<sup>3</sup>, T. M. Dame<sup>4</sup>, D. Koutroumpa<sup>5</sup>, I. Grenier<sup>6</sup>, and J. M. Casandjian<sup>6</sup>

<sup>1</sup> GEPI/Observatoire de Paris, 5 Place Jules Janssen, 92195 Meudon, France  
e-mail: rosine.lallement@obspm.fr

<sup>2</sup> NASA/GSFC, Greenbelt, MD 20771, USA

<sup>3</sup> Henry A. Rowland Department of Physics & Astronomy, Baltimore, MD 21218, USA

<sup>4</sup> Harvard-Smithsonian Center for Astrophysics, 60 Garden Street, Cambridge, MA 02138, USA

<sup>5</sup> LATMOS-IPSL, Université de Versailles Saint Quentin, INSU/CNRS, 11 Bd D'Alembert, 78280 Guyancourt, France

<sup>6</sup> AIM, Université Paris Diderot & CEA Saclay DSM/Irfu/SaP, 91191 Gif sur Yvette, France

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## ABSTRACT

**Aims.** Most models identify the X-ray bright North Polar Spur (NPS) with a hot interstellar (IS) bubble in the Sco-Cen star-forming region at  $\approx 130$  pc. An opposite view considers the NPS as a distant structure associated with Galactic nuclear outflows. Constraints on the NPS distance can be obtained by comparing the foreground IS gas column inferred from X-ray absorption to the distribution of gas and dust along the line of sight. Absorbing columns toward shadowing molecular clouds simultaneously constrain the CO-H<sub>2</sub> conversion factor.

**Methods.** We derived the columns of X-ray absorbing matter  $N_{\text{Habs}}$  from spectral fitting of dedicated *XMM-Newton* observations toward the NPS southern terminus ( $l^{\text{II}} \approx 29^\circ$ ,  $b^{\text{II}} \approx +5$  to  $+11^\circ$ ). The distribution of the IS matter was obtained from absorption lines in new stellar spectra, 3D dust maps, and emission data, including high spatial resolution CO measurements recorded for this purpose.

**Results.**  $N_{\text{Habs}}$  varies from  $\approx 4.3$  to  $\approx 1.3 \times 10^{21}$  cm<sup>-2</sup> along the 19 fields. Relationships between X-ray brightness, absorbing column, and hardness ratio demonstrate a brightness increase with latitude that is governed by increasing absorption. The comparison with absorption data and local and large-scale dust maps rules out an NPS source near-side closer than 300 pc. The correlation between  $N_{\text{Habs}}$  and the reddening increases with the sightline length from 300 pc to 4 kpc and is the tightest with *Planck*  $\tau_{353\text{GHz}}$ -based reddening, suggesting a much larger distance.  $N(\text{H})/E(B-V)_r \approx 4.1 \times 10^{21}$  cm<sup>-2</sup> mag<sup>-1</sup>, close to *Fermi-Planck* determinations.  $N_{\text{Habs}}$  absolute values are compatible with HI-CO clouds at  $-5 \leq V_{\text{LSR}} \leq +25$  to  $+45$  km s<sup>-1</sup> and an NPS potentially far beyond the Local Arm. A shadow cast by a  $b = +9^\circ$  molecular cloud constrains  $X_{\text{CO}}$  in that direction to  $\leq 1.0 \times 10^{20}$  cm<sup>-2</sup> K<sup>-1</sup> km<sup>-1</sup> s. The average  $X_{\text{CO}}$  over the fields is  $\leq 0.75 \times 10^{20}$  cm<sup>-2</sup> K<sup>-1</sup> km<sup>-1</sup> s.

**Key words.** X-rays: ISM – radio lines: ISM – local interstellar matter – ISM: bubbles – dust, extinction – Galaxy: center

## 1. Introduction

The North Polar Spur (NPS) is one of the best-known features in radio continuum and diffuse soft X-ray background maps. It is seen as a  $\approx 15^\circ$  wide arc that runs with varying intensity from  $l, b \sim 25^\circ, 20^\circ$  to  $330^\circ, 75^\circ$  (e.g., Hanbury Brown et al. 1960; Bowyer et al. 1968). As surveys of the diffuse background expanded, it was seen to be one of the most prominent features of the entire sky, although it was joined by a wide region of diffuse emission in the general direction of the Galactic center both above and below the Galactic plane, the X-ray bulge (Snowden et al. 1995, 1997). Radio Loop I (Berkhuijsen et al. 1971; Haslam et al. 1982) lies next to the NPS, and appears to bound the NPS. Both are also seen in polarized radio emission and in total-intensity and polarized microwave emission (Sofue & Reich 1979; Sun et al. 2014; Vidal et al. 2015; Planck Collaboration XXV 2016). Filaments and arcs seen in HI

and extending up to  $+85^\circ$  northern Galactic latitude are also spatially associated with Loop I and the NPS (Colomb et al. 1980; Kalberla et al. 2005). The NPS/Loop I is detected at GeV energies (Casandjian et al. 2009; Ackermann et al. 2014), presumably due to inverse-Compton scattering of starlight by the energetic electrons, combined with pion decay emission from the cold border.

Using Loop I to outline a small circle on the sky, it was then assumed that the NPS was the limb-brightened edge of a superbubble with a radius of  $\approx 100$  pc centered on the Sco-Cen OB association at  $\sim 130$  pc, with the Sco-Cen OB association easily creating and powering the superbubble with both stellar winds and supernovae (de Geus 1992; Egger & Aschenbach 1995). The size and high-latitude extent of Loop I strongly suggests its proximity, and indeed measured distances to the HI arcs, either from stellar light polarization (Heiles 2000) or from absorption studies (Puspitarini & Lallement 2012), are on the order of 100 pc. The HI shells are thought to be shock-compressed interstellar medium (ISM) at the periphery of the Loop I/NPS expanding structure. Sophisticated models of time-dependent evolution of the ISM under the action of winds and supernovae are able to reproduce the Loop I structure and its interaction

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\*\* Based on data obtained using the telescope *Bernard Lyot* at Observatoire du Pic du Midi, CNRS and Université Paul Sabatier, France.

with the cavity surrounding the Sun (the Local Bubble, or LB), and most of the observations (de Avillez & Breitschwerdt 2005; Breitschwerdt & de Avillez 2006).

A different interpretation of the NPS enhancement has been defended over the years (Sofue et al. 1974; Sofue 1994, 2000), based on several arguments on the geometry and difficulties in adjusting models of supernova remnants (SNR) to the radio continuum, X-ray, and HI measurements. According to the author, NPS/Loop I better traces a shock front propagating through the Galactic halo that is assumed to have originated from an intense explosion and/or a starburst at the Galactic center, of energy  $3 \times 10^{56}$  ergs and about 15 million years ago. While the shock can mimic the radio and X-ray NPS, the post-shock high-temperature gas may also explain the observed X-ray bulge around the Galactic center. More recently, Bland-Hawthorn & Cohen (2003) uncovered a 200 pc wide bipolar structure at the Galactic center at mid-infrared wavelengths that is most likely associated with a bursting episode. Interestingly, they also showed that a large-scale structure extrapolated from the central region and extending up to the halo could be seen from the Sun as a loop although it is open ended, provided it is wider than the solar galactocentric radius. The authors suggested that the Snowden et al. (1997) X-ray bulge observed at 0.5–2.0 keV reveals this bipolar structure. Finally, the existence of Galactic nuclear activity and associated large-scale structures has been spectacularly demonstrated with the discovery of the gigantic  $\gamma$ -ray bubbles in the high-energy *Fermi*-LAT data (Su et al. 2010). The *Fermi* bubbles (FBs) and concentric structures at their feet in the Galactic plane are also seen in total-intensity or polarized radio and microwave emission (Carretti et al. 2013), their inner parts are close to the X-ray bulge (see Fig. 6 from Casandjian et al. 2015), and their edges seem to parallel bright arcs in the 1.5–2 keV ROSAT maps (see Fig. 20 from Su et al. 2010).

Following the discovery, several models have been proposed for the FBs: one class of scenarios considers a recent, short outburst activity of the type of active galactic nuclei (AGN), while other models consider less energetic, long-duration Galactic wind models maintained by supernovae. In the former (latter) case the  $\gamma$ -ray emission is mainly of leptonic (hadronic) origin (see, e.g., Crocker et al. 2015; and Sarkar et al. 2015, for further discussion and description of their analytical and numerical Galactic wind models). Outside the bubbles, Su et al. (2010) also identified larger gamma-ray structures, the so-called inner arc that seems to border the low-latitude portion of the northern bubble at Galactic longitude  $+20^\circ$ , and the outer arc at  $+25^\circ$ ;  $+30^\circ$ . Both are seen in polarized microwave emission, and the inner arc is clearly considered as a Galactic center feature. The outer arc, which is very similar to the inner arc, seems to run parallel to the low-latitude part of the NPS (see Fig. 2 of Su et al. 2010).

Based on these geometrical arguments, Su et al. (2010) considered the possibility that Loop I and the northern outer arc are parts of the relics of previous bubbles. On the other hand, while there are other X-ray arcs in the vicinity of the Galactic center, there are no structures similar to the NPS either in the northern Galactic hemisphere on the other side of the FB or at all in the south. This would require a rather asymmetric origin mechanism.

The implications of the two scenarios in terms of Galactic nuclear activity, its temporal variability, and outflow interaction with the halo are drastically different, and today the NPS distance and the link between NPS/Loop I and the Galactic center are still a matter of debate. Planck Collaboration XXV (2016)

have extensively discussed the NPS-Loop I characteristics and origin based on *Planck*-WMAP polarization and listed the various constraints on their distance. As noted by the authors, new evidence for a distance greater than the traditional value of 130 pc has been published. Wolleben (2007) showed that the radio emission from Loop I is strongly depolarized below  $30^\circ$  latitude and attributed this depolarization to fluctuations in the foreground Faraday depth. The required path length beyond the Local Bubble is above 70 pc, which places the front face of the loop beyond the Sco-Cen association. More recently, Puspitarini et al. (2014) failed to identify a large cavity in 3D maps of the nearby interstellar (IS) dust (Lallement et al. 2014) that might be filled with hot gas and produce the NPS emission, in contrast to other X-ray enhancements that have cavity counterparts. Instead, the authors suggested that the near side of the NPS source, that is, the closer bubble boundary if the NPS source is a nearly spherical region, is located beyond 200 pc. Sofue (2015) used the X-ray absorption pattern and inferred that the source is beyond the Aquila Rift main dense part, which is again beyond the Sco-Cen association (but see Sect. 2). On the other hand, Planck Collaboration XXV (2016) concluded based on geometrical arguments and the *Planck* polarization maps that the NPS is not associated with the Galactic center.

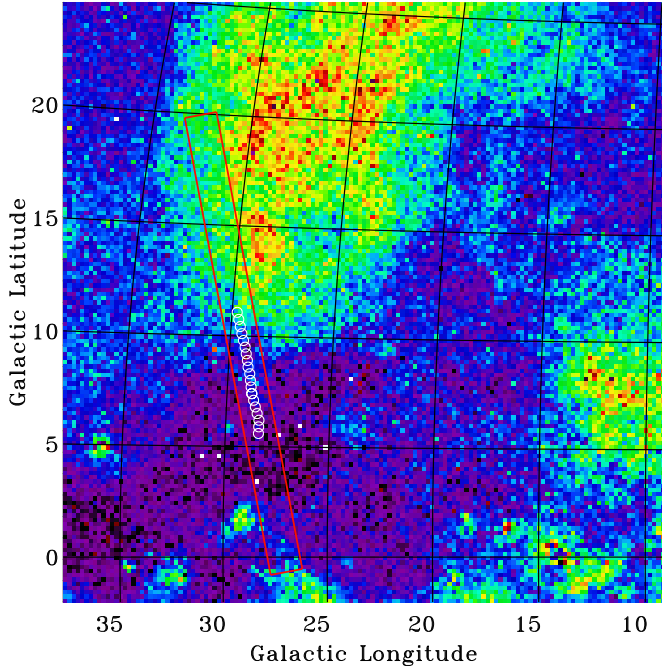
Here we address the question of the NPS distance by means of a detailed comparison between the columns of interstellar gas that are located in front of the NPS source and absorb its X-ray emission, and the distribution with distance of the IS matter derived from stellar absorption data. The former information is obtained by means of spectral fitting of new, dedicated X-ray observations, and the latter from dedicated ground-based spectroscopic data and published dust maps. We complement this study by additional comparisons between X-ray absorption and IS emission data. Dedicated CO measurements at high spatial resolution were performed for this purpose.

In Sect. 2 we describe the new *XMM-Newton* data and their spectral analysis. In Sect. 3 we present the ground-based optical data and compare them and 3D reddening maps with the X-ray absorbing columns deduced from the *XMM* fitting results. In Sect. 4 we present the new CO data and compare X-ray absorbing columns with line-of-sight IS dust and gas based on emission data. In Sect. 5 we summarize the comparisons and draw conclusions.

## 2. X-ray observations of the NPS southern terminus

### 2.1. RASS data

Near the Galactic plane the NPS has an angular width of roughly ten degrees, and at all latitudes it has a low characteristic temperature (on the order of 0.2 keV, Willingale et al. 2003; Miller et al. 2008), hence it produces very little emission above 1.5 keV. ROSAT with its  $2^\circ$  diameter field of view (FOV), large grasp, and strong soft response, and the ROSAT All-Sky Survey (RASS) would seem to be the ideal tools with which to study the structure of the NPS. In what follows we demonstrate that this is not the case and that more spectrally resolved data and higher signal are necessary to draw useful conclusions. Figure 1 shows a RASS map of the southern end of the NPS, the location of our *XMM-Newton* observations, and the location of a more extensive region over which we began a ROSAT study. Figure 2 shows the profiles of the X-ray emission along the  $\approx 20^\circ$  strip shown in Fig. 1 in the ROSAT R4 (0.44–1.01 keV), R5 (0.56–1.21 keV), and R6 (0.73–1.56 keV) bands. From the profile in each band we have subtracted that level of emission seen in the most highly

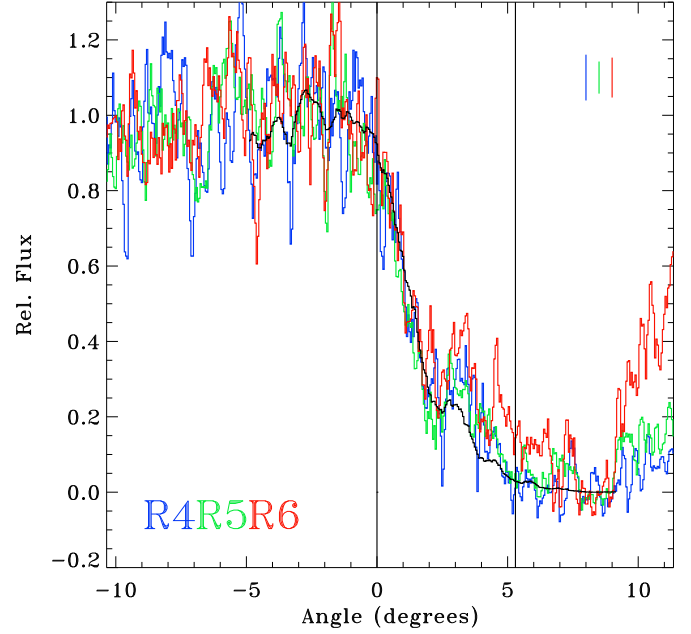


**Fig. 1.** RASS R4+R5+R6 band image of the NPS. The color scale is linear and extends from 0 (black) to  $1.5 \times 10^{-3}$  count  $s^{-1}$  arcmin $^{-2}$  (red). The white circles show the location of the *XMM-Newton* pointings. The red box is the region from which the ROSAT profile shown in the next figure was extracted.

absorbed part of the Galactic plane ( $N(H) > 10^{22}$  cm $^{-2}$ ) and normalized the peak of the NPS emission to unity. There are a number of reasons why this process is naïve, but it demonstrates that within the uncertainties all three bands show essentially the same absorption profile. If the NPS were absorption bounded, we would expect the profile of the higher energy bands to extend to lower latitudes than the profiles of the lower energy bands (as discussed by Sofue 2015). Figure 2 shows that the energy resolution and the count-rate limited angular resolution of the RASS, 12', does not allow us to clearly diagnose whether the edge of the NPS is absorption or emission bounded. Similarly, the uncertainties in the ROSAT data do not allow us to determine the absorbing column as a function of position.

## 2.2. XMM-Newton data

A series of pointed observations with *XMM-Newton* of the southern terminus of the NPS were proposed to determine whether it is emission or absorption bounded (GO program P074189, P.I. K. Kuntz) and to use the comparison between the absorption and the Galactic distribution of cooler, X-ray absorbing material to determine the NPS source location. The positions of the pointings were chosen to 1) cover a range of 3/4 keV surface brightnesses that include the apparent southern terminus of the NPS, 2) lie near the transverse center of the NPS, 3) include a serendipitous XMM pointing at the northern end (shown in Fig. 1 but finally not used in the analysis because of its lower integration time and the loss of the center part because the point source had to be excluded) and avoid a serendipitous pointing with a bright point source at the southern end, and 4) sample the apparent absorption feature near  $b \approx 9^\circ$  (see Fig. 3). Nineteen partially overlapping 0:5 diameter fields provide full coverage from  $b = 5.6$  to  $11:1$  at an average longitude  $l = 29^\circ$ . Images of the surface brightness and spectrum hardness ratio are displayed in Fig. 3



**Fig. 2.** Profile of the emission in the RASS R4, R5, and R6 bands taken from the box shown in the previous figure. In each case the lowest emission level of the profile was removed and the resulting profile was normalized so that the emission in its brightest part (the NPS) was set to unity. The black histogram shows the expected absorption profile. The perpendiculars show the extent over which we obtained the *XMM-Newton* data. The bars in the upper right show the typical uncertainty of the ROSAT data in each band where the absorption changes most strongly.

along with the IRAS 100  $\mu$ m data. There is a global decrease of the X-ray brightness with decreasing latitude and an associated hardening of the spectra. The dust distribution does not monotonically vary with latitude, instead there is an IR maximum at  $b \approx 9^\circ$  and an associated X-ray darkening is clearly visible in Fig. 3 (right panel). As we show in Sect. 4, it corresponds to a molecular cloud visible in the CO maps.

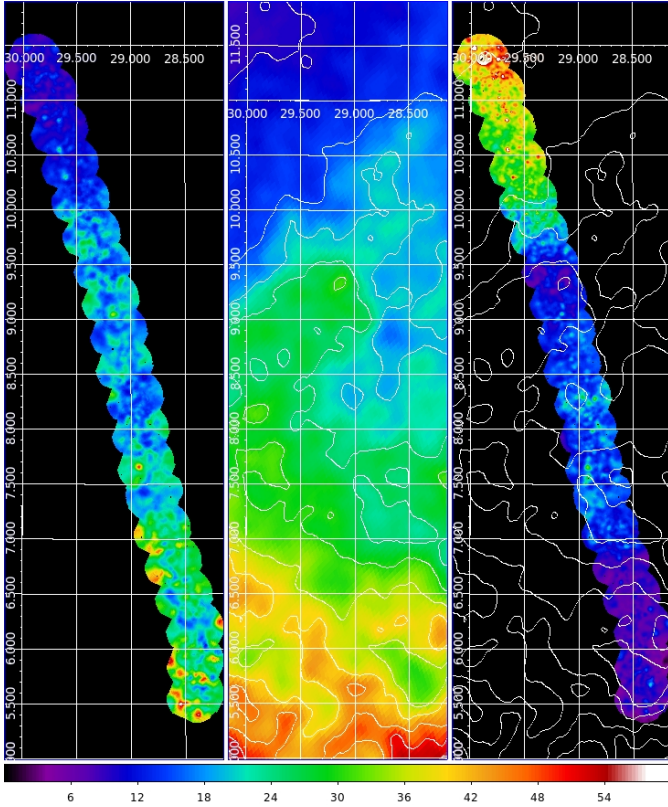
The processing of the XMM data (background modeling and subtraction, point source excision, and exposure correction) followed the procedures outlined in Snowden et al. (2008) based on the calibration of Kuntz & Snowden (2008). The XMM Extended Source Analysis Software<sup>1</sup> procedures are incorporated in the XMM Science Analysis System software<sup>2</sup>.

All MOS1, MOS2, and pn spectra were fitted simultaneously for the 19 fields, that is to say, 57 spectra, along with three RASS spectra to constrain the Local Hot Bubble (LHB) contribution (see Figs. 4 and 7). The models include three thermal emission components (APEC models with Anders & Grevesse 1989, abundances), (i) an unabsorbed LHB component characterized by a single low temperature, (ii) an absorbed hotter component representing the NPS, and (iii) a hot Galactic bulge component. A cosmic X-ray background contribution (CXB) in the form of a power law with a spectral index of 1.46 is also included in the model. The contribution from heliospheric solar wind charge exchange is assumed to be small and included in the LHB contribution.

<sup>1</sup> <ftp://legacy.gsfc.nasa.gov/xmm/software/xmm-esas/xmm-esas-v13.pdf>

<sup>2</sup> <http://www.cosmos.esa.int/web/xmm-newton/what-is-sas>



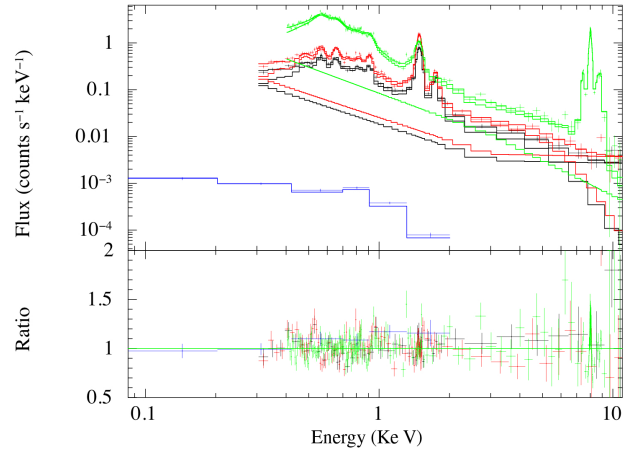


**Fig. 3.** Nineteen NPS pointings: the (0.75–1.3)/(0.4–0.75) keV hardness ratios (*left*), IRAS 100 $\mu$  data (*middle*) and 0.3–1.4 keV brightnesses (*right*). The color scale applies to the three quantities: from 0 to 4 for the hardness ratio, from 0 to 65 M Jsr<sup>-1</sup> for IRAS, and from 0 to 6 cts s<sup>-1</sup> deg<sup>-2</sup> for the 0.3–1.4 keV brightness.

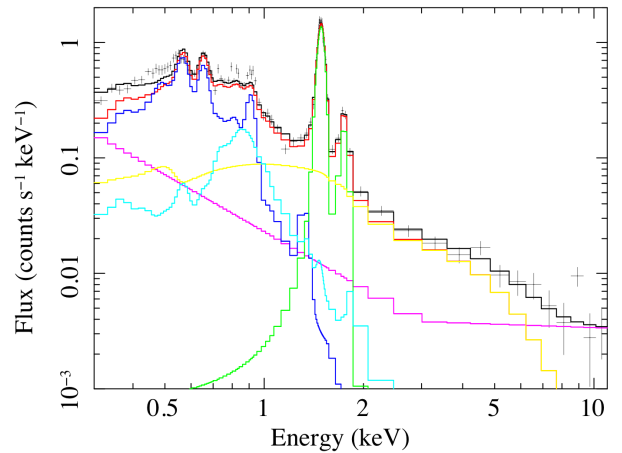
The first step in the X-ray spectral analysis process was to fit the RASS spectra separately to determine the LHB emission temperature and flux ( $kT_{\text{LHB}} = 0.111 \pm 0.006$  keV,  $F_{\text{LHB}} = 1.33 \pm 0.06 \times 10^{-8}$  ergs s<sup>-1</sup> cm<sup>-2</sup> sr<sup>-1</sup>). These parameters were then held fixed in all subsequent fits. Next, the 57 XMM spectra and 3 RASS spectra were simultaneously fit with the temperature for the NPS emission tied to the same value for all spectra and the temperature for the Bulge emission likewise. The fluxes for both the NPS and the Bulge, along with their associated absorption column densities, were allowed to vary individually for all pointings. The fitted temperatures are  $T_{\text{NPS}} = 0.179 \pm 0.002$  keV and  $T_{\text{Bulge}} = 0.702 \pm 0.007$  keV, and the fitted fluxes and absorption column densities are listed in Table 1. The fit had a  $\chi^2$  value of 25 974 for 17 779 degrees of freedom. The fitted NPS absorbing column densities and transmitted intensities are displayed in Fig. 5, along with an appropriate absorption curve (determined by absorbing the NPS thermal spectrum using WebPIMMS<sup>3</sup>). Although there is considerable scatter, the data are reasonably consistent with the distribution expected by absorption of a relatively uniform distant emission component by foreground material for all observation directions. This demonstrates that the NPS brightness decrease toward lower latitudes is primarily governed by the foreground absorption precluding an emission-bounded terminus, and strongly favors an NPS source located farther toward the Galactic plane.

With the large number of both spectra and parameters, the spectral fitting process was slow and made more complicated by

<sup>3</sup> <http://heasarc.gsfc.nasa.gov/cgi-bin/Tools/w3pimms/w3pimms.pl>



**Fig. 4.** Fitted MOS1 (black), MOS2 (red), and pn (green) spectra for the NPS pointing 17 along with the RASS spectrum (blue) for a region with a diameter of one degree centered on the pointing. For each of the EPIC spectra three curves show the fitted broken power law that represents the soft proton background (lower for all), the fitted cosmic spectrum plus fluorescent background lines (middle for all), and the sum of the two representing the best fit to the data. The MOS fluorescent lines are the Al K alpha and Si K alpha at 1.49 and 1.74 keV, respectively. The pn fluorescent lines are the Al K alpha and a variety of Cu, Ni, and Zn lines around 8 keV.



**Fig. 5.** Fitted NPS measured brightness as a function of the fitted absorbing column in the free-flux case. The curve is the model X-ray band-average absorption roughly scaled to the data.

the strong correlation between many parameters such as the flux and absorption column density. However, these preliminary results show that the NPS emission continues toward the Galactic plane as far as the observations extend. Figure 6 shows the fitted values for the NPS flux as a function of Galactic latitude, along with both constant and linear fits to the data. For the third stage of the spectral analysis we fit 1) a tied constant value to the NPS flux ( $\chi^2 = 26031.3$ ), 2) flux values tied with the linear fit results in Fig. 6 ( $\chi^2 = 26024.3$ ), and 3) values tied with the slope reduced by a factor of two ( $\chi^2 = 26024.6$ ). In all three cases the tied flux values were allowed to float, and there were 17 797 degrees of freedom. The fitted values for the ISM column densities for the three cases are listed in Table 2. The fitted values are similar and show the same trends. Differences between the resulting columns give an estimate of the systematic uncertainties associated with the fitting method. Until Sect. 4.2 in this paper we use the results from case 3.

**Table 1.** Observed XMM fields and fit results for freely varying NPS and Galactic Bulge fluxes and absorbing column densities.

EXP	Field	Gal. long (°)	Gal. lat (°)	$N_{\text{H}}(\text{NPS})$ $10^{20} \text{ cm}^{-2}$	$F(\text{NPS})$ unabsorbed $10^{-7} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$	$F(\text{NPS})$ transmitted $10^{-8} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$	$N_{\text{H}}(\text{bulge})$ $10^{20} \text{ cm}^{-2}$	$F(\text{bulge})$ $10^{-7} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$
0760080101	01a	28.4	5.6	40.0 ± 4.5	1.99 ± 1.21	0.75 ± 0.48	90.6 ± 4.9	1.35 ± 0.24
0741890201	02	28.4	5.9	30.5 ± 7.6	0.67 ± 0.75	0.43 ± 0.16	57.7 ± 10.3	0.46 ± 0.31
0741890301	03	28.4	6.2	36.6 ± 3.6	1.73 ± 0.88	0.79 ± 0.40	88.9 ± 5.1	1.03 ± 0.23
0741890401	04	28.5	6.5	41.4 ± 4.3	2.14 ± 1.30	0.77 ± 0.46	101.1 ± 8.0	1.40 ± 0.41
0741890501	05	28.6	6.8	46.8 ± 3.7	4.14 ± 2.13	1.04 ± 0.53	49.9 ± 12.9	0.31 ± 0.23
0741890601	06	28.7	7.1	35.7 ± 3.0	3.35 ± 1.48	1.60 ± 0.71	63.8 ± 4.9	1.10 ± 0.26
0741892201	07a	28.8	7.4	18.2 ± 3.4	1.74 ± 0.88	2.59 ± 1.32	62.3 ± 6.8	1.25 ± 0.43
0741890801	08	28.85	7.7	35.8 ± 3.5	2.78 ± 1.41	1.32 ± 0.67	59.4 ± 5.5	0.88 ± 0.25
0741890901	09	28.9	8	26.3 ± 3.1	1.96 ± 0.93	1.65 ± 0.79	47.4 ± 6.7	0.63 ± 0.26
0741891001	10	28.95	8.3	21.8 ± 1.9	2.34 ± 0.69	2.64 ± 0.78	65.4 ± 3.6	0.87 ± 0.18
0741891101	11	29.05	8.6	29.8 ± 3.2	1.89 ± 0.98	1.28 ± 0.66	32.3 ± 8.1	0.36 ± 0.21
0741891201	12	29.1	8.9	28.0 ± 2.9	1.86 ± 0.85	1.40 ± 0.64	38.9 ± 9.8	0.31 ± 0.24
0741891301	13	29.175	9.2	36.1 ± 2.7	2.96 ± 1.22	1.38 ± 0.57	48.1 ± 11.2	0.33 ± 0.28
0741891401	14	29.25	9.5	18.4 ± 3.0	0.97 ± 0.46	1.39 ± 0.65	50.0 ± 12.2	0.37 ± 0.33
0760080201	15a	29.35	9.8	24.9 ± 2.4	2.37 ± 0.88	2.17 ± 0.81	28.5 ± 7.6	0.30 ± 0.18
0760080301	16a	29.45	10.11	21.5 ± 2.4	2.40 ± 0.84	2.77 ± 0.97	0.0 ± 2.3	0.12 ± 0.03
0741891701	17	29.55	10.43	20.4 ± 1.8	3.67 ± 0.99	4.53 ± 1.23	0.0 ± 2.5	0.14 ± 0.03
0741891801	18	29.65	10.75	16.9 ± 1.8	2.98 ± 0.82	4.70 ± 1.29	0.0 ± 2.9	0.15 ± 0.04
0741891901	19	29.75	11.07	14.7 ± 1.5	2.93 ± 0.69	5.41 ± 1.28	0.4 ± 0.1	0.26 ± 0.02

**Table 2.** Fit results for a unique NPS flux over the  $\approx 5$  deg latitude interval, a flux linearly varying with latitude with the average slope deduced from the free-flux case, and a slope reduced by a factor of two.

EXP	Field	$N_{\text{H}}(\text{NPS})$ cst $10^{20} \text{ cm}^{-2}$	$N_{\text{H}}(\text{NPS})$ lin $\times 0.5$ $10^{20} \text{ cm}^{-2}$	$N_{\text{H}}(\text{NPS})$ lin $10^{20} \text{ cm}^{-2}$
0760080101	01a	42.7 ± 1.3	38.5 ± 1.3	35.3 ± 1.4
0741890201	02	50.1 ± 1.6	46 ± 1.6	43.0 ± 1.6
0741890301	03	41.4 ± 1.3	38.1 ± 1.3	35.5 ± 1.3
0741890401	04	45.2 ± 1.3	39.5 ± 1.3	37.2 ± 1.3
0741890501	05	38.8 ± 1.3	36.4 ± 1.4	34.6 ± 1.4
0741890601	06	30.5 ± 1.2	28.5 ± 1.2	26.9 ± 1.2
0741892201	07a	22.3 ± 1.2	20.5 ± 1.2	19.0 ± 1.3
0741890801	08	33.3 ± 1.2	31.7 ± 1.2	30.5 ± 1.3
0741890901	09	28.7 ± 1.2	27.4 ± 1.2	26.4 ± 1.2
0741891001	10	21.8 ± 1.1	20.8 ± 1.2	20.0 ± 1.2
0741891101	11	32.5 ± 1.2	31.7 ± 1.3	31.1 ± 1.3
0741891201	12	31.2 ± 1.2	30.7 ± 1.3	30.2 ± 1.3
0741891301	13	32.6 ± 1.2	32.4 ± 1.3	32.1 ± 1.3
0741891401	14	31.4 ± 1.2	31.5 ± 1.2	31.2 ± 1.2
0760080201	15a	24.5 ± 1.2	24.8 ± 1.2	24.8 ± 1.3
0760080301	16a	21.0 ± 1.2	21.6 ± 1.3	21.7 ± 1.3
0741891701	17	14.2 ± 1.2	15.1 ± 1.4	15.3 ± 1.3
0741891801	18	13.6 ± 1.2	14.6 ± 1.3	14.9 ± 1.3
0741891901	19	11.6 ± 1.1	12.8 ± 1.2	13.3 ± 1.2

**Notes.** Fit qualities are very similar for the three cases (see text).

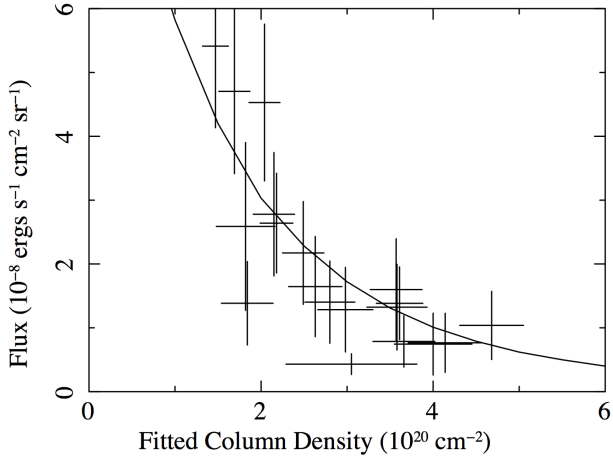
To summarize the XMM spectral results, the best-fit values for the LHB, NPS, and Bulge parameters are  $kT_{\text{LHB}} = 0.111 \text{ keV}$ ,  $kT_{\text{NPS}} = 0.181 \text{ keV}$ ,  $kT_{\text{Bulge}} = 0.705 \text{ keV}$ , and  $F_{\text{LHB}} = 1.34 \times 10^{-8} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$ ,  $F_{\text{NPS}}$  increases linearly between  $1.45 \times 10^{-7}$  and  $2.62 \times 10^{-7} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$  with increasing Galactic latitude, and  $F_{\text{Bulge}}$  generally decreases from  $1.33 \times 10^{-7}$  to  $0.16 \times 10^{-7} \text{ ergs s}^{-1} \text{ cm}^{-2} \text{ sr}^{-1}$  with increasing Galactic latitude. The fitted values for the NPS foreground  $N_{\text{Habs}}$  vary from 4.3 to  $1.3 \times 10^{21} \text{ cm}^{-2}$  in the 19 fields and generally decrease with increasing Galactic latitude, and the Bulge  $N_{\text{Habs}}$  generally decreases from 9.7 to  $0.0 \times 10^{21} \text{ cm}^{-2}$  (the latter value giving an idea of some of the systematics in the fits). The fitted NPS column is equal to or higher than the columns found from spectral analysis of Suzaku and XMM NPS data by Miller et al. (2008) and Willingale et al. (2003) at higher latitude and similar longitude

( $N_{\text{H}}(\text{abs}) = 0.7 \times 10^{21} \text{ cm}^{-2}$  at  $(l, b) = (25, +20)$ ,  $0.5 \times 10^{21} \text{ cm}^{-2}$  at  $(l, b) = (27, +22)$ ,  $0.33 \times 10^{21} \text{ cm}^{-2}$  at  $(l, b) = (20, +30)$ , and  $0.2 \times 10^{21} \text{ cm}^{-2}$  at  $(l, b) = (20, +40)$ ).

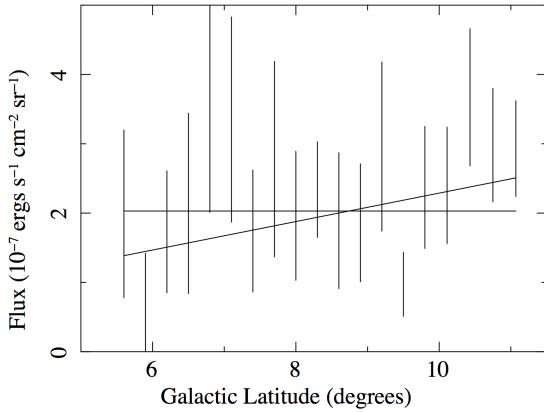
### 3. Comparison between X-ray absorbing columns and $\leq 1 \text{ kpc}$ absorption measurements

#### 3.1. High-resolution NARVAL optical spectra and interstellar KI profile-fitting

We selected a series of nearby target stars located in the direction of the NPS southern terminus for which a HIPPARCOS parallax distance is available. The targets stars are shown in Fig. 8 superimposed on the RASS 3/4 keV map. The targets are early-type stars or conversely, cool G-K stars with a radial velocity

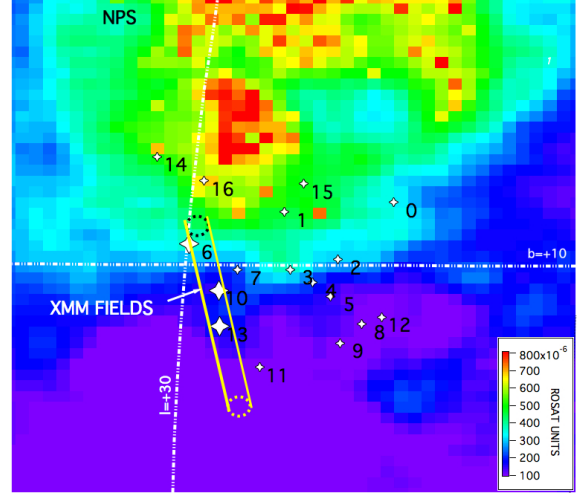


**Fig. 6.** Fitted NPS fluxes for the 19 pointings.



**Fig. 7.** Fitted MOS2 spectrum from the NPS pointing 17 showing the individual components: dark blue – NPS, light blue – Galactic background, yellow – extragalactic background, green – instrumental fluorescent background, red – sum of the above, violet – soft proton background, black curve – sum of all components.

different from the mean interstellar motion in this area. High signal-to-noise ratio and high-resolution spectra ( $R = 80\,000$ ) were recorded with NARVAL, the spectropolarimeter of the *Bernard Lyot* telescope (2 m) at Pic du Midi Observatory, used in the spectrometric mode. Observations were distributed in 2014 in the frame of a dedicated program (P.I. R. Lallement). We used the standard reduction pipeline. Telluric lines were removed using our rope length method as described in [Raimond et al. \(2012\)](#) and using TAPAS synthetic transmittance spectra ([Bertaux et al. 2014](#)). In the case of the strongest telluric lines around the  $7665 \text{ \AA}$  KI line we additionally kept a mask in regions corresponding to the deepest absorptions. For the early-type stars we fit the spectra simultaneously for the  $5889\text{--}5895 \text{ \AA}$  NaI-D doublet and the two  $7665\text{--}7699 \text{ \AA}$  KI lines, allowing for multiple clouds and polynomial functions for the stellar continua and using classical Voigt profiles (see, e.g., [Welsh et al. 2010](#)). The combined use of NaI and KI allowed us to better constrain the cloud parameters and constituted a validity check for the KI absorptions in the region that is contaminated by the strong telluric lines (see Fig. A.1). For three targets that were strongly contaminated we only fit the KI  $7698 \text{ \AA}$  transition (Fig. A.2). Late-type star absorption lines were treated differently depending on the difference between the stellar and IS radial velocities. For large differences the stellar line was



**Fig. 8.** NPS XMM pointing and ground-based targets (white diamonds) superimposed on the ROSAT 0.75 keV map. The three targets that are located in the XMM path are shown as larger signs.

simply treated as a continuum (Fig. A.3). For small differences the stellar radial velocities were determined preliminarily from the whole spectrum, and simulated KI stellar lines were added to the fitting model at the star velocity (Fig. A.4). In all cases a good estimate of the KI column density could be determined. The KI columns were converted into equivalent  $N(\text{H})$  using the empirical formula established by [Wety & Hobbs \(2001\)](#) based on  $\approx 50$  high-resolution spectra (formula 3 in their Table 2). The target characteristics, HIPPARCOS data, fitted KI columns, and equivalent  $N(\text{H})$  columns are listed in Table 3.

### 3.2. First constraints from the target stars along the XMM path

Three of the target stars are located within our XMM fields, hip88841 (star 6), hip89151 (star 10), and hip89472 (star 13), shown in Fig. 8. This allows a comparison between their distance-limited  $N(\text{H})$  columns deduced from KI absorption lines and the measured XMM  $N_{\text{Habs}}$  (see Table 3). For the three targets the optical columns are either significantly below or slightly below the XMM absorbing columns, strongly suggesting that the NPS source is located beyond the stars. Their HIPPARCOS distances are  $135(-8, +9)$ ,  $279(-65, +175)$ , and  $344(-87, +175)$  pc respectively, implying that the largest of the three shortest distances is 257 pc. We conclude that the near side of the NPS source is more than 250 pc distant.

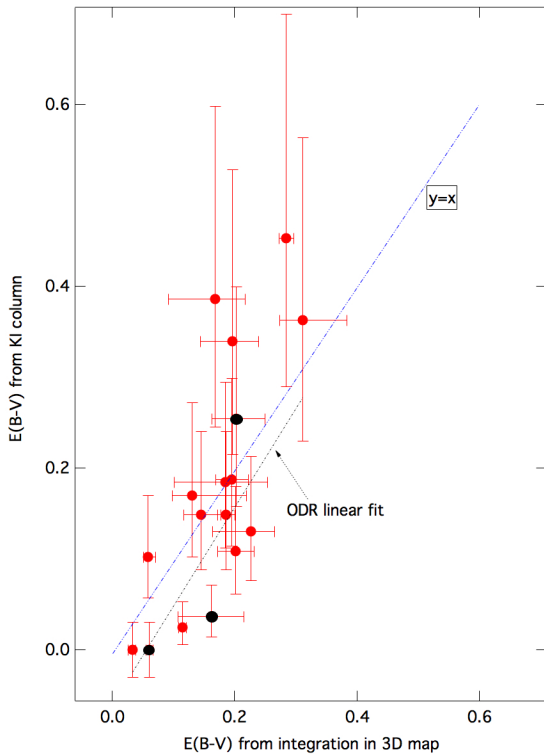
### 3.3. 3D map of nearby dust and absorption measurements

The whole Narval absorption dataset was used to validate the computed 3D map of local dust reported by [Lallement et al. \(2014\)](#) in the NPS southern terminus area. The maps were obtained by inverting individual color excess measurements, and they have two limitations: first, there is a minimum size for the inverted structures because of the limited number of target stars. Second, there is a bias toward weakly reddened stars, and for this reason, there are far fewer targets in the database that are located beyond opaque clouds than in front of them. These limitations influence in a complex way the distance range over which the maps can be used safely, and it is useful here to obtain an independent confirmation of the dust reddening found by inversion, especially for the Aquila Rift area. To do so, we integrated

**Table 3.** Observed stars, their HIPPARCOS numbers and distances, interstellar KI absorbing columns, and estimated line-of-sight gas column and color excess.

#	HIP	glon (°)	glat (°)	sptype	dist pc	distmin pc	distmax pc	$N(\text{KI})$ $\text{cm}^{-2}$	$N(\text{H})$ $\text{cm}^{-2}$	$E(B - V)$ mag	(+)	(-)	ebv(3D) mag	(+)	(-)
0	87633	24.82	11.98	A0	183	169	199	3.26e+11	1.33e+21	0.19	0.07	0.11	0.20	0.03	0.03
1	88148	27.59	11.66	A3	135	126	146	1.34e+11	8.11e+20	0.10	0.04	0.07	0.06	0.01	0.01
2	88376	26.12	10.18	B9	227	197	269	8.42e+11	2.26e+21	0.34	0.12	0.19	0.20	0.05	0.04
3	88671	27.31	9.86	G8II	215	195	240	2.3e+11	1.10e+21	0.15	0.06	0.09	0.15	0.03	0.03
4	88698	26.72	9.46	A3	235	175	355	3.18e+11	1.31e+21	0.18	0.07	0.11	0.19	0.08	0.07
5	88753	26.24	9.03	B9	549	369	1075	9.4e+11	2.41e+21	0.36	0.13	0.20	0.31	0.04	0.07
6*	88841	29.95	10.65	A2	135	127	144	4e+08	3.15e+19	0.00	0.03	0.03	0.06	0.01	0.01
7	88870	28.65	9.85	K2	366	286	510	1.46e+11	8.51e+20	0.11	0.05	0.07	0.20	0.03	0.03
8	88878	25.41	8.16	G5	244	213	285	1.9e+11	9.86e+20	0.13	0.05	0.08	0.23	0.06	0.04
9	89148	25.93	7.57	K0	204	175	244	2.8e+11	1.22e+21	0.17	0.07	0.10	0.13	0.03	0.09
10*	89151	29.10	9.21	A2	279	204	444	4e+10	4.13e+20	0.04	0.02	0.03	0.16	0.05	0.05
11	89680	27.92	6.80	A0	228	192	279	1.04e+12	2.55e+21	0.39	0.14	0.21	0.17	0.08	0.05
12	88740	24.91	8.38	F8	95	90	100	0	0	0.00	0.03	0.03	0.03	0.01	0.01
13*	89472	29.00	8.08	K2	344	257	518	5.28e+11	1.74e+21	0.25	0.10	0.15	0.20	0.04	0.05
14	88149	31.00	13.37	B2V	200	190	211	2.8e+10	3.38e+20	0.02	0.02	0.03	0.12	0.01	0.01
15	87812	27.16	12.55	B2IV-V	417	356	503	1.36e+12	2.96e+21	0.45	0.16	0.25	0.28	0.01	0.01
16	88192	29.72	12.63	B5I	377	313	474	2.3e+11	1.10e+21	0.15	0.06	0.09	0.19	0.01	0.01

**Notes.** We also list the color excess obtained from integration within the 3D local map. Stars located in the XMM path are indicated by an asterisk.



**Fig. 9.** Estimated color excess of the target stars based on their KI lines, using the empirical relationship of Gudennavar et al. (2012), and estimated reddening through integration within the local dust map. A linear fit taking into account errors on both quantities (orthogonal distance regression, ODR) is superimposed. The three target stars along the XMM path are shown in black.

through the computed 3D map the color excess  $E(B - V)$  from the Sun to each target location, and did it separately for their HIPPARCOS shortest, most probable, and largest distances. The resulting color excesses are listed in Table 3. We compared these map-integrated values with the KI-based estimated H columns. Figure 9 shows the comparison and reveals a clear increasing trend, but a large dispersion. Such departures from linearity are expected because of the very low spatial resolution of the map and also because the H/KI ratio varies. The important result here

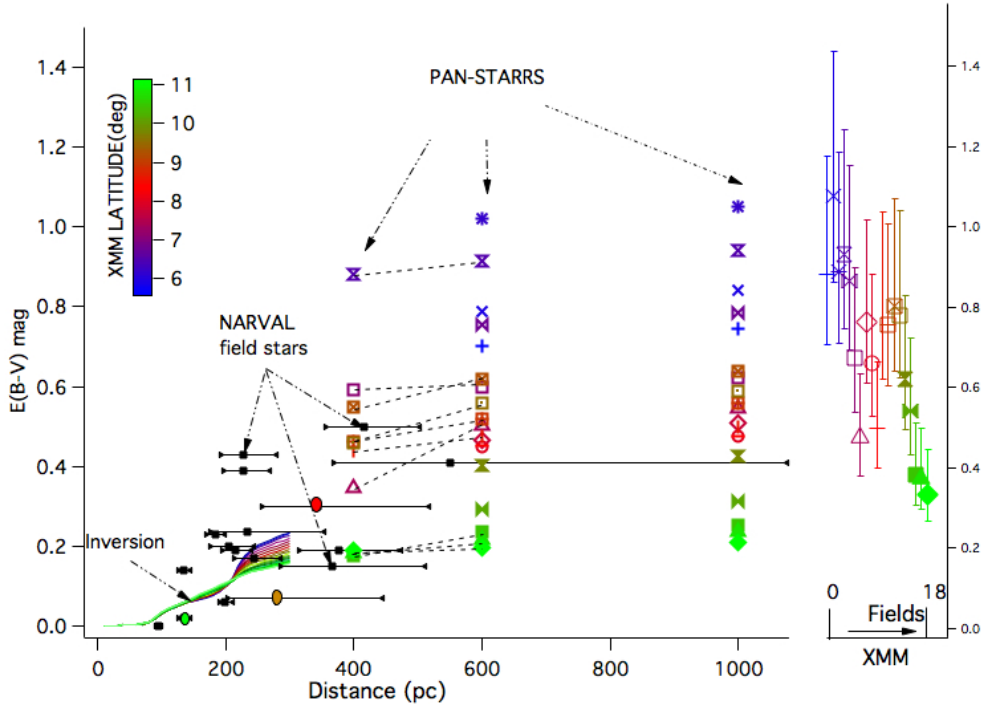
is that within error bars the best linear fit is compatible with the empirical relationship between KI and  $E(B - V)$  found by Gudennavar et al. (2012), as shown in the figure. This implies that in this area and for H columns on the order of those measured to our targets, an integration through the map provides a reasonable order of magnitude of the reddening. Since our targets have well-defined distances up to  $\approx 300$  pc, this means that we do not miss significant structures up to this distance. At larger distance, however, the 3D maps become too uncertain because we lack reddening measurements (see Lallement et al. 2014, for more information).

Based on this validity check, we integrated through the maps from the Sun to 300 pc along the 19 XMM field directions. The integrations are shown in Fig. 10. Given the small angular range covered by the XMM pointings, the reddening does not vary more than by about 30% from lowest to highest reddening, and  $E(B - V)$  does not exceed 0.25, which is the color excess value reached at the lowest latitude. Such reddenings correspond to  $N(\text{H})$  columns on the order of  $1.0 \times 10^{21} \text{ cm}^{-2}$ , well below the XMM measured absorptions at low latitudes. We also show in Fig. 10 the color excess values estimated from KI for the field stars. When both types of data are used, the figure shows a first increase in opacity at about 100–150 pc, followed by a larger increase distributed between 200 and 400 pc. This corresponds to the well-known Aquila Rift clouds. The variability from star to star confirms the complex structure of these clouds, as shown in Fig. 8 of Puspitarini et al. (2014). Nevertheless, the comparison between the reddening achieved at 300 pc along the XMM path based on the local maps and the XMM absorption column densities very strongly suggests an NPS source beyond 300 pc, beyond the second wall revealed by the map and star absorptions.

#### 3.4. Local and Pan-STARRS 3D reddening maps

Green et al. (2015) have produced 3D reddening maps based on the Pan-STARRS (PS) photometric survey, which possess a high degree of angular resolution and map up to 5–10 kpc depending on directions. These maps are used here to derive the reddening as a function of distance along the XMM fields. They become precise only beyond about 300–400 pc and are ideally complementary to the local maps. To do so, we used the online tool built by the authors and extracted the reddening measurements





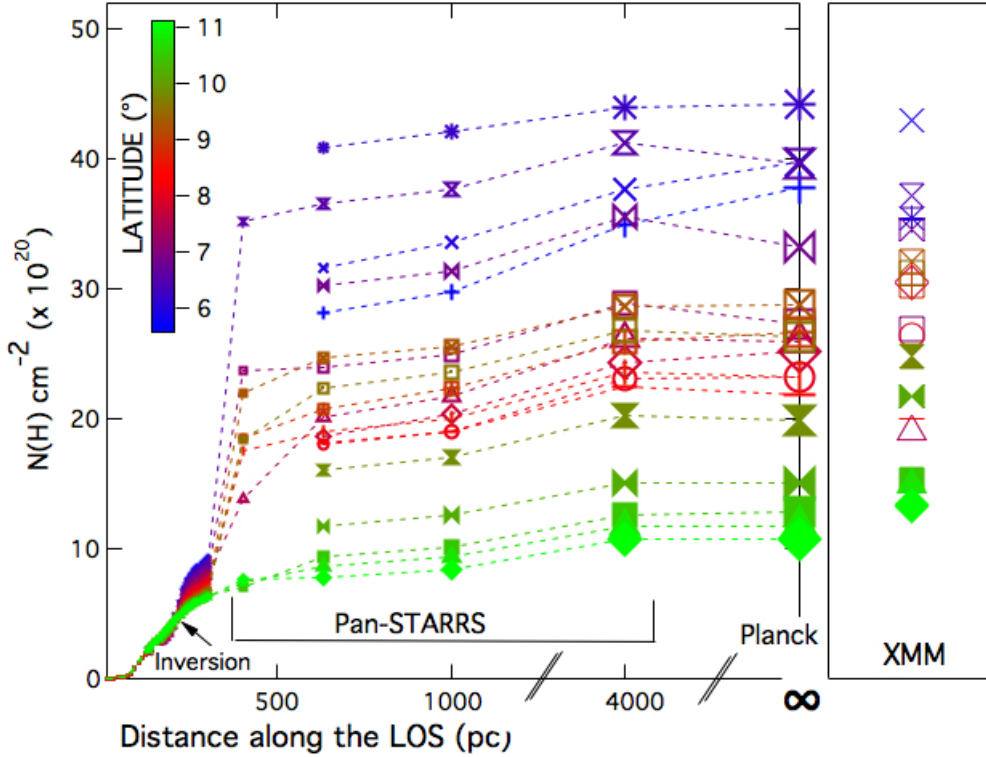
**Fig. 10.** *Left panel:* color excess  $E(B - V)$  estimates for the 19 XMM field directions, colored according to the Galactic latitude: (i) integration in the local-reddening-inverted maps (solid lines); (ii) Pan-STARRS color excess at 400, 600, and 1000 pc from Green et al. (2015), rescaled values using measurements at 4 kpc and *Planck*  $\tau$ -based  $E(B - V)$ s (see text). Colors refer to the Galactic latitude (inserted color scale) and markers are different for each direction. Some measurements are missing at shorter than 1 kpc because their uncertainty is too large). We also show the color excess values based on KI absorption lines for the NARVAL field stars in the NPS terminus region as a function of their HIPPARCOS distance ranges (black dots and error bars). The three targets that are located along the XMM path are indicated by large circles whose colors refer to the latitude. *Right panel:* XMM NPS fitted absorbing columns scaled to  $E(B - V)$  using  $N(H) = 3$  to  $5 \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$ .

and uncertainties for 29 directions separated by  $0.1^\circ$  within each XMM circular field. We averaged only those data that have a good quality flag. At short distance, namely 400 and 600 pc, the number of valid points is small for some fields, and in a few cases there is a large dispersion that is due to uncertainties in the distances of the closest clouds, themselves linked to statistical uncertainties of the photometric method. We conservatively kept measurements in fields possessing at least five well-clustered valid points, and in all fields we excluded two-sigma outliers.

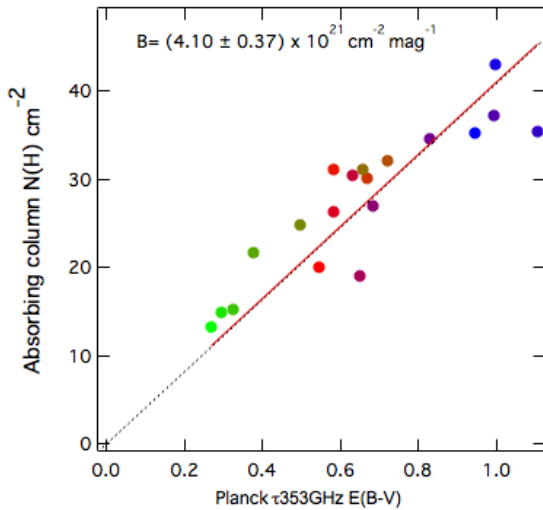
As discussed by Green et al. (2015), the PS reddening in the Aquila region is overestimated, possibly because stellar metallicity gradients are not taken into account or because of peculiar properties of the dust grains. For this reason, we chose to scale the PS  $E(B - V)$ s values to *Planck* measurements based on the 353 GHz dust optical depth  $\tau_{353}$ . To do so, we first correlated the PS reddenings with the *Planck* reddenings, averaged on the same 29 directions in each field, at increasing distances. As expected, the correlation coefficient increases with distance, and at 4 kpc the *Planck* and Pan-STARRS  $E(B - V)$ s are nearly perfectly correlated (Pearson = 0.992). This is expected because at this distance and for the XMM field latitudes the distance to the Plane is  $Z \geq 400$  pc, far above the dust scale height, which is on the order of 130 pc. Beyond 4 kpc the correlation between PS and *Planck* starts to decrease because of an observational bias. In strongly reddened fields and at large distance, the number of target stars becomes insufficient for the statistical analysis and there is a favored selection of the less reddened directions, which results in an average reddening lower than the actual one. For these reasons, we considered the 4 kpc measurement as the best estimate of the reddening over the entire sightline.

We therefore fitted the PS(4 kpc) – *Planck* relationship to a second-order polynomial and applied the derived relationship to the whole set of PS data (the mean ratio between the PS and *Planck* measurements is on the order of 1.25, in agreement with Fig. 12 from Green et al. 2015). Table 3 lists the resulting PS values at 400, 600, 1000 pc and 4000 pc and the *Planck*  $E(B - V)$ s. Figure 10 displays the reddening integrated through the local 3D map as a function of distance, the reddening estimated from the KI absorption in the field stars, and these *Planck*-scaled PS reddening values. As discussed for a long time (see, e.g., Arce & Goodman 1999; Bonifacio et al. 2000; Cambrésy et al. 2001), photometric reddenings are generally lower than those estimated from dust emission, especially for strong extinction, but the use of these various diagnostics traces the evolution of the reddening with distance rather well.

The combination of 3D maps in Fig. 10 reveals a strong reddening jump between 300 pc and 600 pc for latitudes below  $+7^\circ$  (blue signs), with some variations in the jump location from one direction to the other. This wall of dense IS matter is also detected in the spectra of the target stars located beyond 300 pc (small black circles and distance intervals in Fig. 10). For those directions that possess a reliable PS measurement at 400 pc, a small or significant reddening gradient between 400 and 600 pc is evident, as illustrated in the figure by the connecting dashed lines. This suggests that the cloud complex that produces the reddening jump extends beyond 400 pc. On the other hand, the plateau between 600 and 1000 pc allows us to conclude that the far side of the cloud complex is located between 400 and 600 pc and not farther out. To infer the NPS location with respect to this wall of dense structures, we converted the



**Fig. 11.** *Left panel:* from left to right for the 19 field directions, colored according to the Galactic latitude, and as a function of distance: (i) estimated gas column between the Sun and 300 pc from integration in the local reddening maps, (ii) gas columns estimated from Green et al. (2015) Pan-STARRS color excess measurements up to four distances, and (iii) gas columns estimated from *Planck*  $\tau$ -353GHz reddening measurements. See the text and Fig. 10 for the Pan-STARRS reddening scaling. All the reddening values are converted into  $N(\text{H})$  using  $N(\text{H}) = 4. \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$ , the central value of Fig. 10. *Right panel:* soft X-ray absorbing column  $N_{\text{Habs}}$  deduced from global fitting of XMM spectra with fixed NPS flux. The vertical scale is the same as in the left panel.



**Fig. 12.** XMM gas absorbing column (linear case) vs. *Planck*  $E(B - V)$   $\tau_{353}$  (averaged over the XMM FOV). The color coding is for Galactic latitude and is the same as in Figs. 10 and 11.

XMM absorbing gas columns into  $E(B - V)$ s using the broad interval  $N(\text{H}) = 4 \pm 1 \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$ , an interval significantly large to account for inhomogeneity of the dust-to-gas ratio. This interval corresponds to the most representative recent results of Planck and Fermi Collaborations Int. XXVIII (2015), who have taken the contributions of the HI, DNM, and CO-bright phases into account in their analysis. Figure 10 shows the

resulting reddening values and how they compare with the previously discussed distance-limited reddening measurements or estimates. The comparison very strongly suggests that the X-ray absorbing matter corresponds to the totality of the cloud complex that is distributed between 300 and 600 pc. More specifically, the absorptions match the 600 pc (or equivalently 1 kpc) reddening values for a reasonable conversion factor on the order of  $4. \times 10^{21} \text{ cm}^{-2} \text{ mag}^{-1}$ . On the other hand, using the lower  $N(\text{H})$ - $E(B - V)$  factor gives reddenings that are marginally compatible with the 400 pc data (ten directions only). As a consequence, given the uncertainty on the conversion factor and because we do not know the precise outer boundary distance of the cloud complex, we very conservatively conclude that the near side of the NPS source lies beyond at least 300 pc. This result firmly precludes any nearby source, in particular a source associated with the Sco-Cen star-forming region. On the other hand, we note that the NPS near side is potentially as far away as 1 kpc or more, since it is not possible at this stage to determine whether the small reddening increase that is detected between 400 pc and 600 pc is seen in the XMM absorptions, and neither can we determine the next small increase detected between 1 kpc and 4 kpc in some directions, see also Fig. 11.

#### 4. Comparison between X-ray absorbing columns and absorption over larger distances

##### 4.1. Dust

Based on the previous results, we extended our comparisons to larger distances and additionally used emission data as tracers of

**Table 4.** Dust reddening in the XMM directions:  $EBV_{\text{loc}}$  is the color excess integrated from the Sun to 300 pc through the local dust map.

Field	$EBV_{\text{loc}}$ mag	PS-0.4 kpc mag	PS-0.6 kpc mag	PS-1 kpc mag	PS-2 kpc mag	PS-3.16 kpc mag	PS-4 kpc mag	$Planck_{\tau}$ mag
01A	0.23		0.70	0.75	0.81	0.86	0.87	0.95
02	0.23		0.79	0.84	0.91	0.94	0.94	0.99
03	0.23		1.02	1.05	1.10	1.10	1.10	1.10
04	0.22	0.88	0.91	0.94	1.01	1.03	1.03	0.99
05	0.21		0.76	0.78	0.85	0.89	0.89	0.83
06	0.20	0.59	0.60	0.62	0.68	0.72	0.72	0.68
07	0.20	0.35	0.50	0.54	0.59	0.65	0.65	0.65
08	0.19		0.47	0.51	0.55	0.61	0.61	0.63
09	0.19		0.45	0.48	0.53	0.58	0.58	0.58
10	0.18		0.45	0.48	0.52	0.56	0.56	0.55
11	0.18	0.44	0.48	0.50	0.54	0.59	0.59	0.58
12	0.17	0.46	0.52	0.56	0.61	0.65	0.65	0.67
13	0.17	0.55	0.62	0.64	0.70	0.72	0.72	0.72
14	0.17	0.46	0.56	0.59	0.64	0.67	0.67	0.66
15A	0.16		0.40	0.43	0.46	0.50	0.51	0.50
16A	0.16		0.29	0.31	0.34	0.38	0.38	0.38
17	0.16	0.18	0.23	0.25	0.28	0.32	0.32	0.32
18	0.16	0.18	0.21	0.23	0.26	0.29	0.29	0.29
19	0.16	0.19	0.19	0.21	0.25	0.27	0.27	0.27

**Notes.** The Pan-STARRS (PS) color excess values are averaged over the field and scaled to match  $Planck_{\tau}$  at 4 kpc (see text).

**Table 5.** Correlations between the X-ray absorbing columns and color excesses.

$d$ (pc)	1 $\chi^2_{\text{red}}$	1 $B$	1 $\sigma(B)$	2 $\chi^2_{\text{red}}$	2 $B$	2 $\sigma(B)$	3 $\chi^2_{\text{red}}$	3 $B$	3 $\sigma(B)$
600	33.6	48.5	2.3	30.7	50.8	2.2	26.7	53.9	2.1
1000	30.2	46.3	2.1	27.1	48.5	2.0	23.1	51.5	1.9
2000	26.9	43.0	1.9	23.9	45.0	1.7	20.4	47.8	1.6
3160	22.5	41.0	1.6	19.6	42.8	1.5	16.9	45.5	1.4
4000	21.9	41.0	1.6	19.0	42.9	1.5	16.3	45.5	1.4
<i>Planck</i>	19.8	41.0	1.5	16.3	42.9	1.4	13.0	45.5	1.2

**Notes.**  $N(\text{H})_{\text{abs}} = B \times E(B - V)$ . (1) NPS flux varying linearly with latitude; (2) linear variation: slope  $\times$  (0.5); (3) NPS flux constant in the fields.

the total amount of dust and gas along the fields. For the dust, we extended our comparisons to PS at  $\approx 3.2$  and 4 kpc and to *Planck*. For *Planck* we used the color excess  $E(B - V)$  map derived from the dust optical thickness  $\tau_{353\text{GHz}}$  (Planck Collaboration XI 2014), and averaged the reddening  $E(B - V)_{\tau_{353}}$  over the XMM fields of view in the same way as described above for the correlation with PS at 4 kpc. We averaged the PS data and scaled them in the way described in the previous section (averaged and scaled data in Table 4). Figure 11 shows the reddenings from the local map, PS determinations at 400, 600, 1000, and 4000 pc and *Planck*, this time all converted into equivalent  $N(\text{H})$  columns using  $4.0 \times 10^{21} \text{ H cm}^{-2} \text{ mag}^{-1}$ , a value appropriate for the local matter, as stated in the previous paragraph. We also plot XMM NPS absorbing columns, which allows a direct visual comparison between the absorptions and the reddening as a function of distance. The figure shows that despite seemingly random discrepancies for the five lowest latitudes, overall the NPS absorbing columns are on the order of the converted  $E(B - V)$ s for distances as large as 4 kpc and equivalently to the *Planck* values. However, the low level of the reddening increase from the Aquila Rift clouds to large distances, for instance, between 600 pc and 4 kpc, and the model uncertainties on  $N_{\text{Habs}}$  do not allow us to unambiguously separate these two limits.

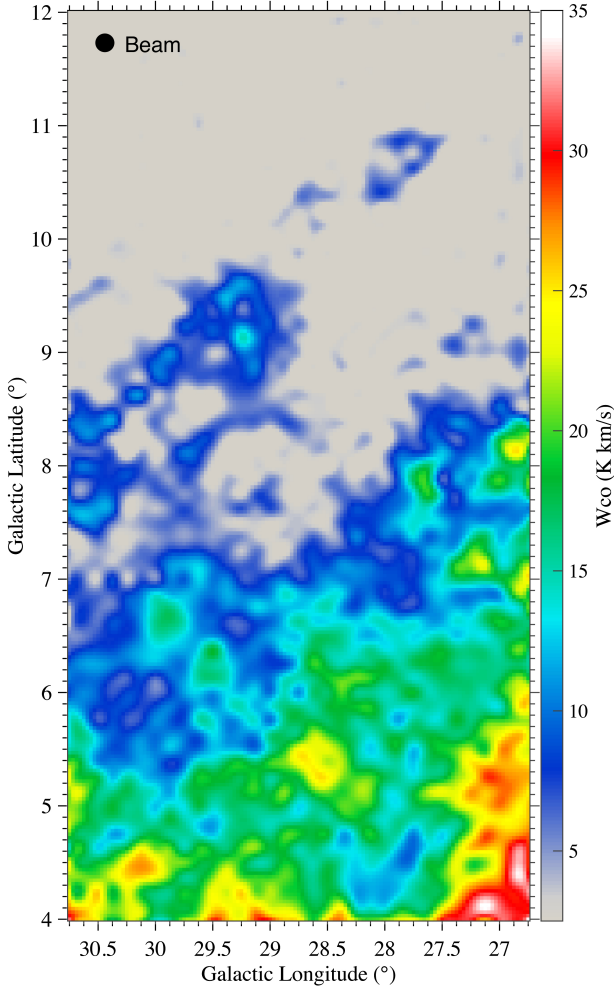
If the NPS source lies beyond all clouds located within a given distance and this distance is the same along the  $5^\circ$  XMM

path, then there must be a similar latitude profile between the reddening and  $N(\text{H})$ . Table 5 compares the standard deviations from proportional linear relationships relating the measured absorbing columns to the PS reddenings at varying distances and to *Planck*. The table shows that for the three fitting methods, the correlation improves for increasing distances and is optimal for the *Planck* reddening. Although part of the effect may be linked to the improved statistics of the PS measure between 600 pc and 3 kpc and the quality of the *Planck* determination, these correlations suggest that most of the dust detected by PS and *Planck* lies in front of the NPS source. Figure 12 shows the correlation between  $N_{\text{Habs}}$  and *Planck*  $\tau_{353\text{GHz}}$ -based  $E(B - V)$ . The slope of the proportional linear relationship is  $4.0 \times 10^{21} \text{ H cm}^{-2} \text{ mag}^{-1}$ , a reasonable value, as we stated before.

## 4.2. Gas

### 4.2.1. New CO survey data

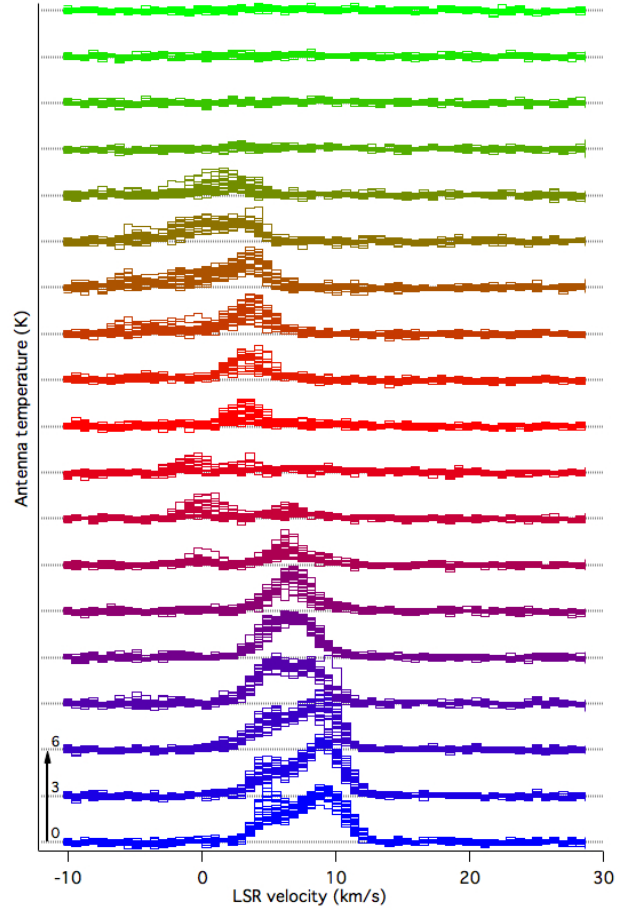
Since the region of the XMM fields was only sampled every quarter degree by the whole Galaxy CO survey of Dame et al. (2001), we undertook new more finely spaced CO observations of the region with the 1.2 m telescope of the Center for Astrophysics, which is the same instrument as was used for the survey of Dame et al. The region  $l = 26.75\text{--}30.75$ ,  $b = 4\text{--}12$  was



**Fig. 13.** Our new CO survey of the XMM fields, integrated from  $-5$  to  $+25$   $\text{km s}^{-1}$ , the full range over which significant emission is detected. The survey was sampled every  $7.5'$  with the  $8.4'$  beam shown in the upper left corner.

sampled slightly better than every beam width (every  $7.5'$ ) with a  $8.4'$  beam. A total of 1536 positions were observed with integration times of  $\approx 1$  min, but they were automatically adjusted based on the instantaneous system temperature to obtain a uniform survey sensitivity of  $0.2$  K in  $0.65$   $\text{km s}^{-1}$  spectra channels. Since existing data in the region revealed that the entire emission was confined to the velocity range  $-5$  to  $+25$   $\text{km s}^{-1}$ , we were able to switch the frequency by  $15$  MHz while keeping the entire emission within the  $64$  MHz band of the spectrometer. The spectra were subsequently folded and fifth-order polynomial baselines were removed. During the observations, the telluric emission line of CO appeared at velocities between  $-45$  and  $-39$   $\text{km s}^{-1}$ , well displaced from the celestial emission, and could be easily removed from the data.

Figure 13 shows the velocity-integrated CO brightness in the XMM measurements area. In addition to the latitudinal decrease, an isolated cloud is located at  $\approx +9^\circ$  latitude, with a marked correlation between CO emission and X-ray absorption in this region, which demonstrates that the NPS source is beyond this cloud. As shown by the individual spectra displayed in Fig. 14, for the entire set of sightline the CO velocity is in the range of between  $-5$  and  $+13$   $\text{km s}^{-1}$ , and the high-latitude cloud has a LSR velocity between  $-5$  and  $+2$   $\text{km s}^{-1}$ , significantly lower than lower latitude molecular material.

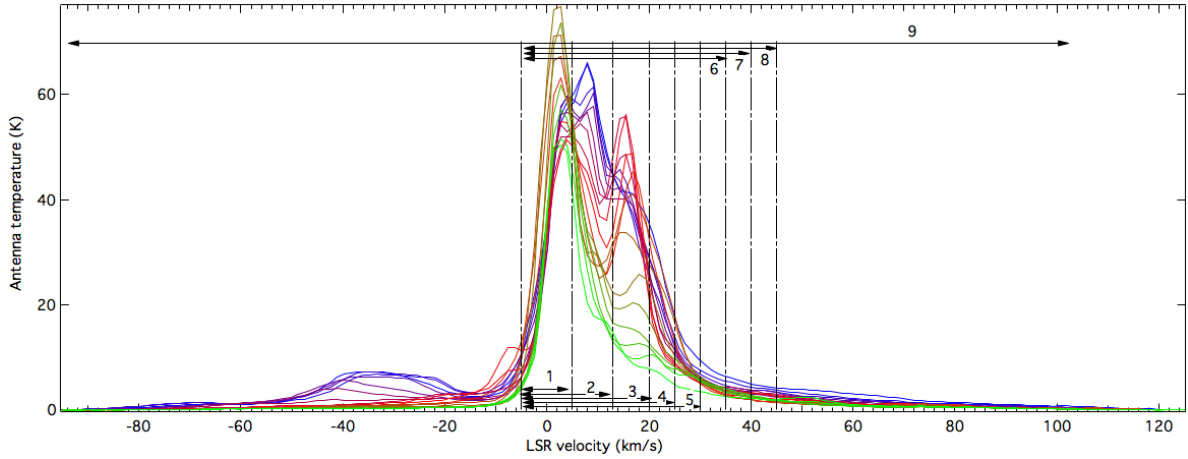


**Fig. 14.** New  $^{12}\text{CO}$  measurements: we show spectra interpolated within the spectral cube for each of the 29 directions distributed in each of the 19 XMM fields. The spectra are displaced by  $3$  K from one XMM field to the other. The isolated shadowing molecular cloud  $+9^\circ$  is characterized by a lower radial velocity than low-latitude clouds.

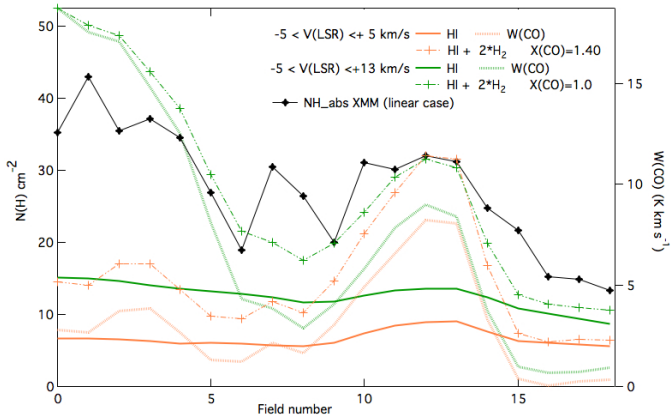
#### 4.2.2. Combined HI and CO data

Here we compare the observed  $N_{\text{Habs}}$  to total columns  $N(\text{HI}) + 2N(\text{H}_2)$  derived from HI  $21$  cm and CO spectra. We used the new  $^{12}\text{CO}$  survey data described above and the recently published HI  $21$  cm EBHIS survey data (Winkel et al. 2016) in a combined study of the IS gas in the directions of the XMM fields. EBHIS and CO spectra were interpolated for the same 29 directions within each XMM field, taking advantage of the high resolution of these two datasets, then resulting spectra were averaged over the FOV. Figure 15 displays the resulting average EBHIS HI spectra, which, as expected, reveal velocity components at lower and higher LSR velocities compared to CO. Based on the main components that appear in the spectra and the figure, we considered a series of LSR velocities intervals: (1)  $-5$  to  $+5$   $\text{km s}^{-1}$ , that is, nearby gas corresponding to the first and most intense component, which includes the  $+9^\circ$  molecular cloud, (2)  $-5$  to  $+13$   $\text{km s}^{-1}$ , that is, the first two components and the entire detected CO, (3)  $-5$  to  $+20$   $\text{km s}^{-1}$ , that is, the extension to the third strong component seen at all latitudes at velocities around  $+15$   $\text{km s}^{-1}$ , (4)  $-5$  to  $+25$   $\text{km s}^{-1}$ , including the apparent extension at high latitude of the previous component. We also considered the following extensions: (5)  $-5$  to  $+30$   $\text{km s}^{-1}$ , (6)  $-5$  to  $+35$   $\text{km s}^{-1}$ , (7)  $-5$  to  $+40$   $\text{km s}^{-1}$ , (8)  $-5$  to  $+45$   $\text{km s}^{-1}$ , and finally (9)  $-100$  to  $+100$   $\text{km s}^{-1}$ . For each velocity interval we were able to integrate the HI and CO profiles and computed the





**Fig. 15.** Averaged HI 21 cm EBHIS spectra for the 19 fields. Vertical lines indicate the limits of the velocity intervals used in Sect. 4.



**Fig. 16.** Comparison between NPS X-ray absorbing gas column for the 19 fields (black signs) and columns from radio emission data: we show field-of-view-averaged (H) columns deduced from HI 21 cm EBHIS data (solid lines) and from new CO data (dotted lines) for the first two velocity intervals (yellow and green). We also plot the total  $N(\text{H})$  from HI and CO assuming the highest  $X_{\text{CO}}$  conversion factor that is compatible with the X-ray absorption and the  $b = +9^\circ$  molecular cloud shadow (see text).

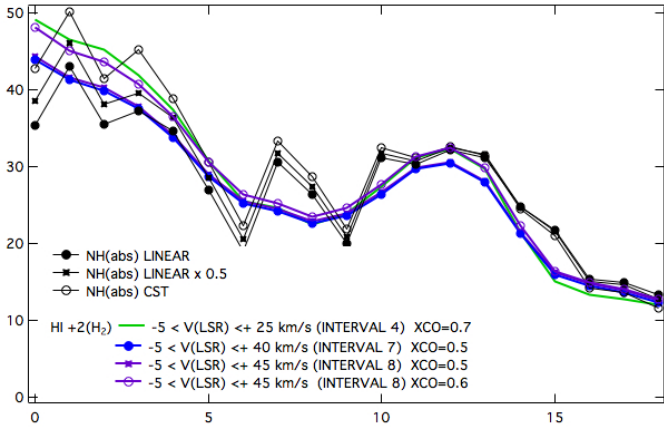
corresponding total gas profiles  $N(\text{H})$  in the 19 fields for a given  $X_{\text{CO}} = N(\text{H}_2)/W(\text{CO})$ .

Figure 16 shows  $W(\text{CO})$ ,  $N(\text{HI})$ , and  $N(\text{HI}) + 2 \times N(\text{H}_2)$  for the first two velocity intervals. There is a well-defined emission maximum in fields 10 to 14 that corresponds to the shadowing molecular cloud at  $b = 9^\circ$  that we previously discussed, whose LSR velocity is around  $-2 \text{ km s}^{-1}$ , that is to say, within this first interval. Its velocity, which is well defined thanks to the CO data, is slightly lower than that of the bulk of the Aquila Rift clouds and lower latitude clouds, but it seems geometrically associated to these lower latitude structures. Thus both the geometry and its low velocity suggest that it is a nearby structure linked to the main Aquila Rift clouds. Given its small distance, and subsequently its small size, it is extremely likely that the NPS source, a very wide structure, encompasses the totality of the shadow, providing a first firm upper limit on the  $X_{\text{CO}}$  factor for this cloud. Figure 16 shows the sum of  $N(\text{HI})$  and the converted  $W(\text{CO})$  for the first velocity interval  $-5, +5 \text{ km s}^{-1}$ , which at high latitude is dominated by the emission from this cloud, here for  $X_{\text{CO}} = 1.4 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$ . With this value,  $N(\text{H}) = N(\text{HI}) + 2N(\text{H}_2)$  reaches the X-ray absorbing H column,

implying that  $X_{\text{CO}} = 1.4 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$  is a strict upper limit to the  $X_{\text{CO}}$  factor for this cloud. On the other hand, it is clear from the  $N(\text{H})$  profiles in the figure that matching the absorption at all latitudes is firmly precluded when we restrict the absorbing matter to this velocity interval, and that addition of gas at higher velocity is required. This is in agreement with the absorption velocities found in optical spectra of stars within 500 pc, showing heliocentric absorptions of between  $-20$  and  $-5 \text{ km s}^{-1}$ , that is, corresponding to LSR velocities in the first and second intervals ( $V(\text{LSR}) - V(\text{Helio}) \approx 15\text{--}17 \text{ km s}^{-1}$ ). Using the second, wider interval ( $-5, +13 \text{ km}^{-1}$ ), it is possible to find a much better match of the XMM absorptions, but this time,  $X_{\text{CO}}$  is more strongly limited and constrained below  $\approx 1.0 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$ , as shown in the figure. This value is much lower than the Galaxy-averaged value of  $\approx 1.8 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$  (Dame et al. 2001), but is on the order of the factors found from combined *Fermi-Planck* analyses of nearby clouds (see, e.g., Planck and Fermi Collaborations Int. XXVIII 2015; or Remy et al. 2015).

Figure 17 shows the extension of the  $N(\text{H})$  computations to more extended velocity intervals. In this figure we included the three fitted columns  $N_{\text{Habs}}$  found for the constant NPS flux case, the linearly varying flux with maximum slope, and the intermediate case with a reduced slope (see Sect. 2). It is clear from Figs. 16 and 17 that the global  $N_{\text{Habs}}$  profiles can be well reproduced for several velocity intervals and  $X_{\text{CO}}$  factors, but that none of the model profiles deduced from HI and CO reproduces their field-to-field variations perfectly well: for a few directions (especially fields 6 and 9) we obtain discrepant results that are impossible to account for by varying the velocity intervals or  $X_{\text{CO}}$  alone. This variability deserves further studies that are beyond the scope of this global analysis.

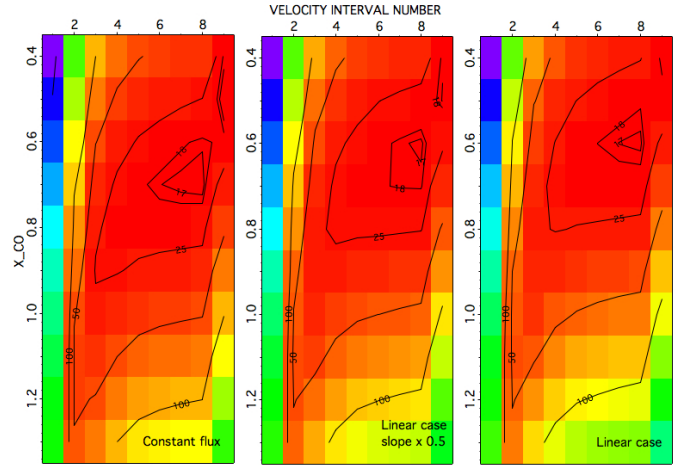
For each of the three hypotheses we used to derive the  $N_{\text{Habs}}$  profiles and for various values of the model parameters describing  $N(\text{HI}) + 2N(\text{H}_2)$ , a residual (sum over the 19 observing points of  $(\text{data-model})^2$ ) can be computed. We first searched for the lowest residual between  $N(\text{H})_{\text{abs}}$  and  $N(\text{HI}) + 2N(\text{H}_2)$  for the three cases by varying  $X_{\text{CO}}$  in steps of  $0.1 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$  and considering all velocity intervals. Figure 17 shows the corresponding total columns of H nuclei based on EBHIS and the new CO survey for the three best-fit solutions, using unweighted data-model deviations (see below). In the three cases  $X_{\text{CO}}$  is found to be very low, on the order of  $0.5\text{--}0.6 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$ , and velocity intervals with an upper limit beyond  $+35 \text{ km s}^{-1}$



**Fig. 17.** Comparison between NPS X-ray absorbing gas column for the 19 fields (black diamonds) and total columns from HI and CO for wider velocity intervals. Solutions that best match the three fitted  $N_{\text{H(}abs)}$  cases are indicated by the same markers as in the corresponding XMM profiles.

are required. We additionally display in Fig. 17 a profile corresponding to the  $-5, +25 \text{ km s}^{-1}$  velocity interval (4) and  $X_{\text{CO}} = 0.7 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$ . It is clear from the figure that such a solution, although not favored by minimization of residuals, is also quite close to the  $N_{\text{H(}abs)}$  profile for the constant flux case. From velocity intervals 3 to 8 there is little change in the HI columns, as Fig. 15 shows, and no variation at all of H<sub>2</sub>, as can be derived from the CO spectra. As a consequence, all profiles are very similar, and it is difficult to preclude a solution, especially given the discrepancies we previously noted for several directions.

Despite these ambiguities, it is informative to formally study the uncertainties on the parameters  $X_{\text{CO}}$  and the velocity range. We repeat that the values of  $N_{\text{Habs}}$  were determined for various hypotheses about the variation of the X ray emission with latitude: constant, linear variations with two different slopes. For a given hypothesis, the formal error bar on  $N_{\text{Habs}}$  derived from the X-ray analysis is very small, on the order of 1. to  $1.5 \times 10^{20} \text{ cm}^{-2}$ , as can be seen in Tables 1 and 2, because the number of data points in the spectra is very large. However, given the uncertainty on the hypothesis on the latitudinal variation of the flux, the order of magnitude of the systematic error bar on  $N_{\text{Habs}}$  can be estimated by the range of  $N_{\text{Habs}}$  values found for the various hypotheses, that is, between 0.2 and  $8 \times 10^{20} \text{ cm}^{-2}$  depending on the direction (see Fig. 17). Accordingly, we estimated the total uncertainty for each direction as the quadratic sum of the statistical error bar for each fit solution and the difference between the two extreme  $N_{\text{Habs}}$  solutions (namely the constant case and the linear case), scaled by a coefficient  $\alpha$  on the order of one. A classical way of estimating the actual total uncertainty it is to determine the magnitude of error bar that provides a sum of weighted residuals equal to the number of data points minus the number of free parameters, the velocity interval number, and  $X_{\text{CO}}$ , which here is  $19 - 2 = 17$ . We therefore adjusted  $\alpha$  in such a way that this condition was fulfilled, which gives  $\alpha = 1.35, 1.40,$  and  $1.47$  for the three cases. The corresponding sum of weighted residuals for the nine velocity intervals and  $X_{\text{CO}}$  between 0.4 and 1.3 is shown in Fig. 18. For the three cases shown in Fig. 18, the residuals are lowest for velocity intervals 7 and 8, that is, for the widest velocity intervals we considered, except for the last one ( $-100, +100 \text{ km s}^{-1}$ ). In the case of random Gaussian errors, the one-sigma domain of model parameters  $X_{\text{CO}}$  and velocity



**Fig. 18.** Weighted residuals between the measured absorbing columns and the HI+H<sub>2</sub> columns in the three cases, as a function of  $X_{\text{CO}}$  and the IS velocity interval.

range numbers would be delimited in Fig. 18 by the iso-contour  $\chi^2 = 18$  ( $\delta\chi^2 = 1$ ). This condition is not fulfilled here, as we discussed above; nevertheless, we can draw some conclusions from the global pattern. The figure shows that three velocity intervals are firmly excluded: intervals 1, 2, and 9. Interval 3 is likewise significantly disfavored, but except for these intervals, the  $\chi^2$  varies very weakly. This is a consequence of the very small amount of atomic gas between  $+25$  and  $+45 \text{ km s}^{-1}$ . According to Galactic rotation models (see, e.g., Vallée 2008), the LSR radial velocity of Sagittarius Arm gas is  $\approx +20, +30 \text{ km s}^{-1}$ , which implies that the NPS source may be located beyond the Arm. Figures 15 and 18 show that velocity interval 9 is precluded essentially because of the significant contribution of gas at high negative velocities ( $\leq -20 \text{ km s}^{-1}$ ). Galactic rotation models locate gas like this at very large distances, well beyond the Galactic center. This suggests that the NPS is, as expected, closer than the center. On the other hand, the amount of gas between  $+45$  and  $+100 \text{ km s}^{-1}$  is also very small, which means that it is difficult to preclude a contribution from far beyond Sagittarius, and an even more distant NPS.

It is also clear from Fig. 18 that values of  $X_{\text{CO}}$  higher than  $0.75 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$  are very unlikely, as the residuals increase very strongly beyond this value. The highest value of  $\approx 0.75 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$  is found for the constant flux case, while for the two cases that are favored by X-ray spectral fitting the upper limit is lower, on the order of  $0.60 \times 10^{20} \text{ cm}^{-2} \text{ K}^{-1} \text{ km}^{-1} \text{ s}$ . This confirms the low CO-H<sub>2</sub> conversion factor associated with the nearby molecular clouds inferred from *Fermi-Planck* analyses.

## 5. Conclusions and discussion

We analyzed the southern terminus of the X-ray NPS with unprecedented detail and spatial resolution by means of a dedicated series of *XMM-Newton* measurements. Spectral fitting to the whole dataset provided the X-ray absorbing columns of IS gas in 19 directions spanning a Galactic latitude interval of 5.5 degrees at  $l \approx 29^\circ$ . The absorbing columns vary from 1.3 to  $4.3 \times 10^{21} \text{ cm}^{-2}$  between  $b = +11^\circ 1'$  and  $b = +5^\circ 6'$ . A shadowing cloud visible in CO maps is clearly detected at  $b \approx +9^\circ$ . These measurements allowed for the first time a detailed comparison with tracers of the IS dust gas and gas distribution, both

distance-limited data and total LOS emission-based data, in particular a dedicated radio  $^{12}\text{CO}$  high-resolution survey and ground-based absorption measurements in support of the XMM program. We obtained the following results:

- 1) There is compelling evidence from surface brightnesses, hardness ratios, and absorbing columns that the southern terminus of the North Spur is fully absorption bounded and consequently that the source extends farther out toward the plane.
- 2) For the three stars located in the XMM fields, estimated IS gas columns  $N_{\text{Habs}}$  deduced from absorption lines are below the X-ray absorbing columns, implying a first determination of the minimal distance to the NPS inner boundary of 260 pc.
- 3) The set of target stars was used to validate local dust maps in the NPS area and to estimate the integrated reddening along the XMM directions up to 300 pc. Reddenings were converted into gas columns that are compared with  $N_{\text{Habs}}$  for the 19 XMM fields. This comparison showed that the distance to the X-ray source front is clearly beyond 300 pc.
- 4) Independently, the comparison with the larger scale Pan-STARRS (PS) 3D dust maps (Green et al. 2015) also implies a minimal distance to the NPS of at least 300 pc, in agreement with evidence from recent studies based on other X-ray data, 3D tomography of dust, or radio polarization (Sofue 2015; Puspitarini et al. 2014; Wolleben 2007). The comparison favors a larger distance, and uncertainties on the calibrations and the conversion factors allow for as much as 4 kpc. The low amount of reddening between 600 pc and 4 kpc (see Fig. 11) precludes a precise upper limit. On the other hand, a large distance is independently favored by the intercomparison of the correlations between  $N_{\text{Habs}}$  and the PS color excesses at various distances. As a matter of fact, the correlation improved when the sightline length was increased from 400 pc to 4 kpc. Part of the effect may be linked to the improved statistics of the PS measure, however, and more work is needed to distinguish the two effects.
- 5) The absorbing column  $N_{\text{Habs}}$  latitude profile deduced from the X-ray spectral fitting was found to correlate more tightly with the reddening deduced from *Planck* dust optical depths than with any distance-limited reddening, suggesting that the emission originates beyond a very large part of the matter that the *Planck* emission traces, that is, potentially several kpc away. In terms of absolute values, the correspondence between the X-ray absorbing columns and *Planck*  $\tau_{353\text{ GHz}}$ -based reddenings results in  $N(\text{H})/E(B-V) = 4.1\text{ cm}^{-2}\text{ mag}^{-1}$ , which matches recent determinations from *Fermi* and *Planck* data joint analyses well (Planck and Fermi Collaborations Int. XXVIII 2015).
- 6) The existence of a shadowing molecular cloud at  $b = +9^\circ$  allows the use of the X-ray absorbing columns and emission data to constrain the CO- $\text{H}_2$  conversion factor  $X_{\text{CO}}$  below  $1.0 \times 10^{20}\text{ cm}^{-2}\text{ K}^{-1}\text{ km}^{-1}\text{ s}$ .
- 7) The combination of X-ray absorbing columns and emission data for the entire XMM path constrains the average  $X_{\text{CO}}$  below  $0.75 \times 10^{20}\text{ cm}^{-2}\text{ K}^{-1}\text{ km}^{-1}\text{ s}$ , with the most probable value as low as  $0.6 \times 10^{20}\text{ cm}^{-2}\text{ K}^{-1}\text{ km}^{-1}\text{ s}$ . These values are unusually low, but  $X_{\text{CO}}$  factors of between 0.6 and  $1.1 \times 10^{20}\text{ cm}^{-2}\text{ K}^{-1}\text{ km}^{-1}\text{ s}$  have been recently derived from combined *Fermi-Planck* studies of local clouds (Planck and Fermi Collaborations Int. XXVIII 2015; Remy et al. 2015).
- 8) Absolute values of X-ray absorbing columns and their latitude profiles are compatible with HI and CO based columns

of gas at  $-5 \leq V_{\text{LSR}} \leq +25$  to  $+45\text{ km s}^{-1}$ , with the broadest interval being favored. The large uncertainty on the velocity interval of the gas absorbing the NPS is due to the very limited amount of gas at positive velocities between  $+25$  and  $+45\text{ km s}^{-1}$ . According to kinematical models of Galactic rotation (e.g., Vallée 2008), at  $l = +29^\circ$  the radial velocity of the Sagittarius Arm gas is on the order of  $+20, +30\text{ km s}^{-1}$ , resulting in a potential location of the NPS in front of or beyond this Arm, depending on the actual Sagittarius velocity range, with the latter case being more likely. Because there is a very small additional amount of gas faster than  $+45\text{ km s}^{-1}$ , a larger (positive) interval and a subsequent even more distant NPS cannot be precluded. However, as discussed in Sect. 4, more work is needed to understand the observed discrepancies and reduce the systematic uncertainties on the absorption profiles before a precise distance can be determined.

The shortest distance to the NPS derived from this study clearly demonstrates the absence of a link with the nearby Sco-Cen star-forming region. The high probability of a much larger NPS distance obtained from comparisons with dust and gas absorption and emission again raises the question of a possible link between the Spur and outflows from the inner Galaxy (*Fermi* bubbles, Galactic wind). Planck Collaboration XXV (2016) disfavor such a link based on the identification of northern and southern polarized emission structures with Loop I secondary arcs and the following geometrical arguments: the strong north-south asymmetry of the NPS, the absence of a pinched structure symmetric about the Galactic plane, and the absence of any trace of interaction between NPS/Loop I and the *Fermi* bubbles. However, it can be argued that NPS-Loop I and the *Fermi* bubbles may trace completely distinct episodes of nuclear activity, in which case geometrical arguments become weaker.

Refined analyses of the XMM spectra might help to constrain the NPS characteristics even more. Although absorbing columns are not expected to vary significantly after the constraints on the component parameters are relaxed, allowance for departures from thermal emission, for non-solar ion abundances and abundance ratios, and the inclusion of a dedicated modeling of the heliospheric charge-exchange emission are expected to place stronger constraints on the emission mechanisms. This is beyond the scope of the present work, which was mainly devoted to comparisons with existing IS data. From the point of view of the ISM distribution, future 3D mapping of Galactic clouds and their distance and velocity assignments may hopefully allow us to take more advantage of this study and better constrain source location of the NPS. In particular, future measurements with *Gaia* are expected to better constrain the cloud distribution from both parallax data and improved reddening estimates. In addition, accurate measurements of Sagittarius gas velocities are expected to allow using our results in a more conclusive way. *Gaia* finally may also be able to identify and locate young star associations in the inner Galaxy that may give rise to a giant super-bubble and a structure similar to NPS-Loop I. In the longer term, high-sensitivity, high spectral and spatial resolution Athena X-ray spectra (Nandra et al. 2013) are expected to shed additional light on the NPS spectral characteristics and the nature of its source.

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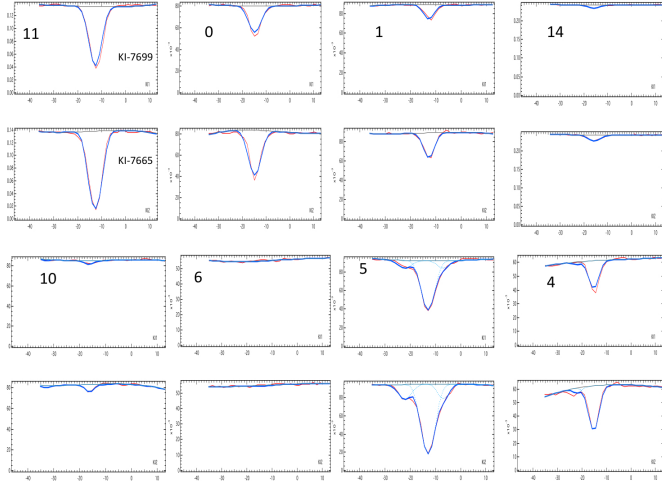
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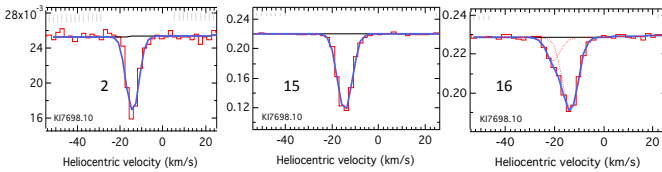


### Appendix A: interstellar KI absorption lines in nearby stellar spectra

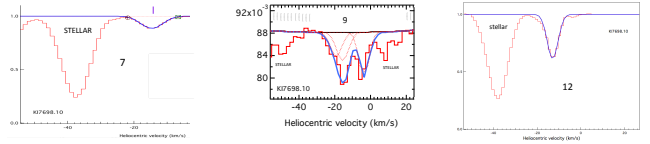
Figures A.1–A.4 display the various model adjustments performed to extract IS KI absorbing columns from the TBL-NARVAL optical spectra of the stars listed in Table 3. The fitting methods vary according to the stellar type, the telluric contamination, and the velocity shift between the stellar and IS KI lines. The resulting columns are listed in Table 3.



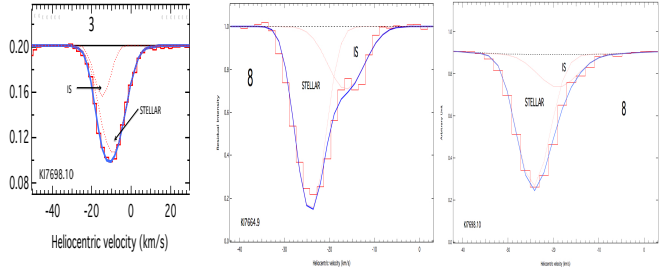
**Fig. A.1.** Determination of the interstellar KI columns along the line of sight to the nearby stars of Table 3. Numbers follow the convention of the table. Velocities are heliocentric. Flux units are arbitrary. For the eight early-type stars shown in the figure we performed a simultaneous fit of the two KI transitions.



**Fig. A.2.** Same as Fig. A.1 for three early-type stars that are strongly contaminated by telluric absorption. KI was determined from KI 7698 Å alone.



**Fig. A.3.** Same as Fig. A.1 for three late-type stars with a large velocity shift between the stellar KI line and the IS KI lines. KI was determined from KI 7698 Å alone. For star 9 there is a redshifted partially overlapping and weak stellar line that is not included in the adjustment.



**Fig. A.4.** Same as Fig. A.1 for two late-type stars with overlapping stellar and IS KI lines. The stellar line shift has been preliminarily determined from the strong sodium lines and was then imposed. The stellar KI line was fit as an artificial IS line with a high temperature. KI was determined from the 7698 Å line alone for star 3 and from the two transitions for star 8 (the two adjustments are shown in separate windows).