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Polyaromatic disordered carbon grains as carriers of the UV bump: Far-UV to mid-IR spectroscopy of laboratory analogs

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ABSTRACT

Context. A multiwavelength study of laboratory carbons with varying degrees of hydrogenation and sp^2 hybridization is required to characterize the structure of the carbonaceous carriers of interstellar and circumstellar extinction.

Aims. We study the spectral properties of carbonaceous dust analogs from the far-ultraviolet to the mid-infrared and correlate features in both spectral ranges to the aromatic/aliphatic degree.

Methods. Analogs to carbonaceous interstellar dust encountered in various phases of the interstellar medium have been prepared in the laboratory. These are amorphous hydrogenated carbons (a-C:H), analogs to the diffuse interstellar medium component, and soot particles, analogs to the polyaromatic component. Thin films ($d < 100$ nm) have been measured in transmission in the vacuum-ultraviolet (VUV; 120–210 nm) within the atmospheric pressure experiment (APEX) chamber of the DISCO beam line at the SOLEIL synchrotron radiation facility. Spectra of these films were further measured through the UV-Vis (210 nm–1 μ m) and in the mid-infrared (3–15 μ m).

Results. Tauc optical gaps, E_g , are derived from the visible spectra. The major spectral features are fitted through the VUV to the mid-infrared to obtain positions, full-widths at half maximum (FWHM), and integrated intensities. These are plotted against the position of the π - π^* electronic transitions peak. Unidentified or overlapping features in the UV are identified by correlations with complementary infrared data. A correlation between the optical gap and position of the π - π^* electronic transitions peak is found. The latter is also correlated to the position of the sp^3 carbon defect band at ~ 8 μ m, the aromatic C=C stretching mode position at ~ 6 μ m, and the H/C ratio.

Conclusions. Ultraviolet and infrared spectroscopy of structurally diverse carbon samples are used to constrain the nanostructural properties of carbon carriers of both circumstellar and interstellar extinction, such as the associated coherent lengths and the size of polyaromatic units. Our study suggests that carriers of the interstellar UV bump should exhibit infrared bands akin to the A/B classes of the aromatic infrared bands, while the circumstellar bump carriers should exhibit bands corresponding to the B/C classes.

Key words. dust, extinction – infrared: ISM – galaxies: ISM – methods: laboratory: solid state – ultraviolet: ISM

1. Introduction

The diversity of carbonaceous solids in interstellar and circumstellar environments in our Galaxy and in extra-galactic lines of sight has been the subject of numerous theoretical, observational, and experimental studies (Draine & Lee 1984; Greenberg & Li 1999; Pendleton & Allamandola 2002; Dartois et al. 2005; Pino et al. 2008; Godard & Dartois 2010; Duley & Hu 2012; Jones 2012a–c). Observations of interstellar extinction were first documented by Trumpler (1930). One prominent spectral feature in extinction curves in the Milky Way is a broad ultraviolet absorption bump consistently centered at 217.5 nm (4.6 μ m⁻¹). The UV bump was first observed in the 1960s by Stecher & Donn (1965), but its carrier is still not fully understood, although it has been generally attributed to electronic transitions in carbon. Fitzpatrick & Massa (1986) first parametrized the UV bump using data from the International Ultraviolet Explorer (IUE) toward 45 reddened OB Milky

Way stars. They found that the position of the bump was mildly variable while its full width at half maximum strongly variable; the lack of correlation between these parameters suggested that graphite was an improbable carrier. In Fitzpatrick & Massa (1988) they included the far ultraviolet (FUV) extinction, finding that a Lorentzian (for the UV bump), a FUV rise, and an underlying linear continuum were necessary components. This study was expanded to 78 stars, published as an atlas (Fitzpatrick & Massa 1990), and further extended to include 329 stars (Fitzpatrick & Massa 2005, 2007). In the latter, stellar atmosphere models were used to eliminate spectral mismatch errors in the curves, and showed that extinction parameters are highly sensitive to local conditions. The exact origin and variability of interstellar extinction and the nature of the dust carriers across different wavelength ranges is still a central question in cosmic dust research.

Astronomical observations in the infrared also play an important role in dust characterization. Unidentified infrared

emission bands (UIBs), also known as aromatic infrared bands (AIBs), are features between 3 and 20 μm (3.3, 6.2, 7.7, 8.6, and 11.2 μm) widely observed in a range of environments in our Galaxy (e.g., Rosenberg et al. 2014; Peeters 2014). These bands are attributed to the infrared fluorescence of polycyclic aromatic carriers, associated to the polycyclic aromatic hydrocarbon (PAH) hypothesis (Leger & Puget 1984; Puget et al. 1985; Allamandola et al. 1985; Puget & Leger 1989; Allamandola et al. 1989). However, alternative carriers have also been proposed, such as amorphous organic solids with mixed aromatic-aliphatic structures due to mixed sp^2 and sp^3 hybridizations. Observations of a single carbon star with a circumstellar disk by Sloan et al. (2007) concluded that its spectra was consistent with hydrocarbon mixtures containing both aromatic and aliphatic bonds. Acke et al. (2010) suggested that a strong UV flux reduces the aliphatic component and emphasizes the spectral signature of the aromatics in the emission spectra towards Herbig Ae stars. Yang et al. (2013) studied the ratio of the CH stretching modes at 3.3 and 3.4 μm in emission towards UIBs, and argued that their observed emitters are predominantly aromatic nanoparticles.

The spectroscopy of laboratory analogs to interstellar carbon dust has helped to unveil the origin of the UV bump. Draine & Lee (1984) proposed early on that crystalline graphite could fit the UV bump, and later on Jones (1990) argued that hydrogenated amorphous carbons (a-C:H) containing small polycyclic aromatic units may be suitable carriers. Observations of the aliphatic bands at 6.85 and 7.25 μm in absorption later revealed the presence of interstellar hydrogenated amorphous carbons (Pendleton & Allamandola 2002; Dartois et al. 2004). Laboratory hydrogenated carbons have since then been considered realistic carriers of the bump (Colangeli et al. 1997; Schnaiter et al. 1998; Furton et al. 1999; Dartois et al. 2004). Laboratory a-C:Hs at different degrees of hydrogenation have been compared to observed astronomical spectra (Dartois et al. 2007; Duley & Hu 2012). Soot nanoparticles prepared via laser pyrolysis (Jäger et al. 2006, 2008) are also candidates but only following the activation of π - π^* transitions by UV irradiation (Gadallah et al. 2011). Indeed, energetic thermal and non-thermal processes can alter carbonaceous dust in the diffuse interstellar medium (ISM). In photo-dissociation regions, photodesorption processes compete with hydrogenation to effectively alter their structure and optical properties. This has been experimentally explored via UV irradiation of thin a-C:H films (Scott & Duley 1996; Mennella et al. 1996, 2002; Alata et al. 2014, 2015), as well as via ion irradiation (Mennella et al. 2003; Godard et al. 2011). Recent optical models of interstellar dust (Jones 2012a–c; Jones et al. 2013) have shown that amorphous carbons extending over a wide range of optical gap energies can be used to fit the observed extinction from the FUV to the infrared. Russo et al. (2014) performed a UV study extending from 190 to 1100 nm on molecular-weight-sorted species prepared by combustion. Bescond et al. (2016) recently studied the optical properties of soots in the visible near-UV from 200–1100 nm and performed an inversion model to extract the size distribution and fractal dimension of soots. Few laboratory spectroscopic studies on carbonaceous dust extending over a wide spectral range exist, due to experimental difficulties in accessing the vacuum-ultraviolet (VUV). In this range, different techniques have been used including electron energy loss spectroscopy (EELS; Fink et al. 1984; Zubko et al. 1996) and synchrotron VUV sources (Colangeli et al. 1993; Gavilan et al. 2016).

For this study, we produced in the laboratory carbonaceous solids of mixed aromatic and aliphatic structures at different degrees of hydrogenation. By measuring them over a large spectral range (from the mid-infrared to the VUV), we expect to unveil relations between vibrational modes and electronic transitions, both observable signatures of interstellar dust, extending our previous findings in Gavilan et al. (2016). We aim to constrain the carbonaceous carriers of the interstellar UV bump and the circumstellar (CSM) UV bump and shed light on the spectroscopic decomposition of UV extinction based on our laboratory data. For the ISM, our work will be compared to the parameters of the interstellar 217.5 nm UV bump (Fitzpatrick & Massa 2005, 2007). For the CSM, our work will be compared against the very few UV observations of circumstellar envelopes, in this case of H-poor post-asymptotic giant branch (AGB) and (hot) R Cor Borealis stars (Drilling et al. 1997; Blanco et al. 1995; Waelkens et al. 1995). Although this CSM is not representative of H-rich carbon stars, which provide most of the galactic carbonaceous dust, we use these UV observations as a first constraint to the circumstellar UV bump.

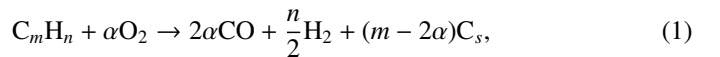
This article is organized as follows. In Sect. 2 we describe the experimental setups and sample preparation. In Sect. 3 we present the analysis of UV and IR bands. In Sect. 4 we discuss our analysis and astronomical implications, and in Sect. 5 we present our conclusions.

2. Experiments

The experimental setups used to prepare the carbonaceous samples have already been described in Gavilan et al. (2016). Here we provide further details relevant to the samples used in this study. Thin carbon films ($d < 100$ nm) deposited on MgF_2 windows were used for measurements between 120–1000 nm, and thicker carbon films ($d > 300$ nm) deposited on IR transparent KBr, CsI, and KCl windows for measurements in the mid-infrared.

SICAL-ICP, a low pressure ($\sim 10^{-2}$ mbar) radiofrequency (R.F.) plasma reactor was used to prepare the hydrogenated amorphous carbon samples. A 13.56 MHz R.F. power is inductively coupled to the plasma by impedance matching via a copper coil surrounding a glass cylinder (e.g., Dworschak 1990). While the a-C:H sample presented in Gavilan et al. (2016) was prepared using methane (CH_4) as the precursor gas, and an R.F. input power of 100 W, the a-C:H presented in this paper was prepared using a plasma fed by toluene (C_7H_8), and an R.F. input power of 150 W. Our goal in this study was to produce a more aromatic a-C:H. The higher R.F. power (150 versus 100 W) has an important influence on the structure of the produced solid hydrocarbon. Higher ion energies lead to preferential sputtering of hydrogen due to the lower binding energies of CH with respect to CC bonds.

We used the *Nanograins* combustion flame setup to produce soot nanoparticle films (Carpentier et al. 2012; Pino et al. 2015). In premixed flames, sooting is the result of competition between the rate of pyrolysis, growth of soot precursors, and oxidation rate (Mansurov 2005). The formation of soot nanoparticles can be described by the following equation,



where m and n are positive integers related to the hydrocarbon gas precursor type, s is a positive integer related to the number of carbon atoms in a solid soot nanoparticle, and α is a positive rational for the oxidation coefficient. Theoretically, for soots to

form in an ethylene flame, Eq. (1) imposes $C/O > 1$. In practice, the sooting regime appears as soon as $C/O \sim 0.7$ at 40 mbar. Soots produced in an ethylene (C_2H_4) flame, first introduced in Gavilan et al. (2016), were prepared under the following conditions: C/O ratio = 1.05, height above the burner (HAB) = 30 mm. In this paper we present soots prepared under varying flame conditions and HAB ranging from 18 to 50 mm. We used the following C/O ratios and pressures: 0.82 (60 mbar), 1 (28 mbar), 1.05 (33–40 mbar), 1.3 (40 mbar). The deposition time of the films varied from approximately one to four minutes for the closest HAB of 25 mm to 1 to 20 s beyond 30 mm.

3. UV and IR spectroscopy

VUV spectra of all samples were recorded in the 120–250 nm range using the APEX branch of the DISCO (dichroism, imaging and spectroscopy for biology) beam line (Giuliani et al. 2009) at the SOLEIL synchrotron. These spectra were completed in the UV-Vis (210–1000 nm) measured with a spectrophotometer (Specord, Analytikjena). The overlapping energy range was used to combine the VUV and UV-Vis datasets. Infrared spectra of the same films were measured using an evacuated fourier transform infrared (FTIR) spectrometer (Bruker Vertex 80V) equipped with a KBr beam-splitter and a HgCdTe (mercury cadmium telluride, MCT) detector working in the 4000 to 600 cm^{-1} (2.5 to 15 μm) at 2 cm^{-1} resolution. The infrared spectra were corrected with a polynomial baseline to subtract the continuum absorption increasing towards higher wavenumbers due to electronic transitions that extend into the infrared. The spectroscopic analysis involved deconvolution of the main electronic bands for the VUV-UV spectra and for the vibrational modes between 3300 and 600 cm^{-1} for the infrared spectra using the IDL MPFIT routine based on the Levenberg-Marquardt non-linear least-squares minimization method (Markwardt 2009). The VUV spectra of a sample soot series are shown in Fig. 2. All VUV-MIR spectra and the respective deconvolved features are included in Appendix A.

3.1. Relationships among the UV fit parameters

UV-VUV spectra with enough S/N were selected and deconvolved using a Lorentzian and three Gaussian functions. We called these the L1, G1, G2, and G3 bands. L1 is the main UV absorption band, falling between 190 to 260 nm, and it is in most cases well-fitted by a single Lorentzian. G1 and G2 are neighboring peaks, appearing around 300 nm and 190 nm respectively. G3 is the large UV absorption at around 110 nm. No additional linear background was required for these fits. The resulting fit UV-VUV parameters are summarized in Table A.1 of the electronic version.

3.1.1. Determination of the optical gap

We used the visible spectral range (350–1000 nm) to determine the optical band gap of these materials, also known as the Tauc gap. This is often used to characterize the optical properties of amorphous semiconductors (Tauc et al. 1966). The Tauc gap is determined only by the π -states (mostly or totally delocalized) and is intrinsically related to the energy gap between the valence and conduction bands (Silva et al. 1998). The Tauc gap has also been used to generalize the optical properties of interstellar amorphous carbons (Duley & Whittet 1992; Jones 2012a). For our samples, the Tauc gap values were obtained by extrapolating

the linear region in an $(\alpha \times E)^{1/2}$ versus E plot, where α is the absorption coefficient obtained from the relation $\alpha = \tau/h$, where τ is the optical depth, h the film thickness, and E is in eV. For disordered carbon materials like those presented in this paper, the empirical linear dependance between α and E is expressed by,

$$(\alpha \times E)^{1/2} = B^{1/2}(E - E_g), \quad (2)$$

where B is a constant attributed to their thermal processing (Jones 2012a). To determine the Tauc gap from the visible-ultraviolet spectra, we chose the same region in the visible tail of the Tauc spectra where $2.4 < E$ (eV) < 3.2 to extract the linear fit extrapolated to zero absorption, and from 2.4–2.7 eV and 2.7–3.2 eV to determine the average systematic error bar typically used in soot extinction analysis (Adkins & Miller 2015). For the soot produced with the $C/O = 1.05$ flame, HAB = 34 mm, the fitting linear energy range was shifted to 2.5–3.3 eV, in order to avoid a negative gap value. While the thickness of the samples is assumed to be ~ 30 –60 nm as in Gavilan et al. (2016), this is taken as constant for all samples as it does not significantly affect the gap value determination, whose systematic uncertainty is dominated by the choice of linear energy range. Figure 1 shows the calculated gap values for each film. For the prepared samples, the value of the Tauc optical gap (E_g) increases from the least hydrogenated soots to the most hydrogenated a-C:H, that is, $0 < E_g < 2$ eV. The H/C ratio of each sample is obtained from infrared spectra and discussed in Sect. 3.4.

For the a-C:H prepared with a toluene precursor, the optical gap was found to be 1.95 eV, lower than the 2.5 eV estimated for the methane-precursor a-C:H of Gavilan et al. (2016). For the soot samples produced with the $C/O = 0.82$ flame, E_g decreases from 0.78 to 0.62 eV as the HAB increases. This is the same trend as for the $C/O = 1$ flame, where E_g decreases from 1.46 to 0.68, for the $C/O = 1.05$ flame with a decrease from 1.47 to 0.01, and for the $C/O = 1.3$ flame with a decrease from 1.6 to 1 eV. Jones (2012a) sequences the family of amorphous carbons based on the optical gap energy, building on previous work by Robertson (1986). Amorphous carbons that are hydrogen rich, called a-C:H, have $E_g \approx 1.2$ –2.5 eV, while the hydrogen poor carbons, called a-C, have $E_g \approx 0.4$ –0.7 eV. For our soot samples $E_g < 1.6$ eV, and $E_g = 1.95$ for our a-C:H. Thus, Robertson's classification can be further improved by our disordered carbon data.

3.1.2. E_g versus $\omega_c(L1)$

In Fig. 3 we display the position of the main UV absorption (fitted by a Lorentzian, called L1) against the optical gap, yielding the relation,

$$E_g[\text{eV}] = 0.76 \pm 0.3 \times \omega_c(L1)[\mu m^{-1}] - 3.1 \pm 0.5. \quad (3)$$

These two parameters have a linear Pearson correlation coefficient $r = 0.91$. The wide variation in the position of the L1 peak from 3.8 to 5.2 μm^{-1} (192 to 262 nm) is consistent with previous experimental works where the UV peak is centered between 180 and 280 nm (Llamas Jansa et al. 2003). Russo et al. (2014) found that the spectral features in the UV-visible range are mainly affected by the higher molecular weight species of combustion flame products. Because L1 is strongly correlated to the optical gap E_g , variations in the position of L1 can be attributed to changes in the nanostructure due to varying H/C, sp^3/sp^2 , and the size of the polyaromatic units, which will be discussed in the following sections. The Tauc optical gap model,

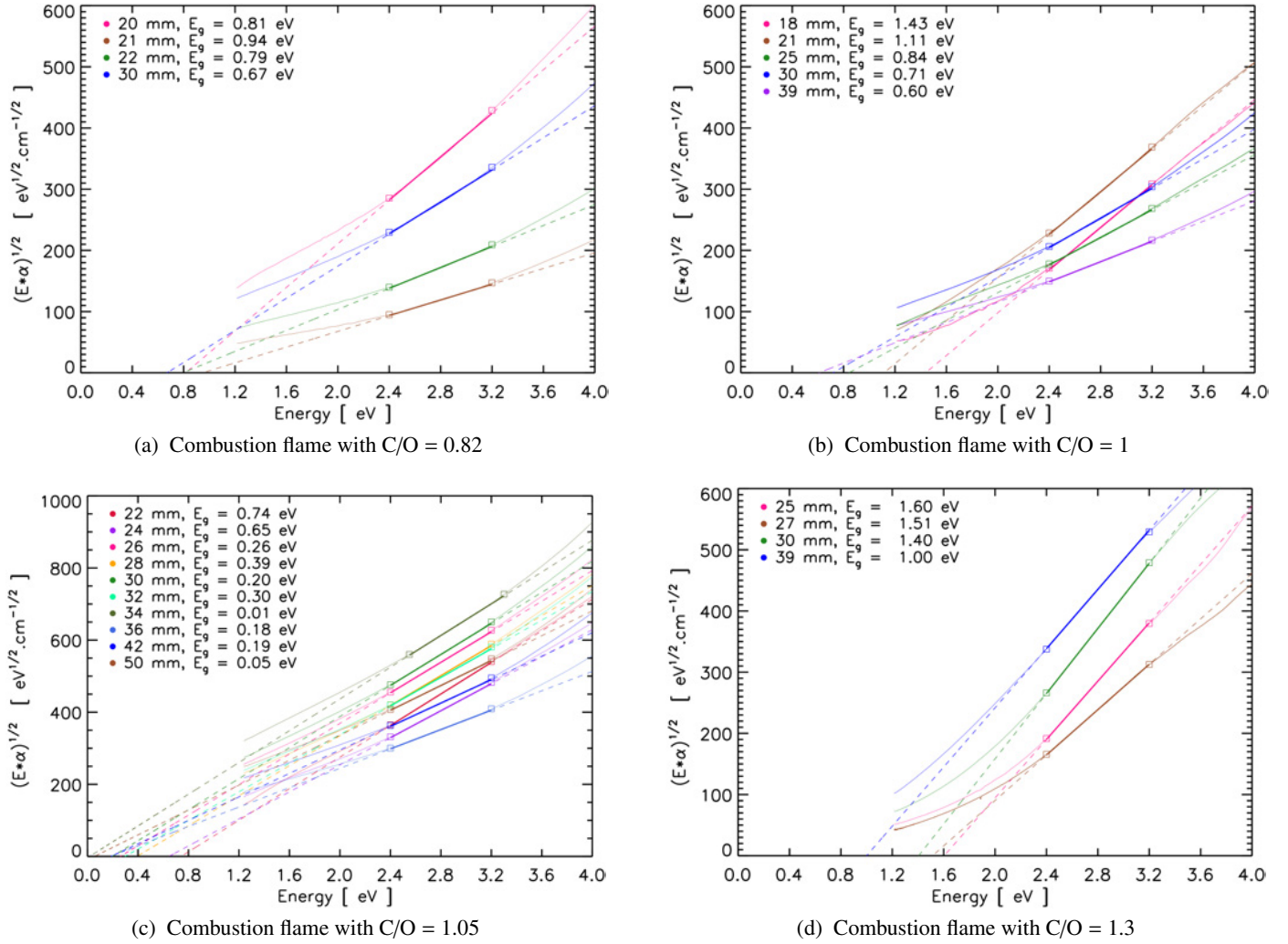


Fig. 1. Tauc optical gap extrapolation using the visible spectra for soot samples prepared with different C/O flame ratios and collected at varying heights above the burner (HAB in mm).

which has been used to give a complete description of the optical properties of carbonaceous materials (Robertson & O'Reilly 1987; Robertson 1991; Jones 2012a), appears to be well-suited for our soot samples.

From now on, our fitting parameters will be compared to the parameters of the interstellar 217.5 nm UV bump (Fitzpatrick & Massa 2005, 2007), that is, $\omega_c = 4.593 \mu\text{m}^{-1}$ with $\sigma = 0.0191 \mu\text{m}^{-1}$ and $FWHM = 0.77\text{--}1.29 \mu\text{m}^{-1}$ with $\sigma = 0.050 \mu\text{m}^{-1}$. We also compare our parameters to the average position of the circumstellar ~ 240 nm UV bump, as determined by observations of the circumstellar envelopes of C-rich stars like V348 Sagittarii (Drilling et al. 1997). For the average circumstellar UV bump, $\omega_c = 4.12 \mu\text{m}^{-1}$, limited from $4 \mu\text{m}^{-1}$ (Blanco et al. 1995) up to $4.4 \mu\text{m}^{-1}$ (Waelkens et al. 1995), and a $FWHM = 1.34 \mu\text{m}^{-1}$ with $\sigma = 0.06 \mu\text{m}^{-1}$ (Drilling et al. 1997). The spectral parameters of our samples will be preferably plotted against the position of L1, $\omega_c(L1)$, strongly correlated to E_g , as L1 (rather than E_g) is the direct spectroscopic parameter.

3.1.3. $\gamma(L1)$ versus $\omega_c(L1)$

As seen in Fig. 4, an anti-correlation is found between the FWHM $\gamma(L1)$ and $\omega_c(L1)$ ($r = -0.79$), while $\gamma(L1)$ versus E_g is also anti-correlated ($r = -0.72$). This result contrasts

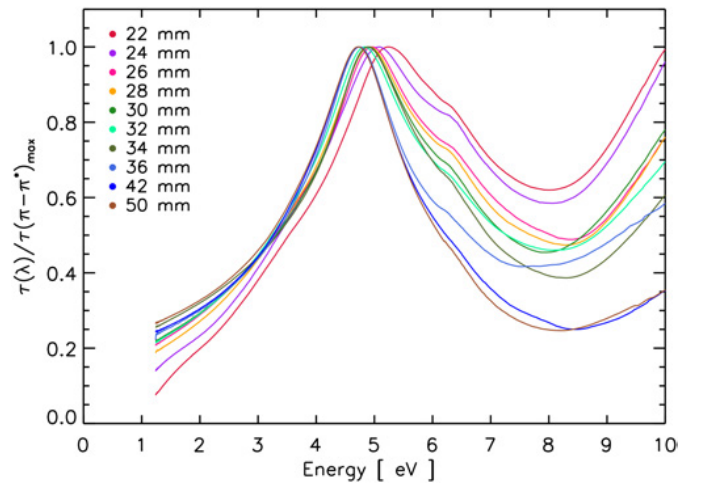


Fig. 2. VUV-Visible spectra of a series of soot nanoparticle films (C/O = 1.05, HAB = 22–50 mm). These have been normalized by the maximum optical depth of their respective $\pi\text{--}\pi^*$ electronic band.

with the non-correlation between both parameters in extinction curves found by Fitzpatrick & Massa (2005). The large scatter in FWHM has been suggested to be due to grain clustering

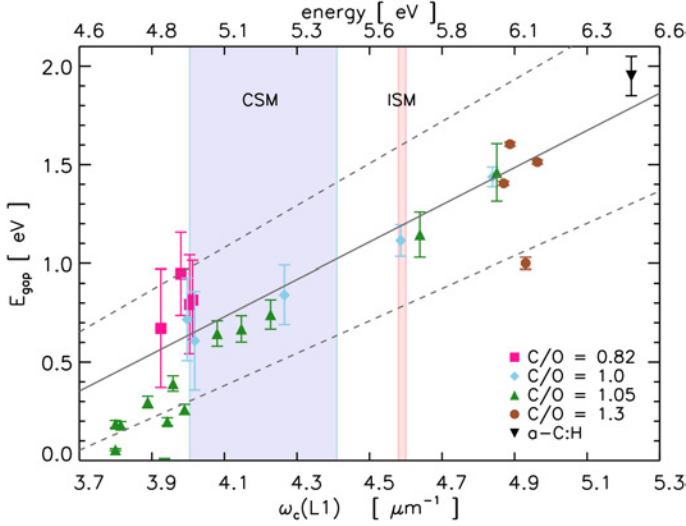


Fig. 3. Tauc optical gap, E_{gop} , for our laboratory samples versus the center of the π - π^* electronic transition, $\omega_c(\text{L1})$. The latter is used as our main spectroscopic parameter in the UV.

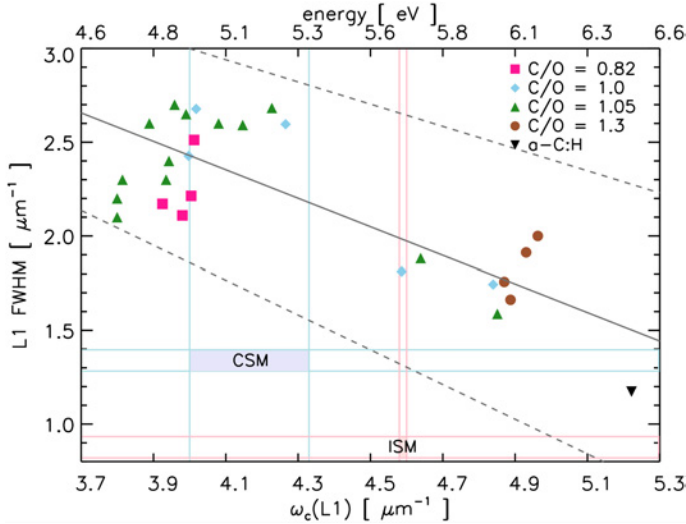


Fig. 4. FWHM of the π - π^* electronic transition, $\gamma(\text{L1})$, versus its center position, $\omega_c(\text{L1})$, showing a clear correlation, absent in interstellar extinction curves (Fitzpatrick & Massa 2005).

broadening (Schnaiter et al. 1998). From all our samples, only one has an L1 peak within the observed ISM FWHM parameter range, the a-C:H prepared with toluene, with $\gamma(\text{L1}) = 1.2 \mu\text{m}^{-1}$, but it is centered at $5.2 \mu\text{m}^{-1}$ and not at the expected $4.6 \mu\text{m}^{-1}$. However, if we consider a different parametrization of the interstellar UV bump (Sect. 4), two soot samples are more representative, C/O = 1 (HAB = 21 mm) and C/O = 1.05 (HAB = 18 mm).

3.1.4. $\omega_c(\text{G1}, \text{G2}, \text{G3})$ versus $\omega_c(\text{L1})$

In Fig. 5 we plot the positions of the main electronic transitions measured in the UV and VUV. G1 and G2 are absorption peaks in the neighborhood of L1. We can see that each electronic transition falls within a particular energy range for all samples: G1 peaks are constrained between 4.8 and $5.7 \mu\text{m}^{-1}$ while G2 peaks fall between 3 and $4 \mu\text{m}^{-1}$. The G1 and G2 transitions rarely appear simultaneously. The position of the G3 peak falls

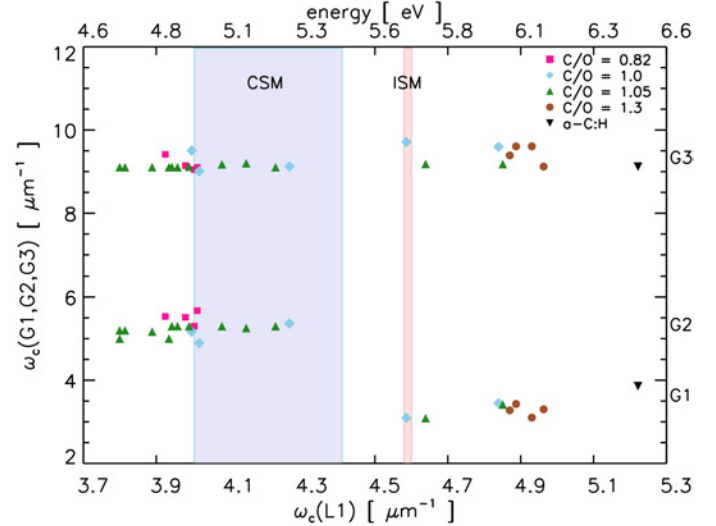


Fig. 5. Positions of the UV and VUV major electronic transitions, $\omega_c(\text{G1})$, $\omega_c(\text{G2})$, $\omega_c(\text{G3})$, versus the center of the π - π^* electronic transition $\omega_c(\text{L1})$. For all samples presented here the positions of these peaks fall at precise energies, with little dispersion, but their relative intensities tend to vary strongly.

consistently at an average of $9 \mu\text{m}^{-1}$. The ratio of the integrated G3 band to the L1 feature averages 2.0 ± 0.8 .

3.2. Relationships among the UV and mid-IR fit parameters

To complement the UV and VUV fit parameters, a detailed band analysis of the C-H stretching mode region (3100 – 2700 cm^{-1}) and bending mode region (1800 – 500 cm^{-1}) was performed. The out of plane (OOP) bending modes in the 900 to 700 cm^{-1} (11 to $15 \mu\text{m}$) are generally used for the classification of the aromatic ring edge structures (Russo et al. 2014), but these are only present in the soots prepared in flames with C/O = 0.82 and 1.05. The mid-infrared attributions are found in previous soot studies (Carpentier et al. 2012; Dartois et al. 2007). Generally, the bending mode region is ~ 1.5 – 4 times more absorbing than the stretching mode region for the soot samples. In a few cases, the integrated intensities of these modes were normalized to the most intense mode of the respective region for intercomparison between samples, to account for thickness variations between samples. The resulting fitted mid-infrared parameters are summarized in Table A.2 of the electronic version.

3.2.1. $\omega_c(\text{sp}^3 \text{ defect})$ versus $\omega_c(\text{arom. C=C stretch.})$

In Fig. 6, we show the strong correlation ($r^2 = 0.89$) between the $\omega_c(\text{aromatic C=C stretch.})$ and $\omega_c(\text{sp}^3 \text{ defect})$ band positions, confirming previous findings (Carpentier et al. 2012). This relation traces differences in shapes and structures of the polyaromatic units in the soot. For polyaromatic carbons, three astrophysical AIB classes are proposed based on their spectra in the 6 to $9 \mu\text{m}^{-1}$ region (Peeters et al. 2002). According to this classification, our laboratory soots correspond to classes B and C of the AIBs.

3.2.2. $\omega_c(\text{sp}^3 \text{ defect})$ & $\omega_c(\text{arom. C=C stretch.})$ versus $\omega_c(\text{L1})$

The sp^3 defect band is assigned to edge defects that can occur due to aliphatic cross-linking and as non-hexagonal rings

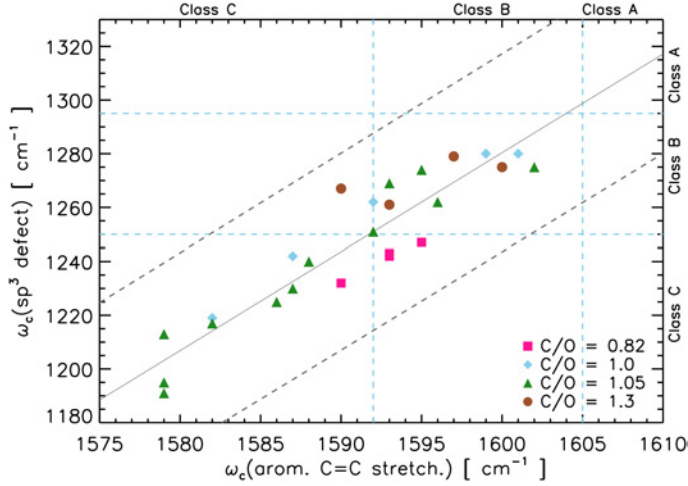


Fig. 6. Correlation between the carbon defect band, $\omega_c(\text{sp}^3 \text{ defect})$, versus the aromatic C=C stretching band, $\omega_c(\text{C}=\text{C})$, verifying previous findings in other soot samples (Carpentier et al. 2012).

carbon defects (Carpentier et al. 2012) and contribute to electronic intra-band states of the material. For our samples, this band appears between 1190–1280 cm^{-1} .

In Fig. 7a we plot the position of the sp^3 defect band versus the center of the L1 peak, which correlates well for all of our samples. From this relation, the L1 position of the CSM bump would correspond to an sp^3 defect located $\sim 1260 \text{ cm}^{-1}$ (i.e., class B/C carrier), while for the ISM bump we expect it to be found $\sim 1280 \text{ cm}^{-1}$ (i.e., class B carrier).

In Fig. 7b we plot the position of the aromatic C=C band versus the center of the L1 peak, also well correlated, as expected from the relation found in Fig. 6. For the a-C:H, the C=C band is found at 1625 cm^{-1} and for the soot with the highest E_g it is found at 1590 cm^{-1} . As the value of E_g decreases (and the H/C ratio of the sample decreases), the position shifts down to 1575 cm^{-1} . The correlation is linear up to L1 positions $\sim 4.7 \mu\text{m}^{-1}$. At higher energies this correlation shows a larger scatter. For the L1 position of the CSM bump, we would expect the C=C vibrational mode to be located at $\sim 1592 \text{ cm}^{-1}$ while for the ISM bump, we expect the defect to be found near 1600 cm^{-1} . It is then not surprising that the sp^3 defect and C=C band positions correlate with the position of L1 as they are all sensitive to the electrons' delocalizations in sp^2 conjugated systems.

3.3. Origin of UV features G1 and G2

3.3.1. $A(\text{G1})/A(\text{C}=\text{O})$ versus $\omega_c(\text{L1})$

G1 is a minor peak appearing for some soots ($\text{C/O} = 1.3$, for all distances from the burner, $\text{C/O} = 1.05$ for 14 and 18 mm) and for the a-C:H. To identify the origin of the G1 peak present in some UV spectra (for flames with C/O ratios = 1, 1.05 and 1.3), we plotted the Lorentzian position versus the ratio of the areas of $A(\text{G1})$ to $A(\text{C}=\text{O})$, showing that this feature appears for a certain type of carbon, whose main $\pi-\pi^*$ band appears at $\omega_c(\text{L1}) > 4.6 \mu\text{m}^{-1}$.

We choose the thickest films, where the carbonyl mode is clearly observed, to identify the origin of G1. These samples correspond to flames with $\text{C/O} = 1, 1.05$. The flame with $\text{C/O} = 1.3$ produced very thin films where oxidation is surface dominated. By plotting the ratio of the G1 feature to the L1 feature in the UV to the ratio of the integrated C=O to the C=C mode, a weak correlation arises. By following the evolution of

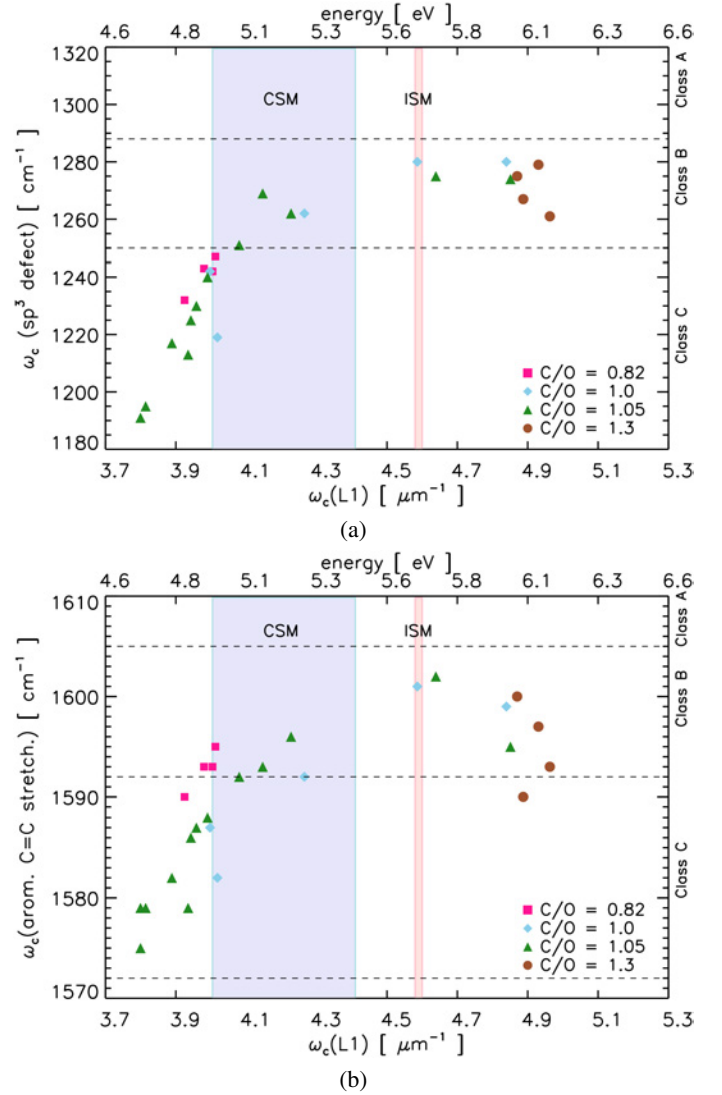


Fig. 7. Aromatic infrared bands (AIBs) in our laboratory soot samples. a) $\omega_c(\text{sp}^3 \text{ defect})$ versus $\omega_c(\text{L1})$. b) $\omega_c(\text{arom. C}=\text{C stretch.})$ versus $\omega_c(\text{L1})$. The laboratory found band positions for the sp^3 defect and aromatic C=C band are compared to the classification of AIBs by (Peeters et al. 2002). Our laboratory samples are classified as Class B and Class C AIBs. The inferred CSM UV bump carriers fall within Class B/C, and the ISM UV bump carriers approaches the A/B class.

the mid-infrared and UV spectra for the flame with $\text{C/O} = 1$ in the series of Figs. A.1–A.7, as soon as the intensity of $\text{C}=\text{O} \ll \text{C}=\text{C}$, the UV G1 peak disappears. The G1 peak can be assigned to the oxidation of the younger soots produced near the burner, as traced by the C=O infrared intensities.

3.3.2. $A(\text{G2})/A(\text{CH}_{\text{arom.}})$ versus $\omega_c(\text{L1})$

G2 is a minor peak appearing for some soots ($\text{C/O} = 1$, beyond 25–39 mm, $\text{C/O} = 1.05$, beyond 22 mm) at $\sim 5.3 \mu\text{m}^{-1}$ (180 nm). This peak appears as the soots are collected farther from the burner, that is, when they are more mature soots with polyaromatic structures or fullerene-like soots (Carpentier et al. 2012). Gadallah et al. (2011) and Llamas Jansa et al. (2003) attributed this peak to carbonyl vibrations. We propose a different attribution. We divided the integrated area of G2 by the integrated area of the aromatic $=\text{CH}$ stretching mode, and confirm that G2 appears only when the intensity of the sp^3 defect $\gg \text{C}=\text{C}$ band

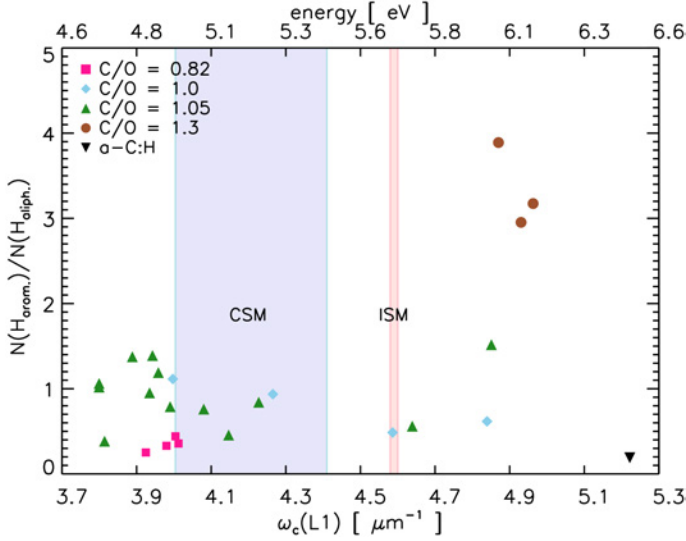


Fig. 8. Aromatic to aliphatic hydrogen ratio versus $\omega_c(L1)$, determined from the integrated infrared $=CH$ and $-CH_{2,3}$ stretching modes, corrected by the ratio of the corresponding oscillator strengths.

and for $\omega_c(L1) < 4.3 \mu m^{-1}$, corresponding to the lower Tauc gap materials, that is, low H/C soots where defects abound.

3.4. H/C ratios

The mid-infrared spectra of different carbonaceous samples reveal the carbon hybridization states and allow us to discriminate between aliphatic and aromatic hydrogen. The C-H stretching peaks in the $3100\text{--}2900 \text{ cm}^{-1}$ allow us to determine the aromatic to aliphatic hydrogen ratio. We take the ratio of the integrated $CH_{aromatic}$ ($\sim 3040 \text{ cm}^{-1}$) to the $CH_{aliphatic}$ modes (CH_2 asymmetric stretching mode, $\sim 2920 \text{ cm}^{-1}$ including the Fermi resonance and the CH_3 asymmetric mode at $\sim 2960 \text{ cm}^{-1}$) normalized by their oscillator strengths, $A(CH_{arom.}) = 1.9 \times 10^{-18} \text{ cm/group}$, $A(CH_{2,as.}) = 8.4 \times 10^{-18} \text{ cm/group}$, and $A(CH_{3,as.}) = 1.25 \times 10^{-17} \text{ cm/group}$ (Dartois et al. 2007). With the exception of a single sample ($C/O = 1$, $d = 18 \text{ mm}$), most spectra show a dominance of methylene ($=CH_2$) over methyl ($-CH_3$) groups. The resulting ratio is shown in Fig. 8 plotted against the central position of L1, resulting in a large scatter. This is also the case for the ratio of the areas of $CH_{arom.}$ to the sum of the areas of ($CH_{aliph.} + CH_{arom.}$). We note in particular the case for soots produced in the $C/O = 1.3$ flame, where the aromatic to aliphatic hydrogen ratio is an order of magnitude greater than in most other materials.

Carbonaceous materials containing hydrogen can be classified in terms of the hydrogen to carbon ratio (Robertson & O'Reilly 1987). To find a representative experimental H/C ratio for our samples, we have considered different combinations of vibrational modes in the infrared spectra. An empirical parameter of the H/C content in our samples is found by taking the ratios of the integrated intensities of the $CH_{arom.}$ mode to that of the aromatic $C=C$ stretching mode ($\sim 1600 \text{ cm}^{-1}$), (taking $A(C=C) \sim 0.175 \times 10^{-18} \text{ cm/group}$, Joblin et al. 1994). We call this the aromatic H/C ratio. Using the empirically found aromatic H/C ratios and the aromatic+aliphatic H/C ratio, that is, $N(CH_{arom.} + CH_{2,3,as.})/N(C=C)$, we calculate X_H , defined as,

$$X_H = \frac{H/C}{1 + H/C}. \quad (4)$$

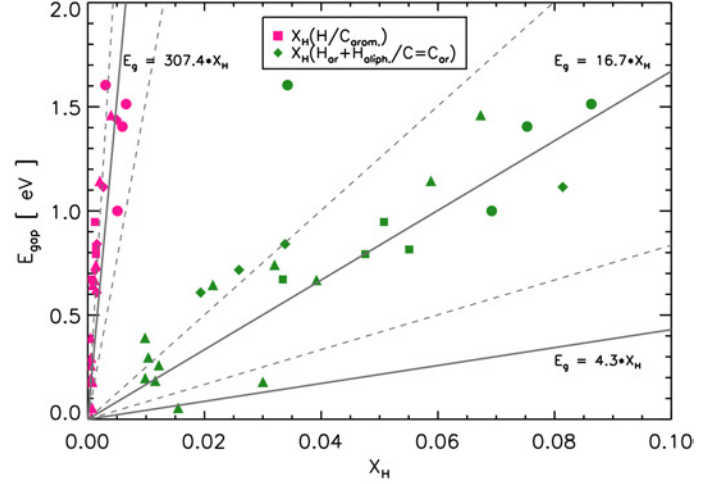


Fig. 9. Empirically found E_g versus X_H . We used the aromatic H/C ratio and aromatic+aliphatic H/C ratio, presented as two linear fits for the hydrogen poor (low E_g) regime (only soots). When both aromatic and aliphatic CH modes are included, our linear fit approaches the coefficient found for amorphous carbon in Tamor & Wu (1990). The shape of the symbols corresponds to those used in Fig. 8.

The result is shown in Fig. 9. Two linear trends are found for the hydrogen poor ($H/C \ll 1$) carbonaceous samples (soots).

Tamor & Wu (1990) provided an empirical linear relationship between the band gap and the H/C ratios of a-C:H materials with $E_g = 1.2\text{--}2 \text{ eV}$, where,

$$E_g(\text{eV}) \approx 4.3 X_H. \quad (5)$$

If we only consider the aromatic H/C ratio, the coefficient is about 100 times larger than the one found by Tamor & Wu (1990). If we include the aliphatic contribution, the fitted slope is still approximately four times larger. Although Eq. (5) is applicable to a-C:H materials with $E_g = 1.2\text{--}2 \text{ eV}$ and our soots have gaps between 0.01 to 1.60 eV, our materials show a linear relationship between the H/C ratio and the gap. This demonstrates that not only knowledge of hydrogen bonding but also knowledge of the carbon coordination is essential to understand the physical and chemical structure of carbonaceous materials. For low gap materials, Tamor & Wu's relationship no longer holds. The evolution of the properties between low gap and high gap materials should be taken into account in optical models of astrophysical carbon.

3.5. Carbon structure

3.5.1. L_a and N_{rings} versus $\omega_c(L1)$

The E_g parameter can be used to unveil the structural properties of aromatic absorbers in amorphous carbons (Robertson & O'Reilly 1987; Robertson 1991), including the average coherence length, L_a (size of the polyaromatic units), and the number of aromatic rings, N_R , for the largest clusters contributing to the electronic density of states, although not necessarily the most abundant ones (Gadallah et al. 2011). An empirical relation between L_a and E_g was proposed by Robertson (1991),

$$L_a(\text{nm}) = \left[\frac{0.77}{E_g(\text{eV})} \right]. \quad (6)$$

A similar empirical relation was proposed for a single graphene sheet where the number of six-fold aromatic rings,

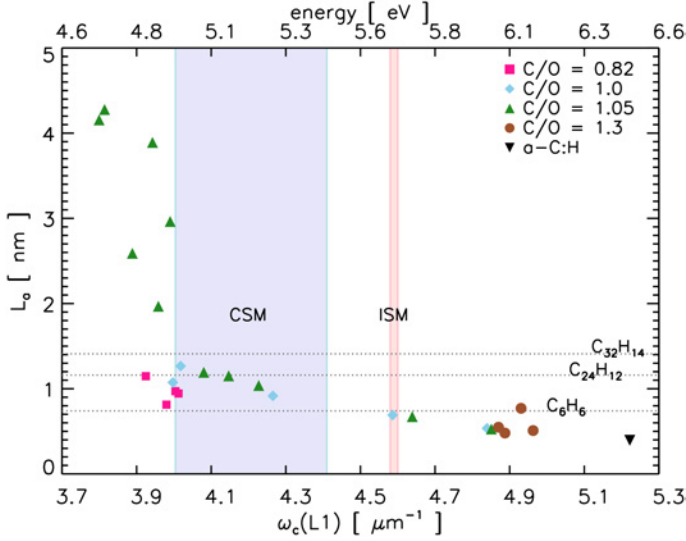


Fig. 10. Average coherence length, L_a , of polyaromatic units of selected laboratory carbons ($L_a < 5$ nm) versus $\omega_c(L1)$. For comparison, we included the coherence length of Benzene (C_6H_6), Coronene ($C_{24}H_{12}$), and Ovalene ($C_{32}H_{14}$) found at <http://pah.nist.gov>.

N_R , can be derived from E_g according to the model by Robertson & O'Reilly (1987), in which the gap energy decreases as a function of the size of the graphene layers via the following relation,

$$N_R = \left[\frac{5.8}{E_g(\text{eV})} \right]^2. \quad (7)$$

We have plotted the E_g derived N_R and L_a values against the position of the main UV bump (also strongly correlated to E_g) as shown in Figs. 10 and 11. When the position of the L1 peak shifts towards the UV, towards a larger band gap, L_a and N_R decrease. For the interstellar UV bump this would correspond to about 25 rings in a polyaromatic unit, and to 65 rings for the circumstellar bump. For the aromatic coherence length, we infer ~ 0.7 nm for the interstellar bump carrier and ~ 1 nm for the circumstellar bump carrier. Our soots consist of $L_a = 0.5$ to 3.3 nm nanoparticles with corresponding band gaps of $E_g = 0.1$ – 1.6 eV. The smallest aromatic length $L_a = 0.4$ nm was found for the a-C:H with $E_g = 2$ eV.

3.5.2. H/C versus L_a and N_{rings}

We use empirical relationships between the number of rings and the H/C ratio of our materials and compare them to purely aromatic carbons, that is, PAHs. In circum-PAHs, PAH with condensed six-membered rings (Violi et al. 2002), the number of rings grows in a sequence as follows: $N_{\text{rings}} = 1, 7, 19, 37, 61, 91$, with corresponding H/C ratios = 1, 0.5, 0.33, 0.25, 0.2, 0.17. To compare the structure of our soots to these PAHs, we will use the relation found in Jones (2012b) to evaluate the radius of the most compact aromatic domains, equivalent to L_a for the circum-PAHs,

$$a_R(\text{nm}) = 0.09 \times (2 \times N_R + \sqrt{N_R} + 0.5)^{0.5}. \quad (8)$$

As show in Fig. 12 our samples are, as expected, much less hydrogenated than circum-PAHs. For these soots, L_a ranges from 0.3 to 10 nm, while the number of rings in a polyaromatic unit spans from approximately one for the most hydrogenated soot

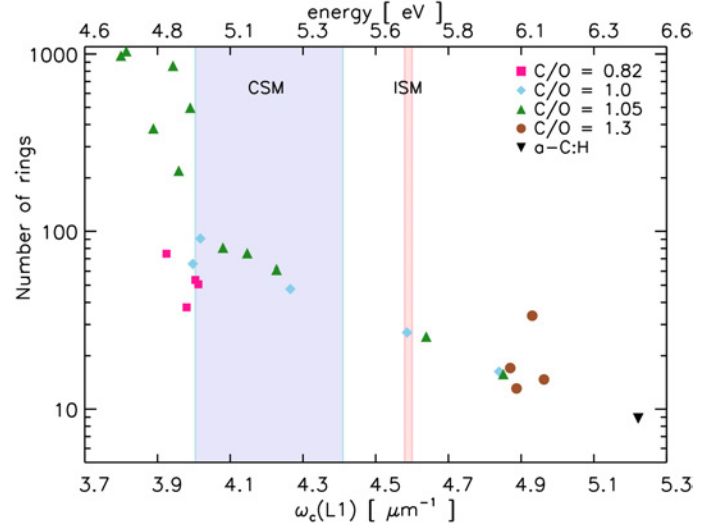


Fig. 11. Number of rings, N_R , in the polyaromatic units of our laboratory carbons versus $\omega_c(L1)$ suggesting $N_R \sim 60$ for the CSM bump and $N_R \sim 25$ for the ISM bump.

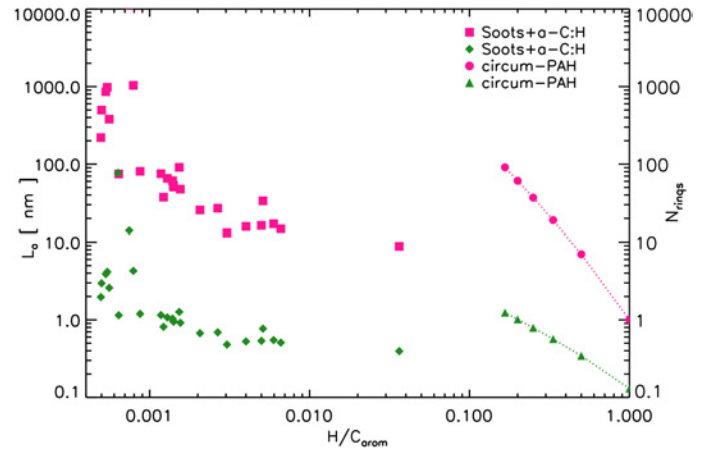


Fig. 12. Average coherence length, L_a , and number of rings, N_R , in the polyaromatic units of our laboratory carbons and circum-PAHs versus the corresponding aromatic H/C ratio.

sample to 1000 for the least hydrogenated. Both L_a and N_R decay exponentially as a function of the aromatic H/C ratio following the same trend as PAHs. Similarly, to find the number of carbon atoms in these polyaromatic units, we use the following equation from (Jones 2012b),

$$N_C = 2 \times N_R + 3.5 \times \sqrt{N_R} + 0.5. \quad (9)$$

As shown in Fig. 13, the polyaromatic units in our soot samples contain between 20 and 100 carbon atoms. These are in the same range as those discussed in Steglich et al. (2010), who proposed, in accordance with semi-empirical calculations, that a distribution of large PAH molecules with a mean size of 50–60 carbon atoms can produce a bump centered at ~ 217.5 nm.

4. Discussion

The transmission measurements from the VUV to the mid-infrared of several carbonaceous samples have allowed us to

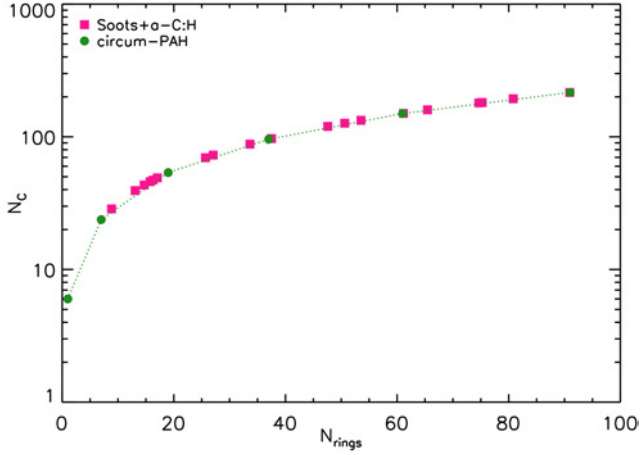


Fig. 13. Number of carbon atoms (N_C) versus the number of rings N_R in the polyaromatic units of the laboratory samples presented here (soots and a-C:H) and circum-PAHs, using Eq. (9).

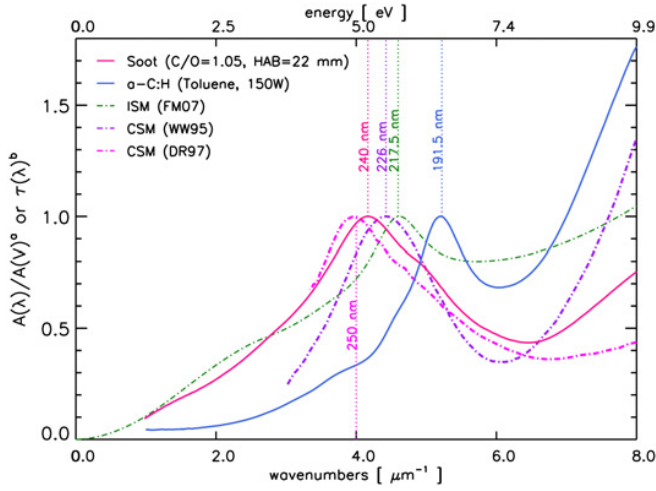


Fig. 14. Normalized laboratory soot and a-C:H visible-VUV spectra compared to the average ISM extinction curve (Fitzpatrick & Massa 2007, FM07), and the CSM extinction curves of the circumstellar envelope of HD 213985 (Waelkens et al. 1995, WW95) and V348 Sagittarii (Drilling et al. 1997, DR97), for which the original spectrum has been heavily filtered. All spectra are re-normalized to the maximum intensity of their π - π^* resonance.

perform a large spectroscopic study covering the main electronic and vibrational transitions. We find the following correlations:

- The position of the UV L1 feature (π - π^*) is correlated to the position of the sp^3 carbon defect, and to the position of the aromatic C=C stretching mode. This reveals a relation between electronic and vibrational band positions in soots.
- The Tauc optical gap, E_g , is correlated to the position of the π - π^* band, $\omega_c(L1)$. Both E_g and $\omega_c(L1)$ are relevant parameters allowing the classification of both soots and a-C:Hs.
- The H/C ratio of these carbonaceous samples is correlated to E_g and therefore to $\omega_c(L1)$. For soots, the empirical relationship between the H/C ratio and the gap differs from that found for highly hydrogenated carbonaceous materials.

The wide variety of our laboratory carbonaceous dust allows us to explore the spectral properties of the observed ISM and CSM extinction curves (e.g., Fig. 14), from the VUV to the IR:

- Spectroscopy in the VUV ($120 < \lambda < 200$ nm) has allowed us to identify an absorption band in the VUV present in all

our samples: the G3 band, with a fitted position centered at $\sim 9 \mu m^{-1}$ (~ 110 nm) and an integrated optical depth ratio of approximately two with respect to L1. G3 is attributed to σ - σ^* transitions and may contribute to the FUV rise identified in Fitzpatrick & Massa (2007).

- For the average ISM bump, centered at 217.5 nm, we infer the following spectral properties of the carrier: $E_g \sim 1.2$ eV, $\omega_c(G3) \sim 9.5 \mu m^{-1}$, $\omega_c(sp^3 \text{ defect}) \sim 1280 \text{ cm}^{-1}$, $\omega_c(C=C) \sim 1600 \text{ cm}^{-1}$.
- For the average CSM bump, centered ~ 226 – 250 nm, we infer the following spectral properties of the carrier: $E_g \sim 0.7$ eV, $\omega_c(G3) \sim 9 \mu m^{-1}$, $\omega_c(sp^3 \text{ defect}) \sim 1250 \text{ cm}^{-1}$, $\omega_c(C=C) \sim 1592 \text{ cm}^{-1}$.

The connection between ultraviolet extinction data and infrared emission has been possible thanks to early datasets from the IUE and the Infrared Astronomical Satellite (IRAS), observing towards the same lines of sight. Cox & Leene (1987) showed the absence of correlation between the mid-infrared excess and the normalized intensity of the UV bump, but later work by Cardelli et al. (1989) showed that their bump normalization was not appropriate. An extended study by Jenniskens & Desert (1993) and later on by Boulanger et al. (1994) revealed a link between mid-infrared emission bands (AIBs) and the UV bump absorption. In contrast, they found no relation between the AIBs and the FUV rise.

Correlations between the positions of UV and infrared absorption bands are found for the carbonaceous samples presented in this paper. These spectral parameters suggest that the carrier of the ISM bump is likely found at the frontier of class B/A of the AIBs, while the carrier of the CSM bump is found at the frontier of class B/C. Based on the Tauc optical gap and using available empirical formulae for carbons, the structure of the ISM UV bump carrier implies an average polyaromatic unit size ~ 0.7 nm and for the CSM UV bump carrier, ~ 1.1 nm, within nanoparticles.

While the position of the interstellar 217.5 nm bump ($4.592 \mu m^{-1}$) is stable, its FWHM varies (Fitzpatrick & Massa 2005). The FWHM of the π - π^* band of our carbonaceous analogs, $\gamma(L1)$, is correlated to the position $\omega_c(L1)$. None of our samples fulfill simultaneously the observed constraints on the UV bump (position, FWHM) in interstellar spectra (Fitzpatrick & Massa 2007). In Fig. 15, we take the average interstellar extinction curve and perform a deconvolution as done for our laboratory samples to explore the effect of the deconvolution on the same basis. A direct comparison of the FWHM of a laboratory soot ($1.88 \mu m^{-1}$) to the newly parametrized FWHM ($1.78 \mu m^{-1}$) from the average interstellar extinction curve indicates that laboratory polyaromatic solid carbons can approximate the average carrier of the UV bump. However, it is more likely than more than one kind of carbon is measured towards a line of sight, and that more hydrogenated carbons contribute strongly to the FUV rise as shown in Gavilan et al. (2016) and in Fig. 14, where our laboratory soot and a-C:H spectra are compared to extinction curves in the ISM and CSM.

5. Conclusions

We have produced laboratory analogs to interstellar carbonaceous dust characterized by their wide degree of hydrogenation and differences in polyaromatic structure. This influences the spectral signatures, in particular the position of the UV bump, attributed to π - π^* electronic transitions. Our carbonaceous analogs

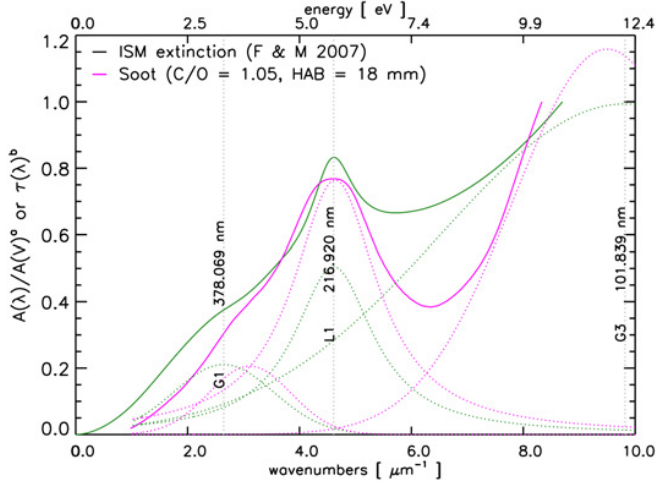


Fig. 15. Deconvolution of the average interstellar extinction curve (Fitzpatrick & Massa 2007) (“normalized to its maximum”) following the deconvolution procedure of laboratory soots described in Sect. 3.1. The L1 peak is centered at $\omega_c = 4.609 \mu\text{m}^{-1}$ with a $\text{FWHM} = 1.779 \mu\text{m}^{-1}$. While the position is unchanged with respect to the parametrization of Fitzpatrick & Massa (2007), the original FWHM value is almost doubled. For comparison, the spectra of a laboratory soot prepared in a flame with $\text{C/O} = 1.05$ and $\text{HAB} = 18 \text{ mm}$ is included, showing a similarly intense and wide $\pi\text{-}\pi^*$ transition, a G1 peak, and a slightly less intense FUV rise.

consist of soots, characterized by low H/C and E_g values. A hydrogenated a-C:H with $E_g \sim 2 \text{ eV}$ was also produced. The UV to mid-infrared spectroscopy has allowed us to find correlations between the electronic transitions and vibrational modes and to deduce structural properties such as the average length ($\sim 0.7\text{--}1.1 \text{ nm}$) of polyaromatic units within nanoparticles.

These laboratory correlations suggest that for the interstellar and circumstellar UV bump carriers, infrared signatures are expected in classes B/A and B/C of the AIBs respectively. More hydrogenated carbons contribute to the FUV rise in agreement with its weak correlation with the UV bump. Observations are then not only sensitive to local environmental conditions but also to the transition of dust from circum- to interstellar environments. As shown by our measurements, disordered polyaromatic carbonaceous grains are viable carriers of the observed UV interstellar and circumstellar extinction.

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References

Acke, B., Bouwman, J., Juhász, A., et al. 2010, *ApJ*, **718**, 558
 Adkins, E. M., & Miller, J. H. 2015, *Phys. Chem. Chem. Phys.*, **17**, 2686
 Alata, I., Cruz-Díaz, G. A., Muñoz Caro, G. M., & Dartois, E. 2014, *A&A*, **569**, A119
 Alata, I., Jallat, A., Gavilan, L., et al. 2015, *A&A*, **584**, A123
 Allamandola, L. J., Tielens, A. G. G. M., & Barker, J. R. 1985, *ApJ*, **290**, L25
 Allamandola, L. J., Tielens, A. G. G. M., & Barker, J. R. 1989, *ApJS*, **71**, 733
 Bescond, A., Yon, J., Ouf, F.-X., et al. 2016, *J. Aerosol Sci.*, **101**, 118
 Blanco, A., Fonti, S., & Orofino, V. 1995, *ApJ*, **448**, 339
 Boulanger, F., Prevot, M. L., & Gry, C. 1994, *A&A*, **284**, 956
 Cardelli, J. A., Clayton, G. C., & Mathis, J. S. 1989, *ApJ*, **345**, 245
 Carpentier, Y., Féraud, G., Dartois, E., et al. 2012, *A&A*, **548**, A40
 Colangeli, L., Mennella, V., Blanco, A., et al. 1993, *ApJ*, **418**, 435
 Colangeli, L., Bussolletti, E., Pestellini, C. C., et al. 1997, *Adv. Space Res.*, **20**, 1617
 Cox, P., & Leene, A. 1987, *A&A*, **174**, 203

Dartois, E., Muñoz Caro, G. M., Deboffle, D., & d’Hendecourt, L. 2004, *A&A*, **423**, L33
 Dartois, E., Muñoz Caro, G. M., Deboffle, D., Montagnac, G., & D’Hendecourt, L. 2005, *A&A*, **432**, 895
 Dartois, E., Geballe, T. R., Pino, T., et al. 2007, *A&A*, **463**, 635
 Draine, B. T., & Lee, H. M. 1984, *ApJ*, **285**, 89
 Drilling, J. S., Hecht, J. H., Clayton, G. C., et al. 1997, *ApJ*, **476**, 865
 Duley, W. W., & Hu, A. 2012, *ApJ*, **761**, 115
 Duley, W. W., & Whittet, D. C. B. 1992, *MNRAS*, **255**, 243
 Dworschak, W. 1990, *Thin Solid Films*, **189**, 257
 Fink, J., Müller-Heinzerling, T., Pflüger, J., et al. 1984, *Phys. Rev. B*, **30**, 4713
 Fitzpatrick, E. L., & Massa, D. 1986, *ApJ*, **307**, 286
 Fitzpatrick, E. L., & Massa, D. 1988, *ApJ*, **328**, 734
 Fitzpatrick, E. L., & Massa, D. 1990, *ApJS*, **72**, 163
 Fitzpatrick, E. L., & Massa, D. 2005, *AJ*, **130**, 1127
 Fitzpatrick, E. L., & Massa, D. 2007, *ApJ*, **663**, 320
 Furton, D. G., Laiho, J. W., & Witt, A. N. 1999, *ApJ*, **526**, 752
 Gadallah, K. A. K., Mutschke, H., & Jäger, C. 2011, *A&A*, **528**, A56
 Gavilan, L., Alata, I., Le, K. C., et al. 2016, *A&A*, **586**, A106
 Giuliani, A., Jamme, F., Rouam, V., et al. 2009, *J. Synchrotron Radiation*, **16**, 835
 Godard, M., & Dartois, E. 2010, *A&A*, **519**, A39
 Godard, M., Féraud, G., Chabot, M., et al. 2011, *A&A*, **529**, A146
 Greenberg, J. M., & Li, A. 1999, *Adv. Space Res.*, **24**, 497
 Jäger, C., Krasnokutski, S., Staicu, A., et al. 2006, *ApJS*, **166**, 557
 Jäger, C., Mutschke, H., Henning, T., & Huisken, F. 2008, *ApJ*, **689**, 249
 Jenniskens, P., & Desert, F. X. 1993, *A&A*, **275**, 549
 Joblin, C., D’Hendecourt, L., Leger, A., & Defourneau, D. 1994, *A&A*, **281**, 923
 Jones, A. P. 1990, *MNRAS*, **247**, 305
 Jones, A. P. 2012a, *A&A*, **545**, C2
 Jones, A. P. 2012b, *A&A*, **540**, A1
 Jones, A. P. 2012c, *A&A*, **542**, A98
 Jones, A. P., Fanciullo, L., Köhler, M., et al. 2013, *A&A*, **558**, A62
 Leger, A., & Puget, J. L. 1984, *A&A*, **137**, L5
 Llamas Jansa, I., Jäger, C., Mutschke, H., & Henning, T. 2003, in *Astrophysics of Dust*, ed. A. N. Witt
 Mansurov, Z. A. 2005, *Combustion, Explosion and Shock Waves*, **41**, 727
 Markwardt, C. B. 2009, *Proc. Astronomical Data Analysis Software and Systems XVIII*, **411**, 251
 Mennella, V., Colangeli, L., Palumbo, P., et al. 1996, *ApJ*, **464**, L191
 Mennella, V., Brucato, J. R., Colangeli, L., & Palumbo, P. 2002, *ApJ*, **569**, 531
 Mennella, V., Baratta, G. A., Esposito, A., Ferini, G., & Pendleton, Y. J. 2003, *ApJ*, **587**, 727
 Peeters, E. 2014, in *The Diffuse Interstellar Bands*, eds. J. Cami, & N. L. J. Cox, *IAU Symp.*, **297**, 187
 Peeters, E., Hony, S., Van Kerckhoven, C., et al. 2002, *A&A*, **390**, 1089
 Pendleton, Y. J., & Allamandola, L. J. 2002, *ApJS*, **138**, 75
 Pino, T., Dartois, E., Cao, A.-T., et al. 2008, *A&A*, **490**, 665
 Pino, T., Carpentier, Y., Féraud, G., et al. 2015, *Highlights of Astronomy*, **16**, 717
 Puget, J. L., & Leger, A. 1989, *ARA&A*, **27**, 161
 Puget, J. L., Leger, A., & Boulanger, F. 1985, *A&A*, **142**, L19
 Robertson, J. 1986, *Adv. Phys.*, **35**, 317
 Robertson, J. 1991, *J. Non Crystalline Solids*, **137**, 825
 Robertson, J., & O’Reilly, E. P. 1987, *Phys. Rev. B*, **35**, 2946
 Rosenberg, M. J. F., Berné, O., & Boersma, C. 2014, *A&A*, **566**, L4
 Russo, C., Stanzione, F., Tregrossi, A., & Ciajolo, A. 2014, *Carbon*, **74**, 127
 Schnaiter, M., Mutschke, H., Dorschner, J., Henning, T., & Salama, F. 1998, *ApJ*, **498**, 486
 Scott, A., & Duley, W. W. 1996, *ApJ*, **472**, L123
 Silva, S. R. P., Robertson, J., Milne, W. I., & Amaratunga, G. A. J., 1998, *Amorphous Carbon: State of the Art*, Proc. 1st Int. Specialist Meeting on Amorphous Carbon (smac ’97) (World Scientific Publishing Co. Pte. Ltd.)
 Sloan, G. C., Jura, M., Duley, W. W., et al. 2007, *ApJ*, **664**, 1144
 Stecher, T. P., & Donn, B. 1965, *ApJ*, **142**, 1681
 Steglich, M., Jäger, C., Rouillé, G., et al. 2010, *ApJ*, **712**, L16
 Tamor, M. A., & Wu, C. H. 1990, *J. Appl. Phys.*, **67**, 1007
 Tauc, J., Grigorovici, R., & Vancu, A. 1966, *Phys. Status Solid B-basic Solid State Physics*, **15**, 627
 Trumpler, R. J. 1930, *PASP*, **42**, 267
 Violi, A., Sarofim, A. F., & Truong, T. N. 2002, *Combustion Science and Technology*, **174**, 205
 Waelkens, C., Waters, L. B. F. M., van Winckel, H., & Daems, K. 1995, *Ap&SS*, **224**, 357
 Yang, X. J., Glaser, R., Li, A., & Zhong, J. X. 2013, *ApJ*, **776**, 110
 Zubko, V. G., Mennella, V., Colangeli, L., & Bussolletti, E. 1996, *MNRAS*, **282**, 1321

(a) 20 mm

(b) 20 mm

(c) 20 mm

(d) 21 mm

(e) 21 mm

(f) 21 mm

(g) 22 mm

(h) 22 mm

(i) 22 mm

(j) 30 mm

(k) 30 mm

(l) 30 mm

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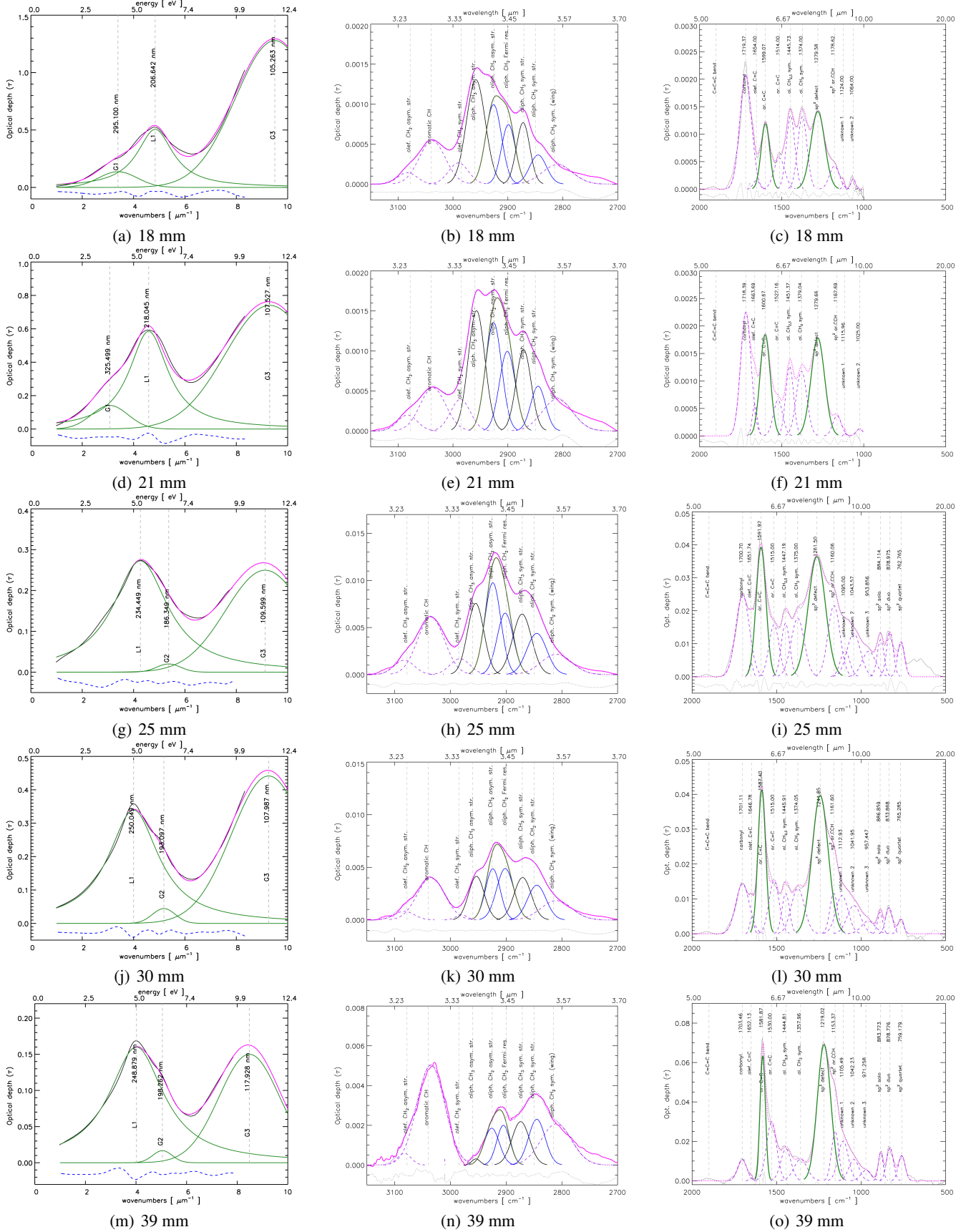
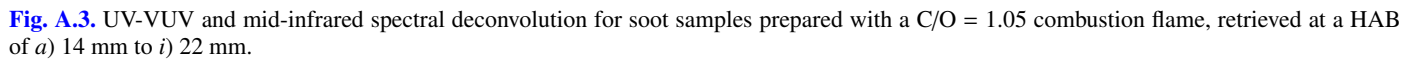


Fig. A.2. UV-VUV and mid-infrared spectral deconvolution for soot samples prepared with a C/O = 1 combustion flame, retrieved at a HAB of a) 18 mm to o) 39 mm.



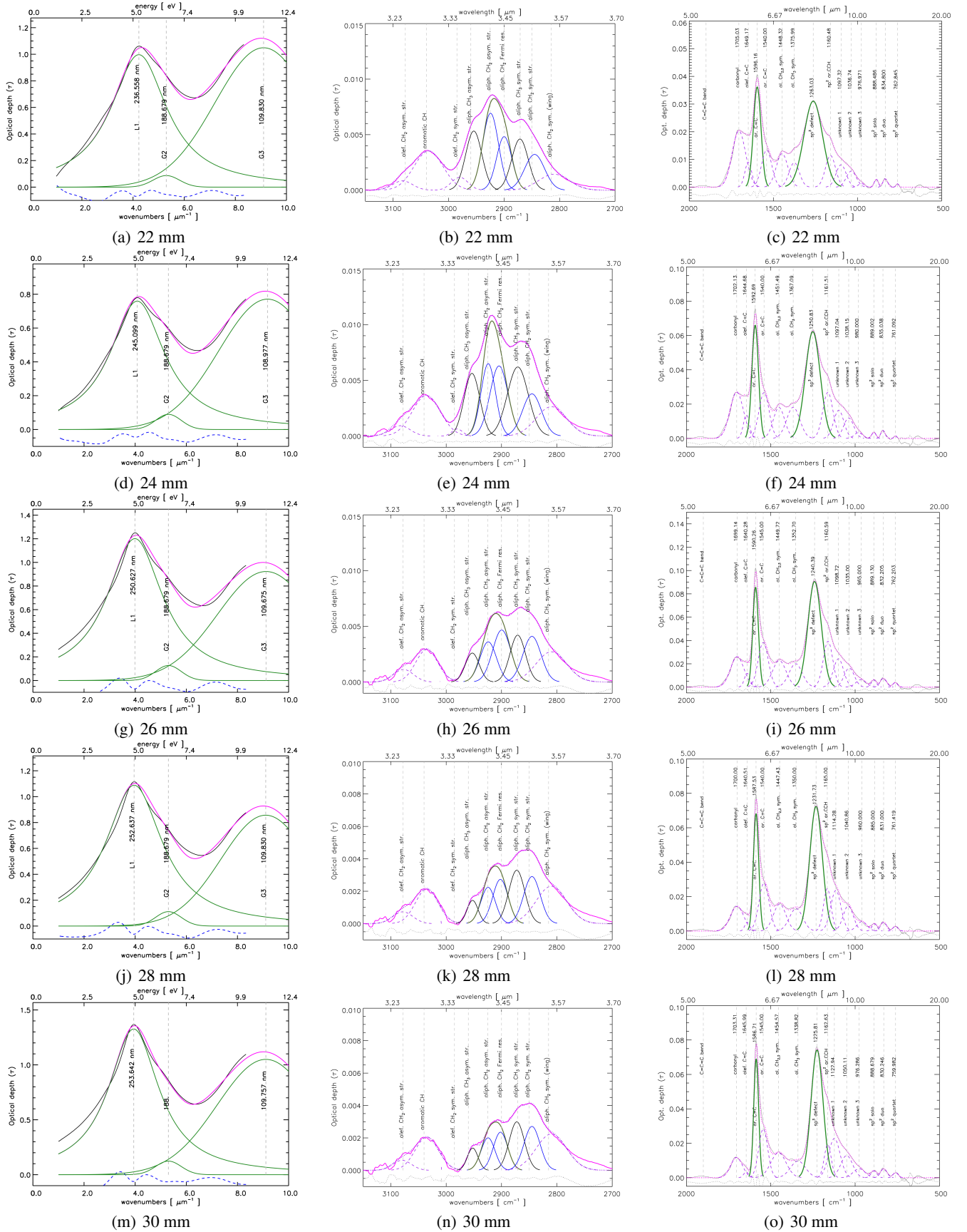


Fig. A.4. UV-VUV and mid-infrared spectral deconvolution for soot samples prepared with a C/O = 1.05 combustion flame, retrieved at a HAB of a) 22 mm to o) 30 mm.

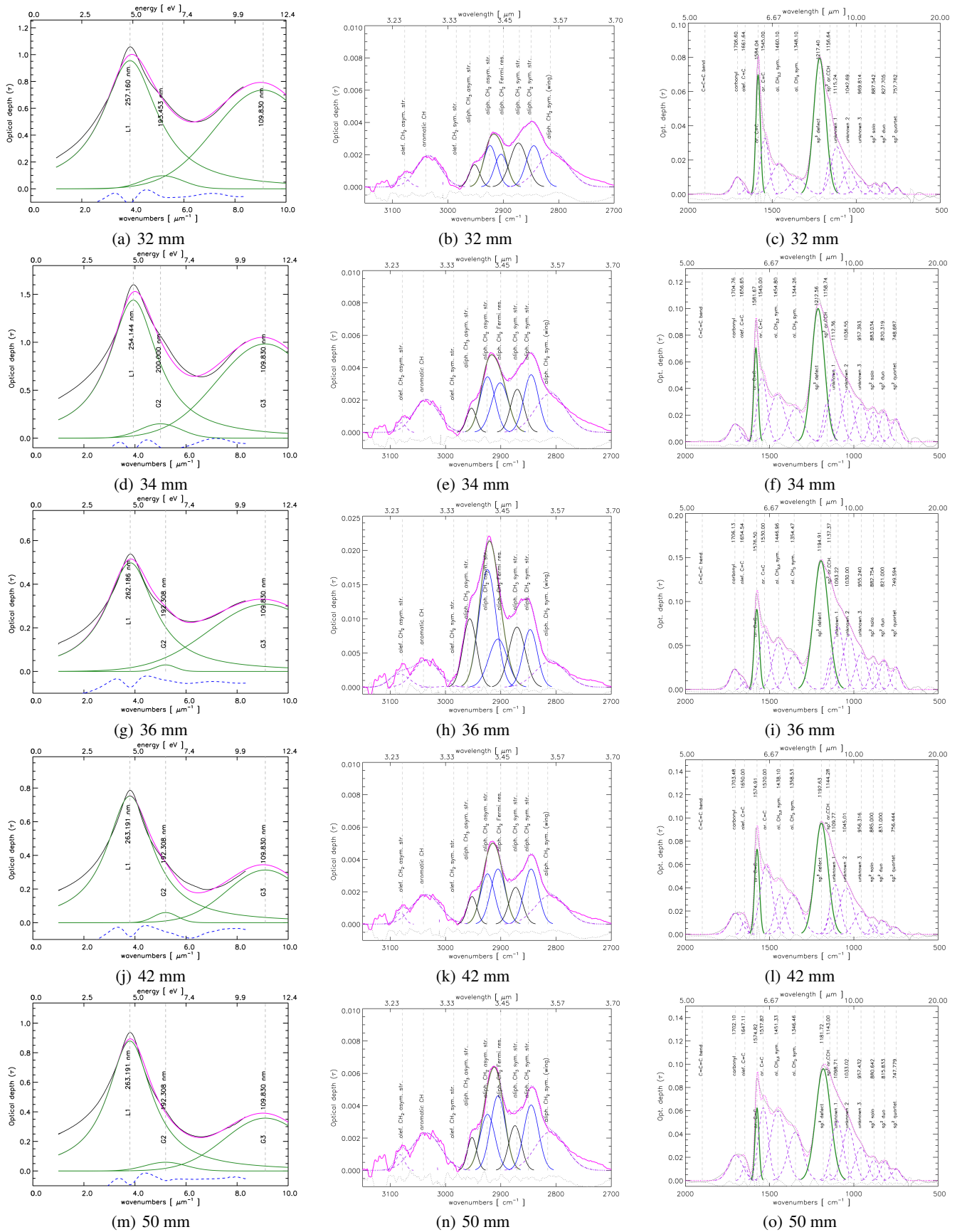


Fig. A.5. UV-VUV and mid-infrared spectral deconvolution for soot samples prepared with a C/O = 1.05 combustion flame, retrieved at a HAB of a) 32 mm to o) 50 mm.

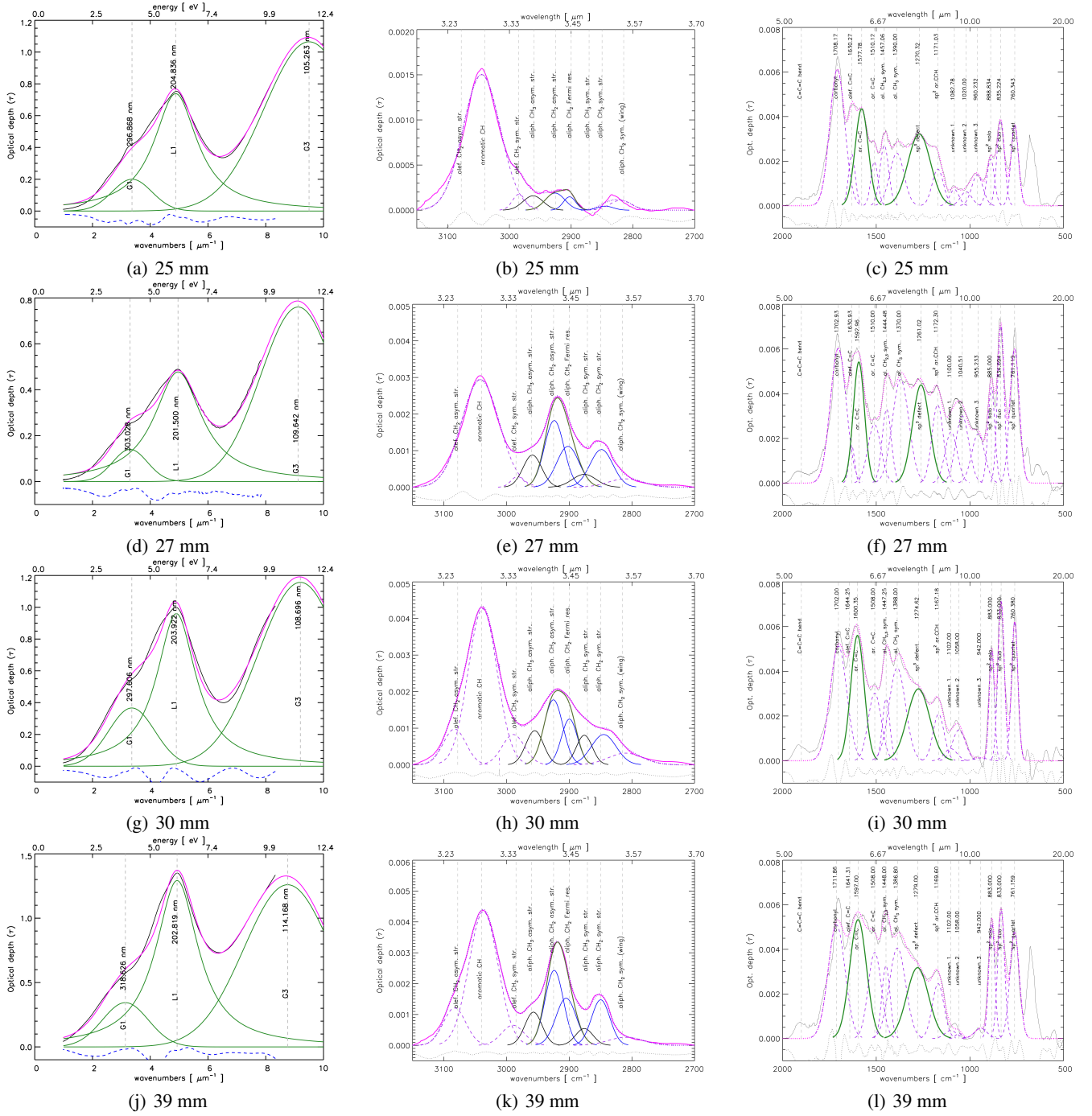


Fig. A.6. UV-VUV and mid-infrared spectral deconvolution for soot samples prepared with a C/O = 1.3 combustion flame, retrieved at a HAB of a) 25 mm to l) 39 mm.

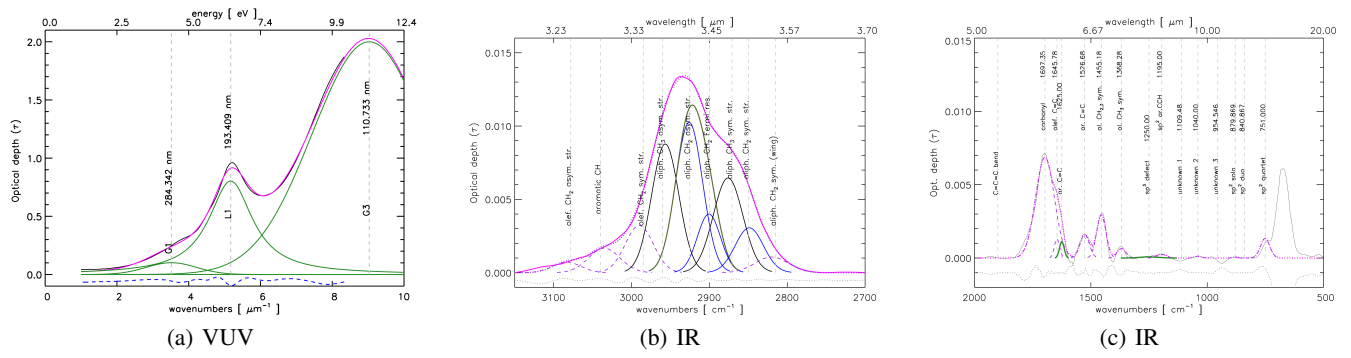


Fig. A.7. UV-VUV and mid-infrared spectral deconvolution for an a-C:H produced with C_7H_8 .

Table A.1. UV and VUV fitting parameters.

Flame C/O	d [mm]	E_g [eV]	$\omega_c(L1)$ [cm ⁻¹]	$\omega_c(G1)$ [cm ⁻¹]	$\omega_c(G2)$ [cm ⁻¹]	$\omega_c(G3)$ [cm ⁻¹]	$\gamma(L1)$ [cm ⁻¹]	$\gamma(G1)$ [cm ⁻¹]	$\gamma(G2)$ [cm ⁻¹]	$\gamma(G3)$ [cm ⁻¹]	$A(L1)$ [cm ⁻¹]	$A(G1)$ [cm ⁻¹]	$A(G2)$ [cm ⁻¹]	$A(G3)$ [cm ⁻¹]
0.82	20	0.82 ± 0.20	40123	NA	56733	91010	25118	NA	9000	43128	24725	NA	1720	50729
	21	0.95 ± 0.21	39800	NA	55122	91437	21094	NA	8000	33800	2997	NA	300	7500
	22	0.79 ± 0.25	40041	NA	52988	90470	22136	NA	12277	32212	6064	NA	638	8196
	30	0.67 ± 0.30	39248	NA	55319	94136	21711	NA	10000	40317	14365	NA	1110	18455
1	18	1.44 ± 0.05	48393	34498	NA	95960	17438	16385	NA	36810	12782	2228	NA	50100
	21	1.11 ± 0.08	45862	30939	NA	97122	18124	13301	NA	47060	16240	1502	NA	39132
	25	0.84 ± 0.15	42653	NA	53663	91242	25958	NA	15091	39000	11118	NA	964	10260
	30	0.72 ± 0.21	39971	NA	51638	95063	24267	NA	12162	35765	12721	NA	993	17451
	39	0.61 ± 0.25	40180	NA	48915	90115	26774	NA	9000	34401	6580	NA	200	8000
1.05	14	1.46 ± 0.02	48509	34165	NA	91815	15879	11705	NA	38870	17806	2000	NA	35240
	18	1.14 ± 0.05	46389	30894	NA	91843	18823	14482	NA	36160	19742	2187	NA	34210
	22	0.67 ± 0.16	41467	NA	52544	92021	25916	NA	10000	45067	18428	NA	799	17124
1.05	22	0.74 ± 0.09	42273	NA	53000	91050	26827	19165	19000	42000	NA	2999	48470	
	24	0.65 ± 0.16	40800	NA	53000	91763	26000	NA	13202	19000	31000	NA	2012	37420
	26	0.26 ± 0.10	39900	NA	53000	91179	26500	NA	8000	19000	50000	NA	1610	45893
	28	0.39 ± 0.08	39583	NA	53000	91050	27000	NA	7254	17436	46147	NA	759	38380
	30	0.20 ± 0.05	39426	NA	53000	91111	24000	NA	8000	19000	50000	NA	1544	52000
	32	0.30 ± 0.11	38886	NA	51692	91050	26000	NA	10000	19000	39000	NA	1367	35000
	34	0.01 ± 0.01	39348	NA	50000	91050	23000	NA	8000	19000	52000	NA	2000	45000
	36	0.18 ± 0.05	38141	NA	52000	91050	23000	NA	7000	22000	18000	NA	500	51700
	42	0.19 ± 0.05	37995	NA	52000	91050	22000	NA	14011	22000	26000	NA	1749	51700
	50	0.05 ± 0.02	37995	NA	50000	91050	21000	NA	7999	21000	29000	NA	600	49350
1.3	25	1.60 ± 0.01	48873	34283	NA	96059	16626	11415	NA	44000	18452	2000	NA	42270
	27	1.51 ± 0.01	49628	33000	NA	91206	20000	9900	NA	35360	14997	1200	NA	36150
	30	1.41 ± 0.01	48700	32769	NA	93890	17578	12228	NA	44000	26098	3000	NA	44990
	39	1.00 ± 0.03	49305	31021	NA	96073	19126	13066	NA	49370	37108	3000	NA	69270
a-C:H	NA	1.95 ± 0.10	52219	38583	NA	91175	11731	16464	NA	39323	12527	966	NA	26057

Table A.2. Infrared fitting parameters.

Flame C/O	Distance [mm]	Band positions		Integrated vibrational modes								
		$\omega_c(\text{C}=\text{C})$ [cm ⁻¹]	$\omega_c\text{sp}^3$ [cm ⁻¹]	$\Delta(\text{CC}-\text{sp}^3)$ [cm ⁻¹]	$\text{CH}_{\text{arom.}}$ [cm ⁻¹]	$\text{CH}_3_{\text{aliph.}}$ [cm ⁻¹]	$\text{CH}_2_{\text{aliph.}}$ [cm ⁻¹]	$\text{CH}_3_{\text{Fermi.}}$ [cm ⁻¹]	C=O [cm ⁻¹]	C=C [cm ⁻¹]	sp ³ [cm ⁻¹]	sp ² solo [cm ⁻¹]
0.82	20	1595	1247	348	0.05	0.10	0.10	0.10	0.63	0.41	1.97	0.40
	21	1593	1243	350	0.03	0.06	0.07	0.08	0.48	0.33	0.98	0.13
	22	1593	1242	351	0.04	0.09	0.07	0.06	0.53	0.39	1.86	0.34
	30	1590	1232	358	0.02	0.04	0.06	0.08	0.41	0.38	1.67	0.26
1	18	1599	1280	319	0.03	0.06	0.04	0.02	0.18	0.08	0.15	NA
	21	1601	1280	321	0.03	0.06	0.05	0.04	0.18	0.14	0.18	NA
	25	1592	1262	330	0.37	0.28	0.35	0.24	2.48	2.86	4.58	0.56
	30	1587	1242	346	0.25	0.15	0.16	0.19	1.44	2.33	4.70	0.23
1.05	39	1582	1219	363	0.32	0.01	0.06	0.06	0.72	2.49	6.79	0.50
	14	1595	1274	321	0.07	0.04	0.04	0.02	0.33	0.20	0.27	NA
	18	1602	1275	327	0.03	0.03	0.05	0.02	0.14	0.16	0.45	NA
	22	1593	1269	324	0.01	0.02	0.03	0.02	0.07	0.14	0.25	NA
1.05	22	1596	1262	334	0.23	0.19	0.25	0.16	2.03	2.02	4.27	0.09
	24	1592	1251	341	0.23	0.20	0.21	0.26	2.70	3.20	6.80	0.11
	26	1588	1240	349	0.15	0.08	0.13	0.19	2.60	3.53	8.86	0.23
	28	1587	1230	357	0.11	0.04	0.07	0.10	1.42	2.69	6.72	0.16
	30	1586	1225	360	0.11	0.03	0.06	0.08	1.03	2.47	6.55	0.17
	32	1582	1217	365	0.11	0.03	0.08	0.06	0.73	2.30	7.06	0.32
	34	1579	1213	366	0.12	0.03	0.12	0.12	1.07	2.19	9.51	1.39
	36	1579	1195	384	0.22	0.31	0.67	0.30	1.71	3.42	16.14	1.80
	42	1579	1191	389	0.12	0.04	0.10	0.11	1.72	2.58	9.08	0.42
	50	1575	1179	396	0.15	0.04	0.12	0.16	2.25	2.47	9.55	0.47
1.3	25	1590	1267	323	0.11	0.01	0.01	0.00	0.73	0.45	0.49	0.20
	27	1593	1261	332	0.22	0.03	0.07	0.05	0.64	0.40	0.50	0.19
	30	1600	1275	326	0.25	0.03	0.07	0.04	0.57	0.50	0.48	0.17
	33	1598	1273	325	0.57	0.06	0.12	0.10	0.69	0.83	0.61	0.48
	36	1601	1276	325	0.59	0.05	0.21	0.08	0.83	1.22	0.90	0.61
a-C:H	39	1597	1279	318	0.25	0.04	0.09	0.06	0.57	0.60	0.46	0.21
	NA	1625	NA	NA	0.08	0.38	0.44	0.15	0.69	0.03	NA	NA