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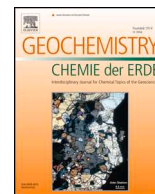
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## Invited Review

## Mineralogy and geochemistry of sedimentary rocks and eolian sediments in Gale crater, Mars: A review after six Earth years of exploration with *Curiosity*

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## ABSTRACT

The Mars Science Laboratory *Curiosity* rover arrived at Mars in August 2012 with a primary goal of characterizing the habitability of ancient and modern environments. *Curiosity* was sent to Gale crater to study a sequence of ~3.5 Ga old sedimentary rocks that, based on orbital visible and near- to short-wave infrared reflectance spectra, contain secondary minerals that suggest deposition and/or alteration in liquid water. The sedimentary sequence in the lower slopes of Mount Sharp in Gale crater preserves a dramatic shift on early Mars from a relatively warm and wet climate to a cold and dry climate, based on a transition from smectite-bearing strata to sulfate-bearing strata. The rover is equipped with instruments to examine the sedimentology and identify compositional changes in the stratigraphy. The Chemistry and Mineralogy (CheMin) instrument is one of two internal laboratories on *Curiosity* and includes a transmission X-ray diffractometer (XRD) and X-ray fluorescence (XRF) spectrometer. CheMin measures loose sediment samples scooped from the surface and drilled rock powders, and the XRD provides quantitative mineralogy to a detection limit of ~1 wt.% for crystalline phases.

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*Curiosity* has traversed > 20 km since landing and has primarily been exploring an ancient lake environment fed by streams and groundwater. Of the 19 drilled rock samples analyzed by CheMin as of sol 2300 (January 2019), 15 are from fluvio-lacustrine deposits that comprise the Bradbury and Murray formations. Most of these samples were drilled from units that did not have a clear mineralogical signature from orbit. Results from CheMin demonstrate an astounding diversity in the mineralogy of these rocks that signifies geochemical variations in source rocks, transportation mechanisms, and depositional and diagenetic fluids. Most detrital igneous minerals are basaltic, but the discovery in a few samples of abundant silicate minerals that usually crystallize from evolved magmas on Earth remains enigmatic. Trioctahedral smectite and magnetite at the base of the section may have formed from low-salinity pore waters with a circumneutral pH in lake sediments. A transition to dioctahedral smectite, hematite, and Ca-sulfate going up section suggests a change to more saline and oxidative aqueous conditions in the lake waters themselves and/or in diagenetic fluids. Perhaps one of the biggest mysteries revealed by CheMin is the high abundance of X-ray amorphous materials (15–73 wt.%) in all samples drilled or scooped to date. CheMin has analyzed three modern eolian sands, which have helped constrain sediment transport and mineral segregation across the active Bagnold Dune Field. Ancient eolian sandstones drilled from the Stimson formation differ from modern eolian sands in that they contain abundant magnetite but no olivine, suggesting that diagenetic processes led to the alteration of olivine to release Fe(II) and precipitate magnetite. Fracture-associated halos in the Stimson and the Murray formations are evidence for complex aqueous processes long after the streams and lakes vanished from Gale crater. The sedimentology and composition of the rocks analyzed by *Curiosity* demonstrate that habitable environments persisted intermittently on the surface or in the subsurface of Gale crater for perhaps more than a billion years.

## 1. Gale crater and the mission goals of the Mars Science Laboratory

The Mars Science Laboratory (MSL) rover *Curiosity* landed in Aeolis Palus on the northwestern plains of Gale crater, Mars in August 2012 and has been actively exploring its surface for more than six Earth years. Gale crater is an impact basin ~155 km in diameter that formed ~3.8 Ga before present during the late Noachian epoch (Thomson et al., 2011) near the Mars equator, on the dichotomy boundary between the southern cratered highlands and northern lowlands (Fig. 1A). The floor of Gale crater is 4.5 km below the Mars datum (Mars average elevation), and Gale existed as a deep depression in the Mars crust (Fig. 1B) at a time when planet-wide morphological evidence suggests that water flowed on its surface (e.g., Craddock and Maxwell, 1993; Howard et al., 2005; Irwin et al., 2005). Therefore, prior to landing it was recognized that there was a high likelihood that water was present in the crater during an early “wet and warm” (at least relatively speaking) period of Mars’ geologic history, at a time when evidence of the first oceans and signs of early life are found on Earth.

The science goals of the Mars Science Laboratory mission are to: 1) assess Mars’ biological potential, 2) characterize the geology of the landing region, 3) explore the past habitability of Mars, and 4) characterize the broad spectrum of surface radiation (Grotzinger et al., 2012). Since landing over six Earth years ago, the MSL Science Team has addressed each of the above goals while investigating processes that were active over much of Mars’ history, from the formation of the crust to modern eolian and atmospheric processes.

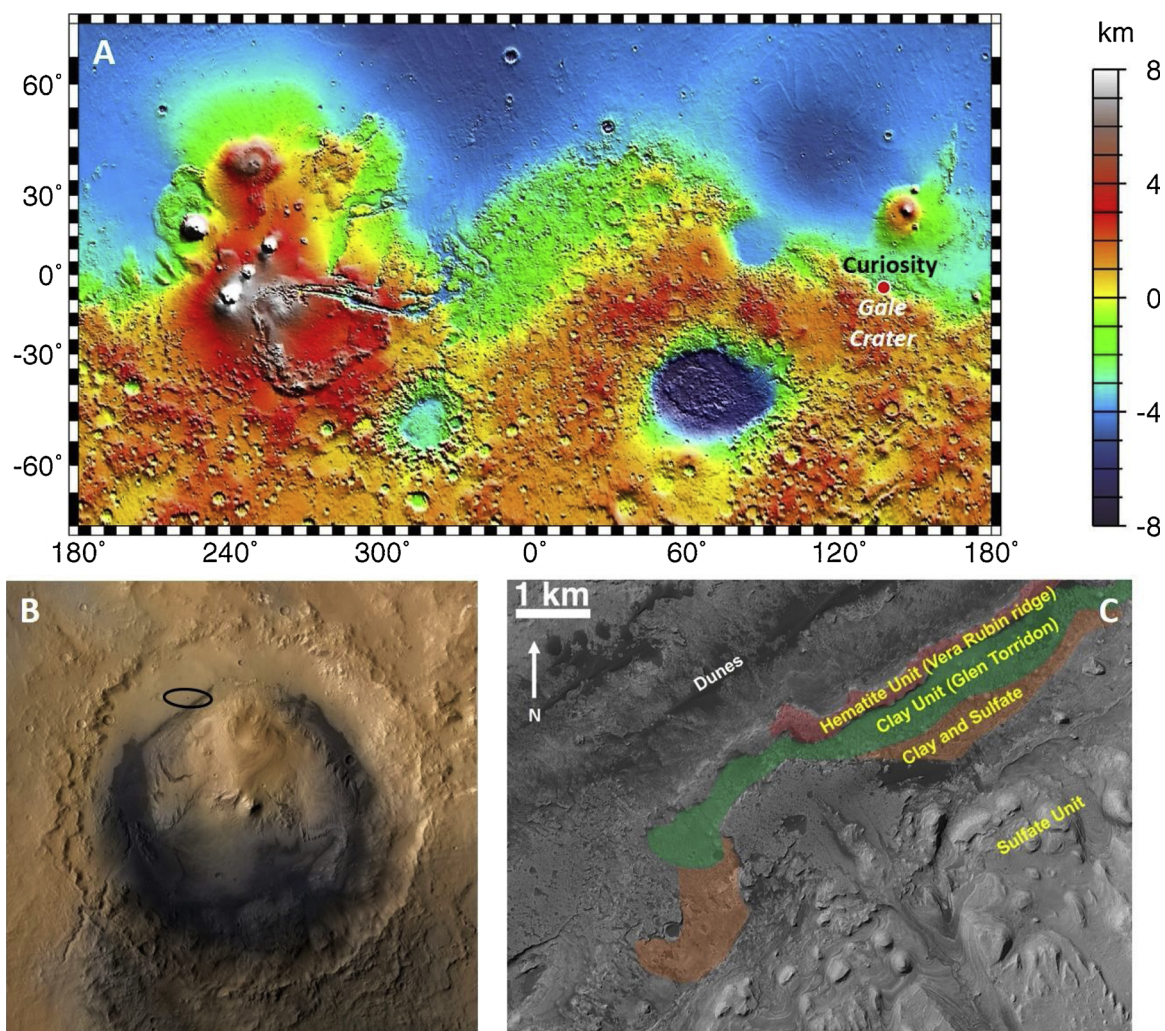
The principal goal of the MSL mission is to find evidence of “habitability” on Mars. In understanding terrestrial habitability and how life originated and radiated on the early Earth, Mars serves as a valuable surrogate; nearly all of Earth’s rocks of this age have been subducted and destroyed as a result of plate tectonics and the remaining rocks from early Earth have been heated to metamorphic temperatures such that the original compositions have been altered (e.g., Fedo, 2000; Mloszewski et al., 2012). Mars never experienced extensive plate tectonics (e.g., Breuer and Spohn, 2003), and there is little evidence that its surface sediments have ever been deeply buried or extensively heated. The opportunity exists on Mars, and especially in Gale crater, therefore, to characterize these ancient sediments that are in much the same condition today as they were when deposited, elucidating the habitability potential of their depositional environments and perhaps even identifying evidence of early life.

Gale is among a class of craters that, after they were formed, were partially or fully filled with sediment, and over eons of time eroded,

exposing original features and any remaining sediments. Recent gravimetric analyses using *Curiosity*’s accelerometers, however, suggest the crater was never completely filled with sediment (Lewis et al., 2019). A key distinguishing feature of Gale is the presence of a ~5 km-high central mound of layered sedimentary rock named Aeolis Mons (informally known as Mount Sharp, named for the American geomorphologist Dr. Robert Sharp; Fig. 1B). Visible and near- to short-wave infrared reflectance spectra collected from orbit reveal that the lowermost slopes of Mount Sharp contain a variety of minerals that are indicative of water-rock interactions, and that these mineral assemblages change as a function of stratigraphic position in a succession of flat-lying, laterally extensive units (Fig. 1C; Milliken et al., 2010; Fraeman et al., 2013, 2016). Phyllosilicate (i.e., layered silicate minerals composed of stacked sheets of silica tetrahedra and Mg-, Fe(II,III)-, and Al-O (OH) octahedra, including the groups smectite, vermiculite, illite, kaolinite, serpentine, micas, and chlorite, commonly called clay minerals) spectral signatures are observed in some stratigraphic units near the base of Mount Sharp, and sulfate-bearing minerals become more prevalent in younger, stratigraphically higher sedimentary units (Milliken et al., 2010). This mineralogical transition suggests that the conditions under which the sediments were deposited changed through time. The broad mineral stratigraphy with sulfate-bearing units overlying phyllosilicate-bearing units has been recognized in similarly aged deposits planet-wide (e.g., Wiseman et al., 2010; Grotzinger and Milliken, 2012), and this mineralogical succession may mark the beginning of the transition from a relatively wet and warm early Mars to a very dry and cold modern Mars (e.g., Bibring et al., 2006; Milliken et al., 2010). The lowest visible units of Mount Sharp are ~3.6–3.8 Ga old, based on crater counting ages and superposition relationships (Thomson et al., 2011). The presence of minerals in these strata that either form in water or as a consequence of aqueous alteration suggests that they preserve evidence of ancient aqueous environments. In addition to phyllosilicate and sulfate mineral spectral signatures, the lower slope of Mount Sharp displays a prominent, local ridge (informally known as Vera Rubin ridge, named for the American astronomer) characterized by strong orbital spectral signatures of the Fe(III) oxide mineral hematite (Fraeman et al., 2013, 2016). The upper portion of Mount Sharp is comprised of dust-mantled cross-bedded units that unconformably overlie portions of sulfate-bearing materials and are likely eolian deposits (e.g., Anderson and Bell, 2010).

The main goal of this paper is to provide an overview of MSL’s major mineralogical and geochemical discoveries with a focus on the mineralogical results from the Chemistry and Mineralogy (CheMin) instrument. We start by describing the capabilities of *Curiosity*’s science





**Fig. 1.** A) Location of Gale crater on Mars Orbital Laser Altimeter (MOLA) elevation map (image credit: NASA/JPL/GSFC); B) Composite of Gale crater and Mount Sharp showing *Curiosity*'s landing ellipse in the northwest plains in black (image credit: NASA/JPL-Caltech/ESA/DLR/FU Berlin/MSSS); and C) HiRISE mosaic of the lower slopes of Mount Sharp and the secondary mineral units identified from orbital visible and near- to short-wave infrared spectroscopy (image credit: NASA/JPL-Caltech/MSSS/JHUAPL/Brown University).

payload and methods utilized to perform the mineralogical and geochemical investigations. We then summarize the geologic and geomorphologic units and materials encountered by *Curiosity* during its traverse. This introduction sets the stage for a more in-depth description of the mineralogy of samples, organized according to their age and depositional environment. These mineralogic details, in concert with geologic context and geochemical measurements, provide insights into the diversity of igneous materials on Mars, modern eolian processes, and the nature of ancient aqueous conditions in Gale crater, leading to the identification of ancient habitable martian environments. Finally, we discuss MSL's most recent investigation of Vera Rubin ridge and look ahead to what the rover may encounter as it climbs Mount Sharp.

## 2. *Curiosity*'s science payload

*Curiosity* carries a suite of scientific instruments to characterize surface and near-surface rocks and sediments (Fig. 2). The elemental composition of rocks and loose sediments is measured remotely (~1–7 m distance) using the ChemCam Laser Induced Breakdown Spectroscopy instrument (LIBS; Maurice et al., 2012, 2016; Wiens et al., 2012, 2015a, 2015b). After more than two years of nominal operations on Mars, ChemCam lost autofocus functionality for LIBS and Remote Micro-Imager (RMI) operation on sol 801. The instrument team

redesigned the flight software to regain full operational capacity and improve the instrument's ability to focus on infinity, which allowed for long-distance RMI mosaics of the upper parts of Mount Sharp and other distant localities (Le Mouélic et al., 2015; Peret et al., 2016). ChemCam can also passively measure spectra in the visible and near-infrared wavelength range. These spectra are useful in identifying Fe-bearing minerals present in rocks and loose sediments (Johnson et al., 2015, 2016, 2017a) and they complement the multispectral imaging capability of Mastcam (Bell et al., 2017; Wellington et al., 2017a). ChemCam has the ability to select autonomously targets using the Autonomous Exploration for Gathering Increased Science (AEGIS) system (Francis et al., 2017). AEGIS selects geological targets in images from the rover's navigation cameras and measures them with ChemCam without input from ChemCam scientists and engineers on Earth. The system has markedly increased the pace of data collection with ChemCam since its implementation in May 2016, in particular improving the statistics on baseline bedrock compositions along the rover traverse.

The Alpha Particle X-ray Spectrometer (APXS), located on the arm of the rover, provides quantitative elemental data from ~2 cm diameter areas of target rocks and soils (i.e., loose, unconsolidated surface materials) through the detection of characteristic X-rays emitted by incident alpha particles and X-rays from  $^{244}\text{Cm}$  sources (i.e., using



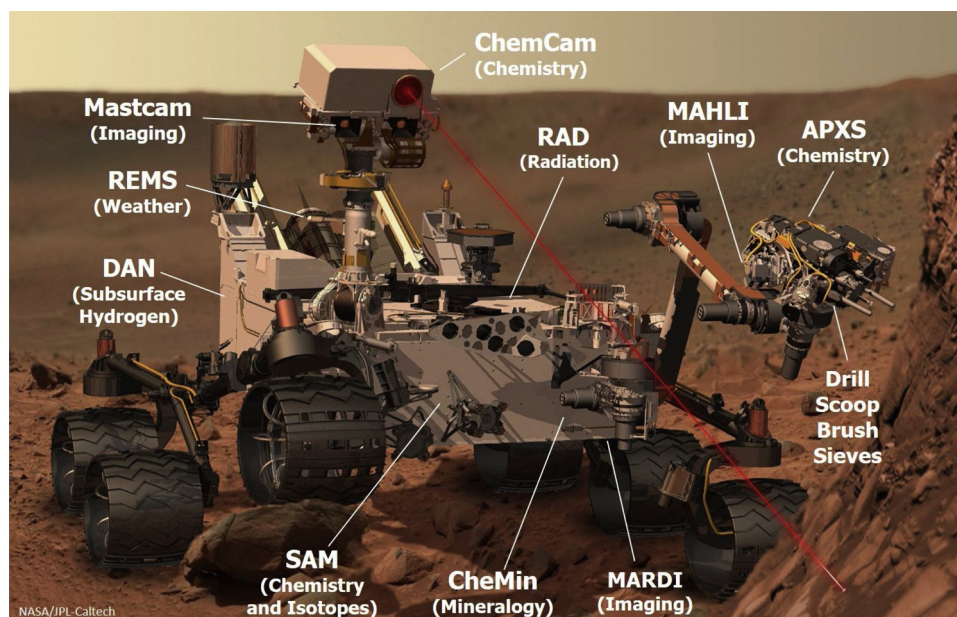


Fig. 2. Schematic of the *Curiosity* rover showing the location of the different science instruments. Image credit: NASA/JPL-Caltech.

particle induced X-ray emission spectroscopy and X-ray fluorescence spectroscopy; Gellert et al., 2006; Campbell et al., 2012). Minerals and X-ray amorphous materials in scooped soil or drilled rock samples are identified and quantified by the CheMin instrument using X-ray diffraction (Blake et al., 2012). Evolved gas analyses from the Sample Analysis at Mars (SAM) instrument of these same samples can inform mineralogy as well as detect and identify organic molecules (Mahaffy et al., 2012). Geomorphological, structural, and sedimentological features are assessed with images of landscapes, outcrops, and rock and sediment textures obtained by the science cameras on the rover: Mastcam (Malin et al., 2017), the Mars Hand Lens Imager (MAHLI, Edgett et al., 2012), the Mars Descent Imager (MARDI, Malin et al., 2017), and the RMI on ChemCam (Maurice et al., 2012). Instruments that interrogate the modern environment include the Dynamic Albedo of Neutrons (DAN) instrument, which uses neutron spectrometry to quantify H abundances in the upper  $\sim 1$  m of the surface as a proxy for water or hydrated minerals (Litvak et al., 2008); the Rover Environmental Monitoring Station (REMS), which measures wind speed/direction, pressure, relative humidity, air and ground temperatures, and ultraviolet radiation (Gómez-Elvira et al., 2012); and the background solar and cosmic radiation detector (RAD; Hassler et al., 2012), which measures ambient particle and radiation fluxes.

Samples provided to the two laboratory instruments inside the rover (CheMin and SAM) are collected, processed and delivered by the Sample Acquisition, Sample Processing, and Handling (SA/SpaH) and Collection and Handling for In-Situ Martian Rock Analysis (CHIMRA) systems (Anderson et al., 2012). Samples can be obtained in two ways: 1) loose soil samples are scooped from the upper few cm of the surface and sieved to  $< 1$  mm and  $< 150$   $\mu\text{m}$  grain sizes. Aliquots of the  $< 150$   $\mu\text{m}$  material are delivered to CheMin and SAM for analysis; SAM can also process the  $< 1$  mm sieved fraction. 2) Powdered rock samples are obtained with a percussion drill. The drill bit is  $\sim 1$  cm in diameter and can penetrate to a depth of  $\sim 6$  cm. For most of the mission to date, the powdered material that reached SA/SpaH and CHIMRA traveled up the drill stem and was sourced from the lower  $\sim 4$  cm of drilled material. The samples were sieved, and aliquots were portioned and then delivered to the two laboratory instruments inside the rover.

An anomaly in the drill feed mechanism occurred on sol 1536 that precluded its use. As a result, beginning with the sample “Duluth,” rocks were drilled by a new method called “feed-extended drilling,” in which the drill remains extended to its full length and is pressed into

the surface by *Curiosity*’s arm (essentially used like a hand drill). With feed-extended drilling, the sample can no longer pass through CHIMRA to be sieved and portioned. Drilled rock samples are now delivered to CheMin and SAM by hovering the drill bit over the instruments’ inlets and reversing the drill rotation to allow sample material to exit the drill bit and drop into the inlets. Engineers at the Jet Propulsion Laboratory (JPL) used the testbed arm and drill to constrain drill and delivery parameters so that CheMin and SAM receive sample volumes similar to those delivered during prior drill operations. The grain size of the powder produced by the drill is typically  $< 150$   $\mu\text{m}$  (Anderson et al., 2012), and any large fragments that could potentially clog the CheMin inlet are kept from entering the instrument by a 1 mm screen in the funnel. Whether scooped or drilled, once the analyses are completed, the remaining material in the drill stem or CHIMRA is dumped onto the martian surface and the APXS instrument is used to obtain a quantitative elemental analysis of the bulk sample and any sieved size fractions. Dump piles are typically  $> 2$  cm in diameter and the alignment of the centers of MAHLI images with the APXS field of view is precise within a few mm such that the dumped material fills the APXS field of view and the underlying surface does not contribute significantly to the APXS-measured chemistry (e.g., VanBommel et al., 2016).

Data collected by the instruments aboard *Curiosity* can be used to characterize the elemental, molecular, isotopic, and mineralogical compositions of samples collected from the upper  $\sim 2$ – $6$  cm of the surface, and the abundance of H in the upper meter of the surface. With this information, the MSL Science Team can infer ancient martian habitability, albeit based on an Earth-centric understanding of the requirements of microbial life (e.g., Hoehler, 2007).

### 3. The CheMin instrument and methods for quantifying mineralogy

Here, we provide a brief overview of the CheMin instrument and the methods used for calibrating the data, identifying mineralogy, quantifying mineral and X-ray amorphous abundances, and calculating crystal chemistry from refined unit-cell parameters. Blake et al. (2012) describe the instrument, its data products, and calibration, and Morrison et al. (2018a, b) discuss the methods for calculating crystal chemistry from unit-cell parameters.

The CheMin X-ray Diffraction (XRD) / X-ray Fluorescence (XRF) instrument produces XRD patterns of scooped soil or drilled rock

powder. Samples have been either sieved to  $< 150 \mu\text{m}$  or a comparable powder has been delivered directly; a  $\sim 75 \text{ mm}^3$  aliquot of material (this volume was more tightly constrained when the drill and CHIMRA operated nominally) is delivered to the CheMin funnel, accessed through a port on the rover deck. Sieved soils and drill powders are delivered through CheMin's inlet funnel to one of the instrument's 27 reusable sample cells. Sample cells are arrayed in pairs on the perimeter of a sample wheel (Fig. 3). The cells are designed to be used more than once. Individual cells are filled and analyzed at the top of the wheel, then rotated  $180^\circ$  and emptied into a sump at the bottom of the instrument. Each dual-cell sample holder is shaped like a tuning fork, with a sample cell attached to each arm. The cells hold the sample in a 7 mm diameter,  $170 \mu\text{m}$  thick volume between two  $\sim 7 \mu\text{m}$ -thick polymer (Mylar or Kapton) windows. A piezoelectric actuator drives the sample holder at its natural resonance and the resulting vibration causes a convective flow of sample material, randomizing grain orientations and minimizing preferred orientation effects. The instrument utilizes a transmission geometry with a microfocus Co X-ray source collimated to a 70  $\mu\text{m}$  diameter X-ray beam that passes through the center of the sample cell (Fig. 4). An X-ray energy-sensitive charge-coupled device (CCD) collects two-dimensional (2D) XRD images over 3–38 hours of analysis. The CCD has a  $600 \times 582$  pixel data collection area, where pixels are  $40 \times 40 \mu\text{m}$ , creating an angular resolution of  $\sim 0.35^\circ/2\theta$ . This angular resolution allows the refinement of plagioclase and olivine unit-cell parameters such that An and Fo values are constrained to a  $1\sigma$  error of  $\sim 2$ -10 (Morrison et al., 2018b). On average, samples are measured for a total of 22.5 h over three separate nights. The CCD detector is operated in single-photon counting mode and can be used to measure the amount of charge generated by each photon (and hence its energy). Diffracted  $\text{CoK}\alpha$  X-ray photons (i.e.,  $\lambda = 1.79027 \text{ \AA}$ ) are identified by their energy and are summed circumferentially to yield a 2D energy-discriminated  $\text{CoK}\alpha$  diffraction pattern (Fig. 4). All detected photons are summed into a histogram that represents an XRF spectrum of the sample. Parts of the XRF system were descope prior to launch such that the XRF data are qualitative, so quantitative compositional data from APXS and/or ChemCam are used to complement mineralogical data from CheMin. The 2D XRD images are converted to 1D diffraction patterns (Fig. 5) for analysis using a modification of the GSE/ADA software (Dera et al., 2013). Initial calibration of the pattern is made with reference to beryl-quartz standards contained in two of five sealed standards on the CheMin sample wheel.

CheMin's mineral detection limit is as low as  $\sim 1 \text{ wt.}\%$ , depending on the crystallinity of the phase and the presence or absence of overlapping peaks (Castle and Treiman, 2019). Quantitative mineral abundances and unit-cell parameters of the major phases (e.g.  $> \sim 5 \text{ wt.}\%$ ) are determined through Rietveld refinement of 1-dimensional XRD patterns. The Rietveld refinement technique (e.g., Rietveld, 1969; Bish and Howard, 1988; Post and Bish, 1989) is used to fit the peak positions, intensities, and breadths by adjusting the scale factors, unit-cell parameters, and full-width half-maxima of peaks generated from crystallographic information files (CIFs). The CheMin team commonly uses the Jade™ software by Materials Data Inc. to perform these refinements. Mathematical relationships between unit-cell parameters and crystal chemistry are used to determine compositions of major phases (feldspar, pyroxene, olivine, spinel, and jarosite) in each sample (Morrison et al., 2018a). Tolerance variations in the machining of the sample cell assemblies cause individual cells to be offset from their ideal diffraction position by  $-25$  to  $-113 \mu\text{m}$  relative to an ideal sample cell-to-CCD distance of 18.5302 mm. To correct for this positioning offset and the resulting shift in 2-theta peak positions, an internal calibration method was developed that is based on the refined cell parameters of plagioclase feldspar, present as a major phase in all but one of the samples measured by CheMin (Morrison et al., 2018b).

Abundances of X-ray amorphous, poorly crystalline, and paracrystalline materials (e.g., volcanic glasses, opaline silica, smectite) are estimated using a modified version of the full pattern fitting program

FULLPAT (Chiperá and Bish, 2002, 2013). FULLPAT uses a least-squares minimization to optimize the fit between measured CheMin patterns and patterns of individual minerals and X-ray amorphous phases measured on CheMin test-bed instruments. X-ray amorphous materials and phyllosilicates in the FULLPAT library include basaltic and rhyolitic glasses, opal-A, allophane, Fe-allophane, ferrihydrite, illite, saponite, montmorillonite, and nontronite.

The composition of X-ray amorphous materials is estimated from mass balance calculations using mineral abundances and the crystal chemistry of major phases calculated from unit-cell parameters derived from Rietveld refinements, the estimated abundance of X-ray amorphous materials from FULLPAT, and the bulk chemical composition measured by APXS of the drill fines that best represent the sample measured by CheMin (e.g., Blake et al., 2013; Morrison et al., 2013; Dehouck et al., 2014; Morrison et al., 2018b). When the drill was operating nominally, APXS data of fines sieved to  $< 150 \mu\text{m}$  were used in the calculations. In the current drill configuration, APXS data of the bulk sample dumped from the drill stem are used. Subtracting the calculated composition of all crystalline phases from the APXS bulk composition yields the composition of the X-ray amorphous component. If the abundance of the amorphous component is underestimated by FULLPAT, its calculated composition may yield negative values for some elemental oxides, which is geochemically impossible. In such cases, the amorphous abundance determined by FULLPAT is increased until the values for all elemental oxides are zero or positive. Note that minor and trace elements in minerals and elements that are present in crystalline phases below CheMin's detection limit ( $\sim 1 \text{ wt.}\%$ ) are incorporated into the calculated amorphous composition, but may not actually reside in that component of the sample. As a result, some elemental oxides, including  $\text{TiO}_2$ ,  $\text{Cr}_2\text{O}_3$ ,  $\text{MnO}$ , and  $\text{P}_2\text{O}_5$ , may be unrealistically concentrated in the amorphous component.

The experimental data records (EDRs, i.e., the raw data) and reduced data records (RDRs, i.e., the processed data) for CheMin diffraction and fluorescence data are available on the Planetary Data System. Data and results for individual samples and sample descriptions

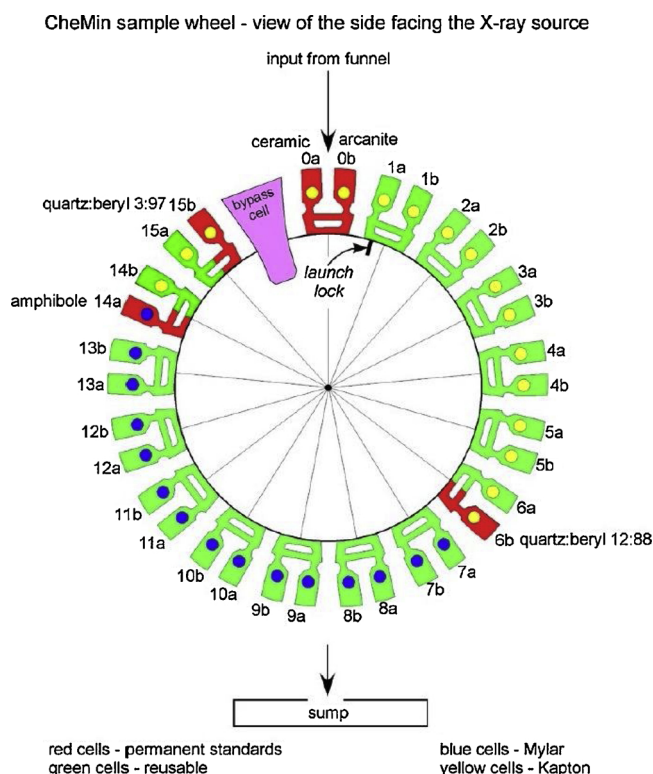


Fig. 3. Configuration of sample cells in CheMin. From Blake et al. (2012).

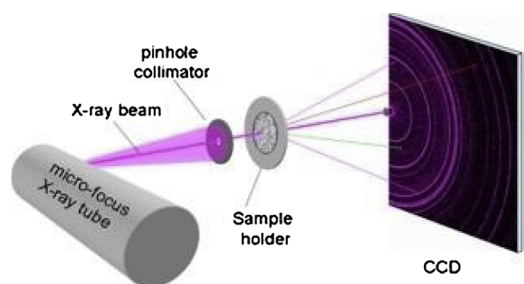


Fig. 4. Illustration of the basic components of the CheMin instrument and the path of the X-rays. From Blake et al. (2012).

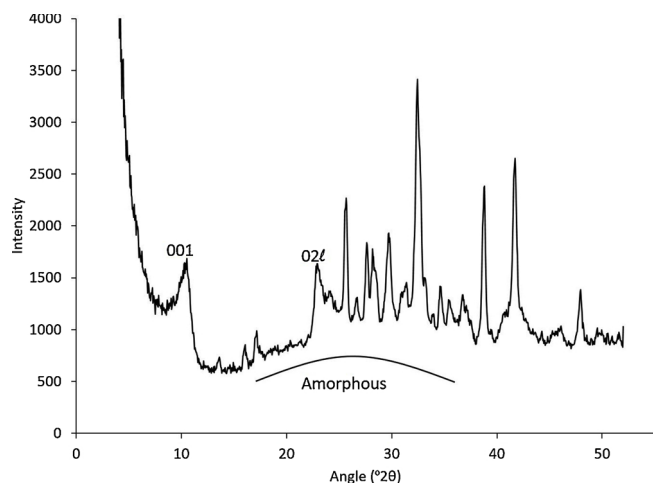


Fig. 5. CheMin 1D XRD pattern of the “Marimba” (MB) sample from the Karasburg member of the Murray formation on Mount Sharp. “Marimba” contains the highest proportion of phyllosilicate of any sample drilled to date (28 wt.%). The (001) and (02l) phyllosilicate peaks are labeled. The X-ray amorphous background is outlined in black.

are also compiled on an Open Data Repository (<https://odr.io/CheMin>), as are companion data from APXS and open-access downloadable pdfs of all publications authored by the CheMin team.

#### 4. An overview of *Curiosity*'s traverse in Gale crater

*Curiosity*'s tenure on Mars is measured in “sols” or martian days (a sol is roughly 39 min longer than an Earth day, and a martian sidereal year is ~668 sols or slightly less than two Earth years in length). As of sol 2300, *Curiosity* has driven over 20 km laterally and gained ~400 m of elevation (Fig. 6). Because the sediments composing Mount Sharp are roughly flat lying, this corresponds to ~400 m of stratigraphic section (Fig. 7). Geologic materials encountered by *Curiosity* during its traverse are almost exclusively sedimentary, and the geochemical compositions of the sedimentary units are largely consistent with low-temperature (e.g., depositional, authigenic, and diagenetic) alteration of rock and mineral fragments derived from a basaltic protolith (McLennan et al., 2014; Siebach et al., 2017; Bedford et al., 2019). However, many examples of float (rocks transported away from their original source region) having primary igneous textures have been identified (Sautter et al., 2014; Cousin et al., 2017a) and their compositions from APXS and ChemCam suggest magmatic diversity in and around Gale crater (Stolper et al., 2013; Sautter et al., 2014, 2015; Cousin et al., 2017a). Some of this float is present as pebbles in the conglomerate facies in the Bradbury group (Williams et al., 2013; Mangold et al., 2016; Fig. 8A shows the Hottah conglomerate where *Curiosity* landed, exposed by the sky crane's landing rockets). The northwest crater rim shows incision from fluvial processes, indicating the nearby crater rim was a likely source of the pebbles (Williams et al., 2013). Some of the igneous float

was probably directly emplaced onto the Gale surface as impact ejecta (Yingst et al., 2013).

Many of the igneous samples are notably feldspar-rich, mainly extrusive trachybasalts/trachyandesites, with some feldspathic cumulates where feldspar phenocrysts (primarily plagioclase) compose up to 80% of the rock. These igneous rocks are thought to have resulted from low pressure fractional crystallization of basaltic melt of the Adirondack Class (from the Mars Exploration Rovers) (Edwards et al., 2017). The similarity of the basaltic end member to much of the Gale sediment bulk compositions suggests that this type of Fe-rich olivine tholeiite is the dominant constituent of the Bradbury group (Edwards et al., 2017; Bedford et al., 2019). However, intriguingly, potassium and sanidine-rich sedimentary units in the Bradbury group (Anderson et al., 2015; Le Deit et al., 2016; Mangold et al., 2016; Treiman et al., 2016) and possible alkaline igneous float (Stolper et al., 2013) show that alkaline igneous rocks were also a significant part of the Gale catchment (Siebach et al., 2017; Bedford et al., 2019) and were present over a larger portion of the southern highlands (Sautter et al., 2015).

The physical sedimentology of the portion of Gale crater traversed by *Curiosity* has been carefully documented by the MSL Sedimentology / Stratigraphy working group, who pored over thousands of images and observations to assemble the stratigraphic column shown in Fig. 7. This figure does not represent a vertical section, but is more akin to a log of the lithologies encountered along the traverse. The mineralogical and geochemical data obtained by *Curiosity* should be viewed and interpreted in the context of this stratigraphic record.

On the plains of Gale crater (i.e., Aeolis Palus), before climbing the slopes of lower Mount Sharp, *Curiosity* investigated primarily fluvial and deltaic deposits (Fig. 8) that compose the Bradbury group (e.g., Grotzinger et al., 2015). The observation of coarse-grained sedimentary deposits, including conglomerate, grain-supported sandstone, and pebbly sandstone with common cross-stratification is consistent with deposition in a fluvial environment (e.g., Williams et al., 2013; Vasavada et al., 2014; Grotzinger et al., 2015; Edgar et al., 2018a). The cobble-sized clasts in the conglomerate identified near the landing site (Fig. 8A) suggest that the ancient rivers that transported these sediments were up to 0.9 m deep with an average velocity of 0.20–0.75 m/s (Williams et al., 2013).

The lower slopes of Mount Sharp, located southwest of the landing site, were the ultimate goal for *Curiosity*. Gale crater is what is termed a “go-to” site in the planetary science community, meaning that the principal science target of the MSL mission was actually outside its 6 × 20 km landing ellipse. However, it was apparent upon landing that a region of high scientific interest where there was a contact between three geomorphological terrain types was nearby. Therefore, from the landing site *Curiosity* drove east to investigate the Yellowknife Bay formation prior to driving southwest toward the lower slopes of Mount Sharp. At Yellowknife Bay, *Curiosity* investigated both sandstone and mudstone facies (e.g., Grotzinger et al., 2014; Edgar et al., 2018a). The Sheepbed mudstone in the Yellowknife Bay formation (Fig. 8B) is interpreted to represent an ancient lake (Grotzinger et al., 2014) and is located at the lowest elevation (and stratigraphic level) investigated by *Curiosity*. If we assume that the strata in Gale crater are flat lying and laterally extensive, the Sheepbed mudstone is the oldest deposit that *Curiosity* investigated. Two samples were drilled from the Sheepbed mudstone in Yellowknife Bay, named “John Klein” and “Cumberland.” These samples were drilled a few meters apart laterally and ~10 cm apart vertically, allowing the mission to investigate local-scale mineralogical and geochemical variations in an area with significant evidence for early-stage diagenesis (Grotzinger et al., 2014; L veill e et al., 2014; Siebach et al., 2014; Stack et al., 2014).

Sandstone beds that dip gently to the south (i.e., toward Mount Sharp) were identified in multiple locations along the 8 km traverse across Aeolis Palus from Yellowknife Bay and are consistent with ancient fluvio-deltaic deposits (Fig. 8C, Grotzinger et al., 2015). Deltas form when rivers flow into standing bodies of water and deposit



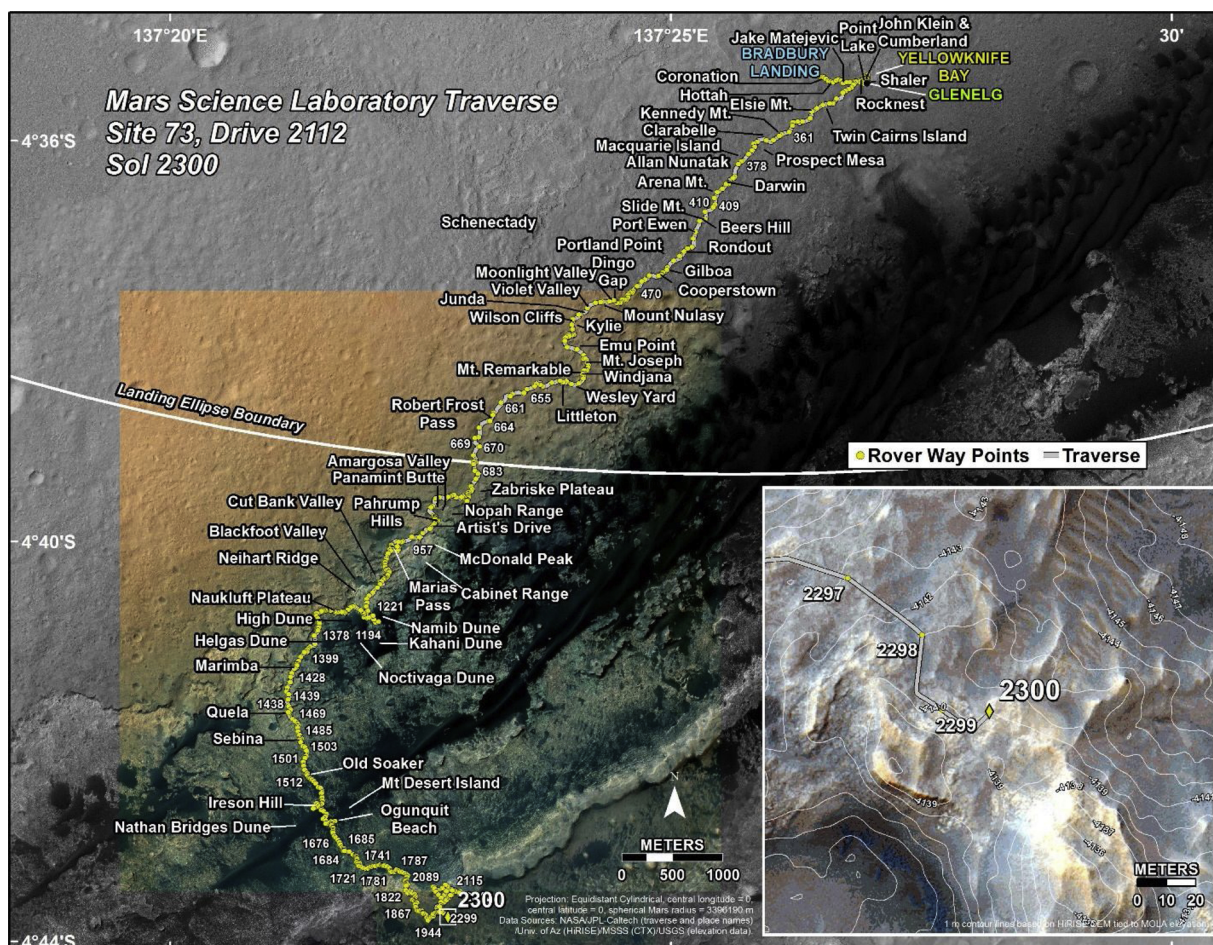


Fig. 6. *Curiosity's* traverse through Sol 2300 with major waypoints listed. Locations and target names along the traverse are unofficial designations and are selected based on themes for each map quadrant. Each quadrant's theme is a location of geologic interest on Earth. For example, there are locations and target names selected from localities in the Northwest Territories in Canada, the Kimberley province in Australia, Namibia, and Scotland. Image credit: NASA/JPL-Caltech/Univ. of AZ/MSSS/USGS.

sediment that was transported by the rivers. The dip of distinct clinoforms (i.e., layered deltaic bodies) indicates that the rivers in the region of *Curiosity's* traverse flowed from north to south, possibly emanating from ice/snow melt in the northern crater rim (Grotzinger et al., 2015). The sample named “Windjana” was drilled from reworked deltaic and eolian sediments in the Kimberley formation (Rice et al., 2017).

*Curiosity* first reached the units that make up lower Mount Sharp in September 2014. This marked the beginning of the MSL Science Team's investigation of the Murray formation, which has continued up to the present time. The Murray formation has been divided into several distinct stratigraphic members based on minor changes in facies associations (Fig. 7). The outcrops that compose the lowermost portion of the Murray formation (the Pahrump Hills member) are dominated by laminated mudstone with mm-to-cm-scale lamination (Fig. 9A,B). Laminated mudstone is common throughout the Murray formation and is consistent with deposition in lakes (e.g., Fedo et al., 2018; Stack et al., 2018), where thick laminations suggest near-shore deposition in shallow water and thin laminations suggest offshore deposition in deep water (Grotzinger et al., 2015). Sedimentary rock in the Pahrump Hills member exhibits exceptional preservation of original depositional fabrics as well as products of early and late diagenesis. Above the Pahrump Hills member, the Murray formation contains a combination of laminated mudstone, meter-scale trough cross-bedding, ripple cross-laminated sandstone (Fig. 9C), and dm-scale cross-stratified sandstone (Fedo et al., 2018). Cross-stratification suggests deposition in higher-energy environments, and the cross-stratified units observed in the Murray

formation are interpreted to have been deposited in both subaqueous (i.e., fluvial) and subaerial (i.e., eolian) environments. Uncommon desiccation cracks (Fig. 9D) suggest drying of sediments and intermittent exposure during lake lowstands (Stein et al., 2018). Isolated observations of NaCl-rich targets (Thomas et al., 2019), rare boron detections (Gasda et al., 2017), and a Mg-sulfate- and Ca-sulfate-rich interval from ChemCam (Rapin et al., 2019) also hint at evaporation. No chloride-rich strata have been found, however, suggesting that desiccation was relatively infrequent when these sediments were deposited.

Twelve samples have been drilled from the Murray formation (Table 1, Fig. 7): “Confidence Hills,” “Mojave2,” “Telegraph Peak,” and “Buckskin” from the Pahrump Hills member, “Oudam” from the Hartmann's Valley member, “Marimba” and “Quela” from the Karasburg member, “Sebina” from the Sutton Island member, “Duluth” from the Blunts Point member, “Stoer” from the Pettegrove Point member, and “Highfield” and “Rock Hall” from the Jura member.

The deposits observed in the Bradbury group and in the Murray formation indicate that Gale crater was the site of a long-lived fluvio-lacustrine environment at ~3.5 Ga (Fig. 10; e.g., Grotzinger et al., 2015; Fedo et al., 2018). Rivers and streams originating from the crater rim flowed into lakes on the crater floor. The paucity of mud cracks suggests lake waters were long-lived. Based on sediment deposition rates on Earth, the 400+ m of primarily lacustrine stratigraphy that *Curiosity* has investigated so far represents a few hundreds of thousands to a few millions of years of time (Grotzinger et al., 2015). After deposition, fluids moved through the sediments both before and after they

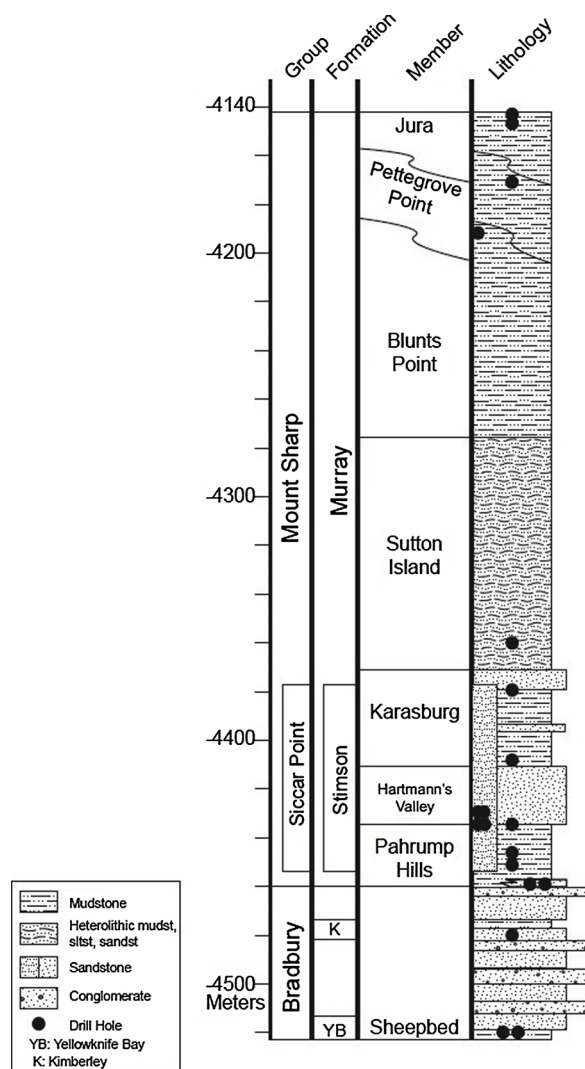


Fig. 7. Stratigraphic column showing the sedimentary units that *Curiosity* has investigated through January 2019. Note that the Vera Rubin ridge topography does not follow elevation. The western leg of *Curiosity*'s traverse up the Vera Rubin ridge showed lower elevations for the Pettegrove Point and Jura members than the eastern leg of *Curiosity*'s traverse down and back up the Ridge. Credit: Sed/strat group of the MSL Science Team.

were lithified (e.g., Martin et al., 2017), suggesting a prolonged history of aqueous diagenesis in surface and subsurface environments at Gale crater, as will be explored in detail later.

In addition to ancient fluvio-lacustrine deposits, *Curiosity* has investigated modern and ancient eolian deposits. *Curiosity* sampled a recently inactive eolian bedform called "Rocknest" (Fig. 11A) close to the landing site (e.g., Blake et al., 2013). *Curiosity* also executed a two-phase sampling campaign in an active eolian basaltic dune field in early 2016 and early 2017 (e.g., Bridges and Ehlmann, 2018; Lapotre and Rampe, 2018). *Curiosity*'s investigation of the Bagnold Dune Field was the first in-situ study of active wind-blown sediments on another planetary surface. Two samples were collected to study sediment transport and sorting across the dune field. The "Gobabeb" sample was collected on sol 1225 from a ripple crest on the Namib Dune, a barchanoidal dune on the northern margin and trailing edge of the dune field (Fig. 11B). The "Ogunquit Beach" sample was collected on sol 1650 from a ripple trough of the Mount Desert Island sand patch near the southern edge of the dune field (Fig. 11C).

The Stimson formation, interpreted as an ancient lithified eolian dune field (Banham et al., 2018), is part of the Mount Sharp and

unconformably overlies the Murray formation (Fig. 12). An unconformity occurs when there is a cessation of sediment deposition, usually with erosion of preexisting sediment prior to continued sediment influx. On Earth, unconformities often occur as a result of tectonic uplift or sea level retreat. Since there is no tectonic uplift in Gale crater and no evidence for lake water breaching the crater rim, this must represent a major change in the climate and depositional regime, including complete drying out of the lake and erosion and removal of a significant portion of the Murray and Bradbury formation sedimentary rocks that had formed in Gale crater. The hiatus in time between the deposition of the Murray and Stimson formations has not been quantified, but it was likely significant because the stratigraphic boundary is abrupt and transitional deposits in relatively wet to relatively dry conditions are not preserved (Banham et al., 2018). Light-toned fracture-associated halos in the Stimson and nearby Murray suggest leaching by late-stage fluids after burial (Frydenvang et al., 2017; Yen et al., 2017; Hausrath et al., 2018). Two of these fracture-associated halos were studied in detail. Samples were drilled from both in the light-toned fracture-associated halos ("Greenhorn" and "Lubango") and in the nearby parent sandstone ("Big Sky" and "Okoruso") to constrain the geochemical processes that formed the halos. Fig. 13 shows an example of a fracture-associated halo observed prior to sampling at a location called Teakettle Junction.

## 5. Mineralogy of samples measured by CheMin

The characterization of the mineralogy and geochemistry of samples from a variety of different depositional environments and time periods provides insight into past and present geologic processes on Mars. All samples measured to date contain igneous minerals, secondary phases, and X-ray amorphous materials. Here we summarize the abundances of minerals and amorphous components, crystal chemistry of major phases from refined unit-cell parameters, and calculated composition of amorphous components from all samples analyzed through sol 2300, the limit of reportable analyses at the time of this writing. Results are grouped according to their depositional environment: Fluvio-lacustrine, ancient eolian, and modern eolian.

### 5.1. Fluvio-lacustrine samples

Most of the rock samples drilled to date were obtained from ancient fluvio-lacustrine deposits (Table 1). Of the 19 rock samples analyzed by CheMin, 13 are mudstones interpreted to be lacustrine deposits and two are fine-grained sandstones interpreted to be reworked fluvial/eolian deposits. The mineralogy of fluvio-lacustrine rocks can tell us about the composition of rocks in the source regions of the sediments, the weathering history of the sediments, and the fluid chemistry of syn- and post-depositional aqueous environments. CheMin has recorded enormous mineralogical variability between fluvio-lacustrine samples that suggests differences in igneous sources, paleoclimatic conditions, and geochemistry of aqueous environments. The mineral abundances of fluvio-lacustrine samples are presented in Table 2, the crystal chemistry of the major minerals in Tables 3 and 4, the abundances of phyllosilicates and X-ray amorphous materials in Table 5, and the calculated composition of the X-ray amorphous materials in Table 6. These results have been reported previously by Vaniman et al. (2014); Treiman et al. (2016); Rampe et al. (2017); Yen et al. (2017); Bristow et al. (2018); Morrison et al. (2018b), Achilles et al. (submitted), and Rampe et al. (2020). Here, we provide an overview of the minerals identified in fluvio-lacustrine samples and summarize the mineralogy of each unit moving up section.

The mineralogy of fluvio-lacustrine samples is variable within the stratigraphic section. Samples contain varying abundances of plagioclase, felsic igneous minerals, mafic igneous minerals, Fe-oxide minerals, phyllosilicates, sulfate minerals, the phosphate fluorapatite, halite, and X-ray amorphous materials (Fig. 14). Felsic minerals include



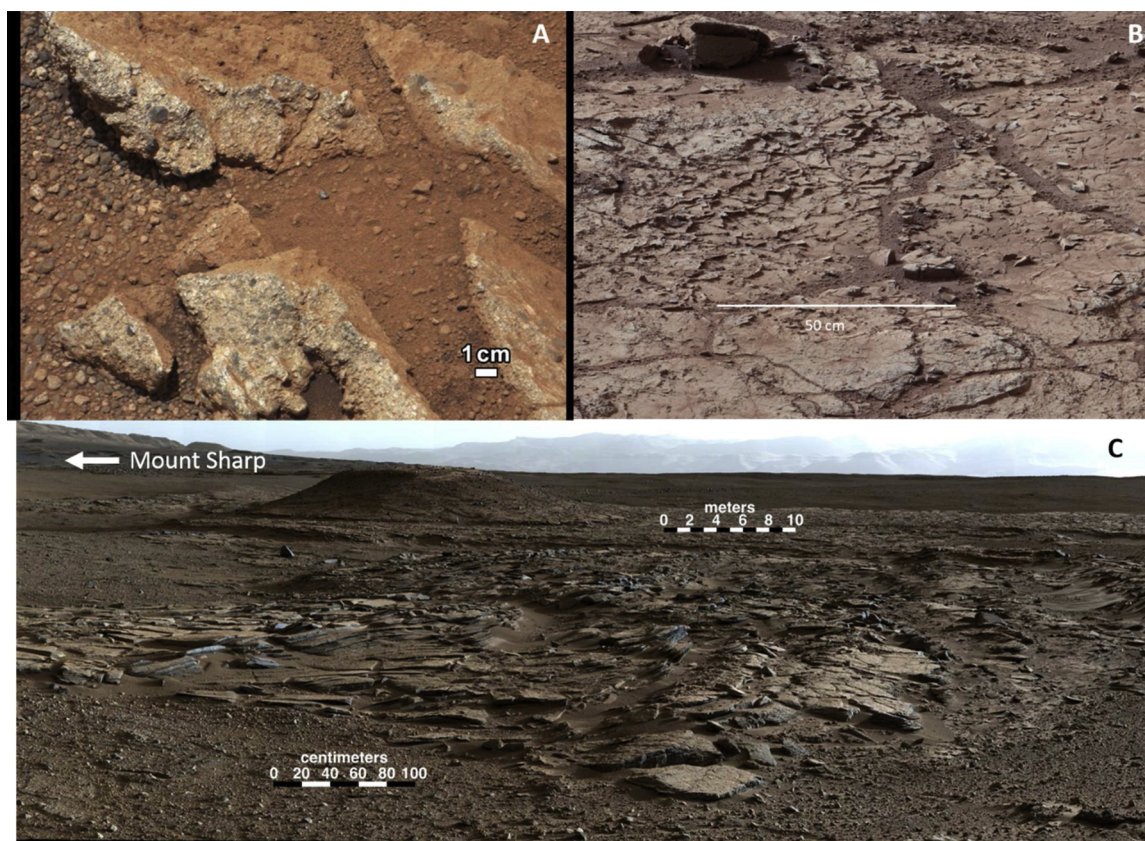


Fig. 8. Images of outcrops on the plains of Gale crater. A) “Link” conglomerate near *Curiosity*’s landing site (sol 27, sequence mcam00129). B) Sheepbed mudstone at Yellowknife Bay (sol 153, mcam00848). C) Dipping sandstone beds at the Kimberley formation (sol 590, mcam02484). Image credits: NASA/JPL-Caltech/MSSS.

alkali feldspar (variety sanidine), tridymite, cristobalite, and trace amounts of quartz. Mafic igneous minerals identified include pyroxene and olivine. The CheMin team has reported on pyroxene compositions previously (Blake et al., 2013; Bish et al., 2013; Vaniman et al., 2014; Treiman et al., 2016; Rampe et al., 2017; Bristow et al., 2018; Morrison et al., 2018b; Rampe et al., 2018). Overlapping peaks of pyroxenes in CheMin X-ray diffraction patterns and the low angular resolution of the instrument, however, preclude confident identification of pyroxene phases or crystal chemistry. We report the published pyroxene subtype compositions for several samples in Table 3, but caution that there is a great deal of uncertainty associated with them and they should not be used for detailed petrologic interpretations.

The most common Fe-oxide minerals include hematite and magnetite, but the Fe-oxyhydroxide mineral, akaganeite [ $\beta$ -FeO(OH,Cl)], has been identified in the Bradbury group on the plains of Gale crater (Vaniman et al., 2014; Treiman et al., 2016) and very recently from the Pettegrove Point and Jura members of the Murray formation (Morrison et al., 2019; Rampe et al., 2020). For samples that contain at least minor amounts of magnetite, the refined magnetite unit-cell parameters are smaller than those of stoichiometric magnetite, suggesting partial oxidation of the Fe(II) toward the Fe(III) defect-spinel maghemite. Alternatively, substitution of smaller cations such as  $\text{Cr}^{3+}$ ,  $\text{Mg}^{2+}$ , or  $\text{Al}^{3+}$  could account for a smaller unit cell, although it is not clear whether sufficient amounts of these cations are present (Morrison et al., 2018a, b).

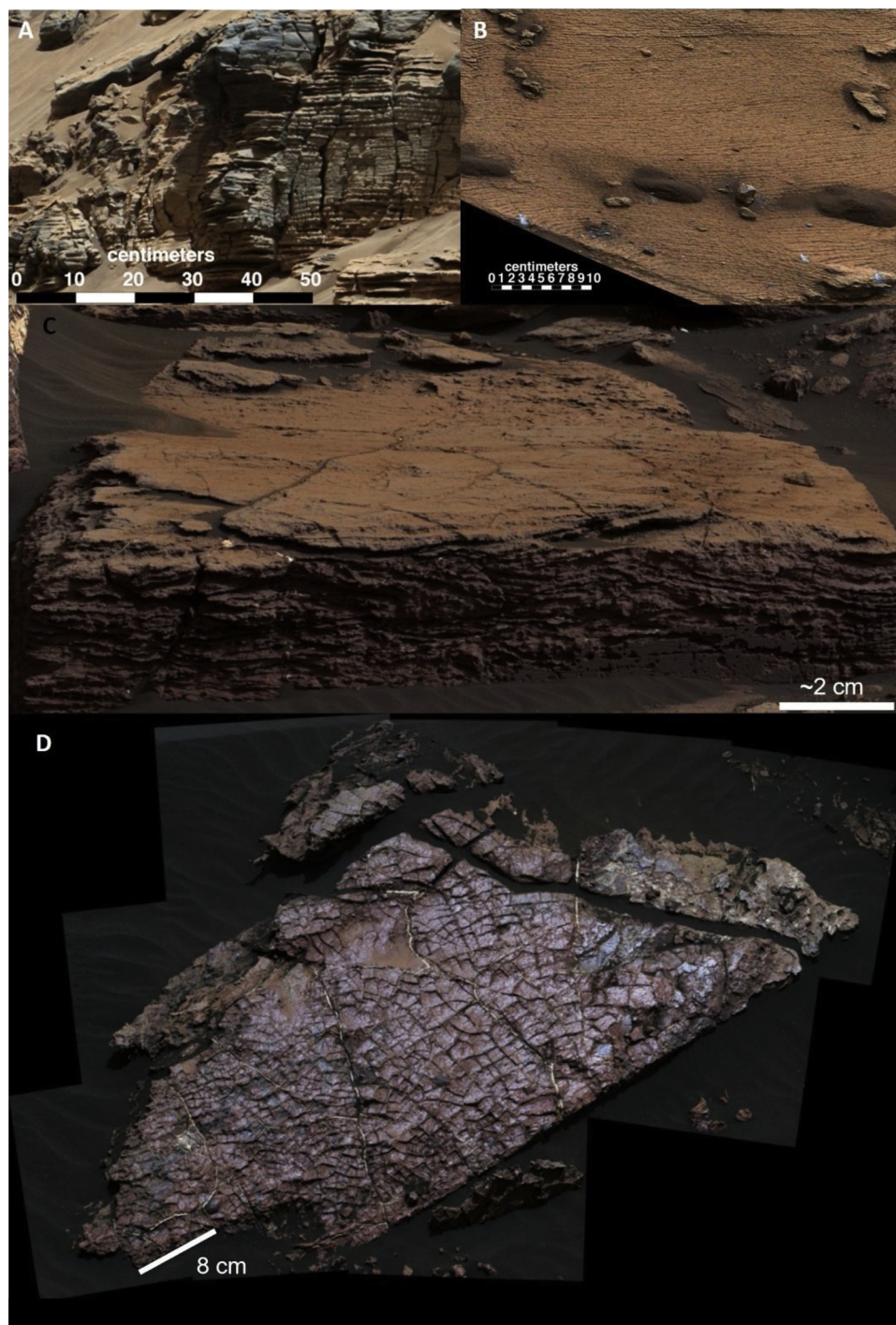
Phyllosilicates are readily identified in CheMin diffraction data by the presence and position of (001) peaks (Fig. 15), which are a measure of interlayer spacing along the c-axis of the crystal. Most patterns show a d(001) of  $\sim 10$  Å, which is consistent with a collapsed (i.e., dehydrated) swelling clay mineral (e.g., smectite) or illite. One sample from Yellowknife Bay, named “Cumberland,” has a d(001) up to  $\sim 13.5$  Å that may indicate intercalated metal-hydroxyl groups in the smectite

interlayer site (Bristow et al., 2015). The low abundance of potassium in most samples suggests little, if any, illite is present (Bristow et al., 2015, 2018). The detection of collapsed smectite confirms the hypothesis that smectite on the martian surface is dehydrated based on orbital spectral data (e.g., Wray et al., 2009; Che et al., 2011). The position of the (02l) diffraction band, which is a measure of the b-axis of the crystal (Fig. 15), suggests a change from trioctahedral smectite to smectite with more dioctahedral character moving up section (Bristow et al., 2015, 2018). SAM evolved gas data show complementary evidence for changes in smectite structure and composition within the stratigraphic section based on the temperature of structural  $\text{H}_2\text{O}$  release (Ming et al., 2014; Bristow et al., 2018). Phyllosilicates in the samples collected from the Hartmann’s Valley member and the members that compose Vera Rubin ridge have d(001) peaks at  $\sim 9.6$  Å, consistent with a completely collapsed smectite or Fe-pyrophyllite (Bristow et al., 2018; Rampe et al., 2020).

The Ca-sulfate minerals anhydrite, bassanite, and gypsum are common (Vaniman et al., 2014, 2018), especially up section, and jarosite is present in minor to trace amounts in many of the samples collected from the Murray formation on the lower slopes of Mount Sharp (Rampe et al., 2017; Bristow et al., 2018; Achilles et al., submitted, Rampe et al., 2020). Minor amounts of fluorapatite were identified in some samples from the Pahrump Hills member at the base of the Murray formation, and trace amounts of halite were identified in one sample from the Karasburg member in the Murray formation (“Quela”).

X-ray amorphous materials are abundant ( $\sim 20$ – $60$  wt.%) in all fluvio-lacustrine samples. X-ray amorphous materials are identified in CheMin diffraction patterns from the presence of a broad hump in the background centered between  $\sim 22$ – $26^\circ 2\theta$  and scattering at low angles (Fig. 5). The identity of the X-ray amorphous materials is difficult to ascertain from XRD data alone because they generally do not have distinct XRD patterns. The relative position of the X-ray amorphous





**Fig. 9.** Images of outcrops from the Murray formation. Thickly (A) and thinly (B) laminated mudstone deposits from the Pahrump Hills (sol 712, mcam03030; sol 792, mcam03445). C) Ripple cross-lamination in the Sutton Island member (sol 1477, mcam07413). D) Desiccation cracks in the “Old Soaker” target (sol 1555, mcam07981). Image credits: NASA/JPL-Caltech/MSSS.

hump can be used to discern relative SiO<sub>2</sub> abundances in amorphous silicates, where high-SiO<sub>2</sub> amorphous silicates (e.g., opal-A, rhyolitic glass) have a hump centered at lower angles, whereas amorphous silicates with lower SiO<sub>2</sub> (e.g., basaltic glass) have a hump centered at higher angles. This trend should not be used, however, to characterize the amorphous component when multiple amorphous materials are present in the same sample, as is the case for samples from Gale crater. XRD patterns of amorphous sulfate, for example, look very similar to amorphous silicates (Morris et al., 2015).

Calculations of the composition of the X-ray amorphous component using CheMin mineral abundances and APXS bulk chemistry of the

post-sieve or bulk powder dump piles demonstrate a large variation in compositions between samples (Table 6). There are many sources of error for the calculated amorphous component (e.g., abundance of crystalline phases, abundance of amorphous material, calculated crystal chemistry of major phases, assumed composition of phyllosilicates), so the calculated amorphous component likely represents the amorphous composition plus the residual error of the calculated composition based on the model mineralogy and mineral chemistry determined by Rietveld refinement. The aliquot of sample delivered to CheMin comes from the same material as the post-sieve or bulk powder dump piles, so the bulk APXS compositions of these dump piles are likely similar to those

**Table 1**  
Information about samples analyzed by CheMin.

Sample (abbreviation)	Sol(s) Collected	Elevation (m)	Depositional Environment
Rocknest (RN)	61, 66, 69, 74, 93	-4516.9	Modern eolian (inactive)
John Klein (JK)	182	-4519.5	Lacustrine
Cumberland (CB)	279	-4519.5	Lacustrine
Windjana (WJ)	621	-4481.5	Reworked eolian and fluvial
Confidence Hills (CH)	759	-4460.3	Lacustrine
Mojave2 (MJ)	882	-4459.4	Lacustrine
Telegraph Peak (TP)	908	-4453.2	Lacustrine
Buckskin (BK)	1060	-4446.8	Lacustrine
Big Sky (BS)	1119	-4434.7	Ancient eolian
Greenhorn (GH)	1137	-4434.5	Ancient eolian (halo)
Gobabeb (GB)	1224	-4423.8	Modern eolian (active)
Lubango (LB)	1320	-4429.0	Ancient eolian (halo)
Okoruso (OK)	1332	-4429.3	Ancient eolian
Oudam (OU)	1361	-4435.5	Reworked eolian and fluvial
Marimba (MB)	1422	-4410.4	Lacustrine
Quela (QL)	1464	-4379.7	Lacustrine
Sebina (SB)	1495	-4360.8	Lacustrine
Ogunquit Beach (OG)	1651	-4300.0	Modern eolian (active)
Duluth (DU)	2057	-4192.5	Lacustrine
Stoer (ST)	2136	-4169.9	Lacustrine
Highfield (HF)	2223	-4147.0	Lacustrine
Rock Hall (RH)	2261	-4143.8	Lacustrine

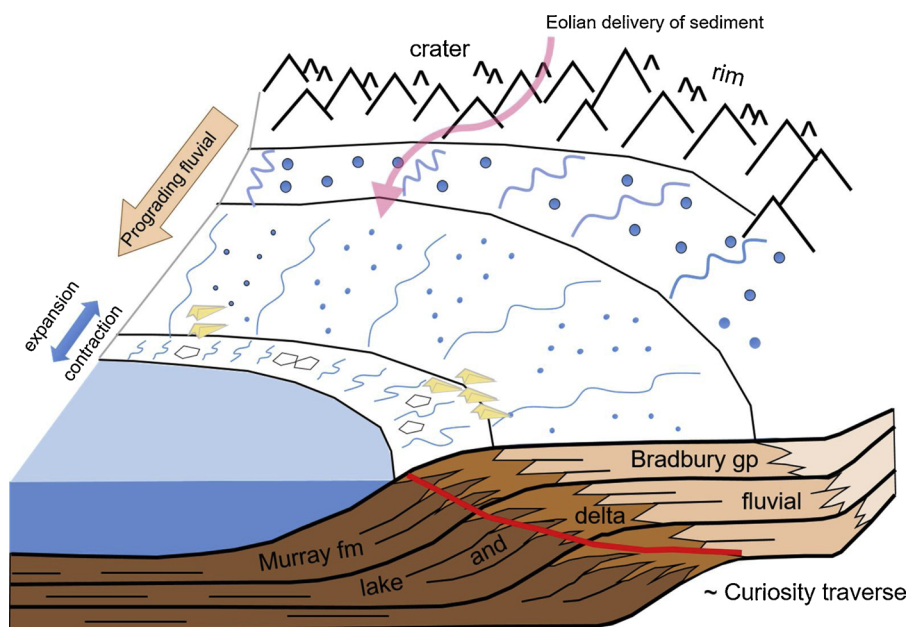
of the samples inside CheMin. Therefore, differences between the composition of the CheMin-analyzed sample and the APXS-analyzed sample are unlikely to be significant sources of error.

Despite some sources of error for these calculations, the results can be used to examine relative elemental enrichments and depletions between samples (e.g., Dehouck et al., 2014; Rampe et al., 2017; Yen et al., 2017). The amorphous component in all samples contains  $\text{SO}_3$ ,  $\text{FeO}_T$ , and  $\text{SiO}_2$ , suggesting amorphous sulfates, nanophase Fe-oxides, and silicates are present in all samples. The amorphous component in many samples contains abundant  $\text{FeO}_T$  (e.g., “Windjana,” “Marimba,” and “Duluth”), suggesting the presence of nanophase Fe-oxides and/or amorphous Fe-sulfates. The amorphous component in some samples is relatively enriched in  $\text{SiO}_2$  (e.g., “Buckskin” and “Oudam”), suggesting the presence of opal-A and/or rhyolitic glass. Aluminum abundances

are relatively low in all samples, suggesting amorphous aluminosilicates like allophane and maskelynite do not make up a significant portion of the amorphous component. None of the calculated compositions of amorphous material is consistent with pure basaltic glass. Although some basaltic glass could be present, low amounts of  $\text{SiO}_2$ ,  $\text{Al}_2\text{O}_3$ ,  $\text{MgO}$ , and/or  $\text{CaO}$  in all samples preclude abundant basaltic glass. This finding is consistent with Mössbauer data collected by the Mars Exploration Rovers in Gusev crater and Meridiani Planum, which demonstrate that Fe is not present in basaltic glass at either landing site (Morris et al., 2006, 2008).

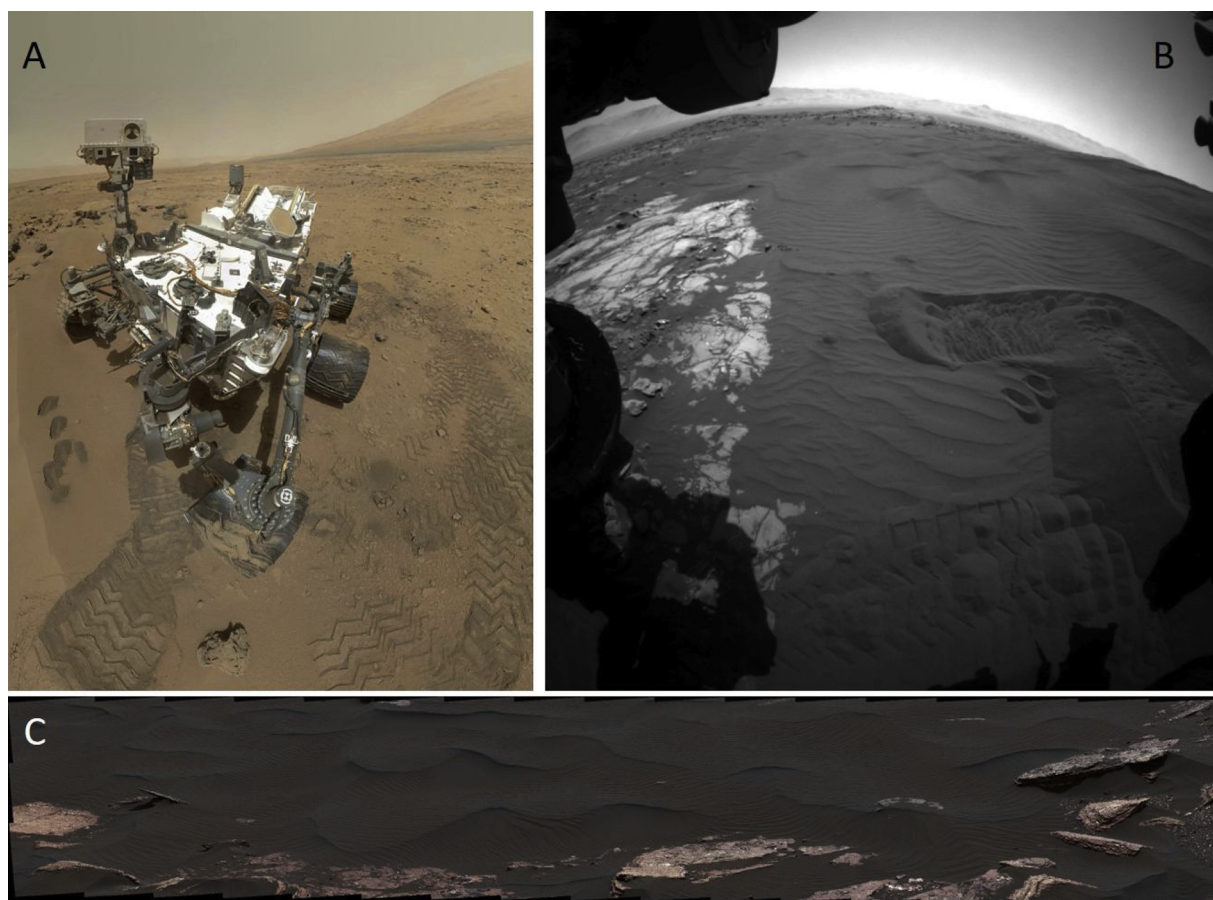
The dominant oxides present in the amorphous component in each sample could provide further information about aqueous conditions. High concentrations of  $\text{FeO}_T$  suggest the presence of an amorphous Fe phase, like ferrihydrite. Ferrihydrite-like materials form by rapid hydrolysis of  $\text{Fe}^{3+}$  in solution (e.g., Schwertmann and Cornell, 2000). Aqueous  $\text{Fe}^{3+}$  can be mobilized at low pH, or it could form by rapid oxidation and hydrolysis of  $\text{Fe}^{2+}$  in solution that was sourced from the alteration of mafic igneous minerals. Elevated  $\text{FeO}_T$  and  $\text{SO}_3$  in the amorphous component could indicate the presence of amorphous Fe-sulfate, which would implicate acidic conditions. Abundant  $\text{SiO}_2$  in the amorphous component suggests the presence of opaline silica. There are no diffraction properties to refine the identification of opal-A, but in some samples, opal-CT has been identified and based on characteristic diffraction peaks in this paracrystalline material (Table 5). Opaline silica can form in a variety of environments, and is an expected alteration product of basaltic sediments because Ca-rich plagioclase and mafic igneous silicates are more susceptible to dissolution than felsic minerals (e.g., McLennan, 2003). The presence of abundant amorphous silica may signify aqueous alteration and dissolution of basaltic minerals or it may be a primary volcanic glass in the case of the “Buckskin” sample.

The abundance of X-ray amorphous materials in the ancient rock samples is especially perplexing because amorphous materials in rocks on Earth mature into crystalline phases over time. Opaline silica, for example, transforms to quartz through a dissolution-precipitation reaction in early diagenetic environments (e.g., Kastner et al., 1977). On Earth, amorphous silica is not usually found in rocks over a few million years old and is absent in rocks over ~145 Ma old (e.g., Tosca and Knoll, 2009). Compositionally complex amorphous materials, similar to those identified in Gale crater, have been identified in modern subglacial sediments formed from basaltic andesite and andesite sources in the Three Sisters Volcanic Complex in Oregon (Smith et al., 2018). These materials are enriched in  $\text{SiO}_2$  (~50–60 wt.%),  $\text{Al}_2\text{O}_3$  (~15 wt.

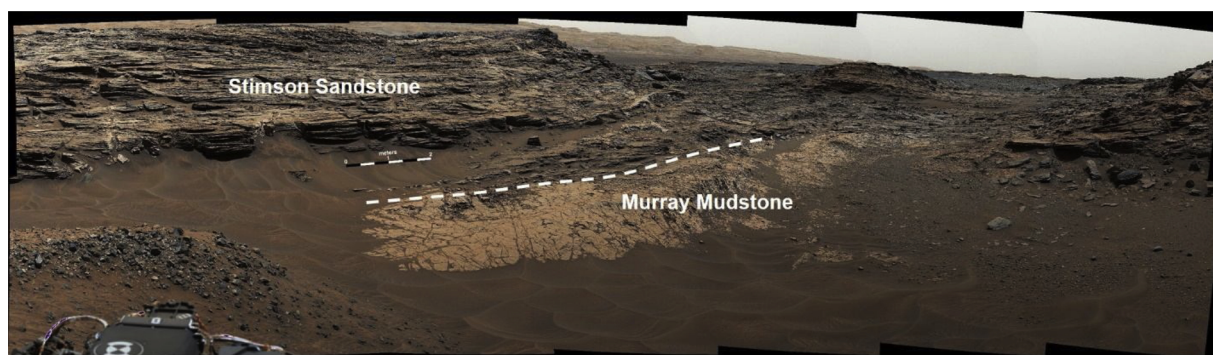


**Fig. 10.** Schematic showing the depositional environments preserved by the Bradbury Group and Murray formation. The Bradbury group is dominated by fluvio-deltaic deposits as rivers and streams emanating from the rim delivered sediments across the plains and into lakes on the crater floor. The Murray formation is dominated by lacustrine deposits. Desiccation features and salts would have precipitated along the lake margin. *Curiosity's* traverse is drawn in red, showing the beginning of the traverse in fluvio-deltaic deposits then transitioning to lacustrine and lake-margin deposits. Credit: C. Fedo. (For interpretation of the references to colour in this figure legend, the reader is referred to the web version of this article.)





**Fig. 11.** *Curiosity's* sampling sites of modern eolian materials. A) *Curiosity* selfie taken by MAHLI showing scoops from “Rocknest” (RN) sand shadow (sol 84). *Curiosity's* selfies are a mosaic of dozens of MAHLI images. B) Front Hazcam image showing wheel scuff and scoops of “Gobabeb” (GB) from Namib Dune (sol 1229, fhaz00323). C) Meter-scale ripples at Mount Desert Island sand patch, from which “Ogunquit Beach” (OG) was scooped (sol 1647, mcam08526). Images credit: NASA/JPL-Caltech/MSSS.



**Fig. 12.** Unconformable contact between the Murray and Stimson formations at the Marias Pass location at the top of the Pahrump Hills member (sol 992, mcam04393). The contact is denoted as a dashed line. The “Buckskin” sample was drilled from this location. Scale bar is 2 m. Image credit: NASA/JPL-Caltech/MSSS.

%), and  $\text{FeO}_T$  (~15 wt.%) and have lesser amounts of CaO, MgO,  $\text{Na}_2\text{O}$ , and  $\text{K}_2\text{O}$ . Transmission electron microscopy of the sediments indicates that much of the amorphous component is secondary in nature, suggesting it formed from water-rock interactions at relatively low temperatures. Although these materials have more  $\text{Al}_2\text{O}_3$  and less  $\text{SO}_3$  than amorphous materials in Gale crater, the formation of compositionally heterogeneous amorphous materials in cold and icy environments may explain some portion of the amorphous component in Gale crater. Their preservation for perhaps billions of years, however, suggests groundwater was intermittent and was not present in sufficient volumes to cause the maturation of amorphous materials to crystalline phases.

### 5.2. Modern eolian samples

The mineralogy of modern eolian sediments can be used to evaluate the extent of sorting and grain segregation in a single bedform and across dune fields, the relative contribution of underlying bedrock to the sediment, and the composition of modern sediment sources. The in-situ measurements by CheMin also allow us to ground truth orbital remote sensing measurements and clarify differences between the in-situ and orbital measurements. CheMin has analyzed three modern eolian sediments. “Rocknest” was scooped from an inactive sand shadow near the landing site, whereas “Gobabeb” and “Ogunquit



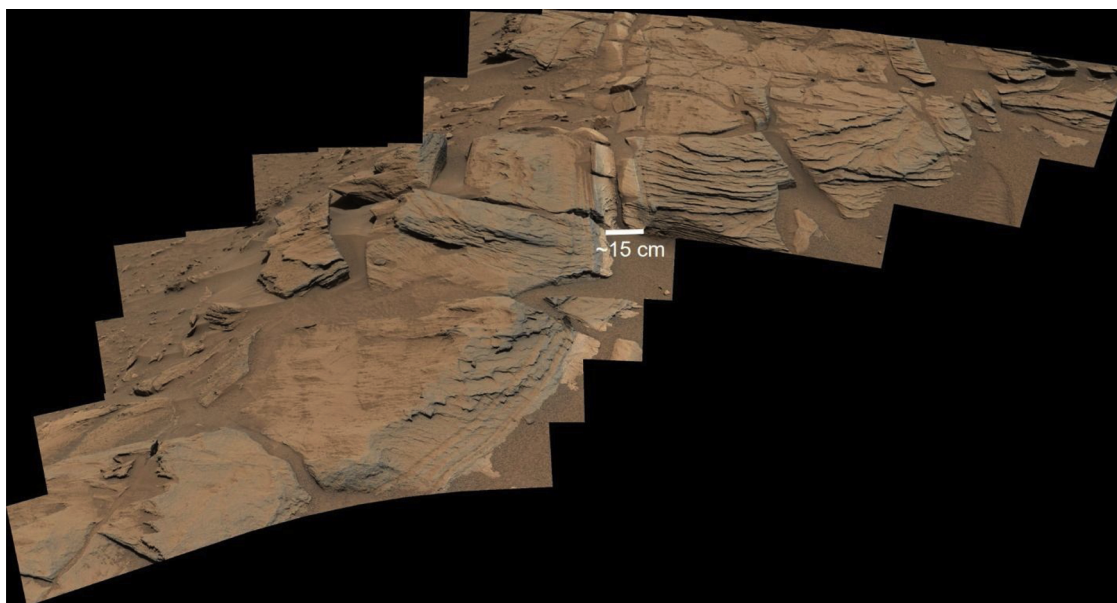


Fig. 13. Mastcam mosaic of a fracture-associated halo observed at Teakettle Junction (sol 747, mcam03215). Image credit: NASA/JPL-Caltech/MSSS.

Beach” were scooped from different locations in the active Bagnold Dune Field. “Gobabeb” was scooped from a barchanoid dune named Namib Dune near the trailing edge of the dune field and “Ogunquit Beach” from a ripple field downwind of Namib Dune. All three samples have similar mineralogy and are dominated by basaltic igneous minerals and X-ray amorphous materials with minor amounts of Fe-oxide minerals (magnetite and hematite are present in all samples and ilmenite is present in “Rocknest”), anhydrite, and quartz (Table 7; Fig. 16). Mineralogical results from the modern eolian samples are reported by Blake et al. (2013); Bish et al. (2013); Achilles et al. (2017); Morrison et al. (2018b), and Rampe et al. (2018).

The basaltic igneous minerals present in each sample include plagioclase, pyroxene, and olivine. Plagioclase constitutes 38–47 wt.% of the crystalline fraction, and the calculated plagioclase compositions (Table 8) from the refined unit-cell parameters for each sample are An<sub>49(4)</sub>, An<sub>63(5)</sub>, and An<sub>48(5)</sub> for Rocknest, Gobabeb, and Ogunquit Beach, respectively. Each sample has ~26–33 wt.% pyroxene. The

presence of both augite and pigeonite were reported previously (Bish et al., 2013; Blake et al., 2013; Achilles et al., 2017; Morrison et al., 2018b; Rampe et al., 2018) but, again, we caution that the low angular resolution of the CheMin instrument and overlapping pyroxene peaks reduces our confidence in these identifications and their calculated crystal compositions from refined unit-cell parameters. Olivine is present in abundances of 18–26 wt.% and the olivine composition is the same between all samples within a 1 $\sigma$  error (Fo<sub>54-60</sub>).

The major constituents of the calculated amorphous components are SiO<sub>2</sub>, FeO<sub>T</sub>, Al<sub>2</sub>O<sub>3</sub>, and SO<sub>3</sub> (Table 9). Thus, like the amorphous materials in fluvio-lacustrine samples, they are not consistent with any one single X-ray amorphous material. When compared to an Adirondack basalt composition in Gusev crater, basaltic glass is inconsistent with the calculated amorphous compositions. The Rocknest amorphous component is too depleted in SiO<sub>2</sub>, the Gobabeb amorphous component is too enriched in SiO<sub>2</sub>, and the amorphous components from all modern eolian samples are too depleted in Al<sub>2</sub>O<sub>3</sub>, MgO, and CaO and

Table 2

Mineral abundances (in wt.%, renormalized to 100% crystalline) from CheMin data of fluvio-lacustrine samples. Plag = plagioclase, Kspar = potassium feldspar, Px = pyroxene, Oliv = olivine, Qtz = quartz, Crist = cristobalite, Trid = tridymite, Magn = magnetite, Hem = hematite, Akag = akaganeite, F-ap = fluorapatite, Anhy = anhydrite, Bass = bassanite, Gyps = gypsum, Jaro = jarosite, Hal = halite. Uncertainties in parentheses are 1 $\sigma$ . Abundances and errors have been rounded to the nearest whole number. The symbol “–” indicates that phase is at or below the detection limit of CheMin.

	JK <sup>1</sup>	CB <sup>1</sup>	WJ <sup>1</sup>	CH <sup>1</sup>	MJ <sup>1</sup>	TP <sup>1</sup>	BK <sup>1</sup>	OU <sup>2</sup>	MB <sup>2</sup>	QL <sup>2</sup>	SB <sup>2</sup>	DU <sup>3</sup>	ST <sup>3</sup>	HF <sup>3</sup>	RH <sup>3</sup>
Plag	44(1)	43(1)	6(1)	38(2)	55(2)	44(2)	43(2)	52(1)	44(3)	44(3)	38(2)	56(1)	45(2)	47(2)	38(4)
Kspar	4(1)	5(1)	26(1)	9(1)	–	8(2)	8(1)	–	8(2)	7(1)	4(1)	7(1)	3(3)	3(1)	–
Px	26(2)	31(2)	42(3)	26(7)	16(2)	13(3)	–	8(2)	2(2)	7(2)	7(4)	10(3)	6(2)	10(3)	17(2)
Oliv	8(0)	3(1)	7(2)	2(1)	1(1)	2(1)	–	–	–	–	–	–	–	–	–
Qtz	–	–	–	1(1)	2(0)	1(0)	–	2(0)	2(1)	1(1)	1(0)	2(1)	1(1)	1(1)	–
Crist	–	–	–	–	–	12(2)	6(0)	–	–	–	–	–	–	–	–
Trid	–	–	–	–	–	–	34(1)	–	–	–	–	–	–	–	–
Magn	9(1)	11(0)	16(1)	6(1)	7(1)	13(1)	7(0)	–	–	–	–	2(1)	1(1)	1(1)	–
Hem	1(0)	1(0)	2(1)	13(1)	7(1)	2(0)	–	26(1)	20(1)	23(2)	24(2)	13(2)	28(1)	20(1)	5(0)
Akag	1(0)	2(1)	1(1)	–	–	–	–	–	–	–	–	–	2(1)	–	11(1)
F-ap	–	–	–	2(1)	4(1)	3(0)	–	–	–	–	–	–	–	–	3(1)
Anhy	4(0)	1(0)	1(0)	–	–	–	2(0)	6(1)	12(1)	9(1)	17(1)	3(1)	6(1)	8(1)	21(3)
Bass	3(0)	2(0)	1(0)	–	–	–	–	–	4(1)	6(0)	4(0)	7(1)	1(0)	3(1)	–
Gyps	–	–	–	–	–	–	–	6(0)	7(1)	1(1)	4(0)	–	4(1)	5(1)	–
Jaro	–	–	–	2(1)	8(2)	2(2)	–	–	2(1)	1(1)	3(1)	–	2(1)	–	4(1)
Hal	–	–	–	–	–	–	–	–	–	1(0)	–	–	–	–	–

<sup>1</sup> From Morrison et al. (2018b).

<sup>2</sup> From Achilles et al. (submitted).

<sup>3</sup> From Rampe et al. (2020).

**Table 3**Crystal chemistry of major igneous phases in fluvio-lacustrine samples calculated from refined unit-cell parameters. Uncertainties in parentheses are  $1\sigma$ .

	Plag	K-spar	Augite	Pigeonite	OPX	Olivine
JK <sup>1</sup>	An <sub>43(4)</sub>	Or <sub>53(18)</sub>		En <sub>60(8)</sub> Fs <sub>35(9)</sub> Wo <sub>5(4)</sub>	En <sub>31(3)</sub> Fs <sub>69</sub>	Fo <sub>60(12)</sub>
CB <sup>1</sup>	An <sub>35(5)</sub>	Or <sub>77(19)</sub>		En <sub>58(9)</sub> Fs <sub>37(10)</sub> Wo <sub>6(5)</sub>	En <sub>27(5)</sub> Fs <sub>73</sub>	Fo <sub>60(21)</sub>
WJ <sup>1</sup>	An <sub>28(45)</sub>	Or <sub>87(5)</sub>	En <sub>50(7)</sub> Fs <sub>9(9)</sub> Wo <sub>42(6)</sub>	En <sub>70(12)</sub> Fs <sub>28(13)</sub> Wo <sub>3(5)</sub>		Fo <sub>72(8)</sub>
CH <sup>1</sup>	An <sub>40(4)</sub>	Or <sub>82(11)</sub>		En <sub>57(19)</sub> Fs <sub>43(19)</sub> Wo <sub>0(6)</sub>		
MJ <sup>1</sup>	An <sub>42(3)</sub>			En <sub>60(18)</sub> Fs <sub>41(19)</sub> Wo <sub>0(6)</sub>		
TP <sup>1</sup>	An <sub>37(3)</sub>	Or <sub>69(11)</sub>		En <sub>58(23)</sub> Fs <sub>43(24)</sub> Wo <sub>0(7)</sub>		
BK <sup>1</sup>	An <sub>41(3)</sub>	Or <sub>77(14)</sub>				
OU <sup>2</sup>	An <sub>40(7)</sub>					
MB <sup>2</sup>	An <sub>39(5)</sub>					
QL <sup>2</sup>	An <sub>39(6)</sub>					
SB <sup>2</sup>	An <sub>42(6)</sub>					

<sup>1</sup> From Morrison et al. (2018b).<sup>2</sup> From Achilles et al. (2020).**Table 4**Magnetite unit-cell parameters and crystal chemistry for fluvio-lacustrine samples<sup>1</sup>. Uncertainties in parentheses are  $1\sigma$ .

Sample	<i>a</i> cell length (in Å) <sup>2</sup>	Chemical Formula
JK	8.372(2)	Fe <sub>2.82(5)</sub> □ <sub>0.18</sub> O <sub>4</sub>
CB	8.369(2)	Fe <sub>2.81(5)</sub> □ <sub>0.19</sub> O <sub>4</sub>
WJ	8.373(1)	Fe <sub>2.83(5)</sub> □ <sub>0.17</sub> O <sub>4</sub>
CH	8.365(3)	Fe <sub>2.79(5)</sub> □ <sub>0.21</sub> O <sub>4</sub>
MJ	8.357(2)	Fe <sub>2.76(5)</sub> □ <sub>0.24</sub> O <sub>4</sub>
TP	8.355(1)	Fe <sub>2.75(5)</sub> □ <sub>0.25</sub> O <sub>4</sub>
BK	8.359(1)	Fe <sub>2.77(5)</sub> □ <sub>0.23</sub> O <sub>4</sub>

<sup>1</sup> From Morrison et al. (2018b).<sup>2</sup> The *a* cell length for stoichiometric magnetite is 8.3969 Å.**Table 5**X-ray amorphous and phyllosilicate abundances in fluvio-lacustrine samples from FULLPAT analyses\*. Uncertainties in parentheses are  $1\sigma$ . The symbol “–” indicates that material was not detected in FULLPAT models.

Sample	Phyllosilicate	Amorphous	Opal-CT
JK <sup>1</sup>	22(2)	28	–
CB <sup>1</sup>	18(2)	31	–
WJ <sup>2</sup>	10(2)	15	–
CH <sup>3</sup>	8(2)	39(15)	–
MJ <sup>3</sup>	5(1)	53	–
TP <sup>3</sup>	–	27(15)	11(2)
BK <sup>4</sup>	–	54	6(1)
OU <sup>5,6</sup>	3(1)	44(11)	7(1)
MB <sup>5,6</sup>	28(3)	40(10)	–
QL <sup>5,6</sup>	16(2)	52(13)	–
SB <sup>5,6</sup>	19(2)	51(13)	–
DU <sup>7</sup>	15(4)	37	–
ST <sup>7</sup>	10(3)	38	–
HF <sup>7</sup>	5(1)	49	4(1)
RH <sup>7</sup>	13(3)	34(8)	–

\* Abundances of amorphous phases for JK, CB, WJ, MJ, BK, DU, ST, and HF were constrained by CheMin crystalline abundances and APXS bulk chemistry, because the FULLPAT-derived values resulted in compositions with negative oxide values.

<sup>1</sup> From Vaniman et al. (2014).<sup>2</sup> From Treiman et al. (2016).<sup>3</sup> From Rampe et al. (2017).<sup>4</sup> From Morris et al. (2016).<sup>5</sup> From Bristow et al. (2018).<sup>6</sup> From Achilles et al. (submitted).<sup>7</sup> From Morris et al. (2019).

are too enriched in SO<sub>3</sub>. This composition suggests that multiple X-ray amorphous materials constitute the X-ray amorphous component, including volcanic and/or impact glass, nanophase Fe-oxides, and amorphous sulfates. There are clear differences in the compositions of

the amorphous components in active eolian samples (“Gobabeb” and “Ogunquit Beach”) and in the inactive eolian sample (“Rocknest”). The amorphous component of the active samples is more enriched in SiO<sub>2</sub>, whereas the amorphous component of the inactive sample is more enriched in SO<sub>3</sub> and Cl. The composition of dust on *Curiosity*’s observation tray as measured by APXS is enriched in SO<sub>3</sub> and Cl (Berger et al., 2016), so this difference in composition between the active and inactive eolian samples is a result of higher dust contents in the inactive sediments (Achilles et al., 2017).

### 5.3. Ancient eolian samples

*Curiosity* drilled two pairs of samples from the Stimson formation, a lithified eolian dune field that unconformably overlies the Murray formation. These samples were selected because of their association with light-toned alteration halos parallel to fractures in the rock (Fig. 13). One sample of each pair was collected from parent rock and the second from within the alteration halo: “Big Sky” (parent rock) was drilled 2 m from “Greenhorn” (alteration halo), and “Okoruso” (parent rock) was drilled 3 m from “Lubango” (alteration halo), and the “Okoruso”/“Lubango” pair was collected ~650 m southwest of “Big Sky”/“Greenhorn.” APXS and ChemCam measurements were also made in parent rock and in the alteration halos to identify geochemical trends associated with the processes that formed the halos. The mineralogical and geochemical trends observed by *Curiosity* are reported in detail by Yen et al. (2017).

CheMin analyses of the parent eolian sandstone samples (“Big Sky” and “Okoruso”) show that they are dominated by igneous minerals, magnetite, and X-ray amorphous materials (Table 10). The parent sandstone contains abundant plagioclase feldspar and pyroxene (modeled as pigeonite and orthopyroxene; Table 11) and lacks olivine. Magnetite constitutes ~15 wt.% of the crystalline component, and it may have formed diagenetically from the dissolution of olivine and partial oxidation of Fe(II)-bearing fluids (Yen et al., 2017; Hausrath et al., 2018). The bulk sample comprises ~20 to 35 wt.% X-ray amorphous materials. The X-ray amorphous component is depleted in SiO<sub>2</sub> relative to a basaltic composition and contains abundant FeO<sub>T</sub> and MgO (Table 12), which is also consistent with alteration of olivine. The amorphous component also contains SO<sub>3</sub>, suggesting nanophase Fe-oxides and amorphous sulfates are present in the parent sandstone. These amorphous materials, in addition to the magnetite, are likely cementing agents. Minor amounts of sanidine, hematite, quartz, fluorapatite, and Ca-sulfate are also present in the parent sandstone.

The alteration halos have a distinctly different mineralogy and geochemical composition from the parent sandstone. The samples drilled from halos show a significant decrease in igneous mineral abundance (particularly for pyroxene), an enrichment in Ca- and mixed-cation sulfates, and a three-fold increase in X-ray amorphous materials. The calculated composition of the X-ray amorphous

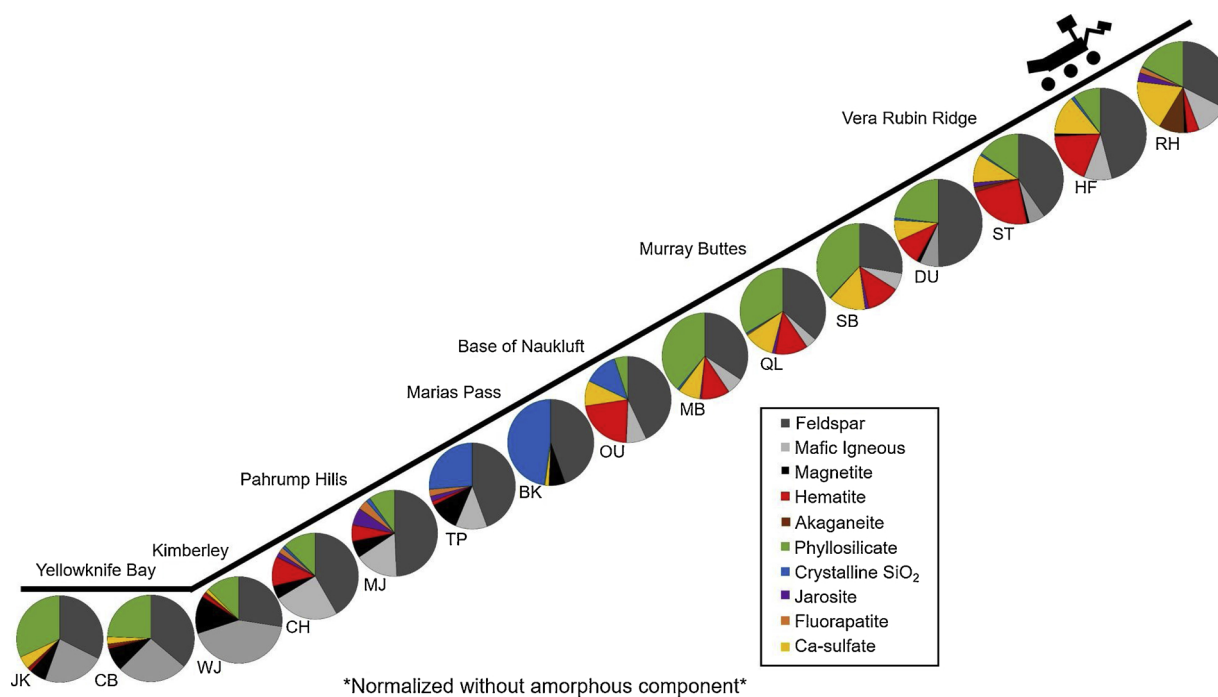
**Table 6**  
Calculated X-ray amorphous component composition for fluvio-lacustrine samples, renormalized to 100%.

	JK <sup>1</sup>	CB <sup>1</sup>	WJ <sup>1</sup>	CH <sup>1</sup>	MJ <sup>1</sup>	TP <sup>1</sup>	BK <sup>1</sup>	OU <sup>2</sup>	MB <sup>2</sup>	QL <sup>2</sup>	SB <sup>2</sup>	DU <sup>3</sup>	ST <sup>3</sup>	HF <sup>3</sup>	RH <sup>3</sup>
SiO <sub>2</sub>	33.0	25.6	3.6	53.2	55.7	53.0	75.9	64.3	45.4	45.4	49.8	34.3	55.0	61.7	38.1
TiO <sub>2</sub>	3.3	3.7	5.2	3.1	2.5	5.9	2.8	2.1	1.8	2.0	2.1	2.6	2.7	1.7	2.8
Al <sub>2</sub> O <sub>3</sub>	0.0	0.0	3.9	5.7	7.2	1.9	0.0	5.4	5.8	4.0	4.4	0.0	5.1	7.9	5.2
Cr <sub>2</sub> O <sub>3</sub>	1.5	1.7	2.4	1.1	0.8	1.7	0.2	0.6	0.6	0.6	0.6	0.8	0.9	0.6	0.9
FeO <sub>T</sub>	20.9	18.0	43.1	13.7	12.3	17.1	4.8	11.2	29.7	20.2	18.8	26.1	3.0	10.2	17.5
MnO	1.0	1.1	2.7	1.0	0.8	1.2	0.2	0.4	0.1	0.4	0.3	0.5	0.4	0.3	0.2
MgO	13.5	17.5	17.7	4.8	5.8	3.5	1.5	8.3	0.0	1.7	2.6	4.0	11.1	6.7	5.4
CaO	9.2	11.1	0.6	0.8	0.0	1.4	2.1	0.0	3.9	7.6	5.8	8.0	4.4	1.5	2.9
Na <sub>2</sub> O	2.3	1.7	1.0	2.1	1.8	2.3	1.2	1.7	2.3	1.9	2.2	0.1	2.0	1.4	2.3
K <sub>2</sub> O	0.9	0.5	0.0	0.0	1.1	0.0	0.8	1.7	0.9	0.8	0.9	1.3	1.2	1.3	0.4
P <sub>2</sub> O <sub>5</sub>	2.9	3.2	3.1	1.0	0.8	1.4	2.2	1.0	1.8	2.1	1.2	2.8	2.2	1.6	0.2
SO <sub>3</sub>	10.0	11.5	13.9	12.3	10.3	9.2	7.8	2.5	6.9	11.8	9.4	17.3	10.6	3.3	23.4
Cl	1.6	4.4	2.8	1.1	0.9	1.4	0.5	0.7	0.8	1.7	2.0	2.2	1.3	1.8	0.7
F	0.0	0.0	0.0	0.1	0.2	0.1	0.0	0.0	0.0	0.0	0.0	0.0	0.0	0.0	0.0

<sup>1</sup> From Morrison et al. (2018b).

<sup>2</sup> From Achilles et al. (submitted).

<sup>3</sup> From Rampe et al. (2020).



**Fig. 14.** Mineral abundances of fluvio-lacustrine deposits determined from CheMin X-ray diffraction data. Abundances of X-ray amorphous components are not included in the pie charts here, but each sample has ~25–60 wt.% X-ray amorphous material. JK = John Klein, CB = Cumberland, WJ = Windjana, CH = Confidence Hills, MJ = Mojave2, TP = Telegraph Peak, BK = Buckskin, OU = Oudam, MB = Marimba, QL = Quela, SB = Sebina, DU = Duluth, ST = Stoer, HF = Highfield, RH = Rock Hall. Image credit: NASA/JPL-Caltech.

component suggests it is enriched in SiO<sub>2</sub>, consistent with abundant opaline silica. The halos are extensively depleted in Mg, Al, Mn, Fe, Ni, and Zn, consistent with a preferential loss of mafic phases as a result of aqueous alteration, and are enriched in Si and S (Table 12). The study of alteration halos by *Curiosity* was especially significant because the halos formed after the deposition and lithification of the Stimson formation, so they represent late-stage aqueous processes at Gale crater.

## 6. Diverse igneous sediment sources in and around Gale crater

CheMin has not analyzed parent igneous rocks, but the igneous minerals found in modern eolian sediments and ancient sedimentary rocks can provide some insight into the igneous history and source rocks near Gale crater. All samples analyzed by CheMin contain igneous minerals, and their variety suggests sediments were derived from both

primitive and evolved igneous sources. The abundance of basaltic igneous minerals (i.e., plagioclase, pyroxene, and olivine) in modern and ancient eolian samples suggests mafic igneous sources for these sediments. Olivine-bearing units identified on the rim of Gale crater from orbital visible and near- to short-wave infrared reflectance spectral data collected by the Compact Reconnaissance Imaging Spectrometer for Mars (CRISM) on the Mars Reconnaissance Orbiter and/or dark mafic sedimentary units in Mount Sharp may be the source for the modern eolian sediments (Ehlmann and Buz, 2015; Achilles et al., 2017; Rudolph et al., 2019). The pyroxene and olivine compositions of the modern eolian sediments derived from refined unit-cell parameters show that olivine and pigeonite have similar molar Mg/(Mg + Fe), whereas the augite is much more enriched in Mg (Achilles et al., 2017). This discrepancy can arise from either multiple mafic igneous sources, chemical zonation within a single igneous source, or partial



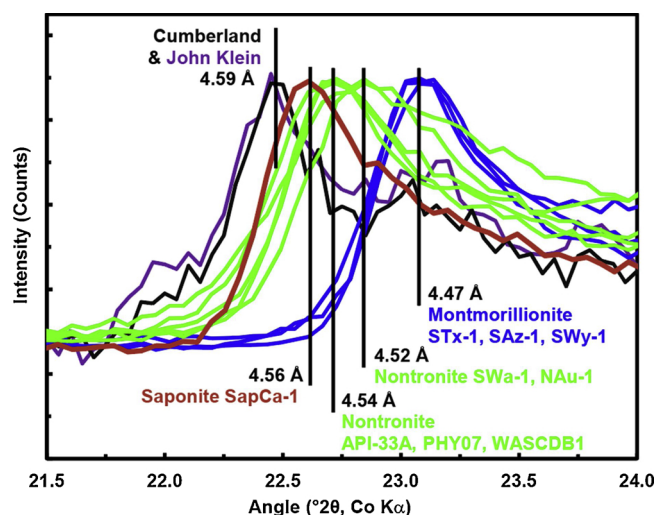


Fig. 15. Position of (02l) peaks in XRD patterns of the John Klein and Cumberland samples measured by CheMin on Mars and of different smectite minerals measured on the laboratory version of CheMin. Vertical lines denote peak maxima and values represent d-spacings. Adapted from Vaniman et al. (2014).

equilibration of a magma with a composition similar to that of the martian meteorite Nakhla (Achilles et al., 2017). However, pyroxene compositions are difficult to constrain from CheMin XRD data and these interpretations should not be considered conclusive. Furthermore, compositional zoning in crystals cannot be inferred from bulk XRD data because results from Rietveld refinements reflect mean values of crystal chemistry, so there are limitations to our interpretations of the igneous petrogenesis of the sediment sources.

Basaltic igneous minerals are found in many of the fluvio-lacustrine mudstone samples, a further indication that mafic igneous sources contributed to ancient sediments in Gale crater, but the mineralogy of two samples, in particular, suggests that considerably evolved igneous sources were present as well. The “Windjana” sample was drilled from reworked fluvial-eolian sediment in the Kimberley formation on the plains of Gale crater. Geochemical analyses of the Kimberley formation show it is enriched in  $K_2O$  (up to 5.6 wt.% from ChemCam; Le Deit et al., 2016), and mineralogical analysis of the “Windjana” sample demonstrates that it contains abundant sanidine (~Or<sub>95</sub>), in addition to pyroxene, olivine, and magnetite (Treiman et al., 2016). The presence of sanidine indicates that “Windjana” was not sourced from a slowly cooled igneous rock because the ordered K-feldspars orthoclase and

microcline form under those conditions. Treiman et al. (2016) speculate that the sanidine formed from potassic volcanic or metasomatic rocks and that the sediments in “Windjana” were derived from multiple igneous sources, because it is unlikely that sanidine and pigeonite are from the same protolith. They suggest that sanidine was sourced from a potassic trachyte or melatrachyte, whereas the pigeonite was from a basalt with a composition similar to the shergottite martian meteorites. This conclusion from CheMin data is consistent with the identification of at least two distinct geochemical sources (potassic and basaltic) of sediments from APXS and ChemCam data of all targets in the Bradbury group (Le Deit et al., 2016; Mangold et al., 2016; Edwards et al., 2017; Siebach et al., 2017; Bedford et al., 2019). Hydrothermal alteration of basalt by potassic fluids, however, can also produce sanidine (Morris et al., 2020), and the CheMin team is investigating the possibility that this process caused the formation of sanidine in Windjana (e.g., Ott et al., 2019).

The most silica-rich targets analyzed to date are in the Marias Pass area at the top of the Pahrump Hills member in the Murray formation, where some ChemCam analyses showed > 90 wt.%  $SiO_2$  (Frydenvang et al., 2017) and up to 6 wt.%  $H_2O$  (Rapin et al., 2018). CheMin analysis of the “Buckskin” sample demonstrated that the lacustrine mudstone in this location contains abundant tridymite, plagioclase, sanidine, magnetite, cristobalite, and opaline silica (Morris et al., 2016). The detection of abundant tridymite was a surprise to the Science Team because Mars is generally a basaltic planet. This assemblage at “Buckskin” is most consistent with silicic volcanism, and tridymite on Earth is commonly found in rhyolites that experienced vapor-phase alteration or acidic hydrothermal alteration at temperatures up to 900 °C (Morris et al., 2016, and references therein). Tridymite formation may not require silicic volcanism; tridymite has been identified in mafic and intermediate volcanic rocks in which amorphous silica from acid-sulfate leaching was altered by high-temperature vapors or contact metamorphism (e.g., Getahun et al., 1996; Del Moro et al., 2011). However, the high concentration of tridymite and the absence of mafic silicates (e.g., pyroxene, olivine) suggest the sediments at Marias Pass were in fact derived from a silicic volcanic source.

The source of the silicic sediments remains enigmatic. Infrared spectroscopic measurements of the Gale crater rim do not reveal any silica-rich source areas. Anhydrous silica minerals would be difficult to detect in visible and near- to short-wave infrared reflectance spectra, but would be detectable in thermal-infrared emission spectra if the outcrop was sufficiently large (e.g., Smith and Bandfield, 2012). The source of the silicic sediments may have been completely denuded; it may be obscured by younger units; its spatial extent may be limited and therefore may not be detectable by remote sensing measurements; or it

Table 7

Abundances of minerals and amorphous component- (in wt.%) for modern eolian sediment samples measured by CheMin. Uncertainties in parentheses are  $1\sigma$ , where the error is applied to the last significant digit(s). The symbol “-” indicates that phase is at or below the detection limit of CheMin.

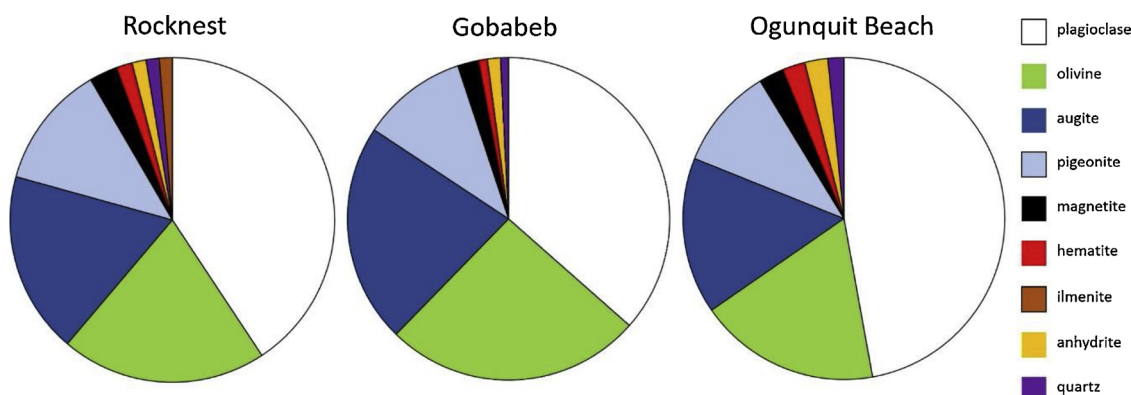
	Rocknest <sup>1</sup>		Gobabeb <sup>1</sup>		Ogunquit Beach <sup>2</sup>	
	Crystalline	Xtal + Amor	Crystalline	Xtal + Amor	Crystalline	Xtal + Amor
Plagioclase	40.7(5)	26.3(7)	36.5(8)	23.7(8)	47.1(10)	24.7(5)
Pyroxene	30.4(25)	19.7(17)	32.6(8)	20.9(12)	25.9(24)	13.6(12)
Olivine	20.5(4)	13.3(4)	25.8(4)	16.8(6)	18.2(7)	9.6(3)
Quartz	1.3(3)*	0.8(2)	0.8(1)*	0.5(1)	1.6(2)*	0.8(1)
Magnetite	2.5(5)	1.8(3)	2.1(2)	1.4(2)	2.5(3)	1.3(2)
Hematite	1.4(3)*	0.9(2)	0.9(1)*	0.6(2)	2.3(3)	1.2(2)
Ilmenite	1.3(5)	0.9(3)	-	-	-	-
Anhydrite	1.6(1)*	1.0(1)	1.3(1)*	0.8(1)	2.3(3)	1.2(2)
Phyllosilicate†	-	-	-	-	-	7.4(30)
Amorphous	-	35(15)	-	35(15)	-	40(15)

<sup>1</sup> From Achilles et al. (2017).

<sup>2</sup> From Rampe et al. (2018).

\* At the detection limit of the CheMin instrument.

† From contamination in CheMin funnel from previous sample (Sebina).



**Fig. 16.** Relative mineral abundances of eolian samples (renormalized without X-ray amorphous component). Abundances from Achilles et al. (2017) and Rampe et al. (2018).

may be in the mostly concealed central uplift peak of Gale crater. Although there is no known pathway to produce tridymite at the relatively low temperatures that these sediments may have experienced as a result of burial ( $< 80$  °C; Borlina et al., 2015), there is geochemical evidence for hydrothermal alteration in the vicinity of Marias Pass (e.g., elevated Ge in complex, late diagenetic veins in the Pahrump Hills member; Berger et al., 2017), and the possible role of hydrothermal fumarolic alteration in the formation of the mineral assemblage in Buckskin is being evaluated (Yen et al., 2020). Furthermore, gray hematite identified in the Oudam sample from Hartmann's Valley and the Highfield sample from Vera Rubin ridge commonly forms at elevated temperatures of  $\sim 80$ – $200$  °C (Catling and Moore, 2003; Rampe et al., 2020).

## 7. Modern eolian processes in Gale crater

The mineralogy of active and inactive eolian deposits is remarkably similar (Achilles et al., 2017; Rampe et al., 2018). All eolian samples measured by CheMin are dominated by plagioclase, olivine, pyroxene, and X-ray amorphous materials. The primary difference between active and inactive eolian sediments is the composition of the X-ray amorphous material. The inactive deposits, represented by the “Rocknest” sample, contain martian dust. Martian dust compositions have been constrained by APXS based on observations of airfall dust on the rover's observation tray (Berger et al., 2016) and by ChemCam, because nearly all first laser shots of rocks and soils measure eolian dust (Lasue et al., 2018). APXS and ChemCam measurements indicate the dust is enriched in  $\text{SO}_3$ , Cl, and Fe relative to bulk soils. The X-ray amorphous component of “Rocknest” is also enriched in  $\text{SO}_3$ , Cl, and Fe relative to the bulk sample composition, suggesting dust contributes to the X-ray amorphous component in the inactive soils (Achilles et al., 2017). SAM evolved gas measurements of “Rocknest” show a greater low-temperature  $\text{H}_2\text{O}$  release than active eolian samples, and ChemCam data demonstrate the fine-grained component of inactive soils is enriched in H, consistent with more adsorbed water on dust particles (e.g., Leshin et al., 2013; Meslin et al., 2013).

Mineral segregation in the active Bagnold Dune Field has been recognized using CRISM orbital data, where barchanoidal dunes at the

**Table 8**

Crystal chemistry of major igneous phases in modern eolian samples calculated from refined unit-cell parameters. Uncertainties in parentheses are  $1\sigma$ .

	Plagioclase	Augite	Pigeonite	Olivine
Rocknest <sup>1</sup>	An <sub>49(4)</sub>	En <sub>47(5)</sub> Fs <sub>17(5)</sub> Wo <sub>36(2)</sub>	En <sub>48(4)</sub> Fs <sub>52(5)</sub>	FO <sub>57(2)</sub>
Gobabeb <sup>1</sup>	An <sub>63(5)</sub>	En <sub>45(4)</sub> Fs <sub>19(5)</sub> Wo <sub>37(2)</sub>	En <sub>48(6)</sub> Fs <sub>50(9)</sub> Wo <sub>3(4)</sub>	FO <sub>54(2)</sub>
Ogunquit Beach <sup>2</sup>	An <sub>48(5)</sub>	En <sub>64(10)</sub> Fs <sub>4(12)</sub> Wo <sub>33(4)</sub>	En <sub>66(11)</sub> Fs <sub>32(15)</sub> Wo <sub>8(6)</sub>	FO <sub>60(3)</sub>

<sup>1</sup> From Achilles et al. (2017).

<sup>2</sup> From Rampe et al. (2018).

**Table 9**

Calculated compositions for the X-ray amorphous component in modern eolian samples.

	Rocknest <sup>1</sup>	Gobabeb <sup>1</sup>	Ogunquit Beach <sup>2</sup>
SiO <sub>2</sub>	34.58	49.53	45.60
TiO <sub>2</sub>	2.12	2.12	2.15
Al <sub>2</sub> O <sub>3</sub>	5.45	7.86	6.89
Cr <sub>2</sub> O <sub>3</sub>	1.39	0.94	1.07
FeO <sub>T</sub>	23.27	19.82	22.33
MnO	1.20	0.89	0.83
MgO	4.05	0.00	2.60
CaO	4.45	3.06	5.07
Na <sub>2</sub> O	3.38	4.46	2.91
K <sub>2</sub> O	1.39	1.18	1.37
P <sub>2</sub> O <sub>5</sub>	2.70	1.91	1.76
SO <sub>3</sub>	14.04	7.02	6.10
Cl	2.96	1.21	1.32

<sup>1</sup> From Achilles et al. (2017).

<sup>2</sup> From Rampe et al. (2018).

upwind margin and crests of linear dunes have stronger olivine signatures and the stoss sides of barchanoidal dunes have stronger high-Ca pyroxene signatures (Seelos et al., 2014; Lapotre et al., 2017). Orbital data also suggest that olivine and plagioclase abundances are anti-correlated (Lapotre et al., 2017). In-situ mineralogical and geochemical measurements using CheMin, APXS, and ChemCam demonstrate that there are compositional differences in dunes and across the dune field that can be attributed to sediment sorting and/or distinct local sediment sources (Cousin et al., 2017b; Johnson et al., 2017a; Rampe et al., 2018; O'Connell-Cooper et al., 2018). CheMin measurements of a sample from a barchanoidal dune on the upwind margin during Phase 1 of the Bagnold Dune campaign (the “Gobabeb” sample from Namib Dune) and a sample from a sand patch near linear dunes downwind from Namib Dune during Phase 2 of the campaign (the “Ogunquit Beach” sample from the Mount Desert Island sand patch) show that the upwind margin contains more olivine and less plagioclase than the sands downwind. CRISM analyses of the sampling locations, however, show that the upwind margin contains more plagioclase and less olivine than the sands downwind (Rampe et al., 2018). This discrepancy is likely the



**Table 10**

Abundances of minerals and amorphous component for ancient eolian samples collected from the Stimson formation<sup>1</sup>. BS = Big Sky, OK = Okoruso, GH = Greenhorn, and LB = Lubango. The “xtal” columns show mineral abundances renormalized to 100 wt.% without the amorphous component. The “bulk” columns show mineral and amorphous abundances. Uncertainties in parentheses are 1 $\sigma$ , where the error is applied to the last significant digit(s). The symbol “–” indicates that phase is at or below the detection limit of CheMin.

	BS xtal <sup>2</sup>	BS bulk	GH xtal	GH bulk	OK xtal	OK bulk	LB xtal	LB bulk
Plagioclase	45.6(22)	36.5(18)	42.1(24)	14.7(8)	41.9(26)	27.2(17)	43.2(24)	11.7(6)
K-feldspar	1.4(8)	1.1(6)	–	–	2.9(10)	1.9(7)	–	–
Pigeonite	21.2(18)	17.0(14)	4.7(14)	1.6(5)	20.8(20)	13.5(13)	5.9(14)	1.6(4)
Orthopyroxene	10.5(22)	8.4(18)	7.6(20)	2.7(7)	11.0(22)	7.2(14)	10.4(21)	2.8(5)
Magnetite	12.9(8)	10.3(6)	17.3(10)	6.1(3)	17.3(10)	11.2(7)	11.1(9)	3.0(2)
Hematite	3.8(6)	3.0(5)	6.0(8)	2.1(3)	1.1(6)	0.7(4)	2.3(6)	0.6(2)
Quartz	1.7(4)	1.4(3)	2.2(6)	0.8(2)	1.4(4)	0.9(3)	3.5(7)	0.9(2)
Anhydrite	1.5(4)	1.2(3)	16.1(10)	5.6(3)	0.8(6)	0.5(4)	12.3(9)	3.3(2)
Bassanite	–	–	4.0(10)	1.4(3)	1.2(6)	0.8(4)	9.0(10)	2.4(3)
Gypsum	–	–	–	–	–	–	2.3(7)	0.6(2)
Fluorapatite	1.4(6)	1.1(5)	–	–	1.6(6)	1.0(4)	–	–
Amorphous	–	20(10)	–	65(20)	–	35(15)	–	73(20)

<sup>1</sup> From Yen et al. (2017a).

<sup>2</sup> Abundances in BS do not include the 1.4 wt.% tridymite contamination from the Buckskin sample.

result of the scooping location and the sediment size analyzed by CheMin. The “Gobabeb” sample from the Namib Dune on the upwind margin was collected from a ripple crest, and APXS and ChemCam data demonstrate ripple crests show higher concentrations of olivine (Cousin et al., 2017b; Johnson et al., 2017a; O’Connell-Cooper et al., 2017). The “Ogunquit Beach” sample collected downwind was from a ripple trough, in which felsic grains are concentrated. CheMin also analyzed the < 150  $\mu$ m grain size, so the mineralogy of the coarser sand size fraction was not captured by CheMin. Furthermore, CheMin did not observe any differences in relative abundances of olivine and high-Ca pyroxene as is suggested from olivine and high-Ca pyroxene parameter maps from CRISM data (Seelos et al., 2014). This observation suggests that segregation between olivine and high-Ca pyroxene occurs in the > 150  $\mu$ m size fraction.

Orbital spectral data and in-situ compositional data from *Curiosity* can be used in a complementary fashion to construct a more complete picture of modern eolian processes and sediment segregation in Gale crater. Bulk mineralogy from CRISM and the concentration of olivine with respect to plagioclase downwind indicates that fine-grained olivine is preferentially transported and concentrated downwind (Rampe et al., 2018). The stronger olivine spectral signatures compared to high-Ca pyroxene in the upwind margin suggests that there is a coarse-grained olivine lag in the barchanoid dunes. The variations in olivine and plagioclase abundances from samples measured by CheMin as well as spectral and chemical variations observed in place by ChemCam and APXS suggest that there are significant mineralogical variations between ripple troughs and ripple crests. The minor and major mineral crystal chemistry derived from CheMin analyses also suggests that the local bedrock has been eroded and contributes to eolian sediments. Contributions of local bedrock were also positively identified in the “Rocknest” soil by ChemCam (Meslin et al., 2013). The presence of minor to trace amounts of hematite, Ca-sulfate, and quartz in the eolian samples collected from the Bagnold Dune Field are likely from the underlying Murray formation, which contains abundant hematite and

**Table 12**

Calculated compositions for the X-ray amorphous component in ancient eolian samples<sup>1</sup>.

	Big Sky	Greenhorn	Okoruso	Lubango
SiO <sub>2</sub>	24.13	63.53	43.91	68.00
TiO <sub>2</sub>	6.59	1.54	3.65	1.57
Al <sub>2</sub> O <sub>3</sub>	0.12	0.00	3.62	0.00
Cr <sub>2</sub> O <sub>3</sub>	3.36	0.69	1.59	0.41
FeO <sub>T</sub>	17.48	10.16	16.67	4.79
MnO	2.64	0.21	1.52	0.13
MgO	14.68	1.46	14.54	0.87
CaO	0.00	5.47	7.94	6.52
Na <sub>2</sub> O	6.03	2.08	3.24	1.19
K <sub>2</sub> O	1.40	0.46	0.00	0.44
P <sub>2</sub> O <sub>5</sub>	1.69	1.77	0.97	1.87
SO <sub>3</sub>	17.08	11.93	0.41	13.77
Cl	4.81	0.70	1.94	0.45

<sup>1</sup> From Morrison et al. (2018b).

Ca-sulfate, especially, above the Pahrump Hills member (e.g., Bristow et al., 2018). The difference in plagioclase crystal chemistry between the “Gobabeb” and “Ogunquit Beach” samples (An<sub>63(5)</sub> and An<sub>48(5)</sub>, respectively) could result from different plagioclase compositions of the underlying rocks or from different sediment sources for the Bagnold Dune Field.

## 8. Ancient aqueous conditions in Gale crater

Evidence for a long-lived lake or lake system in Gale crater is compelling. The physical characteristics and chemistry of the lake (e.g., its depth, salinity, pH, Eh) can be derived from the sedimentary features and the mineralogy and geochemistry of the lithified sediments. The aqueous history of Gale crater sediments investigated to date can be divided into three temporally distinct processes: Deposition (processes which occur at the time the sediment was deposited), authigenesis/

**Table 11**

Crystal chemistry of major igneous phases in ancient eolian samples calculated from refined unit-cell parameters<sup>1</sup>. Uncertainties in parentheses are 1 $\sigma$ .

	Plagioclase	Orthopyroxene	Pigeonite
Big Sky	An <sub>52(5)</sub>	Mg <sub>0.69(7)</sub> Fe <sub>1.31(7)</sub> Si <sub>2</sub> O <sub>6</sub>	Mg <sub>1.44(7)</sub> Fe <sub>0.39(9)</sub> Ca <sub>0.17(4)</sub> Si <sub>2</sub> O <sub>6</sub>
Greenhorn	An <sub>40(6)</sub>	Mg <sub>0.80(8)</sub> Fe <sub>1.16(9)</sub> Ca <sub>0.04(4)</sub> Si <sub>2</sub> O <sub>6</sub>	*
Okoruso	An <sub>39(5)</sub>	Mg <sub>0.86(20)</sub> Fe <sub>1.14(20)</sub> Si <sub>2</sub> O <sub>6</sub>	Mg <sub>1.39(7)</sub> Fe <sub>0.48</sub> Ca <sub>0.13(5)</sub> Si <sub>2</sub> O <sub>6</sub>
Lubango	An <sub>30(8)</sub>	Mg <sub>0.81(10)</sub> Fe <sub>1.19(11)</sub> Si <sub>2</sub> O <sub>6</sub>	Mg <sub>1.54(17)</sub> Fe <sub>0.18(17)</sub> Ca <sub>0.28(6)</sub> Si <sub>2</sub> O <sub>6</sub>

<sup>1</sup> From Morrison et al. (2018b).

\* Unit-cell parameters were not refined because of insufficient abundance.

early diagenesis (processes which occur concurrently with deposition or shortly thereafter, prior to lithification), and late stage diagenesis (processes which occur after lithification, up to the present day). These processes are exquisitely preserved in the Sheepbed mudstone member of the Yellowknife Bay formation where the first two drill samples (“John Klein” and “Cumberland”) were collected.

The Sheepbed mudstone is interpreted as a shallow lacustrine deposit (Grotzinger et al., 2014). The mudstone exhibits abundant evidence for early diagenesis, including erosionally resistant (“raised”) ridges enriched in magnesium (Léveillé et al., 2014), which are interpreted as lithified subaqueous shrinkage cracks (Siebach et al., 2014), and a variety of spheroidal Ca-sulfate-filled voids and cemented concretions, suggesting that groundwater influenced the sediment before it was lithified (Stack et al., 2014). Light-toned Ca-sulfate bearing veins observed in Yellowknife Bay (e.g., Nachon et al., 2014) and throughout nearly the entirety of *Curiosity*'s traverse (e.g., Grotzinger et al., 2015; Nachon et al., 2017; Kronyak et al., 2019) are evidence for late-stage diagenesis resulting from fluids moving through fractures in the rocks. The observation of hydraulic fracturing in the Sheepbed unit suggests that these sediments were buried to a depth of at least 1000 m (Schieber et al., 2017).

The mineralogy of the “John Klein” and “Cumberland” samples (Vaniman et al., 2014) can be compared with the mineralogy of its presumed original source rock at the crater rim (olivine basalt) to evaluate depositional, authigenic, and diagenetic processes in the Sheepbed mudstone. For example, using the mineralogy of the “Rocknest” sand shadow as a proxy for the source rock, the basaltic minerals plagioclase and pyroxene are present in similar quantities in the Sheepbed mudstone, but olivine is absent. Instead, the mudstone contains a significant quantity of the clay mineral Fe-saponite and excesses of magnetite. Except for late-stage Ca-sulfate bearing veins, the mudstone has retained the same elemental ratios as its presumed igneous source minerals, i.e., these inferred changes in mineralogy would have occurred under isochemical conditions (McLennan et al., 2014). A process called “saponitization” would have led to the observed mineral assemblage, in which olivine and X-ray amorphous material reacted with water, close to the time of deposition, to yield an Fe-rich clay mineral (e.g., Fe-rich saponite; Treiman et al., 2014) plus magnetite and possibly hydrogen (Vaniman et al., 2014; Grotzinger et al., 2014; Bridges et al., 2015; Bristow et al., 2015; Tosca et al., 2018). Geochemical modeling indicates that pore waters hosting such mineral reactions would be driven to neutral to alkaline pH (Bristow et al., 2015, 2017; Tosca et al., 2018), a conclusion supported by the hypothesized intercalation of Mg-hydroxy interlayers in smectite clay of the Cumberland sample (Vaniman et al., 2014). Such an authigenic mineral assemblage also requires that pore waters were poorly oxidizing to anoxic (Bristow et al., 2015, 2017; Hurowitz et al., 2017). Another key observation is the relative absence of enrichments in chlorides, sulfates, and carbonates in the lacustrine mudstone matrix in Yellowknife Bay, as shown by mineralogy and bulk geochemical measurements (McLennan et al., 2014; Grotzinger et al., 2014; Vaniman et al., 2014), demonstrating that pore waters were generally free of dissolved salts.

Taken together, the sedimentological, mineralogical, and geochemical data suggest a relatively fresh water lake with circumneutral pH existed during the deposition of Yellowknife Bay sediments ~ 3.7 Ga ago. The kinetics of olivine alteration suggest the presence of a benign aqueous environment for thousands to hundreds of thousands of years (Bristow et al., 2015). Moreover, the mediated oxidation of Fe(II) in olivine to Fe(III) in magnetite, and perhaps smectite, would have provided a potential chemolithoautotrophic energy source for organisms (Vaniman et al., 2014; Grotzinger et al., 2014; Bristow et al., 2015), and SAM and APXS detected the necessary elements for life (i.e., C, H, N, O, P, and S; McLennan et al., 2014; Ming et al., 2014). Thus, the deposits at Yellowknife Bay represent an important milestone achieved in MSL's prime mission, where an ancient habitable martian environment was

first definitively established. Subsequent investigations by *Curiosity* have further broadened our knowledge of the history of habitable environments at Gale crater and are discussed in more detail in the following section of this paper.

As is common in long-lived lake systems on Earth, aqueous conditions and the sediment sources of the lakes occupying Gale crater changed through time. This evolution is reflected by changes in the primary and secondary phases documented in the fluvio-lacustrine samples along *Curiosity*'s route (e.g., Fig. 14). Note that most units *Curiosity* has drilled that are discussed here are stratigraphically below the hematite-, clay-, and sulfate-bearing units identified from orbit and do not present a clear orbital spectral signature for these secondary phases, demonstrating the importance of in-situ mineralogical measurements for characterizing ancient aqueous environments.

CheMin data reveal common occurrence of phyllosilicates in most fluvio-lacustrine samples (Vaniman et al., 2014; Rampe et al., 2017; Bristow et al., 2018). The phyllosilicates are generally consistent with collapsed smectite, but the character of the smectite changes going up section (i.e., going forward in time). At the base of the section in the “John Klein” and “Cumberland” samples, a trioctahedral Fe-smectite occurs (containing Mg and Fe(II) in its octahedral sites) (Vaniman et al., 2014; Treiman et al., 2014; Bristow et al., 2015). Higher in the section, mixed trioctahedral-dioctahedral smectites occur [Mg and Fe(II) are replaced by Al and Fe(III)], resembling montmorillonite and/or nontronite; Bristow et al., 2018). On Earth, dioctahedral smectite often forms from near-surface oxidative weathering, so the change in smectite structure observed over 100 s of m of stratigraphy suggests a gradual change in the weathering environment. The identification of Fe-pyrophyllite in some samples may point toward localized hydrothermal processes (Rampe et al., 2020).

Other secondary minerals also suggest an increase in the degree of oxidation and chemical weathering intensity experienced by lake sediments over time (Bristow et al., 2018; Mangold et al., 2019). For example, hematite is the prevalent Fe-oxide mineral in samples from the Hartmann's Valley, Karasburg, Sutton Island, Blunts Point, and Pettegrove Point members (samples “Oudam,” “Marimba,” “Quela,” “Sebina,” “Duluth,” and “Stoer”), whereas magnetite is more common in samples drilled near the base of the section in the Sheepbed and Pahrump Hills members (samples “John Klein,” “Cumberland,” “Telegraph Peak,” and “Buckskin”), suggesting increasingly oxidative conditions over time (Bristow et al., 2018). This finding is consistent with the results from the ChemCam and Mastcam reflectance spectra acquired along the traverse (Johnson et al., 2016, 2017b,c; Wellington et al., 2017a, b), which exhibit a shift to variable but pervasively ferric spectral features through the upper portion of the Murray formation. The Ca-sulfate minerals gypsum, bassanite, and anhydrite are also more abundant higher in the Murray formation sedimentary matrix (Bristow et al., 2018; Vaniman et al., 2018), and most diagenetic Ca-sulfate veins show a water content consistent with bassanite from ChemCam LIBS analyses, interpreted as a byproduct of gypsum dehydration (Rapin et al., 2016). A ~ 150 m interval stratigraphically below the “Duluth” target in the Blunts Point member showed sporadic concentrations of Mg-sulfate and Ca-sulfate (26–36 wt.% and 30–50 wt.% of the bulk rock, respectively, in some locations) from ChemCam data (Rapin et al., 2019). These sulfate-rich strata may signify episodic evaporation and formation of brines. Unfortunately, *Curiosity* did not drill in this interval because of the anomaly with the drill feed so we did not analyze these rocks with CheMin and do not know the bulk mineral assemblage. These mineralogical changes in the bedrock suggest that the lake waters in Gale crater became more saline and oxidizing over time, or that later episodes of oxidizing, saline groundwater preferentially altered sediments higher in the stratigraphic section.

One hypothesis put forth to explain the changes in Fe-oxide mineralogy in the stratigraphy proposes that the lake present in ancient Gale crater was redox-stratified (Hurowitz et al., 2017). In this model (Fig. 17), shallow lake waters would have been oxidized by ultraviolet



(UV) photolysis and by the exposure of surface waters to photochemically generated atmospheric  $O_2$ . Deeper lake waters would have remained relatively anoxic as a result of the attenuation of UV light in the water column and decoupling from atmospheric  $O_2$ . Overland flow from rivers and streams would have delivered clastic sediments to the lakes, including basaltic and silicic igneous minerals and, perhaps, phyllosilicates. Groundwater would have transported dissolved solutes, including  $Fe(II)_{(aq)}$ , to the lakes. Groundwater flow into shallow, oxidizing lake waters would have resulted in the precipitation of hematite, and groundwater flow into deep, anoxic lake waters would have resulted in the precipitation of magnetite. Relatively high concentrations of the redox sensitive element Mn are found at the base of the Pahrump Hills member, in particular, and coincide with abundant hematite, suggesting some of the shallow lake waters were extremely oxidizing.

A second hypothesis to explain the changes in Fe-oxide mineralogy in the stratigraphy is that hematite formed via diagenetic oxidation of magnetite (Rampe et al., 2017). In this scenario, magnetite would have formed under alkaline, relatively anoxic conditions, either in the lake sediments prior to lithification or in the sediment source region, and acidic and/or oxidizing diagenetic fluids would have caused the alteration of magnetite to form hematite. Acid-sulfate alteration occurred in some of the samples drilled from the Murray formation, evidenced by the presence of minor to trace amounts of jarosite, which forms under a pH range of ~2 to 4 (e.g., Driscoll and Leinz, 2005). The trace element geochemistry of the Pahrump Hills member of the Murray formation also suggests that low pH diagenetic fluids moved through these sediments. Low abundances of Ni and Zn were measured at the top of the Pahrump Hills member in Marias Pass. Concentrations of these elements increased down-section and were highly enriched relative to average Mars crust at the base of the Pahrump Hills section (Fig. 18). Acid-sulfate fluids moving down-section from Marias Pass may have mobilized Ni and Zn, and these elements may have precipitated as discrete sulfates or as adsorbed species onto phyllosilicates and/or nanophase Fe-oxides as the fluids were neutralized down-section (Fig. 19; Rampe et al., 2017). Leaching by acid-sulfate fluids may also explain the opaline silica detected by CheMin and the coordinated strong enrichments in H and Si observed by ChemCam (Rapin et al., 2018) at the top of the Pahrump Hills member (in the “Telegraph Peak” and “Buckskin” samples) and the Hartmann’s Valley member (in the “Oudam” sample). In the presence of acidic fluids, Si is relatively resistant to mobilization compared to other common elements (e.g., Mg, Ca, Na, K, Al) so that it becomes passively enriched as other elements are transported from the system (e.g., Ming et al., 2008; Squyres et al., 2008). The source of the acid-sulfate fluids is unknown, but could be from oxidation of Fe-sulfides, dissolution of Fe-sulfates in groundwater, or interaction between magmatic volatiles and groundwater (Rampe et al., 2017). Fe-sulfides have not been detected with confidence above detection limits by CheMin, but SAM evolved gas data from the “Rocknest” eolian deposit indicate the presence of trace sulfides (McAdam et al., 2014).

The abundance of diagenetic features in the Yellowknife Bay member of the Bradbury formation and the Murray and Stimson formations extends morphologically to fluidized pipes (Rubin et al., 2017) and decimeter-size concretions (Wiens et al., 2017) which, along with diagenetic chemistry and mineralogy, indicates a rich history of groundwater in Gale crater. Differences in the geochemistry of some of

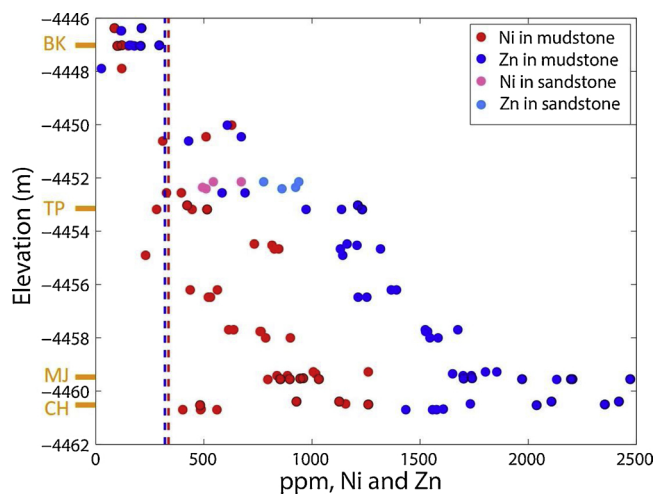


Fig. 18. Ni (in red) and Zn (in blue) abundances vs. elevation from APXS measurements of brushed (points with dark outlines) and unbrushed targets in the Pahrump Hills member. Elevations of the four samples drilled from the Pahrump Hills member are denoted in orange. Pink and light blue points represent Ni and Zn abundances, respectively, collected from a sandstone interval. All other points were measured from mudstone targets. Red and blue vertical lines represent average Mars crustal compositions of Ni and Zn, respectively (Taylor and McLennan, 2009). Yellow markers on the y-axis represent drill locations, where CH = Confidence Hills, MJ = Mojave2, TP = Telegraph Peak, and BK = Buckskin. Adapted from Rampe et al. (2017) (For interpretation of the references to colour in this figure legend, the reader is referred to the web version of this article.).

these features suggest groundwater had variable compositions. For example, the raised ridges at the base of the stratigraphic section in Yellowknife Bay and dendritic aggregates at the base of the Pahrump Hills member (Fig. 20A) are enriched in Mg, suggesting that early-stage Mg-rich fluids, perhaps from the alteration of the abundant mafic minerals in that portion of the section, moved through the sediments before they lithified (Léveillé et al., 2014; Nachon et al., 2017). The apparent precipitation of Mn oxides from groundwater in the Bradbury group and the Murray formation suggests that fluids were strongly oxidizing in some locations (Lanza et al., 2014, 2016; Gasda et al., 2019).

Pervasive Ca-sulfate veins suggest that the late-stage post-lithification diagenetic fluids that moved through the fractured bedrock were enriched in CaO and  $SO_3$ , and also fluorine in some locations (e.g., Nachon et al., 2014, 2017; Forni et al., 2015, 2016; Rapin et al., 2016). A detailed investigation of the morphology and composition of a complex suite of veins in the Pahrump Hills demonstrates three separate post-depositional fluid episodes each with a distinct fluid composition (Kronyak et al., 2019). Dissolution of sulfate-rich horizons may have led to the nearly pure Ca-sulfate veins seen throughout the Gale stratigraphy (Schwenzer et al., 2016). Millimeter-scale lenticular crystal molds whose morphology is reminiscent of gypsum (Fig. 20B) suggest that, in some locations, Ca-sulfate precipitated from early diagenetic fluids at the sediment-water or sediment-air interface (Kah et al., 2015a, b, 2018), but the absence of Ca-sulfate minerals from CheMin analysis of the “Mojave2” sample drilled from a crystal-mold-

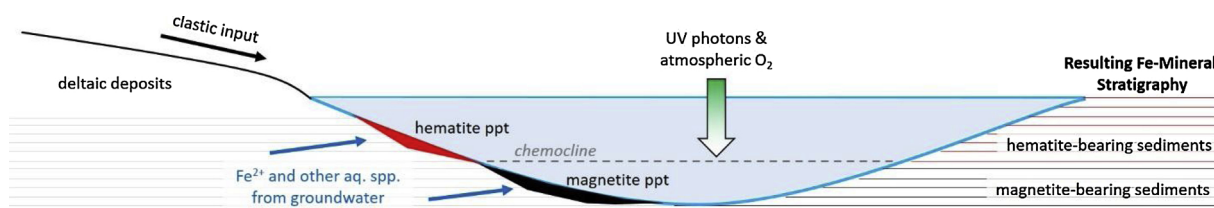


Fig. 17. Redox-stratified lake model, adapted from Hurowitz et al. (2017).

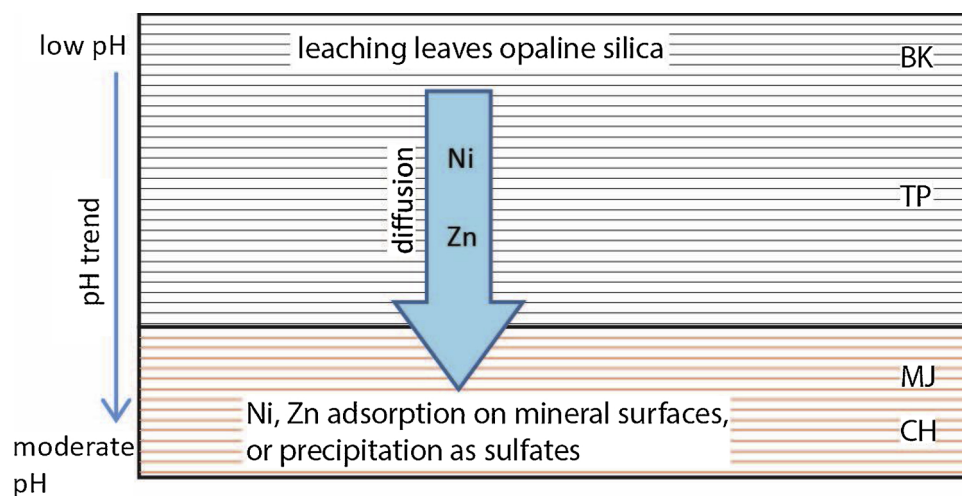


Fig. 19. Diagenetic alteration model to explain the mineralogy and geochemistry of the Pahrump Hills member, adapted from Rampe et al. (2017). CH = Confidence Hills, MJ = Mojave2, TP = Telegraph Peak, BK = Buckskin.

rich mudstone indicates the Ca-sulfate was dissolved by later diagenetic fluids (Rampe et al., 2017; Kah et al., 2018).

Early diagenesis in the Stimson formation and the formation of magnetite cement likely at the expense of olivine is consistent with leaching from near-neutral fluids (Yen et al., 2017; Hausrath et al., 2018). The mineralogy and geochemistry of the fracture-associated halos suggest multiple fluid episodes at depth were associated with late-stage diagenesis of the Stimson and Murray formations (Yen et al., 2017; Hausrath et al., 2018). Acidic, sulfate-rich diagenetic fluids may have caused the initial leaching of Mg, Al, Mn, Fe, Ni, and Zn, the preferential dissolution of pyroxene, and the abundance of Ca-sulfate and amorphous silica (Yen et al., 2017; Hausrath et al., 2018), while later alkaline fluids could have caused the observed addition of Si and precipitation of trace carbonate detected by SAM (Yen et al., 2017). The mobilization of Ni, Zn, and Mn seen in the Pahrump Hills section may be related to leaching observed in the Stimson formation.

In cases where visual evidence of a secondary mineral phase is present (e.g., the concretions in the Yellowknife Bay formation; Stack et al., 2014, and in the Murray formation; Sun et al., 2019), it is possible to determine whether they are authigenic, early or late-stage diagenetic. However, if the mineral occurs in the matrix of the sedimentary rock, it can be difficult to determine which minerals formed in the lake waters and which formed from groundwater, but K-Ar age-dating of jarosite in the Mojave2 sample near the base of the Pahrump Hills member has put one constraint on diagenetic processes in Gale crater.

SAM data indicate that the jarosite in Mojave2 formed  $2.12 \pm 0.36$  Ga ago (Martin et al., 2017). This astounding result suggests groundwater was present in Gale crater at least 1.5 billion years after the rivers and lakes in Gale crater deposited the sediments of the Bradbury group and Murray formation. Although the presence of liquid water on the surface and near-subsurface may have been intermittent, a 1.5-billion-year history of liquid water at Gale crater has important implications for habitability and the potential for the existence of life on Mars.

The absence of certain secondary minerals can also help constrain ancient aqueous environments and diagenesis in Gale crater. Minerals like zeolites commonly precipitate from alkaline fluids, and specific zeolite minerals can be used to identify formation environments (e.g., closed alkaline, saline lakes, burial diagenetic deposits, and hydrothermal deposits; e.g., Ming and Mumpton, 1989). The absence of zeolites in Gale crater indicates the lake waters were not sufficiently alkaline and did not have high concentrations of Si and Al in solution. In basin environments on Earth, smectite in buried sediments commonly alters to illite, chlorite, or a mixed-layer phase (e.g., corrensite) through burial diagenesis (e.g., Denoyer de Segonzac, 1970). The abundance of smectite, but the absence of chlorite in the mudstones drilled by *Curiosity* demonstrates sediments in Gale crater did not experience pressures and temperatures associated with burial > 2 km (Borlina et al., 2015). The presence of amorphous silica, but the absence of abundant quartz further demonstrates that groundwater was present intermittently (Tosca and Knoll, 2009).

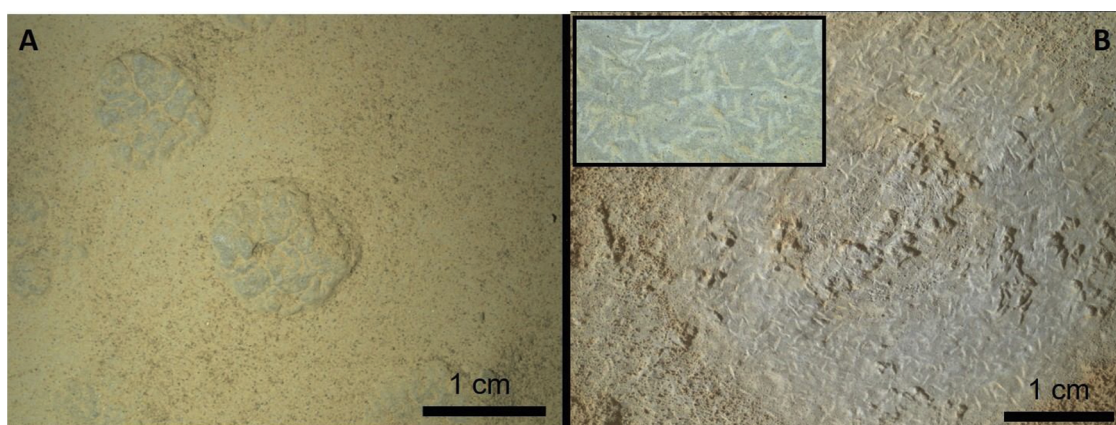


Fig. 20. MAHLI images of diagenetic features in the Pahrump Hills member. A) Mg-Ni-bearing irregular concretions from the base of the Pahrump Hills (sol 810, mhli0447). B) Lenticular crystal pseudomorphs in the target "Mojave2" (sol 880, mhli0462). Inset shows close-up of crystal pseudomorphs (sol 809, mhli00942). Image credits: NASA/JPL-Caltech/MSSS.

Although the geochemistry suggests hydrothermal alteration in certain locations along *Curiosity's* traverse (e.g., based on Ge enrichments; Berger et al., 2017; Yen et al., 2020), and there is some mineralogical evidence for localized hydrothermalism (Rampe et al., 2020), CheMin has not detected abundant or widespread hydrothermal phases, like serpentine, chlorite, and prehnite, which have been detected from orbit in other locations on Mars (e.g., Ehlmann and Edwards, 2014). The absence of carbonate minerals above the detection limit of CheMin (~1 wt.%) in these samples suggests that either the lake waters were not in equilibrium with the CO<sub>2</sub>-rich early martian atmosphere or that the pCO<sub>2</sub> of the early martian atmosphere was lower than predicted by climate models that use CO<sub>2</sub> to warm the martian surface above the freezing point of water (Bristow et al., 2017). The relative paucity of phosphate minerals detected by CheMin, on the other hand, may be related to their detection limit in CheMin rather than their absence in Gale crater sediments. For example, all samples contain ~0.5-1.3 wt.% P<sub>2</sub>O<sub>5</sub> from APXS measurements, so if all P<sub>2</sub>O<sub>5</sub> was present in a single phase, it should be detected by CheMin. This suggests that there are multiple phosphate minerals in Gale crater such that they are below the instrument's detection limit, phosphates are X-ray amorphous, and/or phosphate is chemisorbed to mineral or mineraloid surfaces (e.g., Rampe et al., 2016). Finally, the absence of abundant and pervasive evaporite minerals, contrary to what was observed in the Burns formation in Meridiani Planum by the Mars Exploration Rover *Opportunity* (e.g., Grotzinger et al., 2005; McLennan et al., 2005), indicates that lake waters in Gale crater were not especially saline. Although Ca-sulfate minerals are abundant in some samples from the Murray formation, some of these are clearly late diagenetic fracture fill. We do not know how much is present in the matrix and could have precipitated from saline lake waters. Petrography of returned samples would be necessary for this evaluation. These constraints on aqueous environments from quantitative mineralogy allow us to evaluate the habitability of early Gale crater.

## 9. Evidence for habitable environments at Gale crater

*Curiosity* has found evidence for a long history of liquid water at Gale crater. Liquid water was present for perhaps millions of years on the surface (based on the extent and thickness of fluvio-lacustrine deposits) and spanned over a billion years in the subsurface through radiometric dating of secondary jarosite (Martin et al., 2017; though groundwater was likely present intermittently). SAM-derived D/H ratios of the structurally bound OH in phyllosilicates in Yellowknife Bay are half those of the modern martian atmosphere, but much higher than those expected for very early Mars, suggesting an extended history of water loss from the martian surface (Mahaffy et al., 2015). Furthermore, <sup>36</sup>Ar/<sup>38</sup>Ar ratios measured by SAM from the modern martian atmosphere suggest Mars has experienced significant atmospheric loss over time (Atreya et al., 2013), consistent with data measured of the upper martian atmosphere from orbit with the Mars Atmosphere and Volatile EvolutioN (MAVEN) mission (e.g., Jakosky et al., 2017). Models of CO<sub>2</sub> gas loss by the same sputtering process that caused the loss of Ar suggest that a half bar or more of CO<sub>2</sub> could have been lost over Mars' history (Jakosky et al., 2017). This heavier CO<sub>2</sub>-rich early martian atmosphere may have allowed for the persistent lakes at Gale crater.

The mineralogy and geochemistry of the ancient lake sediments indicate that some of these lacustrine environments had low salinity. The Sheepbed mudstone lacked chlorides and sulfate in the matrix (at least at levels above ~1 wt.%) and had low S and Cl (Grotzinger et al., 2014; McLennan et al., 2014; Vaniman et al., 2014), suggesting lake waters were relatively fresh. Calculations of pore water salinity using bulk geochemistry and the hydration state of smectite suggest the last groundwater to interact with the sediment was hyposaline with ~0.1-0.5 mol/kg NaCl (Fukushi et al., 2019). Other portions of the section had abundant sulfate minerals, suggesting some fluids in Gale crater

were enriched in soluble sulfate salts (e.g., the lake and/or groundwater that deposited the matrix Ca-sulfate detected in the "Oudam," "Marimba," "Quela," and "Sebina" samples from the Karasburg and Sutton Island members; Bristow et al., 2018; Achilles et al., submitted). At least some of the diagenetic fluids were also acidic (e.g., the groundwater from which jarosite precipitated in "Mojave2"; Rampe et al., 2017). The observation of sulfate-bearing and acidic groundwater does not preclude these environments from being habitable. On Earth, microorganisms thrive in evaporite settings, where Ca-sulfate minerals precipitate (e.g., Demergasso et al., 2004), and diverse photosynthetic eukaryotes, chemolithotrophic bacteria, and a variety of heterotrophs exist in acid-sulfate mine drainage in Rio Tinto, Spain (e.g., Fernández-Remolal et al., 2005). In most instances, however, these microorganisms stemmed from ancestors that evolved under clement conditions, so the acidic and sulfate-rich diagenetic environments in Gale crater may not have been ideal locations for the genesis of microbial life (e.g., Squyres and Knoll, 2005).

*Curiosity* has also discovered all the chemical ingredients necessary to support microbial life. The essential elements C, H, N, O, P, and S have been found in deposits throughout the traverse. SAM detected organic molecules in samples from the Sheepbed mudstone and from the Pahrump Hills member. Chlorobenzene and dichloroalkanes were detected up to 300 parts per billion in the Cumberland sample at relatively low temperatures between 200 and 450 °C (Freissinet et al., 2015). Thiophenes, two isomers of methylthiophene methanethiol, and dimethylsulfide are present in the Confidence Hills and Mojave2 samples as evidenced from high-temperature evolution (> 600 °C) (Eigenbrode et al., 2018). The diversity of organic molecules observed at high temperature could result from the protection by sulfate minerals or from the incorporation of sulfur into the molecules, which likely protected the organic molecules from degradation by ionizing and oxidizing conditions at the martian surface (Eigenbrode et al., 2018). Total organic carbon (TOC) contents of these samples is ~0.01-0.1 mmol C (Eigenbrode et al., 2018). For comparison to Earth, saline lake sediments in the Antarctic Dry Valleys have TOC contents of ~15–300 mmol C (Matsumoto et al., 1984, 1989). The origin of the organic molecules found at Gale crater is not known, but could include martian sources (e.g., igneous, hydrothermal, atmospheric, or biological) or exogenous sources such as meteorites, comets, or interplanetary dust particles (Freissinet et al., 2015; Eigenbrode et al., 2018). The variety of temperatures at which organic molecules are released may indicate different origins of the compounds (e.g., adsorbed vs. trapped within minerals). SAM evolved gas analyses show the presence of carbonate minerals and organic molecules in trace abundances based on the observation of CO<sub>2(g)</sub> releases (e.g., Ming et al., 2014; Sutter et al., 2017). SAM data also led to the first detection of nitrate on Mars (Stern et al., 2015), and it is present in all sediments (Sutter et al., 2017). Phosphorus has been detected by APXS in nearly every target, but the speciation is not well constrained. Fluorapatite has been identified by CheMin in samples from the Pahrump Hills (Rampe et al., 2017) and in eolian sandstone (Yen et al., 2017) and inferred from simultaneous enrichments of F and P in ChemCam targets in the Bradbury group and Murray formation (Forni et al., 2015, 2017). Phosphorus enrichments were recognized in fracture-associated halos (Yen et al., 2017), but CheMin did not detect crystalline phosphate in samples collected from the halos, suggesting the phosphorus is present as an X-ray amorphous phase or adsorbed onto mineral or mineraloid surfaces (e.g., Rampe et al., 2016).

A variety of potential energy sources would have been available for martian microorganisms in ancient Gale crater. Minerals in various states of redox could have provided energy for chemolithoautotrophs (Grotzinger et al., 2014), such as the terrestrial Fe-oxidizing bacteria *Pseudomonas Sp HerB*. Magnetite has both Fe(II) and Fe(III), so Fe redox coupling is a potential source of energy in samples from the Yellowknife Bay formation (Bristow et al., 2015) and the Pahrump Hills member. Although sulfide minerals have not been identified above the detection



limit of CheMin, jarosite identified in the samples from the Pahrump Hills may have formed from the alteration of Fe-sulfide (Hurowitz et al., 2017; Rampe et al., 2017), indicating S redox coupling may be another source of energy in Gale crater lakes and groundwater. On Earth, photolithotrophic sulfur-oxidizing purple bacteria use the S redox potential as an energy source (e.g., Trüper, 1982). The nearly global distribution of oxychlorine at Mars' surface also raises the interest of perchlorate-reducing bacteria such as *Dechlorimonas agitata*, *Haloferax denitrificans*, or *Paracoccus halodentrificans*, which use the chlorine redox cycle as a source of energy (e.g., Okeke et al., 2002; Nozawa-Inoue et al., 2008; Lynch et al., 2019). With respect to recent habitability, the detection of a seasonal methane cycle in Gale crater with the SAM instrument is especially intriguing, because extant subsurface methane-producing bacteria is one of many explanations for this signal, including UV degradation of exogenous organic material (Webster et al., 2018).

## 10. Vera Rubin ridge and beyond

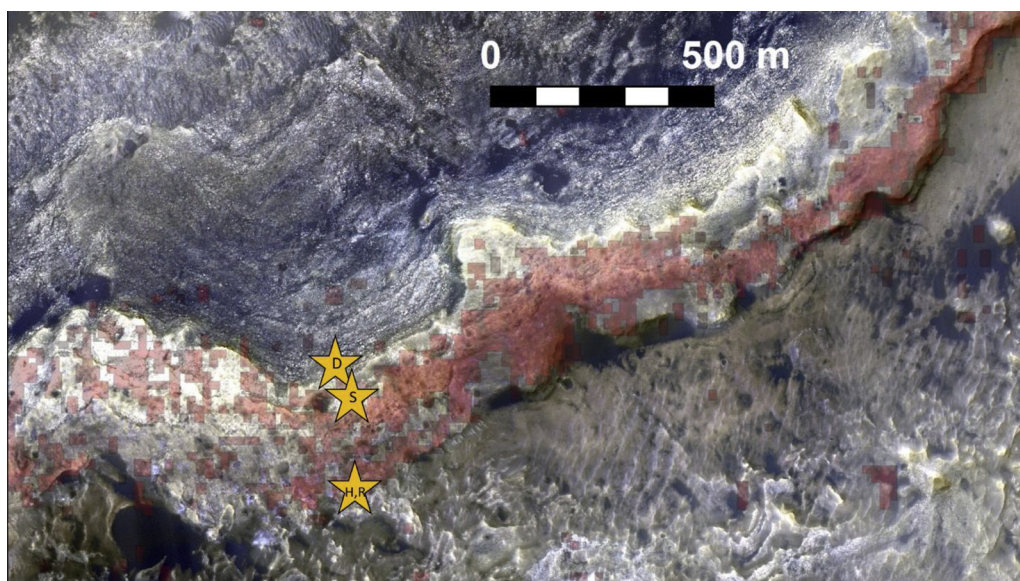
At the time of writing, after more than 2300 sols of rover operations, *Curiosity* had just concluded the investigation of Vera Rubin ridge (formerly called the hematite ridge), having completed drill campaigns at localities called “Duluth,” “Stoer,” “Highfield,” and “Rock Hall” (Fig. 21). These drilling campaigns were significant because *Curiosity* had not successfully delivered a drilled sample to its laboratory instruments CheMin and SAM since November 2016. These were also the first four samples produced with the new feed-extended drill technique designed by the rover engineers at JPL after a problem with the drill feed mechanism precluded the nominal drilling process.

While the engineers worked on the drill, the Science Team explored Vera Rubin ridge using the science cameras and geochemical instruments ChemCam and APXS. Vera Rubin ridge is divided into two members, the Pettegrove Point member constitutes the lower part of the ridge and the Jura member constitutes the upper part. The sedimentology of the entire Vera Rubin ridge is similar to rocks stratigraphically below it, indicating the depositional environments did not change, and mudstone with fine horizontal lamination is ubiquitous, suggesting deposition in a lake environment (e.g., Edgar et al., 2018b; Fraeman et al., 2018). Orbital spectroscopy shows that portions of Vera Rubin ridge, particularly in the Pettegrove Point member, have a strong hematite spectral signature based on bands at 550 and 860 nm (Fraeman et al., 2013, 2016; Fig. 21), consistent with reflectance spectra from ChemCam and Mastcam (Johnson et al., 2017b; c;

Wellington et al., 2017b). CheMin analysis of the “Stoer” sample, drilled from the Pettegrove Point member, identified the most hematite of any sample drilled to date (Morris et al., 2019; Rampe et al., 2020). The overall Fe content of the bedrock, however, is not higher on Vera Rubin ridge than elsewhere in the Murray formation, suggesting that the ridge may be a product of diagenesis (e.g., Frydenvang et al., 2019; Thompson et al., 2019).

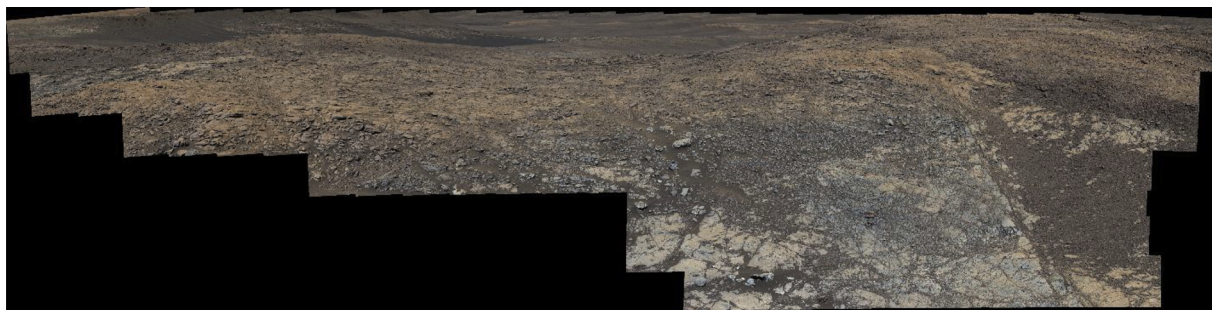
Observations from the rover show visual and geochemical heterogeneities between and within the two members that make up the ridge that suggest diagenetic fluids altered these rocks (e.g., Fraeman et al., 2018; Thompson et al., 2018). Most notably, rocks of the Jura member show a diversity of colors, with both red and gray outcrops (Fig. 22). CheMin analysis of a sample from the gray Jura (“Highfield”) indicates the gray color is from abundant gray hematite (Rampe et al., 2020). Gray hematite has a coarser grain size than red hematite, where grains  $> \sim 3\text{--}5\ \mu\text{m}$  appear gray (Catling and Moore, 2003; Morris et al., 2020). Gray hematite typically forms at elevated temperatures of  $\sim 80\text{--}200\ ^\circ\text{C}$  (Catling and Moore, 2003), and the discovery of this mineral on Vera Rubin ridge may indicate some of the diagenetic fluids were warm (Rampe et al., 2020). In the gray Jura outcrops, ChemCam has observed subhedral to euhedral mm-scale crystals of nearly pure Fe (presumably as oxides) interspersed in Ca-sulfate veins and nodules and in the bedrock (Fig. 23A,B). In other places, local enrichment of Fe is seen on millimeter scales surrounded by lighter-toned bedrock depleted in Fe, indicating localized Fe mobilization in the Jura member (Fig. 23C; L'Haridon et al., 2019). The relatively coarse hematite grains in some locations and the relatively low dust content on the ridge may cause the strong hematite signature from orbit (e.g., Fraeman et al., 2020). The Science Team is currently evaluating the results from the Vera Rubin ridge campaign and will publish these results in a special issue of the Journal of Geophysical Research – Planets.

Beyond Vera Rubin ridge lie the other mineral units identified from orbit, the clay-bearing unit and the sulfate unit. *Curiosity* began the investigation of the clay-bearing unit (unofficially named “Glen Torridon”) in January 2019. A working group devised a traverse through this unit and has identified potential locations to drill (Fig. 24). The traverse was designed to investigate different morphologic sub-units, geologic contacts with Vera Rubin ridge and the sulfate unit, and a location with a strong phyllosilicate signature from orbit (Bennett et al., 2018). Orbital visible and near- to short-wave infrared spectral reflectance data from Glen Torridon are consistent with ferric smectite (Fox et al., 2018). Data from CheMin and SAM will be used to further characterize the phyllosilicate(s) present and determine whether the

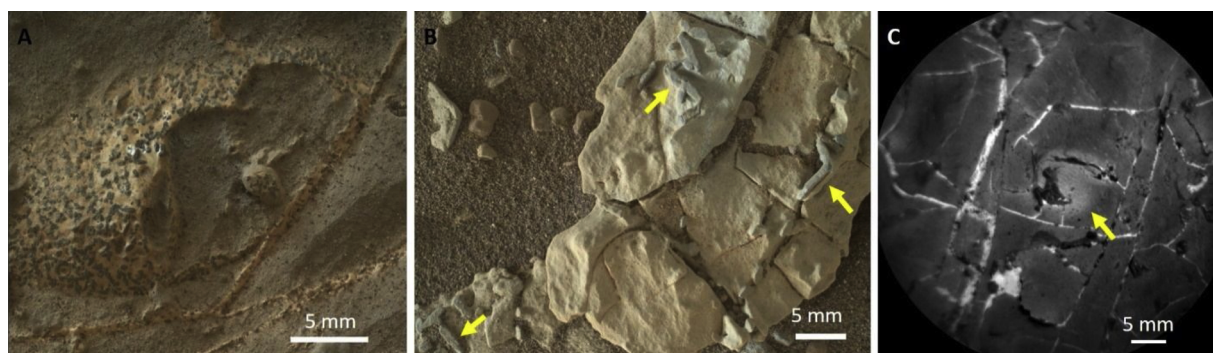


**Fig. 21.** Vera Rubin Ridge from orbit. The image is a HiRISE mosaic merged with CRISM detections of hematite. Red colors represent hematite mapped using orbital visible and near-infrared data from CRISM. The stars mark drill locations near and on the Vera Rubin Ridge, where D = Duluth, S = Stoer, H = Highfield, and R = Rock Hall. Highfield and Rock Hall were drilled  $\sim 30$  m apart laterally, so their drill locations are denoted with one star. Image credit: NASA/JPL-Caltech/JHUAPL/A. Fraeman (For interpretation of the references to colour in this figure legend, the reader is referred to the web version of this article.)

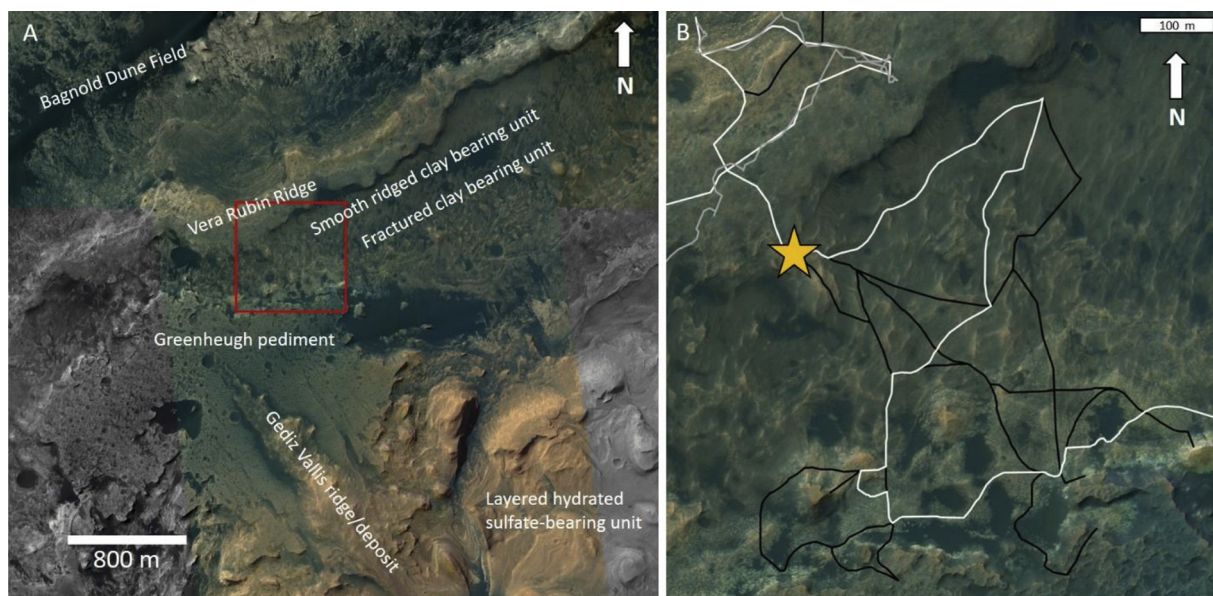




**Fig. 22.** Mastcam mosaic of a portion of the Jura member of the Vera Rubin Ridge, showing a variation in gray outcrop in the foreground to red outcrop in the background (sol 1931, mcam10081). Image credit: NASA/JPL-Caltech/MSSS (For interpretation of the references to colour in this figure legend, the reader is referred to the web version of this article.).



**Fig. 23.** Evidence for Fe mobilization in the Jura member of the Vera Rubin Ridge. A) MAHLI image of subhedral to euhedral Fe-rich crystals within a Ca-sulfate matrix (sol 2217, mhli00732). B) MAHLI image of Fe-rich linear features (sol 1922, mhli00152). C) ChemCam RMI image of Fe-rich dark materials associated with Ca-sulfate-filled fractures. Fe-rich dark materials in the center of the image are surrounded by lighter-toned Fe-poor bedrock. Arrow points to edge of light-toned halo surrounding Fe-rich material (sol 1934, ccam01934).



**Fig. 24.** A) HiRISE mosaic of the foothills of Mount Sharp and units that *Curiosity* will explore in current and future campaigns. The location of figure B is denoted by the red box. B) White path is the proposed route through the clay bearing unit. Black lines are other paths vetted for mobility. Gray line on the Vera Rubin Ridge shows *Curiosity*'s traverse through sol 2084. The star denotes *Curiosity*'s location on sol 2300. Adapted from Fox et al. (2019) (For interpretation of the references to colour in this figure legend, the reader is referred to the web version of this article.).

weathering intensity continues to increase going up Mount Sharp (Bristow et al., 2018; Mangold et al., 2019). Clay minerals can also become concentrated by sediment sorting and provide long-term preservation of organic molecules, due to high surface areas and charged interlayer sites (e.g., Bonaccorsi, 2011). As such, the clay-bearing unit is

a promising target for SAM investigation of organic molecules. The plan for the next few years is to explore the units that made Gale crater such an attractive landing site, further explore the evolution of ancient aqueous environments at Gale crater, and evaluate their habitability.

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