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Identification of a new spectral signature at 3 μm over Martian northern high latitudes: implications for surface composition.

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ABSTRACT

Mars northern polar latitudes are known to harbor an enhanced 3 μm spectral signature when observed from orbit. This may indicate a greater amount of surface adsorbed or bound water, although it has not yet been possible to easily reconcile orbital observations with ground measurements by Phoenix. Here we reprocessed OMEGA/Mars Express observations acquired during the Northern summer to further characterize this 3 μm absorption band increase. We identify the presence of a new specific spectral signature composed of an additional narrow absorption feature centered at 3.03 μm coupled with an absorption at $\lambda \geq 3.8 \mu\text{m}$. This signature is homogeneously distributed over a bright albedo open ring surrounding the circumpolar low-albedo terrains between $\sim 68^\circ\text{N}$ and 76°N and $\sim 0^\circ\text{E}$ and 270°E . This location includes the Phoenix landing site. This feature shows no time variability and can be confidently attributed to a seasonally stable surface component. All together, the stability, spectral shape and absence of significant correlation with other signatures in the 1 – 2.5 μm range discard interpretations relying on water ice or easily exchangeable adsorbed water. The exact full spectral shape cannot be easily reproduced by pure minerals samples, although sulfates, notably lowly hydrated Ca-sulfates, provide interesting comparisons. A modification of the chemical or physical properties of the soil, potentially involving additional sulfates contaminants, or modification of the hydration state of sulfates, and/or modification of their grains size, seems a plausible explanation to this observation, which may then indicate geologically recent water alteration at high northern latitudes.


1. Introduction

Mars ubiquitous wide 3 μm spectral absorption has been first identified using ground observations and associated with the presence of hydrated minerals at the surface of the red planet (Sinton, 1967). This feature was then observed from orbit by e.g. Mariner and Phobos missions (Pimentel et al., 1974; Bibring et al., 1990). The 3 μm variability over the surface was attributed to either difference in absorbed water content, or compositional variations (Murchie et al., 2000). The first global mapping of this spectral signature was obtained through observations of the Observatoire pour la Minéralogie, l'Eau, les Glaces et l'Activité (OMEGA) instrument (Jouglet et al., 2007; Milliken et al., 2007). Either superficial adsorbed water independent of composition, bound structural water within minerals, intermediate cases (e.g., water stored in the interlayer regions of phyllosilicates), or a mix of these possibilities, were considered to explain this feature (Milliken et al., 2007). Whatever its origin, as this spectral feature is mainly associated with OH/H₂O, this 3 μm band was used to estimate a water weight % content of the surface (Jouglet et al., 2007; Milliken et al., 2007). These studies reveal that the water content of the Martian soil increases poleward by a factor greater than two, with the highest levels reached in the northern latitudes (≥ 8 wt. % H₂O, up to

10% to 15%). While the absolute value of this weight % is model-dependent (Jouglet et al., 2007; Milliken et al., 2007; Liu et al., 2020), the relative increase corresponds to a robust latitudinal trend of the band depth. This polar increase of hydration was mainly attributed to an increase of superficial weakly bound adsorbed water (Jouglet et al., 2007; Poulet et al., 2008, 2010). This interpretation was notably supported by a tentative identification of seasonal variations of the 3 μm band depth, suggestive of exchanges between the surface and the atmosphere (Jouglet et al., 2007).

In-situ measurements from the Phoenix lander offered observational "ground truth" constraints about the water content of soils at northern polar latitudes. Phoenix reported the detection of various water-related minerals such as carbonates (2–5%) and perchlorates (< 1%), and an overall soil water content of 1–2% (Boynton et al., 2009; Hecht et al., 2009; Smith et al., 2009; Poulet et al., 2010; Sutter et al., 2012) which is well below the value of 10–11% retrieved using OMEGA data at 3 μm (Jouglet et al., 2007; Milliken et al., 2007; Audouard et al., 2014). Both measurements were argued to be reconcilable if we consider that the 3 μm feature is due to superficial adsorbed water contained within the first few hundred micrometers sampled by OMEGA, and not in the first centimeters sampled by Phoenix (Poulet et al., 2010). However, water from samples heated by Phoenix was not released at low, adsorbed water-compatible temperatures (Smith et al., 2009). Observed temperatures of released, about 300°C and 700°C, are compatible with the presence of several hydrous minerals or phases within soils (Smith et al.,

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2009). In the absence of a significant amount of adsorbed water, the discrepancy between OMEGA and Phoenix measurements may thus be related to either the presence of a very superficial hydrated coating with tightly bound - and not adsorbed - water, or to an improper association between 3 μm band depth/shape and water %.

In-situ measurements of the soil water content were also obtained at equatorial latitudes by the Mars Sample Laboratory (MSL) rover Curiosity. Various instruments (the Sample Analysis of Mars, SAM, the Chemistry and Mineralogy, CheMin, and the Chemistry Camera, ChemCam) lead to the conclusion that most of the water (about 2 %) contained in the martian regolith at MSL landing site is likely not weakly bound adsorbed water, but rather more strongly bound water within the amorphous phase (Leshin et al., 2013; Bish et al., 2013; Meslin et al., 2013). Based on these results, Meslin et al. (2013) suggested that the spatial variations of the surface water content observed from orbit may be partly related to the variations of abundances of hydrated amorphous components.

A reassessment of the 3 μm band in the OMEGA dataset then similarly questioned the interpretation of this band by adsorbed water (Audouard et al., 2014): using a larger dataset and applying a stricter atmospheric filtering on the data, these authors did not confirm the seasonal variations previously reported from OMEGA data at northern high latitudes. Moreover, the surface 3 μm feature was found to be extremely stable with respect to changes in relative humidity. It was then suggested that the 3 μm band at high northern latitudes is also not linked with adsorbed water, and that the increased 3 μm band may be related to a specific polar alteration process involving ice and modifying the amount of tightly bound water within minerals or amorphous phases (Audouard et al., 2014).

Another piece of the puzzle came from laboratory measurements on analogs and meteorites: these measurements revealed that increasing the amount of adsorbed water usually creates an absorption feature at 1.9 μm in addition to the one at 3 μm , while this 1.9 μm band is not observed over most of Mars unlike the ubiquitous 3 μm band, suggesting again that Mars 3 μm band may not be related to adsorbed water (Pommerol et al., 2009; Pommerol et al., 2011; Beck et al., 2015). Note however that under northern polar latitudes, above 60°N, the 1.9 μm band is observed almost everywhere, with an increasing latitudinal trend similarly to the 3 μm feature (Poulet et al., 2008), which is this time consistent with the hydrated water interpretation for the 3 μm polar increase (Poulet et al., 2008). Alternatively, as the 1.9 μm band is typically an indicator of the presence of hydrated minerals such as phyllosilicates (Poulet et al., 2005), the presence of this 1.9 μm may also suggest a specific widespread presence of hydrated minerals at high northern latitudes. Actually, the presence of hydrated minerals such as phyllosilicates or sulfates also leads to a significant co-increase of the 1.9 μm and 3 μm band (Jouglet et al., 2007, figure 10).

Remote observations of the high northern latitudes of

Mars indeed reveal the presence of other spectral features related to hydrated minerals, in particular sulfates, suggesting the presence of water alteration (Langevin et al., 2005a; Poulet et al., 2008; Massé et al., 2012; Carter et al., 2016). These identifications are based on features located in the 1 to 2.5 μm wavelength range. Near-IR data first revealed an area about $\sim 240^\circ\text{E}$ and $\sim 85^\circ\text{N}$ with strong signatures compatible with calcium-rich hydrated sulfates such as gypsum (Langevin et al., 2005a). Note that it has been reported afterward that perchlorates can display similar spectral signatures (Hanley et al., 2015). Weaker signatures of these sulfates (and possibly perchlorates) have then been observed all over the northern circumpolar dune fields surrounding the perennial cap (Massé et al., 2010, 2012). Another candidate for some OMEGA hydration features observed at high northern latitudes is zeolites (Poulet et al., 2010). Large areas of the northern high latitudes actually possess a distinct spectral shape in OMEGA data which may be related to some sort of widespread water-limited alteration (Horgan and Bell, 2012). Such a widespread limited weathering potentially producing hydrated phases may also explain data collected by the Thermal Emission Spectrometer (Wyatt et al., 2004; Michalski et al., 2005).

If hydrated phases concentration within soil is low, such as expected from a limited weathering, then they may not be easily identified in remote observations. Some low-abundance minerals that have been detected through in-situ measurements at northern high latitudes, as e.g. Mg-perchlorates and carbonates at the Phoenix landing site (Boynton et al., 2009; Hecht et al., 2009) have not yet been detected by orbital observations. This may be due to orbital detection threshold limitations that can be up to 6% for carbonates (Jouglet et al., 2007; Poulet et al., 2010). Similarly, the hydrated amorphous phase observed by Curiosity (Leshin et al., 2013; Bish et al., 2013; Meslin et al., 2013) is not identified from orbit. This further suggests that low amount of hydrated minerals or amorphous hydrated phases may be widely distributed at high latitudes, without being easily identifiable by OMEGA and the Compact Reconnaissance Imaging Spectrometer for Mars (CRISM), notably within the most widely used 1 to 2.5 μm spectral range. They could however take part in the spectral shape and deepness of the 3 μm H₂O/OH feature, which shows up more rapidly compared to features in the 1 to 2.5 μm range (e.g. Bishop, 2019).

Overall, northern latitudes show complex spectral behaviors related to hydrated mineralogy and/or adsorbed water, some of which are still not uniquely interpreted nor related to lander observations. In this study, we reprocess the OMEGA observations of the north polar regions of Mars to further constrain the origin of the polar increase of the 3 μm feature.

2. Dataset and methods

2.1. OMEGA data

The OMEGA experiment onboard the Mars Express orbiter is a visible and near-infrared imaging spectrometer composed of three different channels covering the 0.38–5.1 μm

spectral range, respectively called "V" (0.38–1.05 μm), "C" (0.93–2.73 μm) and "L" (2.55–5.1 μm) channels (Bibring et al., 2004). Observations of the Martian surface using the three channels were conducted from 2004 to 2010. Additional observations are still ongoing with limited spectral coverage, but in this study, we will only deal with pre-2010 measurements. It corresponds to a dataset of 9646 hyperspectral cubes covering most of the Martian surface with a typical spatial sampling of 1 km, varying from 350 m to 5 km depending on the position of Mars-Express on its elliptical orbit. Repeated observations of the same region have been frequently obtained over the mission, in particular at high latitudes where time sampling can be about 10° of L_s (Langevin et al., 2005b, 2007) thanks to the quasi-polar orbit of Mars Express.

While the response of the C channel remains stable from 2004 to 2010, the onboard calibration (OBC) of OMEGA reveals that the L channel response has significantly varied over the mission (Jouglet et al., 2009). In this paper, we use the adapted instrument transfer function (ITF) that accounts for these variations, as described in Jouglet et al. (2009). Most of our study will however focus on data acquired during MY 27 northern summer (from orbits 923 to 1224, i.e. over L_s 98° – 137°), a timeframe over which the ITF was nominal.

As the thermal emission from the planet is of importance for wavelengths larger than 3 μm , we remove this thermal component from each OMEGA data cube using the method described in Jouglet et al. (2007), and thereafter used in other studies based on OMEGA observations (e.g. Audouard et al., 2014). This method relies on the comparison between the observed reflectance at 5 μm , where thermal emission dominates for temperatures greater than about 200 K, and the reflectance at 2.4 μm where the thermal contribution is negligible. We first estimate the theoretical thermal-free reflectance at 5 μm using the measured reflectance at 2.4 μm , assuming a typical spectral shape extracted from Erard and Calvin (1997). We then calculate the appropriate thermal additional contribution that reproduces the observed reflectance at 5 μm . This thermal contribution can then be removed at all wavelengths assuming Planck law. An example is shown in Figure 1.

We also correct the OMEGA spectra from the atmospheric gas absorption using the method described in Langevin et al. (2005a) and Jouglet et al. (2007). Each OMEGA spectrum is divided by a transmission spectrum of the martian atmosphere computed from a comparison between the base and the top of Olympus Mons, that is scaled to the appropriated atmospheric column height using the 2 μm atmospheric CO_2 absorption band (Figure 1).

OMEGA data are processed using a Python module developed for this study ("OMEGA-Py"). This module notably includes a re-implementation of the most recent release ("SOFT10") of the official OMEGA software (IDL routines) available with OMEGA data. It also contains several additional data reduction functions and graphic tools.

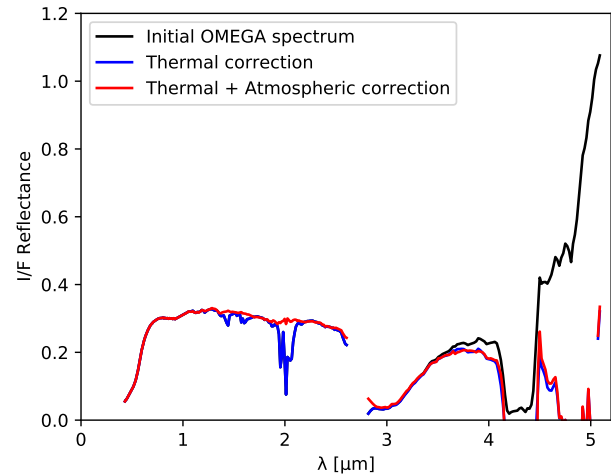


Figure 1: Thermal and atmospheric corrections steps applied to OMEGA spectra in this study. The black spectrum is an example of initial OMEGA reflectance spectrum (averaged over 9 pixels), the blue spectrum is corrected from the thermal component, and the red one is also corrected from the atmospheric absorption. See text for details.

2.2. Data filtering

OMEGA has provided a large number of observations: we can thus apply strong filtering on the dataset while keeping a near global coverage of the polar regions. This data filtering is based on criteria established and detailed in previous studies (e.g. Langevin et al., 2007; Poulet et al., 2007; Ody et al., 2012; Vincendon et al., 2015) that are summarized below.

- To prevent possible effects related to viewing geometry (Pommerol and Schmitt, 2008a), we restrict our study to near-nadir pixels (emergence angle $< 10^\circ$) with an incidence angle $< 70^\circ$.
- As the OMEGA IR detector has been designed and calibrated to operate at temperatures below 80 K, we remove the observation lines that have been acquired with higher temperatures.
- Depending on the hyperspectral cube number, some lines at the beginning and the end of the images have to be removed as they are used for instrument calibration or are usually corrupted (see the OMEGA software documentation for more details).
- Near-saturation of some pixels can occur when observing a surface element brighter than expected (e.g., sun-lighted tilted surfaces). Similarly to the higher quality level criteria defined in Vincendon et al. (2015), we use a conservative threshold of 500 DN for the spectel #40 (corresponding to $\lambda = 1.486 \mu\text{m}$), and reject pixels with lower value to prevent non-linearity effects.
- An instrumental issue perturbs 128 pixels wide observations from orbit 513 to 3283 (see OMEGA software)

Description	Spectral indice
1.5 μm BD (Poulet et al., 2007)	$1 - \frac{R(1.50) + R(1.51)}{R(1.30) + R(1.71)}$
1.9 μm BD (Langevin et al., 2005b)	$1 - \frac{R(1.93)}{0.74R(1.86) + 0.26R(2.14)}$
2.4 μm BD	$1 - \frac{\bar{R}(2.43, 2.45, 2.46)}{\bar{R}(2.27, 2.29, 2.30)}$
Wide 3 μm BD (Jouglet et al., 2007)	$1 - \frac{1}{N_\lambda} \sum_{\lambda=2.9\mu\text{m}}^{3.7\mu\text{m}} \frac{R(\lambda)}{\text{Cont}(\lambda)}$ where $\text{Cont}(\lambda) = R(2.35) + \left(\frac{R(3.7) - R(2.35)}{1.35} \right) (\lambda - 2.35)$
Narrow 3 μm BD	$1 - \frac{\bar{R}(2.96, 3.01, 3.03)}{0.67\bar{R}(2.88, 2.90, 2.92) + 0.33\bar{R}(3.19, 3.21, 3.23)}$
4 μm BD	$1 - \frac{\bar{R}(3.94, 3.96)}{\bar{R}(3.69, 3.71, 3.73)}$

Table 1

Spectral estimators used in this study. $R(\lambda)$ is the reflectance I/F at the given wavelength λ in μm , and $\bar{R}(\lambda_1, \dots, \lambda_n)$ is the average reflectance from the n spectels corresponding to wavelengths $\lambda_1, \dots, \lambda_n$. For the wide 3 μm BD, $\text{Cont}(\lambda)$ is the used spectral continuum defined linearly between 2.35 μm and 3.7 μm , and N_λ is the number of spectels between 2.9 μm and 3.7 μm .

as columns #81 to #96 present 44 corrupted wavelengths. Even if it is possible to recover part of the information, we decided to remove these columns from the study to focus on the more reliable part of the images. Between orbits 2124 and 3283, an additional random noise also affects columns #68 to #128, which are also removed from our study.

- The 3 μm band is very sensitive to the presence of water ice at the surface or in the atmosphere (Vincendon et al., 2011). We remove all pixels with a water ice 1.5 μm absorption (see Table 1) $> 1\%$. We also compute the ice cloud index (ICI) defined in Langevin et al. (2007) as the ratio of the reflectance at $\lambda = 3.40 \mu\text{m}$ over the reflectance at $\lambda = 3.53 \mu\text{m}$ to detect the presence of clouds and remove pixels with an ICI < 0.8 (Audouard et al., 2014).

In addition, 3 μm band depth estimators are known to exhibit a strong dependency to the surface albedo (Jouglet et al., 2007; Pommerol and Schmitt, 2008b), so a specific attention should be taken when comparing regions of different albedo. Thus, in the following, we usually apply restrictions on the value of the reflectance at $\lambda = 2.26 \mu\text{m}$ when studying the spatial evolution of a given criterion.

2.3. Band depth estimators

The spectral estimators used and developed in this study are summarized in Table 1, with their explicit formula.

We first consider three band depths estimator derived from previous studies at 1.5 μm (water ice), 1.9 μm (hydrated minerals or adsorbed water) and 3 μm . For the 1.5 μm and 1.9 μm band depth (BD), we compute the absorption at the center of the band relatively to a continuum defined

linearly between two reference points at the extremities of the band. Wavelengths are selected to avoid atmospheric absorption (see subsection 2.1). E.g., this is why the reflectance at 1.93 μm is used in the computation of the 1.9 μm BD instead of the reflectance at 1.94 μm (Langevin et al., 2005a). For the quantification of the absorption related to the full 3 μm band observed everywhere on Mars, we have computed the 3 μm normalized integrated band depth under a continuum defined linearly between 2.35 μm and 3.7 μm , as described in Jouglet et al. (2007). This band depth is hereafter referred as the *wide* 3 μm BD.

We have then defined three new estimators which appeared necessary during the development of this study. A new estimator focus on the central part of the 3 μm band: it makes it possible to locate a narrow sub-feature within the wide 3 μm band. It is called hereafter the *narrow* 3 μm BD: continuum reference points are taken at 2.9 and 3.2 μm instead of 2.9 and 3.7 μm for the usual, previously used wide 3 μm BD. As most wavelengths used for this criterion are localized around the bottom of the 3 μm band, reflectance values are low due to the strong absorption, and we average consecutive wavelengths to increase the signal-to-noise ratio (SNR).

Another spectral feature identified during this work is located at about 4 μm . In the case of this 4 μm band, as we can only see one part of the band because of the strong atmospheric absorption for wavelength larger than 4.06 μm , we compute a new 4 μm BD estimator using a linear continuum equals to the average reflectance value at $\sim 3.7 \mu\text{m}$.

Finally, we will also explore possible correlations with another minor band at 2.4 μm for which a new estimator has also been developed, using the calculation principles of the 4 μm BD (Table 1).

3. Results

3.1. Phoenix landing site area

As discussed in section 1, Phoenix provided unique ground truth measurements in the north polar region that can be compared to information derived from orbital data.

Figure 2 illustrates the evolution of surface spectral reflectance over a latitudinal band covering a relatively uniform bright albedo area and crossing the Phoenix landing site at 234°E / 68°N. We can see that the wide 3 μm BD progressively increases with latitude, as previously reported (Jouglet et al., 2007; Milliken et al., 2007). In panel (b) of the same figure, we show spectra from three different latitudes (52°N, 62°N & 72°N). We can see that the spectral modifications are not progressive with latitudes, in particular between 2.8 and 4.0 μm . Indeed, we observe that the shape of the northern spectrum differs from the two others (blue versus yellow and green on Figure 2): the maximum of absorption of the 3 μm band is shifted from $\sim 2.9 \mu\text{m}$ to $\sim 3.0 \mu\text{m}$, and a decrease of the reflectance is observed beyond $\sim 3.8 \mu\text{m}$. This difference of behavior is enhanced in the spectral ratios (Figure 2b): while the ratio obtained south of Phoenix (black) reveals an increase of the 3 μm band with a typical band shape similar to the actual band prior ratioing (minimum at $\lambda < 2.9 \mu\text{m}$), the ratio obtained poleward, over the Phoenix area (red), exhibits a strong but narrow absorption band centered on $\lambda \sim 3.03 \mu\text{m}$ (hereafter referred as the *narrow 3 μm band*) along with a wide shallow absorption band starting at $\lambda \sim 3.6 \mu\text{m}$.

We have then mapped the spatial distribution of this narrow 3 μm band in the surroundings of Phoenix using the new associated spectral criteria defined in Table 1 and subsection 2.3. This mapping (Figure 2d) reveals that this band is observed only above 68°N. The transition between the area without signature ($< 66^\circ\text{N}$) towards that with the signature ($> 68^\circ\text{N}$) is abrupt, which is very different from the progressive latitudinal gradient observed for the wide 3 μm band (Figure 2c).

3.2. Spatial distribution

Outside the Phoenix landing site region, we computed the narrow 3 μm BD over the whole OMEGA dataset that contains observations with the 3 channels operating (i.e. up to orbit 8485 and the loss of the C channel) to map the spatial distribution of the newly identified signature at $\lambda \sim 3.03 \mu\text{m}$. It turns out that the signature is only present in the north polar regions. Figure 3 shows composite maps of the northern hemisphere (down to 50°N), and we observe on panels (e) and (f) that the narrow 3 μm band is spread over an open ring, crescent-shaped area around the North pole. This open ring area, hereafter referred as the *3 μm northern ring* or *ring*, corresponds to bright regions at latitudes comprised between $\sim 68^\circ\text{N}$ and 76°N and longitudes between $\sim 0^\circ\text{E}$ and 270°E . The lower (southern) transition to the ring can occur in bright terrains, without changes in terms of surface albedo. Between 30°E and 115°E, the transition is located close to the frontier with low-albedo terrains. The upper (northern) limit of the ring (around 76°N) corresponds to

the delimitation between the high-albedo northern terrains and the dark regions around the perennial north polar cap (cf Figure 4a). Similarly, no detection has been observed in the southward extension of this low-albedo region in Vastitas Borealis, between 280°E and 360°E, while latitudes are comparable to the 3 μm northern ring area. The new detection thus seems to require two conditions for observation: bright albedo and latitudes greater than $\sim 68^\circ\text{N}$.

Figure 4 shows the spatial distribution of the detection (narrow 3 μm BD values larger than 0.27) over a map of geological units from Tanaka et al. (2005). We observe that it is spread across two main Early Amazonian geological units (referred as Vastitas Borealis marginal and interior units in Tanaka et al. (2005)). This suggests that the signature is associated with the superficial composition of the soil overlying both units (OMEGA sensing depth is a few hundreds of micrometers only).

We show in Figure 5 latitudinal profiles of band depth averages between 45°N and 75°N for two areas crossing the 3 μm northern ring ("A" and "B"). Four spectral estimators are shown: wide 3 μm BD, 1.9 μm BD, narrow 3 μm BD, and 4 μm BD. The values of spectral estimators are averaged over longitude (i.e. image lines), considering only the pixels with a surface reflectance at 2.26 μm between 0.3 and 0.35 to avoid possible interference due to albedo variations. We observe that the wide 3 μm BD and 1.9 μm BD continuously increase with the latitude. On the other hand, both the narrow 3 μm and 4 μm bands present an abrupt increasing step from low to higher values over a few degrees of latitude only.

Actually, as shown in Figure 3, the spatial distribution of the 4 μm band is similar to that of the narrow 3 μm one. We can see that the correlation between both is excellent. This strongly suggests that both are connected, i.e. that the phenomena ongoing over the 3 μm northern ring is responsible for both absorption features.

3.3. Correlation with other spectral features

We have averaged several latitudinal spectral ratios ($\sim 70^\circ\text{N}$ over $\sim 63^\circ\text{N}$) extracted all around the 3 μm northern ring to isolate the mean spectral signature of this newly detected area (see red spectrum in Figure 6).

We first compare this signature with another spectral ratio taken south of the ring ($\sim 60^\circ\text{N}$ over $\sim 50^\circ\text{N}$). This ratio shows that the 3 μm increase from $\sim 50^\circ\text{N}$ to $\sim 60^\circ\text{N}$ has a different spectral shape compared to that observed at northern latitudes over the ring. We observe that the main differences between these two spectral ratios reside in the 2.8 – 4 μm spectral range. The increase of absorption at 3 μm is noticeably larger between 63°N and 70°N than between 50°N and 60°N, while the latitudinal range is smaller, suggesting an enhancement of the 3 μm increase over the ring. As previously discussed for the Phoenix region (subsection 3.1), we observe again that the spectral shape between both also significantly differs. Over the 3 μm northern ring, we observe a narrow absorption centered on 3.03 μm along with what may be the beginning of a wide absorption at wavelengths larger than 4 μm (starting at $\lambda \sim 3.6 \mu\text{m}$).

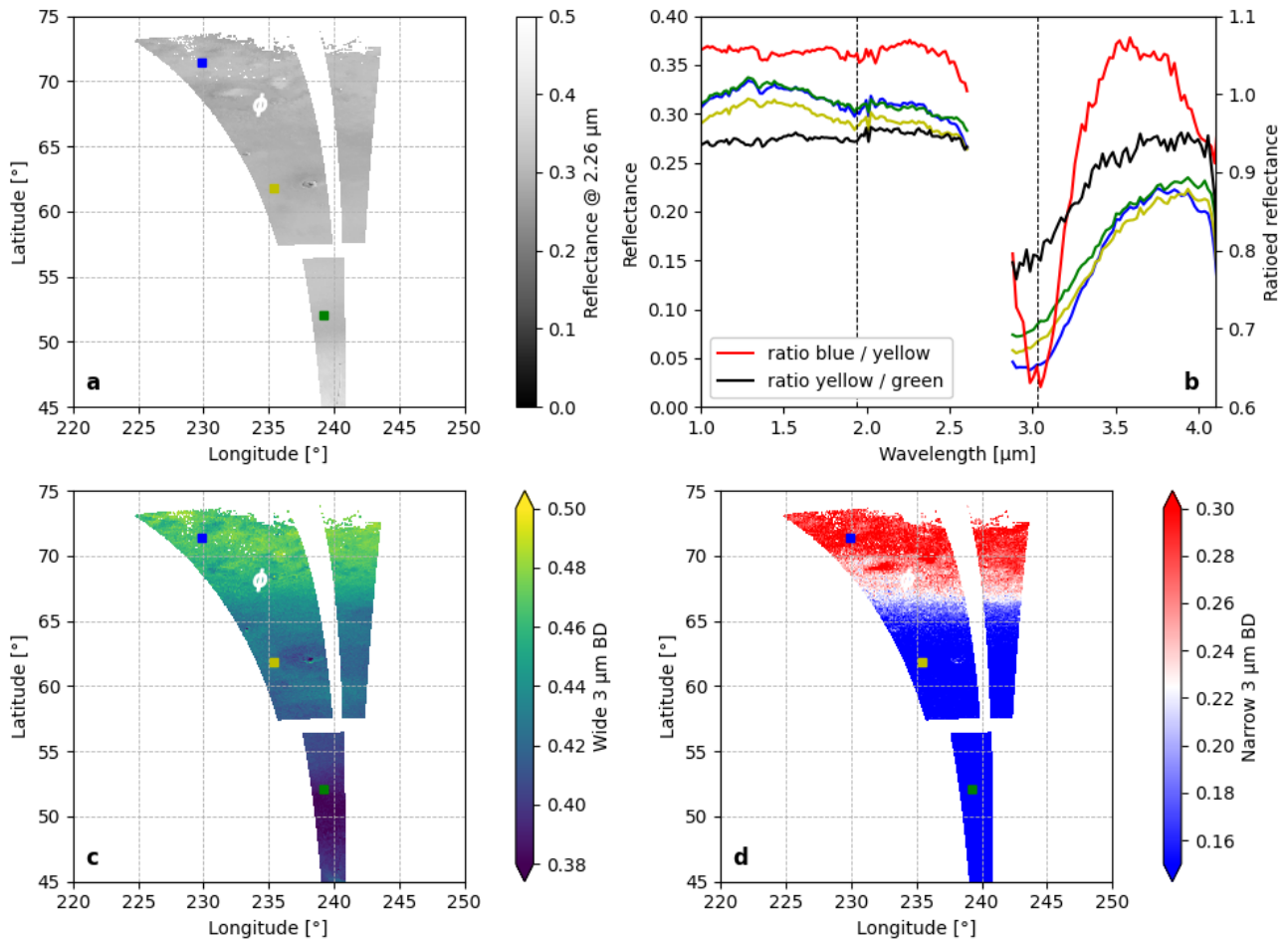


Figure 2: Maps of OMEGA cubes #3 & #4 from orbit 979 ($L_s = 105^\circ$, local time ~ 3 p.m.) showing the surface reflectance at 2.26 μm (a), the wide 3 μm BD (c) and the narrow 3 μm BD (d). The white " ϕ " indicates the position of the Phoenix lander. The blue, yellow, and green spectra of the panel (b) are averaged on a square of 3×3 pixels whose position are indicated by the colored squares on panels (a), (c) & (d). Black dotted lines correspond to wavelengths of 1.94 and 3.03 μm . We observe on panel (b) that the evolution between 60°N and 70°N (red line) is very different from the one between 50°N and 60°N (black line) with a narrow signature centered around 3 μm . Which results in a shift of the position of the maximum of absorption of the widespread Martian 3 μm band from 2.9 μm (yellow & green lines) to 3.0 μm (blue line). We also observe an absorption at $\lambda > 3.6$ μm specific to the red ratio and associated blue spectrum.

Southward, we observe a wider 3 μm absorption centered at $\lambda \sim 2.9$ μm and no 4 μm absorption (cf Figure 6c).

Having a closer look at the 1 – 2.6 μm spectral range, we can observe the presence of some very small absorption features (typically about one percent of absorption in the ratios). The main features that can be identified in the red spectrum are localized at 1.37, 1.9, 2.4 and 2.6 μm . The 1.37 μm and 2.6 μm band are associated with atmospheric water vapor (Encrenaz et al., 2005). As we compute spectral ratio across about ten degrees of latitude, the presence of these water vapor bands in the spectra is consistent with the latitudinal increase of the water vapor content characteristic of this season during which water vapor is released from the cap (e.g. Smith, 2006; Melchiorri et al., 2007; Trokhimovskiy et al., 2015). Note also that the solar incidence angle increases as we go poleward, which results in longer atmospheric path lengths that strengthens differences in ratios. The 1.9 μm

and 2.4 μm bands are observed for various hydrated minerals as they result from the stretching and bending vibrations of H_2O molecules and -OH groups (Gendrin et al., 2005; Bishop, 2019). However, the mapping of these absorption bands in the North polar regions (Figure 3b&d) and their latitudinal profiles in the bright regions (Figure 5) show that their intensity tends to increase with latitude, but without being correlated to the 3 μm northern ring. In addition, regions that exhibit the strongest absorption at 1.9 μm and 2.4 μm are not located in the ring either (see Figure 3 and also Poulet et al. (2008)).

3.4. Comparison with the Southern hemisphere

Figure 6 also presents a typical ratio from the southern polar regions (blue spectrum). As no narrow 3 μm band has been detected in the Southern hemisphere during mapping, the location of the used spectra to calculate the ratio

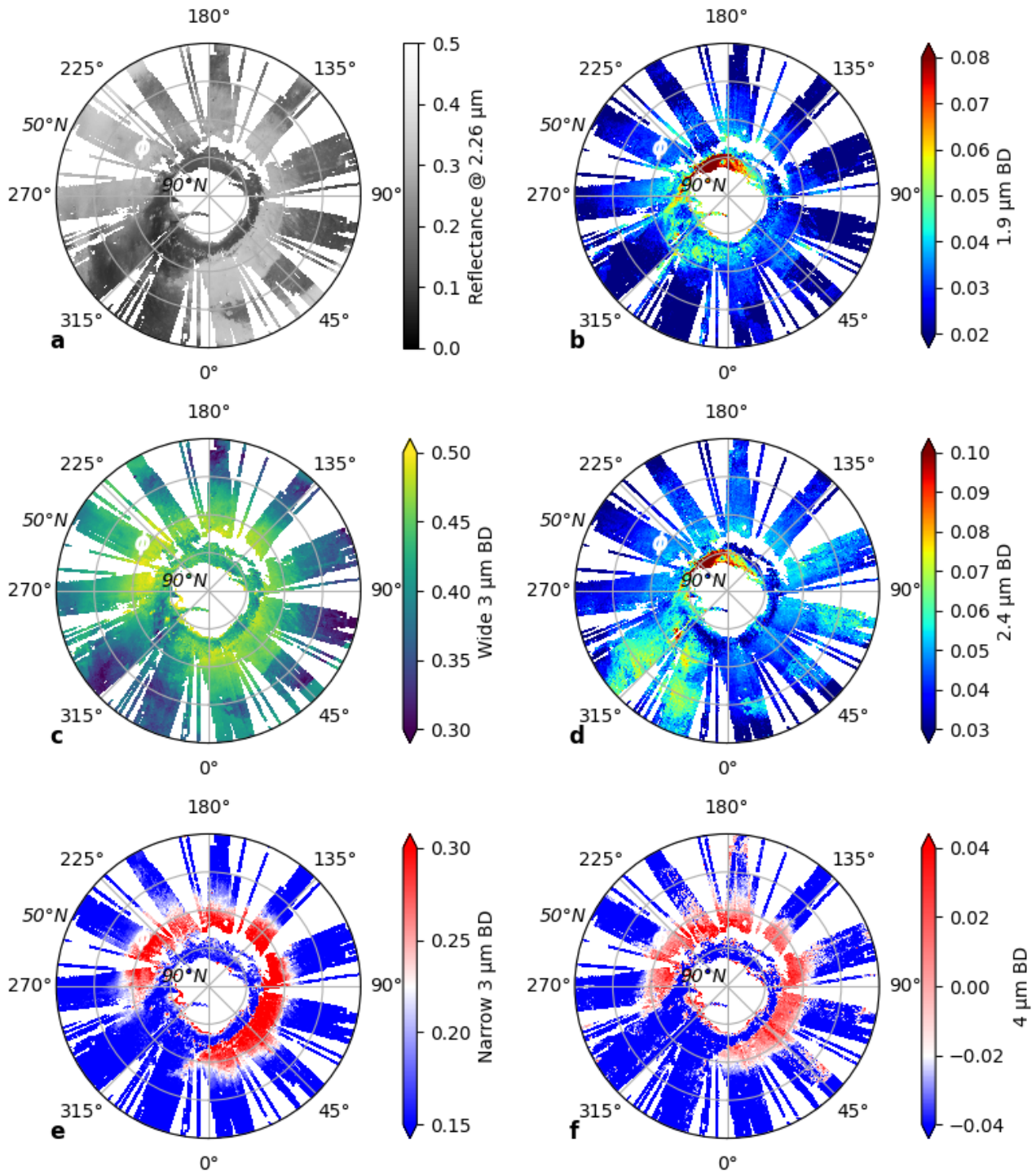


Figure 3: Maps of the martian north polar regions, showing the surface reflectance at 2.26 μm (a), the 1.9 μm BD (b) the wide 3 μm BD (c), the 2.4 μm BD (d), the narrow 3 μm BD (e) and the 4 μm BD (f). The white " ϕ " indicates the position of the Phoenix lander. These composite maps are assembled from OMEGA nadir observations during the northern summer of MY 27 ($L_s = 98^\circ - 132^\circ$), when the response of the C-channel is nominal, i.e. from orbit 923 to 1223 (Jouglet et al., 2009).

We observe on panels (e) and (f) the presence of an open ring structure between 68°N and 76°N (referred as the 3 μm northern ring): this surface distribution is specific to the narrow 3 μm BD and 4 μm BD criteria. This region is not correlated with surface albedo variations although it occurs only in bright albedo terrains (a). It is nearly anti-correlated with the 1.9 μm (b) and 2.4 μm (d) absorption features that tend to be observed mainly in dark terrains. The wide 3 μm BD is also shown for comparison in panel (c).

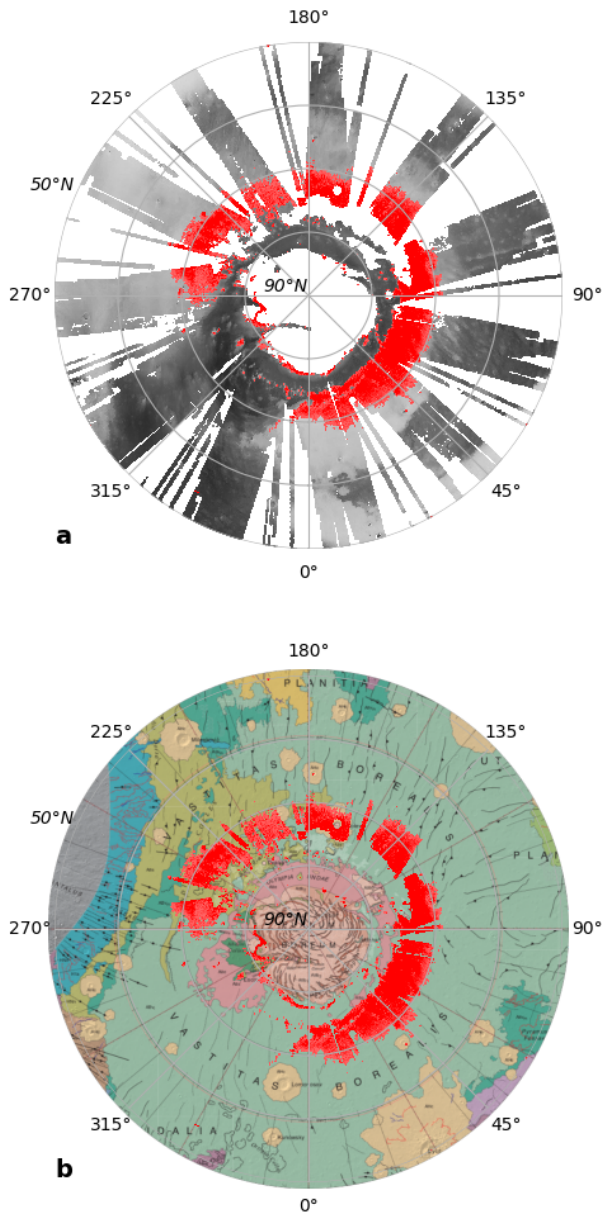


Figure 4: Overplot of the 3 μm northern ring, defined as the region displaying a narrow 3 μm BD larger than 0.27 (red) over a map of the surface reflectance at $\lambda = 2.26 \mu\text{m}$ (a), and a geological map of the Martian north regions from Tanaka et al. (2005) (b). We observe that the narrow 3 μm signature is distributed over two main geological units dated Early Amazonian (see Tanaka et al. (2005) for the geological units legend).

($\sim 74^\circ\text{S}$ over $\sim 62^\circ\text{S}$, blue spectrum) were chosen to be as close as possible to the conditions of observation of the signature in the Northern hemisphere in terms of albedo and latitude. This spectral ratio confirms that the spectral band shape observed over the 3 μm northern ring is specific to the northern high latitudes: we can see that the maximum of absorption of the 3 μm band occurs at $\sim 2.9 \mu\text{m}$ as for the black spectrum, and that there is no absorption at 4 μm .

We can notice that several absorption features are present in the 1 – 2.6 μm spectral range in the southern hemisphere, and that they also differ from the north. This will not be discussed here, but information can be found in Poulet et al. (2008) and Carter et al. (2016).

3.5. Time stability

We have looked for time variations of this signature, as a function of years, season, day, or hour.

Due to the Mars Express orbit and our restrictions applied on the emergence and incidence angle (cf subsection 2.2), all the observations used here to compute maps from the Martian northern summer of MY 27 have been acquired in the middle of the afternoon (around 3 p.m.). Figure 7 shows latitude profiles above the Phoenix landing region derived with 5 successive OMEGA observations taken over ~ 15 days (L_s 105°– 112°). We can see that apart from some small variations due to longitudinal spatial discrepancies between the observations, the narrow 3 μm BD profile remains stable.

We have also compared OMEGA observations acquired in the morning (~ 9 a.m.) and evening (~ 6 p.m.) over different Martian Years (MY 28 to 30). Morning and evening observations imply differences in the observation geometry, essentially in terms of incidence angle: comparing both is not trivial. We did not identify significant variation in the narrow 3 μm BD criteria, despite a wide amplitude in parameters varying with time such as surface temperature.

Overall, available observational constraints suggest that the spatial distribution and spectral shape of this signature are stable with time.

4. Discussion

4.1. Potential atmospheric biases

OMEGA observations of the surface are obtained through the atmosphere. Atmospheric constituents could create an absorption at 3 μm or a modification of the 3 μm band depth or shape. There are three possible atmospheric contributors in this wavelength range: water vapor, water ice clouds/fog, and dust.

Water vapor concentration is at its highest value during northern summer close to the polar cap, with a latitudinal gradient (Smith, 2004, 2006). A spectral ratio between high and low latitudes could thus highlight more water vapor, as discussed in subsection 2.3. However, the impact of water vapor on the 3 μm band was previously shown to be negligible considering the expected depth of water vapor features (Jouglet et al., 2007).

Atmospheric dust is known to modify significantly the depth of the 3 μm feature (Audouard et al., 2014). Dust is however not associated with a narrow 3 μm spectral feature such as observed here (Vincendon et al., 2009). Moreover, the amount of dust in the atmosphere, its variability, and its impact on surface spectra has been studied in details for the MY27 northern summer OMEGA observations of the high latitudes (Vincendon et al., 2007): the optical depth of dust

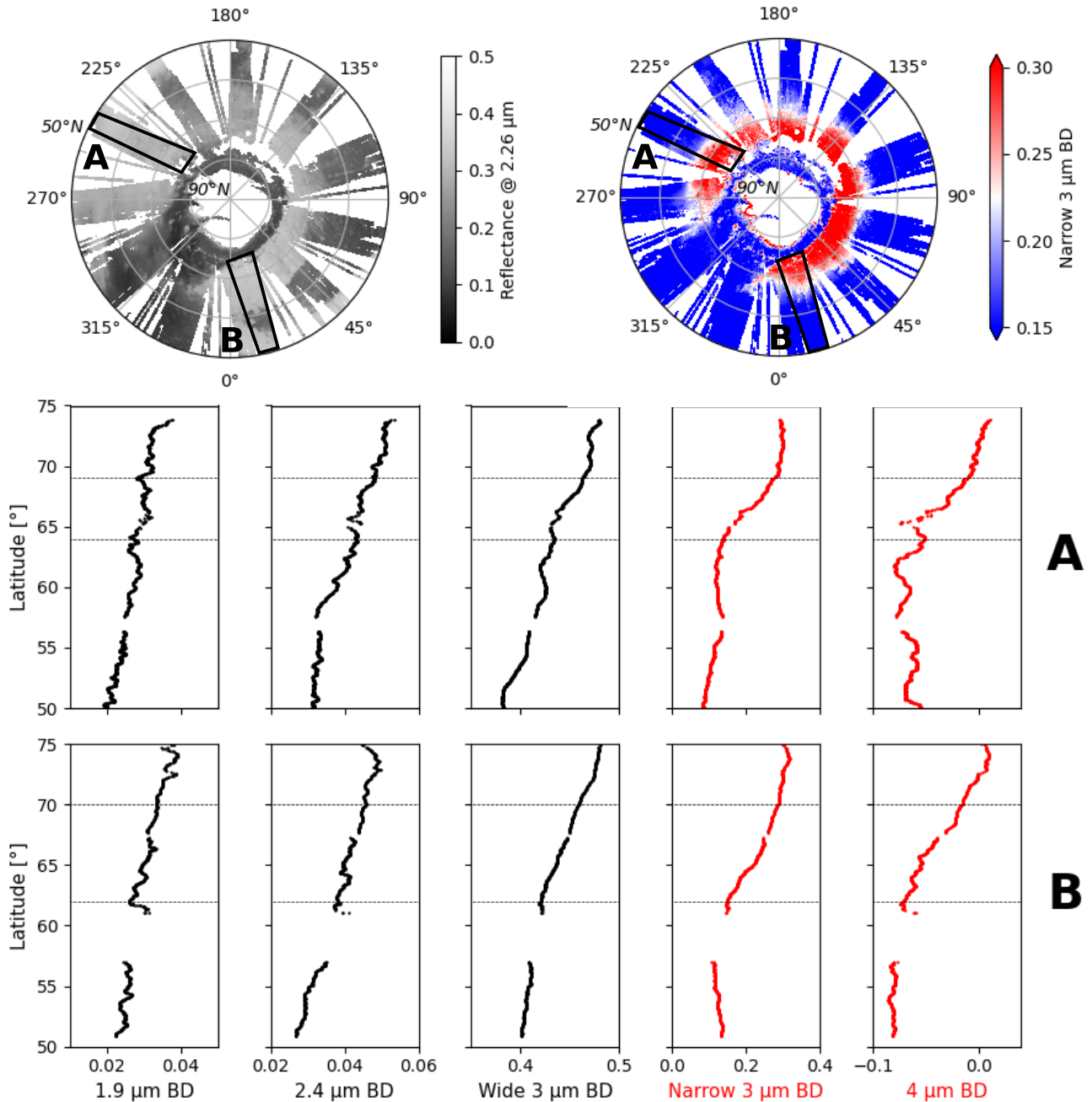


Figure 5: Latitudinal variations of the 1.9 μm BD, 2.4 μm BD, wide 3 μm BD, narrow 3 μm BD and 4 μm BD between 50°N and 75°N, from two bright regions identified on the albedo and narrow 3 μm BD maps. A: OMEGA cubes 3 & 4 of orbit 979, $L_s = 105^\circ$, MY 27, local time $\sim 15\text{h}$, longitude $\sim 230^\circ\text{E} - 243^\circ\text{E}$ (see Figure 2). B: OMEGA cubes 2 & 3 of orbit 941, $L_s = 100^\circ$, MY 27, local time $\sim 15\text{h}$, longitude $\sim 5^\circ\text{W} - 20^\circ\text{E}$. Each point is averaged in longitude (image lines), considering only the pixels with a surface reflectance at 2.26 μm between 0.3 and 0.35 to avoid albedo effects in the latitudinal trends of the band depth estimators. Profiles are then smoothed using a moving average. We observe on panel A that if the 1.9 μm , 2.4 μm and wide 3 μm bands show a continuous increase between 50°N and 75°N (black lines), the narrow 3 μm band behaves differently with a step profile: the narrow 3 μm BD \sim is at 0.12 below 64°N, and jump to ~ 0.3 above 69°N. And a similar two-component trend is also observed for the 4 μm band. For the region of panel B, we observe a less sharp evolution between the two narrow 3 μm band populations, as the transition between the two regimes occurs from $\sim 62^\circ\text{N}$ to $\sim 70^\circ\text{N}$.

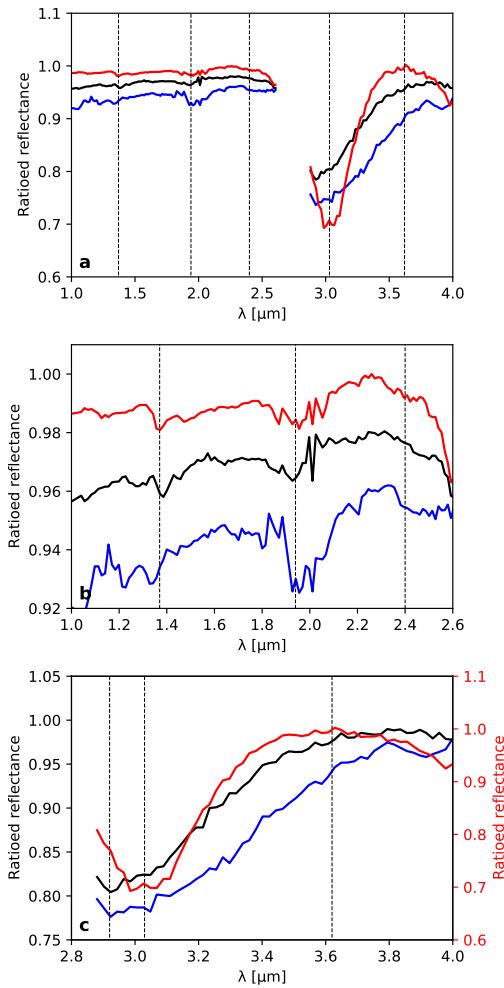


Figure 6: (a) A typical averaged spectral ratio corresponding to the 3 μm northern ring identified in this study (red) is compared to a ratio obtained outside the ring between 50°N and 60°N (black), and to a ratio from southern bright terrains between 62°S and 74°S (blue). Black dotted lines correspond to wavelengths of 1.37, 1.94, 2.4, 3.03, and 3.62 μm . The red spectrum is the average of 4 ratioed spectra taken along the narrow 3 μm northern ring (longitudes $\sim 15^\circ\text{E} / 135^\circ\text{E} / 240^\circ\text{E} / 265^\circ\text{E}$). Each one is the ratio of the average of 200 spectra in a bright region at latitude $\sim 70^\circ\text{N}$ over a similar region in albedo at the same longitude and at latitude $\sim 63^\circ\text{N}$, then normalized with the reflectance at 2.26 μm before averaging ratios from different longitudes. The black spectrum is the average of 3 spectra acquired at different longitudes ($\sim 15^\circ\text{E} / 240^\circ\text{E} / 265^\circ\text{E}$). Each one is the ratio of 200 pixels from a bright region at $\sim 60^\circ\text{N}$ over a region with a similar albedo at the same longitude at $\sim 50^\circ\text{N}$. The blue spectrum is the average of 2 spectra acquired at different longitudes ($\sim 140^\circ\text{E} / 205^\circ\text{E}$). Each one is the ratio of 200 pixels from a bright region at $\sim 74^\circ\text{N}$ over a region with a similar albedo at the same longitude at $\sim 62^\circ\text{N}$. (b) & (c) Zooms on the signatures of the C and L channels respectively. On panels (a) & (b) the blue and black spectra have been offset for clarity. On the panel (c) the y-scale of the red spectrum has been adjusted independently for clarity.

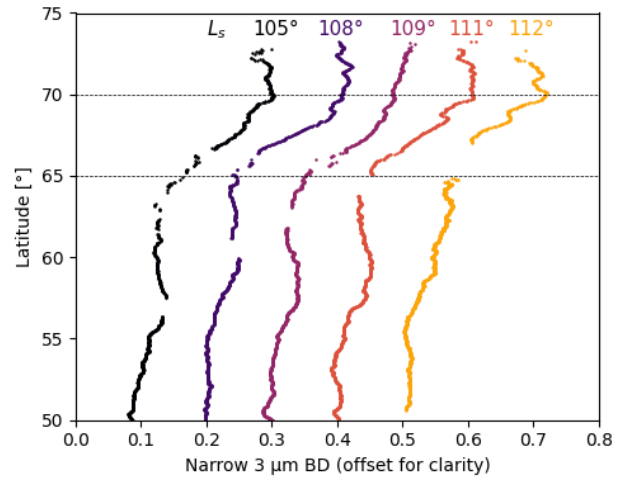


Figure 7: Profiles of the latitudinal evolution of the narrow 3 μm BD (offset for clarity) from several observations of the Phoenix landing site region at different L_s (orbits 0979, 1001, 1012, 1023 & 1034). Each point is averaged in longitude between 235°E and 240°E (image lines), considering only the pixels with a surface reflectance at 2.26 μm between 0.3 and 0.35 to avoid albedo effects in the latitudinal trends of the band depth estimators. Profiles are then smoothed using a moving average.

decreases by a factor of 2-3 from the summer solstice to mid-summer to reach a low value of about 0.3 for which we do not expect any major impact on the 3 μm band (Audouard et al., 2014). Additionally, we do not observe any major modification of the detected spectral feature over that period of changing optical depth (Figure 7). Overall, it thus seems highly unlikely that the observed spectral feature over the ring can be explained by atmospheric dust.

While water vapor and atmospheric dust can be confidently discarded, atmospheric water ice can create a strong and relatively thin spectral absorption near 3 μm (Vincendon et al., 2011; Stcherbinine et al., 2020). A particular attention was therefore paid to remove water ice clouds through the combination of the 1.5 μm BD and ICI water ice criteria (see subsection 2.2). Nevertheless, as the 3 μm band is stronger than the 1.5 μm , some thin water ice clouds modifying the 3 μm band could remain unfiltered. We illustrate this possible effect on Figure 8 with a typical water ice cloud spectral ratio derived on Mars with CRISM data (Vincendon et al., 2011). We can see that a 3 μm atmospheric water ice signature with a significant depth comes with other fainter features at 1.5 μm and 2 μm that are not observed over the 3 μm northern ring (Figure 6). In addition, the 3 μm absorption feature due to water ice clouds differs from the observed narrow 3 μm band, as ice absorption peaks at $\sim 3.1 - 3.2 \mu\text{m}$. Water ice clouds with smaller particle sizes ($r_{\text{eff}} \sim 0.1 \mu\text{m}$), such as those observed frequently at high altitude (Stcherbinine et al., 2020; Luginin et al., 2020; Liuzzi et al., 2020), may produce a narrower 3 μm absorption band without significant signatures at 1.5 or 1.9 μm (Vincendon et al., 2011,

figure 4). However, such clouds do not have a 4 μm absorption feature and their 3 μm absorption is centered on $\sim 3.1 \mu\text{m}$ instead of $3.03 \mu\text{m}$. Overall, observed spectral properties are thus not easily explained by atmospheric water ice. Additionally, as mentioned in subsection 3.5, the spatial distribution and depth of the narrow 3 μm signature are stable over several mid-afternoon observations taken a few degrees of L_s apart (cf Figure 7) and do not show major variation between morning (~ 9 a.m.) and evening (~ 6 p.m.) observations (subsection 3.5), while water ice clouds could be expected to vary (Szantai et al., 2021).

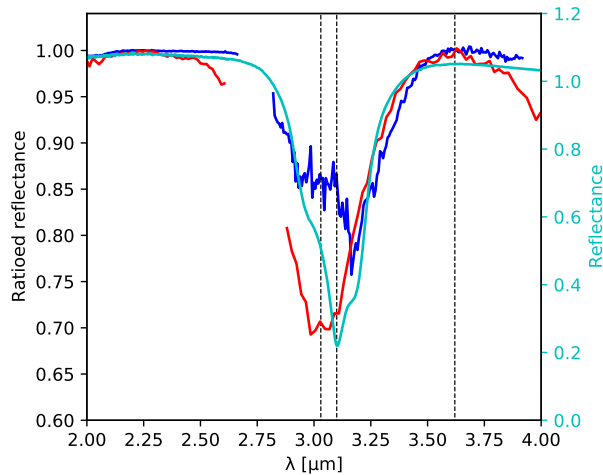


Figure 8: Comparison between the typical spectral ratio associated with the 3 μm northern ring identified in this study (red, see Figure 6 for details) and typical small-grained ice transmission spectra. A spectral ratio of a Martian water ice cloud obtained with CRISM is shown in dark blue (adapted from Vincendon et al. (2011)), and a spectrum of a 0.74 μm thick film of crystalline water ice measured under a temperature of 145K (cyan) from Trotta (1996), available on SSHADE (Schmitt et al., 2018; Schmitt and Trotta, 1996), is shown in light blue. We observe that the position of the absorption band differs between the spectra: about $3.03 \mu\text{m}$ for the narrow 3 μm BD, and between 3.1 and $3.2 \mu\text{m}$ for the small-grained atmospheric or surface water ice. In addition, such small-grained water ice with no weak signatures at 1.5 and $2 \mu\text{m}$ (see also Vincendon et al. (2011)) does not present any $4 \mu\text{m}$ absorption either, unlike our red spectrum. Black dotted lines correspond to wavelengths of 3.03 , 3.1 , and $3.62 \mu\text{m}$.

4.2. Exposed surface water ice

The discovery of a northern polar ring, reported here, is only observed poleward of a clear boundary located at $\sim 68^\circ\text{N}$. Its additional 3 μm band is frequently linked with water and presents some first-order similarities with water ice (Figure 6). As atmospheric water ice has been discarded previously, it is worth investigating the potential implication of surface water ice.

At such latitudes, the perennial subsurface ice of the permafrost is close to the surface (Bandfield and Feldman, 2008) and was exposed by the Phoenix robotic arm did (Arvid-

son et al., 2009) after digging the ground. However, the permafrost distribution in the polar regions (Bandfield and Feldman, 2008, Figure 8) does not match the area detected here, especially between $\sim 200^\circ\text{E}$ and 270°E , where the permafrost remains very close to the surface down to 50°N , while the spectral transition toward the ring-shaped area occurs around 68°N (see subsection 3.1). Actually, water ice buried a few cm below the surface would be missed by OMEGA as it can only sense the first few micrometers. When some subsurface water ice was exposed in a trench dug by the Phoenix robotic arm, observations showed that it quickly sublimated (Arvidson et al., 2009), so permafrost is unstable at the surface. However, local exposure of pergelisol outcrops have been reported in the northern mid-latitudes on pole-facing slopes (Dundas et al., 2018; Harish et al., 2020; Dundas et al., 2021). These are subpixel compared to OMEGA resolution, and we could expect such subsurface ice exposures to be more frequent in the northern high latitudes. Dundas et al. (2021) however shows that it is not the case for the type of exposure they analyzed. Nevertheless, we have tried to simulate the spectral signature that OMEGA would have observed by looking at such subpixel permafrost exposures (Figure 9). We simulated the impact of a spatial mix between an observed spectrum of large-grained perennial water ice (particles sizes $\sim 1 - 10 \mu\text{m}$) and a typical ice-free high latitude spectrum from Figure 6. We observe that even for a few percent of water ice coverage, the spectral ratios exhibit $1.5 \mu\text{m}$ and $2 \mu\text{m}$ absorption bands with intensities comparable to the 3 μm absorption, along with a bluer spectral slope in the $1 - 4 \mu\text{m}$ range. However, neither the blue spectral slope nor sufficient $1.5 \mu\text{m}$ or $2 \mu\text{m}$ absorption regarding the intensity of the 3 μm feature are present in the 3 μm northern ring. Overall, these spatial and spectral considerations contradict a subpixel permafrost exposures explanation.

Small-grained seasonal surface water ice could also persist throughout summer, e.g. in some subpixel shadowed areas. Seasonal water ice is known to be present and stable over shadowed pole-facing slopes earlier and later in the season compared to the seasonal cap (Carrozzo et al., 2009; Vincendon et al., 2010). Actually seasonal water ice has been observed in the Phoenix region during the northern summer (Seelos et al., 2008; Cull et al., 2010). Additionally, the presence of water ice frost that forms during nighttime and remains in the early morning has also been observed by the Viking Lander 2 Camera 2, the Phoenix lander Surface Stereo Imager (SSI) and CRISM observations of the region from $L_s \sim 109^\circ$ (Cull et al., 2010). Up to $L_s \sim 109^\circ$ the Phoenix landing site exposed surface is however ice-free throughout the day (Cull et al., 2010). Our observations were performed in the mid-afternoon (~ 3 p.m.) from $L_s 98^\circ$ to 137° , and no surface water ice have been observed during the afternoon hours under Phoenix latitudes between $L_s \sim 55^\circ$ and $L_s \sim 156^\circ$ (Cull et al., 2010), except well-known perennial patches of water ice in craters. Additionally, observations of water ice or frost in the shadows of pole-facing slopes or craters with CRISM or OMEGA

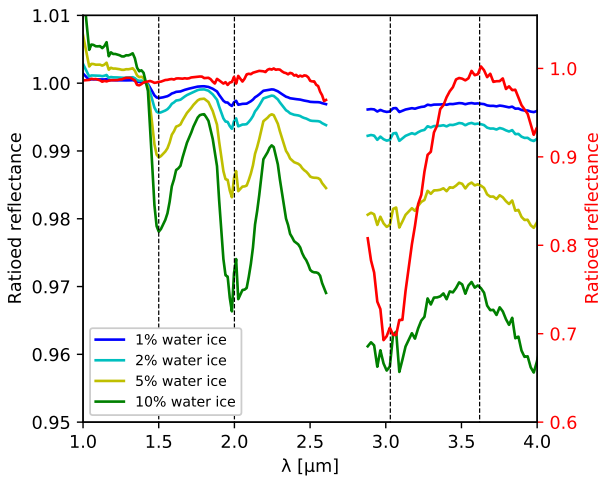


Figure 9: Ratios of spectra simulated with a spatial mix between an OMEGA water ice spectrum and a reference spectrum below the 3 μm northern ring ($\sim 63^\circ\text{N}$ / 241°E) for various amount of water ice, over the same reference spectrum, compared to the typical spectrum of the narrow 3 μm band (red, see Figure 6 for details). The use water ice spectrum is an average of 9 spectra of the edge of the North polar cap (cube 0965_2, $\sim 83^\circ\text{N}$ / 133°E). We observe that even for small amounts of surface water ice within the observed region (1 – 2%), noticeable 1.5 μm and 2 μm absorption band appears in the spectral ratio, along with a blueish spectral slope. Black dotted lines correspond to wavelengths of 1.5, 2.0, 3.03, and 3.62 μm .

(Carrozzo et al., 2009; Vincendon et al., 2010; Harish et al., 2020) exhibit a 1.5 μm absorption of a few percent depth along with a stronger 2 μm absorption feature, while we do not detect these features. Very small water ice particles are needed to produce a 3 μm band without detectable feature at $\lambda = 1.5$ μm : an optical path length of only 4 μm within ice is supposed to be sufficient for OMEGA to detect ice at $\lambda = 1.5$ μm according to Vincendon et al. (2010). We show in Figure 8 the spectral properties of a 0.74 μm -thick surface water ice measured at 145K from Schmitt and Trotta (1996) is compared to the spectral signature identified here. We observe that even if the width of the two 3 μm absorption bands are similar, there is a mismatch between their respective positions with a peak at 3.03 μm compared to 3.13 μm for the water ice signature. The ice spectrum also lacks the 4 μm absorption observed in OMEGA data.

Thus, several issues challenge our attempts to explain the spectral feature of the 3 μm northern ring surface ice, either as large- or small-grained.

4.3. Adsorbed water

Water adsorption refers to weak bonds between water molecules and a mineral surface, that can occur (Milliken et al., 2007) through exchanges with atmospheric vapor (Jouglet et al., 2007). As discussed in section 1, it has been suggested that adsorbed water may be the main explanation for the high latitude increase of the 3 μm band (Jouglet et al.,

2007; Poulet et al., 2008, 2010), although this interpretation has been questioned (Audouard et al., 2014).

The increase of the 3 μm spectral feature related to the water adsorption process comes generally along with an increase of the 1.9 μm band ((Pommerol et al., 2009, Figure 16) and Beck et al. (2015)). However, we observe that the narrow 3 μm and the 4 μm bands identified in this study over the 3 μm northern ring are anti-correlated with the areas where the 1.9 μm band is high (Figure 3). Additionally, the spectral ratio typical of the detected area does not contain a significant 1.9 μm band (Figure 6), although the 1.8 – 2.1 μm range may contain a shallow absorption (< 1%). This does not favor the adsorbed water hypothesis.

We compare in Figure 10 the spectral ratio of the 3 μm northern ring to spectral ratios of a Volcanic tuff sample with different levels of hydration from Pommerol et al. (2009). This sample was chosen because it does not present a strong absorption at 1.9 μm , but a variation of the 3 μm band. We have however checked that the spectral behavior of adsorbed water estimated with this sample was typical of other samples presented in (Pommerol et al., 2009). We can see that the adsorbed water spectral shape does not compare well with the spectral ratio of the 3 μm northern ring. Indeed, we can see that the 3 μm band associated with water adsorption is significantly wider than the narrow 3 μm absorption and that the adsorption spectra are unable to reproduce the decrease of reflectance between 3.7 and 4 μm . The 3 μm northern ring spectral signature thus seems to be related to another process, and adsorbed water does not seem to be the main explanation for the 3 μm band increase observed poleward 68°N .

On the other hand, we observe that the spectral ratio obtained southward of the 3 μm northern ring, that is between 50°N and 60°N , provides a reasonable 3 μm shape match with adsorbed water (Figure 10), along with a small 1.9 μm band that is also consistent with adsorbed water (Figure 6). The observed spectral evolution between 50°N and 60°N may thus reveal an increase in the amount of adsorbed water south of the ring. However, the observed minimum of the 3 μm band is at ~ 2.9 μm in the observed ratio, which is not comparable to adsorbed water. Additionally, as discussed previously (section 1), adsorbed water is difficult to reconcile with the temperatures required to release the water from the soil measured by Phoenix (Smith et al., 2009) and the lack of temporal variability of the 3 μm band (Audouard et al., 2014). The fact that such an increase of adsorbed water would not be observed at the higher latitudes of the ring is also puzzling. Alternatively, an increasing amount of hydrated minerals with latitude would also result in such a spectral ratio with 3 μm and 1.9 μm bands of comparable shapes (Jouglet et al., 2007; Beck et al., 2015). The spectral ratio obtained southward of the 3 μm northern ring may thus indicate the presence of an enhanced hydration within minerals or amorphous phases south of the ring.

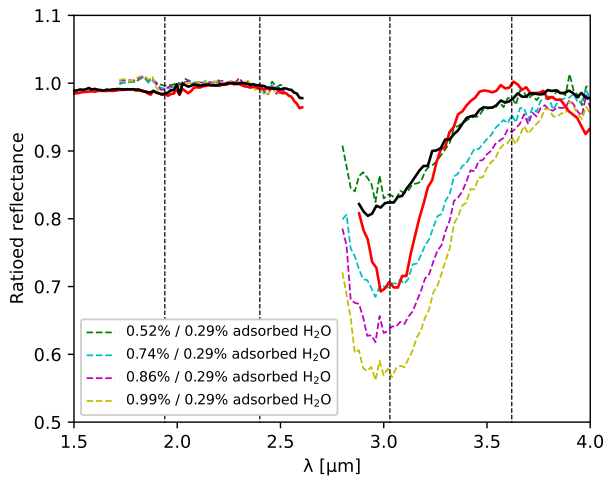


Figure 10: Comparison between high latitude OMEGA spectral ratios and ratios of laboratory materials with different amount of adsorbed water. The spectral ratio identified in this study (red solid line) and the ratio obtained outside the 3 μm northern ring between 50°N and 60°N (black solid line) (see Figure 6 for details) are compared to colored dotted spectra corresponding to ratios of spectra acquired on the same Volcanic tuff sample at different amounts of adsorbed H₂O from Pommerol et al. (2009) and available on SSHADE (Schmitt et al., 2018; Pommerol and Schmitt, 2007). Black dotted lines corresponds to wavelengths of 1.94, 2.4, 3.03, and 3.62 μm .

4.4. Surface physico-chemistry

As discussed in previous sections, the spectral signature identified over the 3 μm northern ring does not seem to be associated with atmospheric, icy or adsorbed water phenomena. Consequently, it must relate to some sort of modification of the physical or chemical properties of the surface.

Studies of the mineralogy of the north polar regions have revealed that the highest amounts of hydrated minerals detectable in the near-IR seem to be located in the low-albedo regions surrounding the perennial polar cap (Langevin et al., 2005a; Poulet et al., 2008; Massé et al., 2010, 2012). This area also presents a high density of dune fields (Tanaka and Hayward, 2008; Massé et al., 2012). These polar hydrated minerals have been mostly identified as Ca-sulfates, and more specifically gypsum (Langevin et al., 2005a; Horgan et al., 2009; Massé et al., 2010, 2012). The highest concentration is located in the Olympia Undae deposit (Langevin et al., 2005a). Perchlorates are an alternative candidate for the spectral shapes typical of these latitudes (Massé et al., 2010; Hanley et al., 2015). While these low albedo terrains are not included in the area of the 3 μm northern ring (see previous discussion), this indicates that sulfates, in particular Ca-sulfates, and possibly perchlorates, are abundant and widespread in the northern polar regions. The Phoenix lander is on the contrary located within the 3 μm northern ring. Phoenix also provided indication that sulfates (notably Mg-sulfates) and, with lower concentrations, perchlorates, participate in the composition of high altitude soils with typical concentrations of a few weight % (Boynton et al., 2009; Hecht et al.,

2009; Smith et al., 2009; Toner et al., 2014).

The observed 4 μm absorption in our OMEGA spectra shows some similarities with the bluer wing of the wide 4.4 – 4.7 μm absorption of sulfates due to SO₄ vibrations harmonic (Bishop et al., 2014; Bishop, 2019). The lack of any significant signature in the 1 – 2.5 μm spectral range may then suggest a weakly-hydrated sulfate (Bishop et al., 2014). Such sulfates are known to be present on Mars since a small amount of anhydrite has been detected in Gale Crater by MSL (Bish et al., 2013). We compare the spectrum identified in this study over the 3 μm northern ring to such lowly hydrated Ca-sulfates in Figure 11. The 4 μm band shape, the relative depth of the 4 and 3 μm bands, and the narrow 3 μm band width can be reproduced by the anhydrite sample. However, the 3 μm band minimum in this anhydrite sulfate sample is not located at the same wavelength as in our observed spectrum. The spectrum of the 3 μm northern ring is also compared to another anhydrite sample that contains minor admixtures (Bishop et al., 2014): this lowers the overall quality of the fit, but we can see that a subtle contamination of this sulfate by other mineral components shifts the minimum of the 3 μm band to longer wavelengths, bringing it closer to our observation.

Studies of the distribution of gypsum across the north circumpolar dune field showed evidence for eolian processes that may have scattered the gypsum away from a few dune field reservoirs (Horgan et al., 2009; Massé et al., 2012). The smallest grains could have been transported further away, and could have accumulated in the bright area, as bright areas on Mars usually correspond to areas where erosion of fines and dust is less efficient. Exposed to southern, warmer conditions, gypsum grains could have undergone desiccation toward bassanite and then even to anhydrite (Poitras et al., 2018). Alternatively, such desiccation could have occurred at the source and could have then facilitated fracturing, producing smaller and wind transportable grains. These desiccated sulfate grains could have been transported further away by winds compared to larger gypsum grains.

We compare in Figure 12 two gypsum spectra obtained for two different grain sizes. While the gypsum spectrum in itself provides a poor fit to our observed spectrum, the ratio of the small grain size over the larger grain size provides a relatively satisfactory explanation for the narrow 3 μm band of the ring. This may indicate that reducing the grain size of calcium-sulfates can produce the spectral evolution observed at the 3 μm northern ring. The 4 μm feature is however absent of the gypsum ratio of two sizes, so reducing the size does not seem to be enough. Additionally, this interpretation would require calcium sulfates with larger grains to be present southward of the identified area, as our spectrum is a ratio between 70°N and 63°N, while observational evidence for gypsum have been obtained essentially poleward. Southward, as discussed in subsection 4.3, terrains located between $\sim 50^\circ\text{N}$ and $\sim 65^\circ\text{N}$ do show enhanced hydration spectral features at 1.9 μm and 3 μm that may be related to hydrated phases such as hydrated minerals. This could indicate that soils south of the 3 μm northern ring actually con-

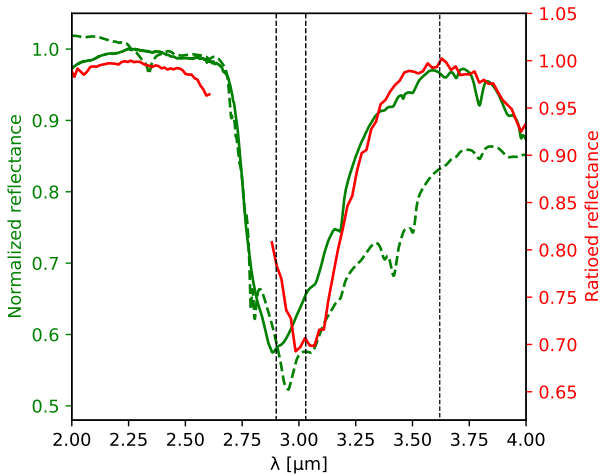


Figure 11: Comparison between a typical spectrum of narrow 3 μm band (red, see Figure 6 for details) and normalized spectra of Anhydrite (green, solid line GDS42, dashed line JB641A), a Ca-sulfate, RELAB spectra downloaded from the USGS spectral library (Kokaly et al., 2017) and PDS spectral library. We observe that the 3 μm band of the GDS42 anhydrite (solid line) peak at $\sim 2.9 \mu\text{m}$, which do not match with the narrow 3 μm band of the red spectrum. However, this Ca-sulfate can reproduce the 4 μm absorption of our typical signature, without any absorption in the 2 μm spectral range. The other anhydrite sample (dashed line) presents a wider 3 μm absorption band tightly shifted to redder wavelength but without 4 μm feature. If the position of the 3 μm absorption is here centered on $\lambda \sim 3.0 \mu\text{m}$, so closer to the observed narrow 3 μm band than the GDS42 sample, the wider shape of the band and the lack of 4 μm absorption do not match with the observed spectral signature of the northern regions. Black dotted lines correspond to wavelengths of 2.9, 3.03, and 3.62 μm .

tain more altered phases compared to equatorial latitudes, as previously proposed (see section 1). The 3 μm northern ring would then correspond to a latitude transition beyond which these phases are modified.

While a sole modification of the grain size does not seem to explain all observed features (notably, the 4 μm band, see previous paragraph), one could imagine that, if sulfates are already present south of the 3 μm northern ring, increasing the latitude can modify the hydration level of sulfates and produce the observed spectral features in ratios. To test this hypothesis, we compare in Figure 12b our observation to a ratio of two samples of jarosite - a hydrous sulfate of potassium and ferric iron - from two different locations on Earth. Both are described as jarosite mixture with similar grain sizes lower than 45 μm . One of the two samples appears more hydrated from the 1.9 μm band depth. Ratioing both results in a spectral shape that presents some similarities with our observation, with both a narrow 3 μm band centered at the appropriate wavelength and a decrease toward 4 μm . This example illustrates that increasing the amount of hydration of sulfates may result in a spectral ratio consistent

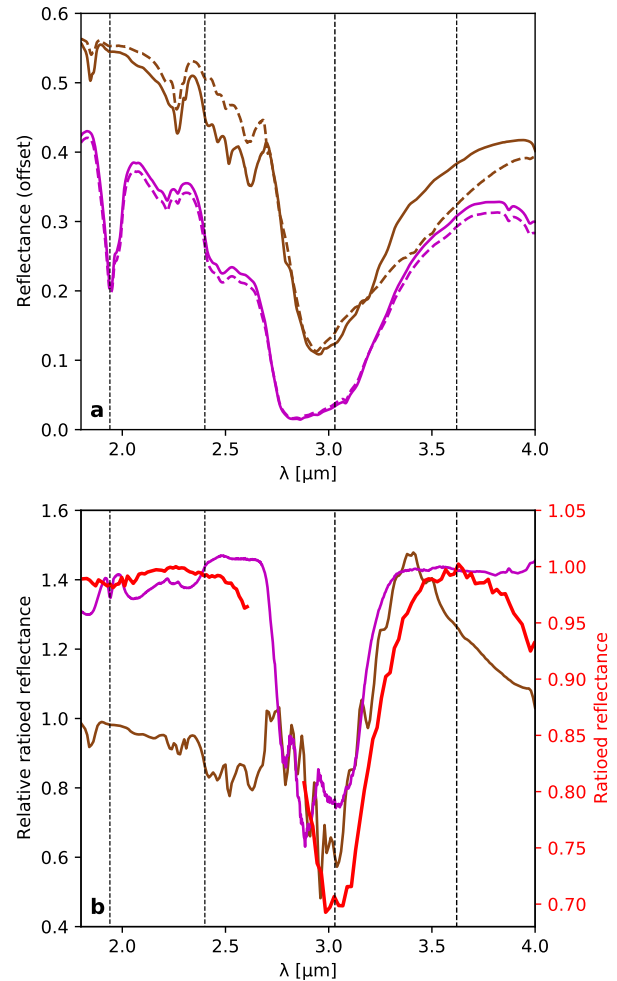


Figure 12: (a) Spectra of Jarosite (brown, offset of 0.1 for clarity) from two different places: Rublev (solid line, JBA78) and St Leger (dashed line, JB701); and spectra of Gypsum (purple, JBE64) for two grain size: < 45 μm (solid line) and 45–90 μm (dashed line). RELAB spectra downloaded from the PDS spectral library. (b) Ratio of the jarosite from Rublev over the one from St Leger (brown) and of gypsum < 45 μm sample over 45–90 μm sample (purple, rescaled by a factor 4 for clarity) compared to the narrow 3 μm band spectrum (red, see Figure 6 for details). We observe that the ratio of the two samples of jarosite exhibits a narrow 3 μm band along with an absorption at 4 μm . If it does not provide a good match for the entire spectrum, it shows that under some conditions a mineral can experiment alteration that results in the apparition of a narrow 3 μm band. Similarly, the ratio of the two spectra of gypsum also seems to exhibit a 3 μm absorption feature that can be compared to the narrow 3 μm band, but without any 4 μm absorption. Black dotted lines correspond to wavelengths of 1.94, 2.4, 3.03, and 3.62 μm .

with our observations.

Overall, while the exact mechanism responsible for the spectral signature of the ring is not firmly identified, available constraints point toward a possible contribution of sulfates in a context of hydration state change and/or grain size

change. As our observed 3 μm northern ring is superficial and overlays different geological units dated from early-Amazonian (cf subsection 3.2), we are probably observing a relatively recent process. This process may be currently acting and somehow related to the seasonal ice cap (Audouard et al., 2014) that cover the area every winter, although this association is not clear as the cap extent at southern latitudes, down to near 50°N (Appéré et al., 2011). Alternatively, this modification of the surface may be related to a longer timescale. Mars high latitudes are currently covered by a mantle that is thought to have been deposited during past glaciations over the last million years (Head et al., 2003). In both hemispheres, this ice-rich mantle is considered to be present and uniform poleward $\sim 60^\circ$, while being only partially preserved equatorward (Head et al., 2003). Sulfates formation can be associated with ice on Earth and this has been proposed as an explanation for sulfates formation on Mars at high northern latitudes (Massé et al., 2010) or near the equator (Niles and Michalski, 2009), either through authigenic formation within ice or associated with atmospheric ice condensation and precipitation (Massé et al., 2010). Widespread spectral features possibly related to sulfates have been observed at high southern latitudes (Poulet et al., 2008; Wray et al., 2009; Ackiss and Wray, 2014; Carter et al., 2016), and widespread spectral signatures potentially linked with limited water alteration have been observed at high northern latitudes in this study (see subsection 4.3 and previous paragraph) and previously (Wyatt et al., 2004; Michalski et al., 2005; Horgan and Bell, 2012). The 3 μm northern ring identified in this study north of $\sim 65^\circ\text{N}$ may then highlight a modification of such a widespread sulfates contamination. This modification could be a hydration state increase occurring once a latitude threshold is crossed in the northern hemisphere where the amount of water vapor and the overall atmospheric pressure are higher.

5. Conclusion

In this paper, we report and characterize the detection of a new spectral feature defined by a deep and relatively narrow 3 μm band, centered at 3.03 μm , that differs from the usual Mars 3 μm band, coupled with a shallow but significant wide 4 μm feature. This spectral signature appears to be specific to the bright regions located at high northern latitudes. More precisely, the detection occurs over a wide, open ring area surrounding the cap between $\sim 68^\circ\text{N}$ and $\sim 76^\circ\text{N}$ and extending from $\sim 0^\circ\text{E}$ to $\sim 315^\circ\text{E}$. This area includes the Phoenix landing site and is referred as the "3 μm northern ring" in this study.

A detailed study of the spectral properties observed in this ring (3 μm band shape, association with other spectral features, time variability) leads to rule out atmospheric and water-ice explanations. The spectral properties cannot either be explained by an increase in the amount of weakly bound adsorbed water.

The ring appears to possess a specific surface signature that may relate to some sort of modification of the physical

or chemical properties of the ice-free soils. The ring occurs north of terrains where a progressive latitudinal increase of surface hydration (1.9 μm and 3.0 μm bands) is observed. The spectral signature of the ring does however not correlate with signatures in the 1 – 2.5 μm spectral range, despite the presence of specific hydrated signatures at similar latitudes in nearby dark terrains.

First comparisons with spectra from online libraries suggest that sulfates, notably Ca-sulfates, may play a role. Ca-sulfates are known to be widespread in the north polar regions, but they have not been yet detected in this area from orbit or by Phoenix. A change in the level of hydration of sulfates and/or a grain size modification may contribute to the observed signature and its latitudinal distribution. The presence of sulfates over these high-latitudes regions may then highlight ice-related formation pathways of these minerals on Mars.

Acknowledgements

The OMEGA/MEx data are freely available on the ESA PSA at <https://archives.esac.esa.int/psa/#!Table%20View/OMEGA=instrument>. The source code of the *OMEGA-Py* Python module is freely available on GitHub at <https://github.com/ASTcherbinine/omegapy>.

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