



**HAL**  
open science

## **LESSER: a catalogue of spectroscopically selected sample of Lyman- $\alpha$ emitters lensed by galaxies**

Xiaoyue Cao, Ran Li, Yiping Shu, Shude Mao, Jean-Paul Kneib, Liang Gao

► **To cite this version:**

Xiaoyue Cao, Ran Li, Yiping Shu, Shude Mao, Jean-Paul Kneib, et al.. LESSER: a catalogue of spectroscopically selected sample of Lyman- $\alpha$  emitters lensed by galaxies. *Monthly Notices of the Royal Astronomical Society*, 2020, 499, pp.3610-3619. 10.1093/mnras/staa3058 . insu-03667032

**HAL Id: insu-03667032**

**<https://insu.hal.science/insu-03667032>**

Submitted on 21 Jul 2023

**HAL** is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire **HAL**, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.

# LESSER: a catalogue of spectroscopically selected sample of Lyman- $\alpha$ emitters lensed by galaxies

Xiaoyue Cao,<sup>1,2</sup> Ran Li<sup>1,2</sup> ,<sup>1,2</sup>★ Yiping Shu<sup>1,3</sup> ,<sup>3</sup> Shude Mao,<sup>2,4</sup> Jean-Paul Kneib<sup>5,6</sup> and Liang Gao<sup>1,2</sup>

<sup>1</sup>*School of Astronomy and Space Science, University of Chinese Academy of Sciences, Beijing 100049, China*

<sup>2</sup>*National Astronomical Observatories, Chinese Academy of Sciences, 20A Datun Road, Chaoyang District, Beijing 100012, China*

<sup>3</sup>*Institute of Astronomy, University of Cambridge, Madingley Rd, Cambridge, CB3 0HA, UK*

<sup>4</sup>*Department of Astronomy, Tsinghua University, Beijing 100084, China*

<sup>5</sup>*Institute of Physics, Laboratory of Astrophysics, Ecole Polytechnique Fédérale de Lausanne (EPFL), Observatoire de Sauverny, CH-1290 Versoix, Switzerland*

<sup>6</sup>*Aix Marseille Université, CNRS, LAM (Laboratoire d'Astrophysique de Marseille) UMR 7326, F-13388 Marseille, France*

Accepted 2020 September 30. Received 2020 September 29; in original form 2020 July 27

## ABSTRACT

We introduce the LEnSed laeS in the EBOSS suRvey (LESSER) project, which aims to search for lensed Lyman- $\alpha$  emitters (LAEs) in the Extended Baryon Oscillation Spectroscopic Survey (eBOSS). The final catalogue contains 361 candidate lensing systems. The lens galaxies are luminous red galaxies (LRGs) at redshift  $0.4 < z < 0.8$ , and the source galaxies are LAEs at redshift  $2 < z < 3$ . The spectral resolution of eBOSS ( $\sim 2000$ ) allows us to further identify the fine structures of Lyman- $\alpha$  ( $\text{Ly}\alpha$ ) emissions. Among our lensed LAE candidates, 281 systems present single-peaked line profiles while 80 systems show double-peaked features. Future spectroscopic/imaging follow-up observations of the catalogue may shed light on the origin of diverse  $\text{Ly}\alpha$  line morphology, and provide promising labs for studying low-mass dark matter haloes/subhaloes.

**Key words:** gravitational lensing: strong – techniques: spectroscopic – galaxies: elliptical and lenticular, cD – dark matter.

## 1 INTRODUCTION

The so-called spectroscopic method is a very effective way to search for galaxy-scale strong lensing systems (Warren et al. 1996; Willis, Hewett & Warren 2005; Willis et al. 2006; Bolton et al. 2004). The basic idea is to scan the spectra of foreground galaxies for abnormal emission lines that actually come from background source galaxies. Detection of such ‘abnormal’ emission lines provides an evidence that two galaxies are nearly aligned along the line of sight, thus are likely a strong lensing system. Based on the spectroscopic method, the Sloan Lens ACS Survey (SLACS, Bolton et al. 2008), the Baryon Oscillation Spectroscopic Survey (BOSS) Emission-Line Lens Survey (BELLS, Brownstein et al. 2012), and the SLACS for the Masses Survey (S4TM, Shu et al. 2015, 2017) have achieved many successes, extending the number of galaxy–galaxy strong lensing systems to  $\sim 200$ , and providing valuable constraints on the galaxy formation theory and cosmological models (Treu et al. 2006; Koopmans et al. 2006, 2009; Gavazzi et al. 2007; Jiang & Kochanek 2007; Bolton et al. 2012b; Chen et al. 2019). Recently, Cheng (2020) finds another  $\sim 100$  possible lensed [O II] emitters through the spectroscopic method by searching the full BOSS data base.

Shu et al. (2016a) initiated the BELLS for the GALaxy-Ly $\alpha$  EmitteR sYstems Survey (the BELLS GALLERY Survey, hereafter) project, which aims at detecting galaxy-scale gravitational lenses with high-redshift Lyman- $\alpha$  emitters (LAEs) as the sources. They find 187 possible galaxy-LAE systems from millions of spectra

in the BOSS data base (Alam et al. 2015), 21 most promising candidates among them have follow-up imaging observations with the *Hubble Space Telescope* (*HST*), and 17 candidates are confirmed as real lensing systems (Shu et al. 2016b). With the help of lensing magnification, these lensed LAEs provide information on the size and morphology of sub-kiloparsec scales clumps in LAEs (Shu et al. 2016b; Ritondale et al. 2019b; Cornachione et al. 2018), and a promising lab for constraining the identity of dark matter (Li et al. 2016, 2017; Ritondale et al. 2019a).

Encouraged by the success of the BELLS GALLERY survey, we apply the similar searching technique to the Extended Baryon Oscillation Spectroscopic Survey (eBOSS, Dawson et al. 2016). The eBOSS survey is part of the Sloan Digital Sky Survey-IV (SDSS-IV) project (Blanton et al. 2017), which started in 2014 using the 2.5-m Sloan telescope (Gunn et al. 2006) at the Apache Point Observatory in New Mexico. The main difference between the method in this work and that in Shu et al. (2016a) is the treatment to the Ly $\alpha$ -emission profile. Shu et al. (2016a) did not classify the line profile of lensed LAEs, although the spectroscopic resolution of BOSS ( $\sim 2000$ ) provides such ability. In this work, we explicitly classify our lensed LAE candidates into single- and double-peaked types.

Previous spectroscopic observations have shown that those double-peaked LAEs are quite common at  $z \sim 3$  (Yamada et al. 2012; Tapken et al. 2007). Theoretically, internal gas dynamics of the LAE with double-peaked Ly $\alpha$  emission is different from the single-peaked one (Neufeld 1990; Ahn 2004; Dijkstra, Haiman & Spaans 2006; Orsi, Lacey & Baugh 2012; Gronke, Bull & Dijkstra 2015; Gronke & Dijkstra 2016; Yang et al. 2017). Therefore, a complete sample that includes both types of lines can provide a more comprehensive picture of the LAEs.

\* E-mail: ranl@bao.ac.cn

Hereafter, we refer to this project as ‘LESSER’ (LEnSed laeS in the EBOSS suRvey). This paper is organized as follows. We introduce the spectroscopic identification pipeline for lensed LAEs in Section 2. The catalogue of candidates and their statistical properties are shown in Section 3. Finally, in Section 4 we discuss prospects for future applications of our sample.

## 2 THE SEARCHING PIPELINE

The spectroscopic method of identifying lens systems has been well described in previous works (Bolton et al. 2004; Brownstein et al. 2012). Shu et al. (2016a) extended this method to search for lensed LAEs. In this paper, we apply the same technique as that used in Shu et al. (2016a) to the eBOSS data (Ahumada et al. 2019), with modifications to distinguish between single- and double-peaked Ly $\alpha$  emission lines. We briefly describe the pipeline below.

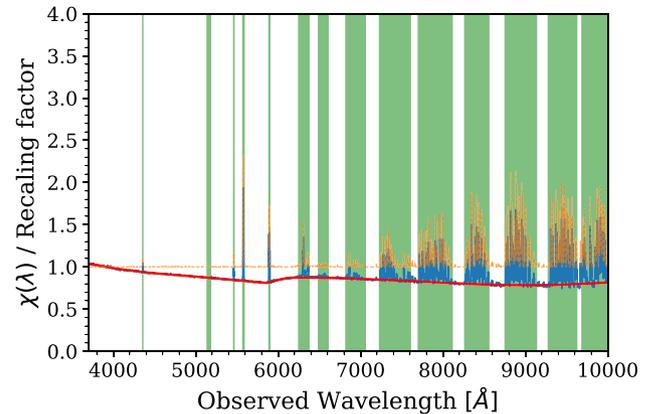
We start with selecting the galaxies tagged as ‘luminous red galaxy’ by the eBOSS pipeline (LRG, Prakash et al. 2016). The spectra of LRGs are optimal places to search for abnormal emissions from background sources for two reasons: first, LRGs are massive early-type elliptical galaxies (ETGs), thus can act as efficient lenses; Second, the spectra of LRGs typically contain few emission lines, and their continuum can be well modelled by the template fitting (Blanton et al. 2017; Bolton et al. 2012c). There are 1571129 LRG spectra selected in this step.

For each LRG spectrum, we mask the regions that correspond to the empirically determined emission, absorption, and sky-line features.<sup>1</sup> We use the galaxy template generated by eBOSS pipeline (Bolton et al. 2012c; Blanton et al. 2017) to fit the masked spectrum, then we subtract the best-fitting LRG spectra from the original LRG spectrum. The residual reflects the ‘features’ that cannot be captured by the galaxy template of LRG, which may contain ‘abnormal’ emission features from another background object.

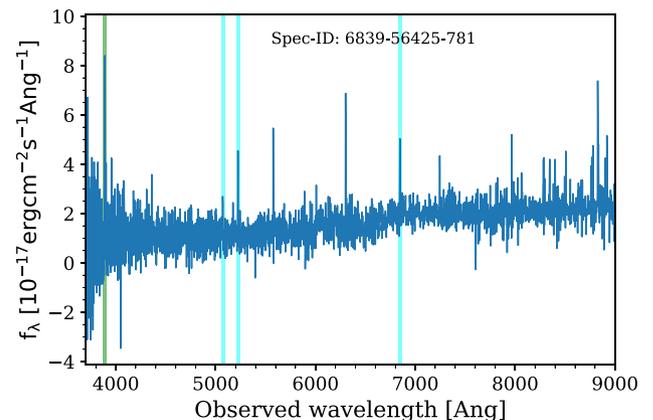
To detect the emission from background source galaxies, we apply an error-weighted match filter on the residual spectrum with a Gaussian kernel, the velocity dispersion of which is set to 150 km s<sup>-1</sup>. We confine our scanning wavelength range to [3700–4800] Å (roughly at redshift  $z \simeq 2$ –3 for Ly $\alpha$  emission), so that the only prominent foreground contamination emission lines is the [O II] 3727 doublet. We retain the detected emissions (‘hit’) with S/N (signal-to-noise ratio) greater than 6. Note that, the skylight subtraction procedure in the eBOSS official pipeline underestimates the noise level at the wavelength location with strong air-glow emission. We correct this effect empirically using sky-subtracted sky-fibre spectra (Brownstein et al. 2012; Shu et al. 2016a). Our noise re-calibration result is shown in Fig. 1. There are 4995 hits detected in this step.

We bin the ‘hits’ by observed wavelength. The overpopulated bins are likely to be related to the sky-line signal, thus hits in these bins are pruned. Similarly, we also prune the hits that are overpopulated at rest-frame wavelength, which are associated with imperfect galaxy template fitting. 2813 hits are removed in this step.

The [O II] 3727 emission from the low-redshift emitters may appear in our searching range (3700–4800 Å) as a contaminator. From the observational point of view, if the [O II] 3727 emission is shown, the H $\alpha$ , [O III] 5007, and H $\beta$  emissions will likely appear simultaneously. Therefore, for each hit, we inspect secondary emission lines at wavelength corresponding to H $\alpha$ , [O III] 5007, and



**Figure 1.** The blue line shows the root mean square (rms) spectrum of the sky-subtracted sky fibres, which empirically indicate the sky-noise level. If the eBOSS pipeline estimates the sky-noise level perfectly, the blue line would be a horizontal line with an ordinate of  $\sim 1$ . The solid blue line is noisy around the green shadow region because of the influence of sky emissions. The red line represents the baseline level of the blue line, which is obtained from B-spline fitting. Divide the blue line by red line; we can estimate how much sky noise is underestimated, as shown by the orange line.



**Figure 2.** This example shows the case where an [O II] emission is misidentified as Ly $\alpha$ . The shaded green marks the location of hit, while shaded cyan regions correspond to the position of H $\alpha$ , [O III] 5007, and H $\beta$  emission assuming hit is [O II] emission instead Ly $\alpha$ . The legend represents SDSS plate-mjd-fibre.

H $\beta$ , assuming that the hit is [O II] 3727. If the quadrature sum of S/N values at these three wavelength locations is higher than  $2.5\sigma$ , the hit is identified as [O II] emission and rejected. An example of the rejected hits is shown in Fig. 2. There are 921 hits removed by this criterion.

Some ‘fake hits’ may still exist in our candidates. We list the source of the spurious emission and our visually removal strategy as follows:

- (i) Visually inspect the results of galaxy template fitting, remove the spurious emission detection due to bad fit.
- (ii) Check the spectra of ‘neighbour fibres,’ which are geometrically near to the fibre with ‘hit’ detection (‘hit-fibre’) in the SDSS plate. If there is a stronger emission appearing in ‘neighbour fibre’ and its wavelength is the same as the ‘hit-fibre.’ Then the hit is considered as a ‘false-positive’ detection induced by neighbouring fibres. Technically, this is called ‘ccd cross-talk’.

<sup>1</sup>We list the wavelength of all those line region in a table that can be found on [https://github.com/caoxiaoyue/search\\_lensed\\_lae/blob/master/paper\\_material/line\\_info.txt](https://github.com/caoxiaoyue/search_lensed_lae/blob/master/paper_material/line_info.txt)

(iii) The outlier induced by cosmic rays can be identified by checking each subexposure spectrum (15 min per exposure). The emission that does not appear in all exposures is due to cosmic rays.

(iv) Visually checking the CCD raw data to ensure the detection is not from CCD defects.

(v) Remove the hits that have a line width less than the spectral resolution of the instrument ( $\sim 70 \text{ km s}^{-1}$ ).

(vi) Remove repetitive candidates that have been independently observed multiple times.

We have 527 hits left after these cuts.

In order to classify our candidates into double and single-peaked types, We first apply a Gaussian match filter with a smaller dispersion of  $75 \text{ km s}^{-1}$  to the spectrum to reduce the noise; then we search the blueward side of the main peak for another subpeak with  $S/N > 2.5$ , candidates with another blueward bump are tagged as ‘double-peaked LAE’.

Previous observation has revealed the  $\text{Ly}\alpha$  emission of LAEs typically shows positive skewness (Yamada et al. 2012; U et al. 2015; Fuller et al. 2020). The asymmetric  $\text{Ly}\alpha$  line profile of LAEs can be well described by the skew-normal function with the formula form in U et al. (2015, see also, Mallery et al. 2012),

$$f(\lambda) = Ae^{-0.5 \times ((\lambda-x)/\omega)^2} \left[ \int_{-\infty}^{s(\lambda-x)/\omega} \exp(-t^2/2) dt \right] + c, \quad (1)$$

where  $s$  is the skewness. The line centre is given by  $\lambda_0 = x + \omega \delta \sqrt{2/\pi}$  and the line width is given by  $\sigma = \omega \sqrt{1 - 2\delta^2/\pi}$  and  $\delta = s/\sqrt{1+s^2}$ . U et al. (2015) systematically investigate the morphology of  $\text{Ly}\alpha$  line profile, based on 304 LAE samples at  $3 < z < 7$  in the Great Observatories Origins Deep Survey and Cosmic Evolution Survey field. They find the  $\text{Ly}\alpha$  emission of LAEs has a median skewness value of  $\sim 1.5$ . To increase the purity of our final catalogue, we apply a cut on the shape of line profile, removing candidates whose main-peak profiles have negative skewness. We have 361 candidates left in our final catalogue.

### 3 CANDIDATE CATALOGUE AND RESULTS

In total, we find 361 lensed LAE candidates from the eBOSS data base, 281 candidates of which are single-peaked, and 80 candidates are double-peaked. The detailed information of our candidates is listed in Tables 1 and 2. Since the eBOSS data base also includes previous BOSS data, we mark candidates that were already found by the BELLS GALLERY project using star symbols in the last column of Tables 1 and 2.

In Figs 3 and 4, we show the spectra in wavelength range  $3700 < \lambda < 4800 \text{ \AA}$  for our single- and double-peaked candidates, respectively. The shadow regions mark the location of the  $\text{Ly}\alpha$  emissions. To show the fine structure of  $\text{Ly}\alpha$  emissions more clearly, we present the zoom-in views around the  $\text{Ly}\alpha$  emission for single- and double-peaked candidates in Figs 5 and 6, respectively. The blue bars in Figs 5 and 6 show the raw residual spectra of our lensed LAE candidates. To quantify the  $\text{Ly}\alpha$  line morphology, we fit the raw-residual spectra with the skew-normal model shown in equation (1), using the Nested-Sampling tool – PYMULTINEST (Buchner et al. 2014), which is a PYTHON wrapper of the FORTRAN code – MULTINEST (Feroz, Hobson & Bridges 2009). The best-fitting values of skewness, width and the total line flux are shown in Tables 1 and 2.

On the left-hand part in Fig. 7, we show the redshift distribution of the lens and the source in our final catalogue. The histogram in the top and right-hand panels show the 1D marginalized probability density function (PDF) for the lens redshift and the source redshift respectively. We perform the Gaussian kernel density estimation

(KDE) for each redshift histogram to derive the smoothed distribution of the redshifts. The lenses in our sample distribute in redshift range  $0.2 < z < 1.0$ , which peaks at  $z = 0.55$ , while the sources distribute in redshift range  $2 < z < 3$  and seems to have double peaks for both single- and double-peaked samples. On average we find that about 20 per cent of our lensed LAE candidates have the double-peaked line profile; the fraction varies significantly with redshift, from  $\sim 45$  per cent at  $z = 2$  to  $\sim 10$  per cent at redshift  $z = 3$ , as shown by the right-hand part in Fig. 7. These values agrees broadly with previous studies for field LAEs, which show the fraction of double-peaked LAEs is around 20–50 per cent (Yamada et al. 2012; Kulas et al. 2012; Trainor et al. 2015; Sobral et al. 2018).

Bolton et al. (2008) empirically show that the probability of a candidate being a real lensing system significantly increases as the predicted Einstein radius becomes larger. To help future follow-up observations, we calculate the predicted Einstein radius ( $\theta_E$ ) of our candidates, assuming the lens follows a singular isothermal ellipsoid model (SIE, Kormann, Schneider & Bartelmann 1994),

$$b_{\text{sie}} = 4\pi\sigma_v^2/c^2 D_{\text{LS}}/D_S, \quad (2)$$

where  $\sigma_v$ ,  $D_{\text{LS}}$ , and  $D_S$  are the velocity dispersion of the lens galaxy, the angular distance between the lens and source, and the angular distance between the observer and source. The  $\theta_E$  for each candidate is also listed in Tables 1 and 2. The blue line in Fig. 8 shows the probability density distribution of  $\theta_E$  for single-peaked candidates, which follows the same distribution as the double-peaked type, which implies single- and double-peaked LAEs are lensed in a similar way. Lensing does not introduce additional selection biases to different types of LAE sources, any statistical difference between different types of LAEs inferred from our lensed samples should also exist in the unlensed case.

We further check whether our candidates already exhibit some lensed features in the images from the DESI Legacy Imaging Surveys (Dey et al. 2019).<sup>2</sup> We classify our candidates into grades A, B, and C by visual inspection. Our classification criteria are as follow:

- (i) Grade A: multiple images or an arc with a counter-image surrounding the main lens is observed. The lensed feature and main lens have different colours. The Einstein radius inferred from the image is smaller than 5 arcsec.
- (ii) Grade B: a secondary image or an arc of a different colour from the lens object is observed, but no obvious counter-image can be observed.
- (iii) Grade C: no discernable lensing feature exists.

Fig. 9 shows all our grade-A candidates. A full atlas that includes all our candidates is available in the online journal. In total, we have 14 grade-A (4 new), 25 grade-B (12 new), and 322 grade-C (213 new) candidates.

### 4 DISCUSSION AND SUMMARY

In this work, we present the LESSER strong lensing catalogue, which is constructed through a systematic search for galaxy-LAE lensing systems from the eBOSS data base. The catalogue including 361 candidate lensing systems. The lens galaxies of these systems are LRGs within the redshift range  $0.4 < z < 0.8$ , and the source galaxies are LAEs at redshift  $2 < z < 3$ . We examine the line profiles of these candidate LAEs and classify them into two subsamples, the ones with single-peaked line profiles (281) and those with double-peaked line profiles (80).

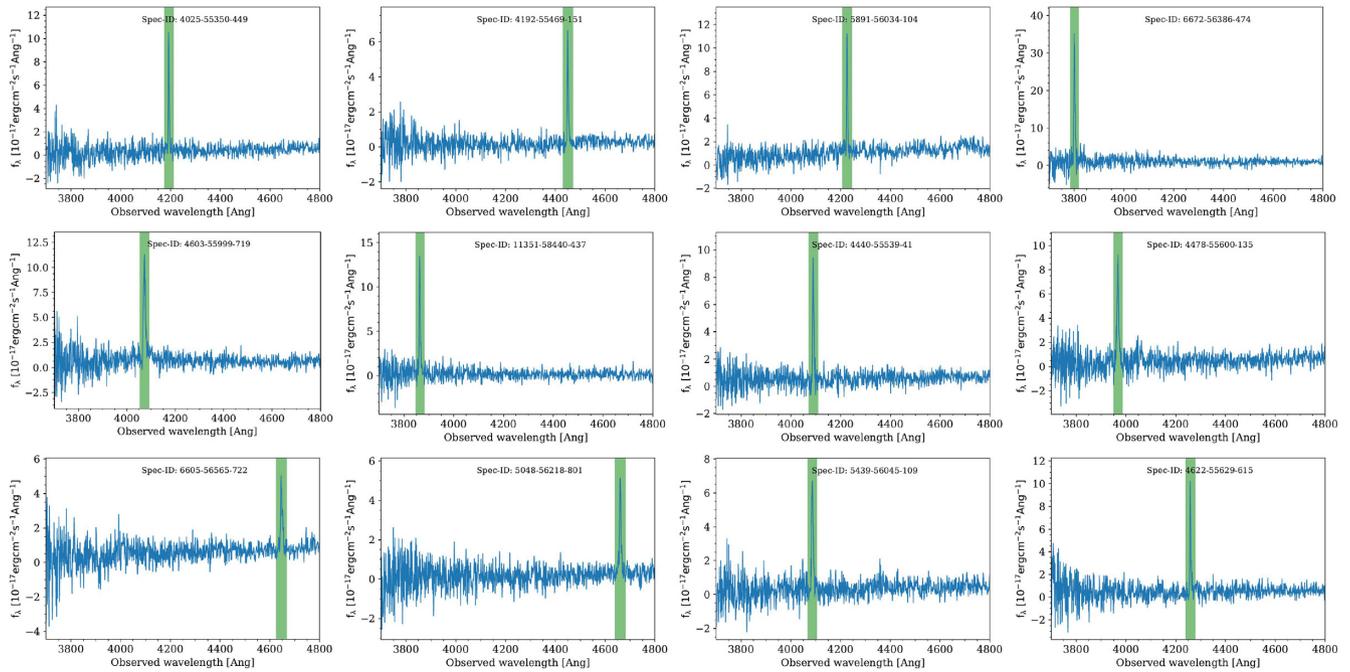
<sup>2</sup>Webpage: <http://legacysurvey.org/>

**Table 1.** Catalogue of single-peaked candidates. From left to right: the SDSS plate-fibre-mjd, right ascension and declination of lens galaxies, redshifts of lens galaxies and background LAEs, *i*-band SDSS cmodel magnitude of lens galaxies, the line flux, width and skewness of Ly $\alpha$  emission quantified by the skew-normal fit, the flux and width are in units of  $10^{-17}$  erg cm $^{-2}$  s $^{-1}$  and  $\text{\AA}$ , respectively. The velocity dispersions of lens galaxies in units of km s $^{-1}$  and associated predicted Einstein radii in units of arcsecond. The last column gives the grade of our candidates based on colour images from the DECaLS survey. Candidates already found by the BELLS GALLERY project are marked with star symbols in the last column. We only list part of the catalogue here for conciseness, a full catalogue is available in the online journal.

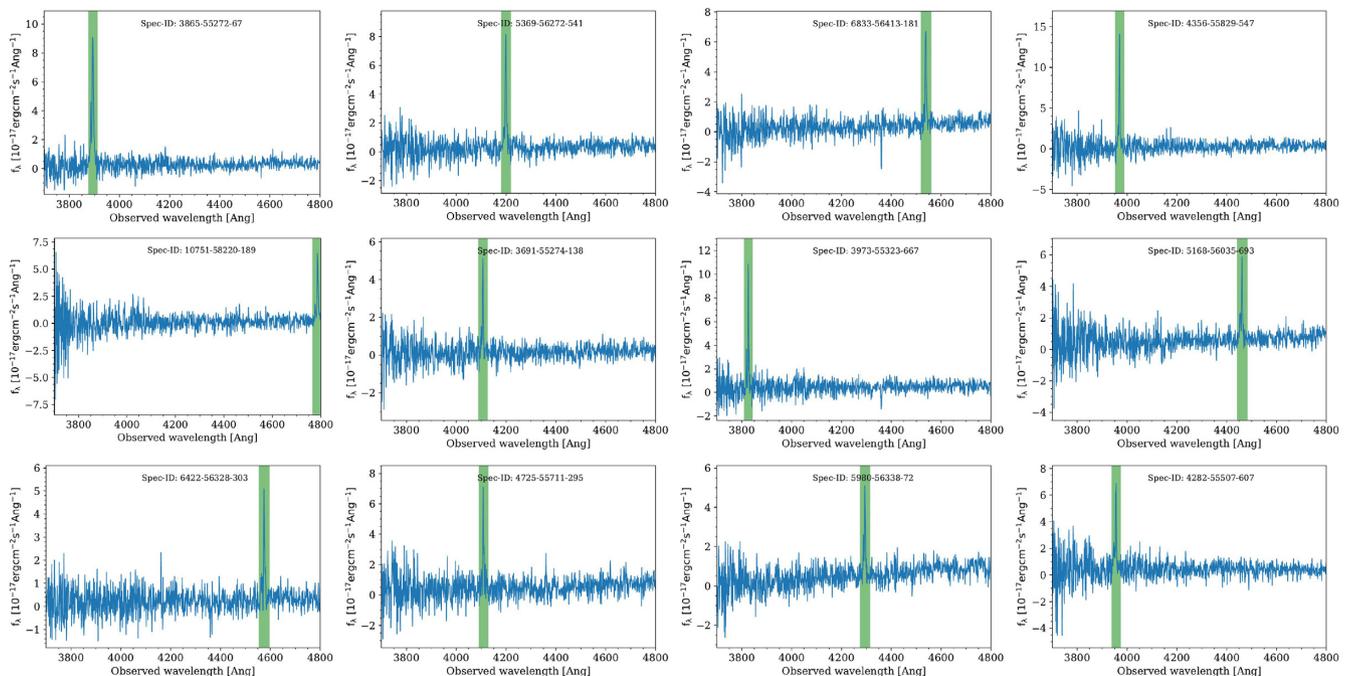
Name	RA	Dec.	$z_l$	$z_s$	mag $_i$	$F_{\text{Ly}\alpha}$	Width	Skewness	$\sigma_{\text{eBOSS}}$	$\theta_E$	Grade
4025-55350-449	14:31:33.30	-00:44:54.17	0.44	2.45	19.87	50.63	2.10	1.97	136.10	0.38	C*
4192-55469-151	21:20:47.64	-00:09:11.33	0.55	2.66	19.96	29.36	2.21	1.60	232.59	1.03	C*
5891-56034-104	11:39:39.00	+16:51:58.52	0.28	2.48	17.89	47.57	2.09	2.82	206.90	1.00	C*
6672-56386-474	12:01:59.03	+47:43:23.19	0.56	2.13	19.30	221.87	3.10	1.45	226.79	0.90	C*
4603-55999-719	09:07:34.40	+42:55:28.56	0.64	2.35	19.93	108.51	5.24	1.69	139.33	0.33	B
11351-58440-437	10:13:23.75	+28:12:00.10	0.71	2.18	20.87	89.00	3.05	1.93	167.30	0.43	C
4440-55539-41	07:42:49.68	+33:41:48.99	0.49	2.36	19.47	50.18	2.63	1.94	249.24	1.20	C*
4478-55600-135	08:30:52.17	+21:28:58.13	0.62	2.26	19.91	74.63	4.00	0.56	229.63	0.90	C*
6605-56565-722	02:01:21.40	+32:28:29.67	0.40	2.82	18.33	30.60	3.68	3.15	258.23	1.44	B*
5048-56218-801	22:28:25.77	+12:05:03.95	0.53	2.83	19.90	27.98	2.62	1.54	241.72	1.14	C*
5439-56045-109	13:44:08.02	+16:21:27.99	0.50	2.36	19.29	51.70	3.77	1.38	204.02	0.80	A*
4622-55629-615	11:10:40.43	+36:49:24.39	0.73	2.50	19.87	46.59	2.30	4.00	353.54	1.99	A*
6281-56295-811	00:29:27.39	+25:44:01.78	0.59	2.45	19.22	37.75	1.88	1.85	212.65	0.81	C*
6727-56369-967	15:16:41.22	+49:54:40.78	0.55	2.87	19.02	29.64	2.08	1.81	270.91	1.42	C*
4976-56046-672	15:42:04.99	+40:19:54.69	0.44	2.50	18.78	42.53	3.68	3.46	242.52	1.21	A*
6147-56239-163	23:17:31.62	+11:54:50.92	0.54	2.64	19.63	39.67	6.15	1.53	272.55	1.41	C*
4399-55811-149	02:37:40.64	-06:41:12.99	0.49	2.25	19.23	38.16	2.24	1.90	0.00	0.00	C*
5813-56363-341	09:18:07.86	+45:18:56.78	0.52	2.34	19.35	44.35	2.35	1.89	80.08	0.12	B*
4615-55618-399	11:31:10.98	+35:50:20.72	0.47	2.30	19.76	37.94	2.87	1.04	102.04	0.20	C*
6497-56329-271	14:13:58.21	+29:32:40.41	0.45	2.72	19.41	17.28	1.76	1.34	158.87	0.52	C*
4315-55503-703	01:13:00.58	+02:50:46.21	0.62	2.61	19.46	29.16	2.64	2.99	697.90	8.65	B*
4205-55454-645	22:45:05.93	+00:40:18.35	0.70	2.54	19.57	27.96	4.69	1.98	112.39	0.21	A*
6435-56341-855	11:10:27.11	+28:08:38.46	0.61	2.40	19.91	33.31	2.74	2.89	193.52	0.66	C*
6659-56607-751	10:29:01.23	+47:28:22.64	0.57	2.83	19.11	18.76	2.22	2.34	213.10	0.86	C*
5463-56003-121	14:29:54.80	+12:02:35.58	0.55	2.83	19.60	24.19	3.59	2.98	339.91	2.22	A*
6840-56685-377	12:41:44.24	+60:24:41.06	0.53	2.84	19.19	20.18	2.30	2.04	270.38	1.43	C*
4209-55478-171	23:14:18.03	-01:10:27.03	0.52	2.27	19.85	25.44	2.01	1.42	279.95	1.46	C*
5730-56607-697	09:18:59.21	+51:04:52.60	0.58	2.40	19.71	26.27	2.30	8.69	294.41	1.56	A*
3802-55514-227	08:05:30.94	+38:09:23.36	0.56	2.27	19.00	33.51	2.29	1.92	304.99	1.67	C
3786-55563-480	10:43:48.35	-02:12:36.06	0.77	2.74	19.92	17.22	2.11	3.16	850.00	11.51	C
6521-56537-324	23:52:12.72	+21:45:26.18	0.51	2.92	19.74	15.63	2.10	1.82	196.01	0.76	C*
5176-56221-583	08:43:00.14	+19:47:42.63	0.70	2.55	19.81	17.47	1.79	2.68	384.25	2.44	C*
3952-55330-944	15:46:11.14	+28:05:42.59	0.30	2.71	17.95	21.08	2.11	3.95	215.58	1.07	C*
5984-56337-186	12:41:32.92	+25:32:58.91	0.54	2.71	19.39	26.98	3.20	1.38	193.60	0.72	C*
3683-55178-662	07:59:30.25	+44:12:45.56	0.47	2.89	19.76	11.64	1.84	1.27	192.75	0.76	C*
11316-58396-95	21:15:36.90	+04:52:46.79	0.89	2.70	21.30	16.38	1.97	1.39	89.45	0.11	C
5068-55749-829	22:13:21.15	+09:02:41.06	0.63	2.67	19.80	17.82	1.99	1.36	224.06	0.89	C*
3673-55178-795	07:48:31.92	+47:10:14.40	0.44	2.51	19.17	17.86	2.91	2.40	139.77	0.40	C*
5366-55958-259	11:15:44.57	+08:47:17.16	0.55	2.89	19.55	16.40	2.19	2.07	235.26	1.07	C*
6725-56390-772	14:43:17.84	+51:07:21.03	0.55	2.95	18.73	16.33	2.82	9.40	169.93	0.56	B*
4619-55599-680	11:25:58.59	+33:56:47.11	0.48	2.92	19.46	15.45	1.98	0.15	196.24	0.78	C*
3929-55335-927	16:11:57.06	+22:14:40.00	0.47	2.27	19.85	24.12	2.54	0.80	336.16	2.20	C
4091-55498-227	21:48:43.04	+04:44:36.11	0.66	2.70	19.85	10.73	1.63	1.00	375.63	2.46	C*
4236-55479-117	02:17:42.02	-00:22:06.47	0.61	2.48	19.47	14.71	2.80	1.85	218.18	0.84	C*
7162-56605-731	01:36:44.88	-06:46:56.69	0.67	2.95	19.90	12.04	2.12	3.08	224.43	0.89	C*
8308-57417-911	07:56:50.15	+51:06:48.34	0.51	2.27	20.41	24.16	4.32	0.13	308.25	1.79	C
5069-56211-141	22:21:51.54	+07:24:00.24	0.75	2.51	19.65	14.69	1.79	1.41	308.02	1.49	C*
4030-55634-117	14:18:15.73	+01:58:32.23	0.56	2.14	19.35	33.88	2.96	2.61	219.28	0.85	C*
8735-58133-17	02:04:57.62	+06:03:59.14	0.72	2.77	19.99	33.16	2.42	2.09	270.16	1.22	C
6297-56218-817	22:33:33.17	+27:20:37.54	0.53	2.70	19.86	14.67	2.19	1.94	195.12	0.73	C*
5314-55952-733	09:35:51.05	+08:51:30.31	0.46	2.56	19.55	18.56	1.74	1.32	293.60	1.75	C*
5712-56602-365	09:15:26.16	+58:51:55.32	0.78	2.83	19.82	13.19	2.01	2.08	318.15	1.62	C*
3960-55663-695	15:21:05.07	+25:41:10.01	0.60	2.43	19.65	21.22	2.59	5.56	180.85	0.58	C
10740-58223-219	15:46:12.04	+34:52:54.49	0.93	2.79	20.83	17.64	2.74	6.79	850.00	10.17	C
8849-57457-315	11:17:51.25	+36:23:43.68	0.68	2.47	20.47	27.77	5.27	0.68	245.18	1.00	B
3946-55659-954	15:54:58.12	+26:46:30.73	0.29	2.44	17.79	15.78	2.24	0.56	208.70	1.00	C*
5896-56047-358	14:10:26.40	+21:28:52.30	0.31	2.95	18.18	15.15	2.87	3.81	137.85	0.44	C*

**Table 2.** Similar to Table 1, but for double-peaked candidates. ‘width’ and ‘skewness’ are the characteristics of the red peak.

Name	RA	Dec.	$z_l$	$z_s$	$\text{mag}_i$	$F_{lya}$	Width	Dkewness	$\sigma_{eBOSS}$	$\theta_E$	Grade
3865-55272-67	14:43:20.73	+33:42:12.10	0.58	2.20	19.56	85.19	3.18	2.75	201.93	0.71	A*
5369-56272-541	11:16:34.56	+09:15:03.05	0.55	2.45	19.41	58.31	2.75	2.56	256.14	1.22	B*
6833-56413-181	12:26:56.46	+54:57:39.05	0.50	2.73	18.88	38.18	2.29	1.72	237.51	1.12	C*
4356-55829-547	23:42:48.68	-01:20:32.54	0.53	2.27	19.58	71.30	1.65	1.83	286.28	1.52	C*
10751-58220-189	14:51:58.54	+35:05:44.62	0.75	2.94	20.55	37.42	1.88	3.79	258.86	1.11	C
3691-55274-138	08:15:35.86	+45:21:33.57	0.56	2.38	20.12	28.03	1.93	0.69	309.53	1.74	C*
3973-55323-667	12:46:58.13	+37:27:00.16	0.50	2.15	19.35	52.14	1.92	1.84	93.35	0.16	C
5168-56035-693	15:02:56.13	+38:20:17.89	0.45	2.67	18.74	29.18	1.74	3.00	230.71	1.09	C*
6422-56328-303	11:41:54.71	+22:16:28.89	0.59	2.76	19.73	28.30	2.32	5.22	267.86	1.33	C*
4725-55711-295	15:58:21.52	+26:58:29.77	0.58	2.38	19.66	37.51	2.34	2.88	285.37	1.46	C*
5980-56338-72	12:33:18.42	+25:23:12.13	0.39	2.53	18.68	25.00	1.79	5.57	220.82	1.04	C
4282-55507-607	23:33:11.12	+02:23:10.93	0.47	2.25	19.93	42.59	1.76	0.22	125.11	0.30	B*
6160-56190-854	23:46:17.29	+09:09:31.44	0.18	2.71	17.36	27.63	1.61	7.59	211.96	1.14	C*
3676-55186-623	07:48:16.74	+44:56:02.84	0.54	2.19	19.05	42.73	2.93	1.81	206.27	0.77	C*
4473-55589-231	07:50:57.45	+22:10:35.84	0.45	2.76	19.33	21.50	2.03	1.84	207.03	0.88	C*
7050-56573-273	01:45:26.41	-05:00:35.59	0.57	2.72	19.95	17.73	2.31	2.25	285.54	1.53	B*
7166-56602-783	23:53:56.11	-08:32:38.86	0.48	2.60	19.28	16.98	2.27	2.23	236.98	1.12	C*
6817-56455-897	13:28:08.68	+61:57:17.63	0.53	2.20	19.63	35.31	3.33	9.74	283.12	1.46	C*
6014-56072-764	14:28:23.29	+23:58:39.22	0.47	2.79	19.37	17.94	1.75	1.74	227.36	1.06	C*
5705-56194-259	00:59:48.32	+16:06:29.49	0.59	2.83	19.98	13.48	1.82	2.47	203.82	0.77	C*
4657-55591-602	01:51:50.89	+14:30:37.36	0.56	2.25	19.19	46.44	2.08	4.76	255.66	1.17	C*
6506-56564-729	23:02:17.65	+30:27:38.29	0.56	2.81	19.79	14.94	2.00	2.19	92.72	0.16	C*
4723-56033-445	15:28:51.87	+31:02:33.89	0.74	2.15	19.81	64.99	3.05	2.05	428.00	2.70	B*
3606-55182-695	01:51:45.36	-00:00:01.31	0.49	2.36	19.72	22.52	1.88	3.92	241.93	1.13	C*
6732-56370-859	14:50:01.80	+48:18:32.21	0.52	2.23	19.47	18.10	2.13	5.82	226.33	0.95	C
4644-55922-885	09:21:33.07	+36:37:17.40	0.56	2.68	19.91	17.10	1.84	1.59	219.14	0.91	C*
5960-56097-209	21:50:34.64	+24:50:16.53	0.50	2.28	19.10	23.86	2.44	1.59	201.49	0.77	C*
6459-56273-547	09:57:55.84	+24:09:44.70	0.58	2.10	19.66	56.58	4.14	3.40	110.56	0.21	C
4707-55653-211	13:25:25.47	+39:39:14.54	0.51	2.83	19.82	13.96	3.59	3.79	199.61	0.79	C*
7585-57190-637	22:34:03.96	+22:10:47.60	0.85	2.76	21.07	8.11	1.36	2.34	259.16	1.01	C
4398-55946-379	02:02:41.41	-06:46:11.33	0.50	2.75	19.40	17.74	2.35	3.96	159.53	0.50	C*
4235-55451-195	02:09:58.61	-00:12:32.73	0.47	2.17	19.68	19.97	1.81	4.19	144.71	0.40	C
7120-56720-417	12:29:10.35	+66:09:24.48	0.51	2.41	19.43	41.45	2.61	8.92	196.17	0.74	C*
6205-56187-591	00:42:32.96	+11:31:13.74	0.48	2.61	19.83	13.71	2.60	4.55	180.82	0.65	C
5004-55711-151	17:32:25.89	+26:42:40.33	0.50	2.20	19.95	15.67	2.75	5.18	176.88	0.59	C*
7033-56565-181	23:55:25.25	-05:21:38.63	0.49	2.24	19.29	27.11	2.37	0.95	294.36	1.65	C*
3923-55325-932	16:00:07.30	+14:51:12.23	0.34	2.37	18.13	18.23	3.14	4.84	193.75	0.82	C*
5359-55953-51	11:05:00.96	+07:37:30.47	0.56	2.81	19.01	12.96	1.64	5.96	615.40	7.23	B
6497-56329-843	14:16:13.97	+30:14:57.05	0.57	2.39	19.49	8.86	2.11	0.39	303.54	1.67	C*
5172-56071-193	14:48:17.66	+39:49:25.15	0.40	2.66	18.30	18.89	2.57	9.18	227.95	1.11	C*
6315-56181-654	16:50:03.85	+49:42:49.82	0.39	2.69	18.14	9.42	2.32	9.95	279.80	1.69	C*
6785-56487-647	15:45:25.18	+59:36:25.99	0.71	2.73	19.86	12.54	1.90	7.77	154.63	0.40	C
3876-55245-223	15:07:07.74	+30:47:56.83	0.62	2.50	19.87	11.60	2.62	5.22	155.21	0.42	B
4005-55325-923	13:05:21.79	+02:27:24.53	0.63	2.13	19.43	30.27	3.58	4.53	206.58	0.70	C
6177-56268-193	00:04:08.06	+14:11:50.71	0.58	2.90	19.81	12.55	1.48	1.26	173.64	0.57	C
5860-56046-853	13:05:52.60	+20:30:29.08	0.63	2.64	19.82	15.77	2.51	7.81	208.01	0.77	C
5782-56272-549	09:50:59.07	+20:23:10.61	0.65	2.21	19.77	22.94	2.28	1.56	281.30	1.29	C
3945-55648-275	15:24:18.57	+18:03:58.60	0.60	2.26	19.44	20.38	2.34	8.23	262.28	1.19	C
7037-56570-769	00:21:23.23	-04:29:26.17	0.56	2.15	19.77	43.13	3.07	1.22	231.63	0.95	C*
6408-56329-327	11:18:16.50	+26:26:48.58	0.45	2.73	19.15	8.72	1.96	8.95	175.24	0.63	C
6629-56365-311	13:55:59.42	+44:42:43.37	0.51	2.91	19.74	6.58	1.91	2.58	186.38	0.69	C
6460-56334-699	10:07:53.34	+34:14:19.42	0.54	2.90	19.06	10.30	2.09	6.19	196.88	0.76	C
4446-55589-635	08:19:36.04	+32:16:45.41	0.70	2.80	19.92	10.35	2.54	5.28	172.29	0.50	C
6111-56270-647	00:10:16.45	+17:58:52.16	0.56	2.28	19.63	11.77	1.67	8.75	261.03	1.23	C
5461-56018-391	14:21:40.44	+14:37:49.00	0.52	2.71	19.72	12.42	3.44	7.45	254.85	1.27	C
4208-55476-395	23:04:34.50	-00:10:58.19	0.56	2.82	19.78	8.56	1.70	2.24	184.78	0.65	C
5403-55979-51	12:35:36.97	+10:49:23.02	0.62	2.28	19.41	28.33	3.37	5.09	366.06	2.28	C
6316-56483-84	16:19:28.74	+49:20:55.93	0.29	2.76	17.95	17.19	2.89	4.23	242.92	1.38	C
4397-55921-505	02:00:25.74	-07:56:31.93	0.26	2.49	17.78	16.44	2.50	5.23	238.05	1.34	C
5648-55923-862	00:13:21.42	+11:40:39.61	0.35	2.18	18.34	20.36	2.00	5.62	223.09	1.07	C



**Figure 3.** Spectra of 12 single-peaked candidates are shown here for demonstration. The emission lines identified as Ly $\alpha$  are marked by shaded green. The text in each panel shows the SDSS ‘plate-mjd-fibre.’ More spectra figures are available in the online journal.



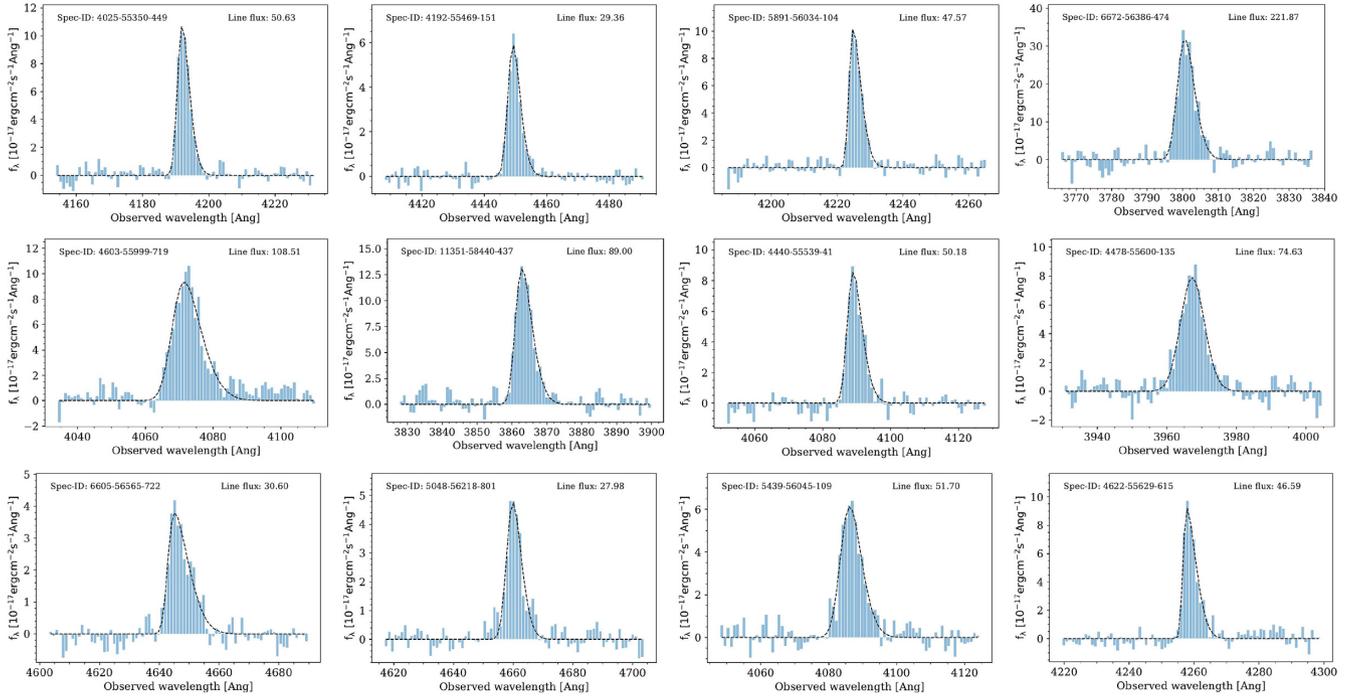
**Figure 4.** Similar to 3, but for double-peaked candidates.

Our lensed LAE candidates with different types of Ly $\alpha$  emission could be potentially beneficial for the study of both background sources and foreground lens galaxies with future follow-up observations. We describe them in detail as follows.

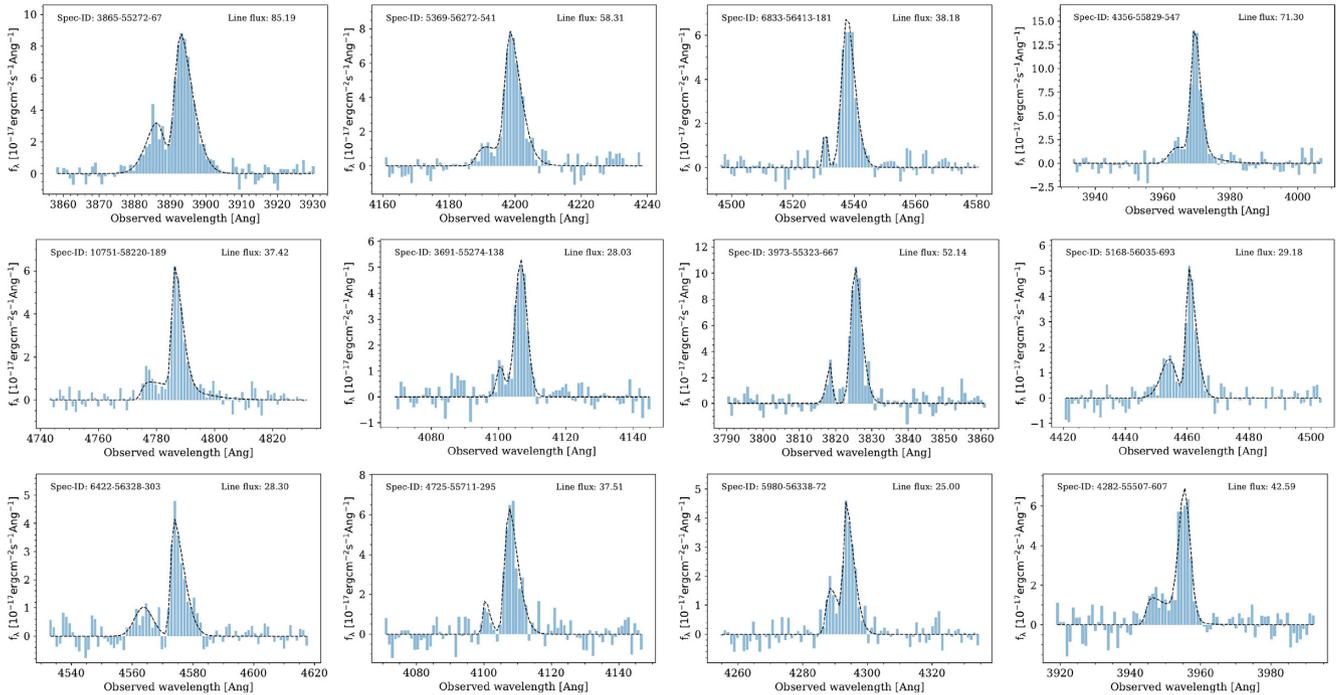
#### 4.1 Science of the background source

A typical lens in our sample can provide a magnification of  $\sim 10$ , boosting both the total flux and the apparent size of the background

LAEs. With future high-resolution imaging follow-up, these lenses can offer new possibilities to study the LAEs in detail. The science potential of the lensed LAEs has been demonstrated in previous work of Shu et al. (2016a), who obtained the morphology of LAEs down to the scale of  $\sim 100$  pc region using follow-up Hubble imaging (see also Cornachione et al. 2018), which cannot be achieved outside the strong lensing systems. Base on the same data set, Ritondale et al. (2019a) built the non-parametric models of the background LAEs, and found these clumps are usually very elliptical, revealing



**Figure 5.** This figure shows zoomed spectra around the Ly $\alpha$  for 12 single-peaked candidates. The blue bars represent residual spectra (observed spectra with the best-fitting composite spectra subtracted), and the black dot lines show our best skew-normal fit to the Ly $\alpha$  emission. More zoomed spectra figures are available in the online journal.

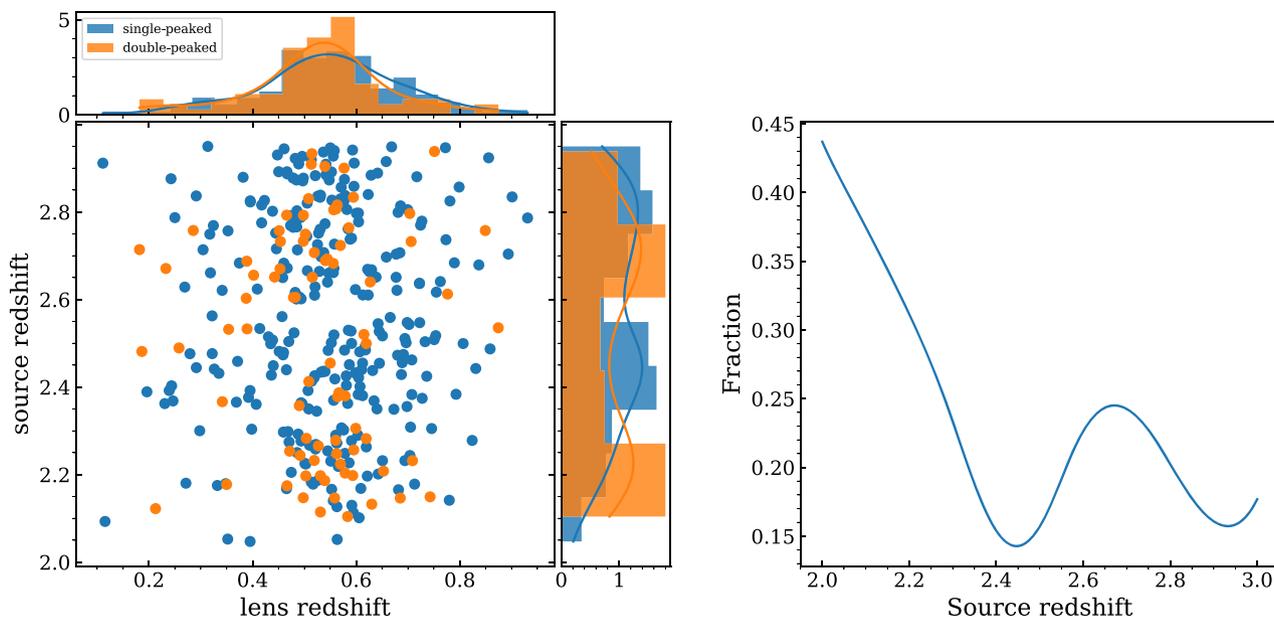


**Figure 6.** Similar to 5, but for double-peaked candidates.

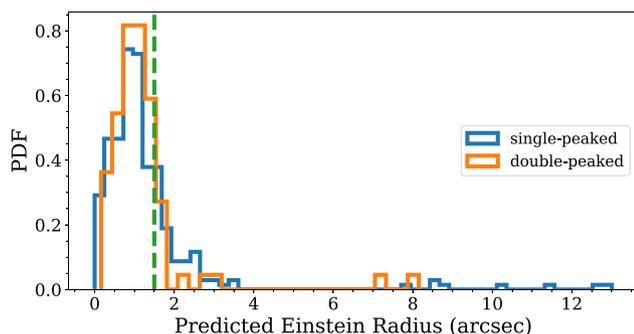
typical configuration of star formation discs. And this phenomenon may indicate that the escape rate of Ly $\alpha$  photon is irrelevant with the inclination of the star formation disc. Thus, some anisotropic escape model for LAEs (such as Verhamme et al. 2012) may not be favoured.

One intriguing question of LAEs is why they have different line profiles. Although many theoretical effort has been made

to explain the diversity of Ly $\alpha$  emission in LAEs (Hayashino et al. 2004; Verhamme, Schaerer & Maselli 2006; Verhamme et al. 2012; Dijkstra & Kramer 2012; Orsi et al. 2012; Gronke & Dijkstra 2016), no consensus has emerged, and more observational evidence is needed to discriminate different theoretical models. If high-resolution imaging can be obtained for our sample, they will form a promising lab for this purpose.



**Figure 7.** The redshift distribution of lens/source galaxies in our sample. Left: the bottom left panel shows the redshift distribution of lens/source galaxies in the 2D plane via scatter diagram. Double-peaked candidates are marked with orange dots, while blue dots represent single-peaked candidates. The orange and blue histograms in the top panel are the 1D marginalized PDF of lens redshift for double- and single-peaked candidates, respectively. The smooth version of PDF estimated via Gaussian KDE assuming the ‘Scott’ bandwidth scenario is shown by the orange and blue lines, where orange represents single-peaked candidates and blue indicates double-peaked candidates. Similarly, the 1D marginalized PDF of source redshift for double- and single-peaked candidates are shown in the right-hand panel. Right: the fraction of double-peaked candidates as a function of redshift, which is evaluated from the smooth PDF of source redshift using the Gaussian KDE.



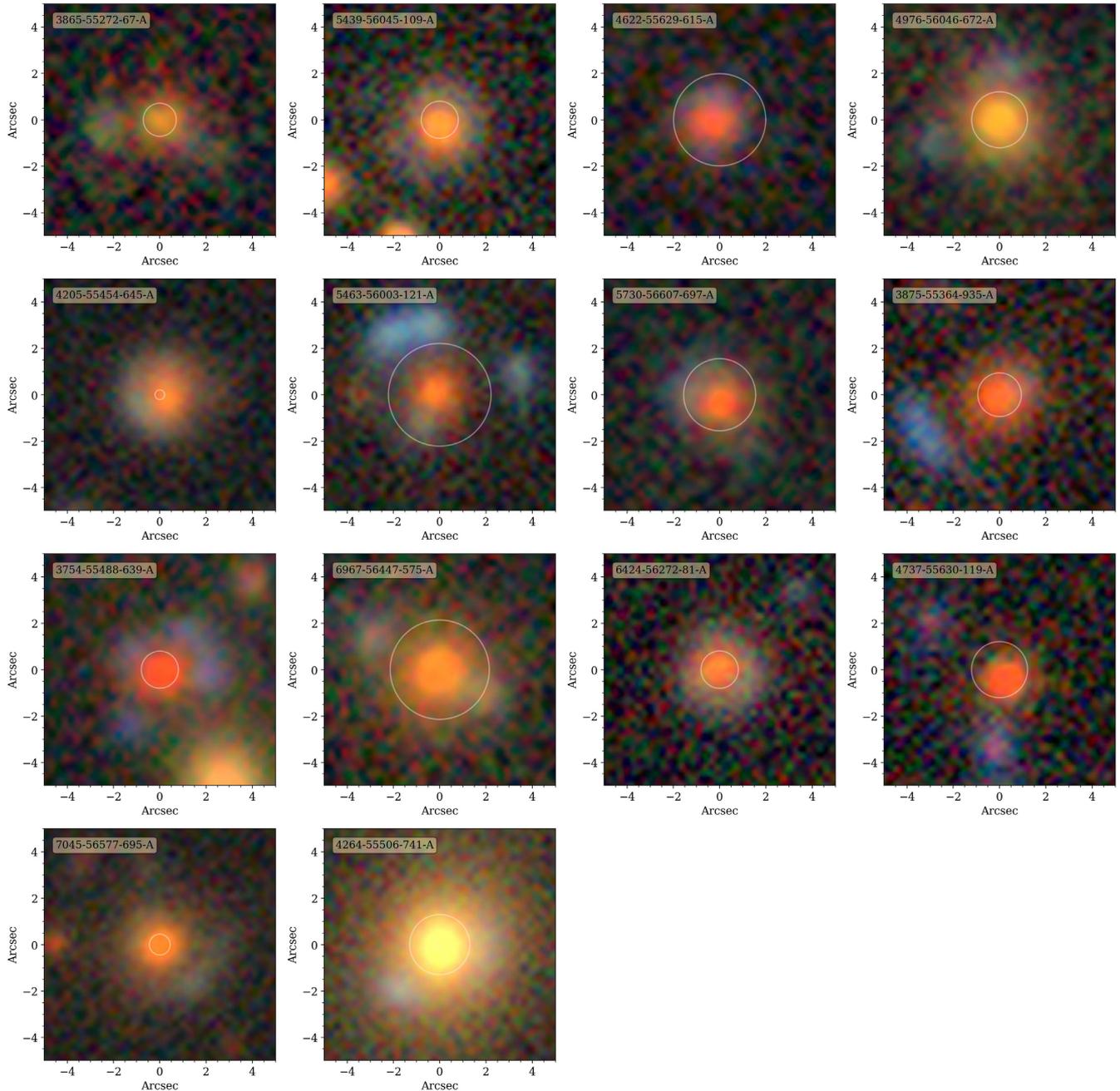
**Figure 8.** The PDF of the predicted Einstein radii assuming the SIE model. Two types of lensed LAEs are shown by different colours. The green dashed line marks the radius of eBOSS fibre (1.5 arcsec). Note that the eBOSS pipeline sometimes generates unreliable estimations for velocity dispersion of target galaxies. Thus some systems in our catalogue may have anomalously large velocity dispersions and unreliably large Einstein radii.

#### 4.2 Application for foreground lens science

Massive ETGs are the end-product of hierarchical galaxy formation. Strong gravitational lensing system can constrain the inner density distributions of the ETGs, which contains crucial information of the galaxy evolution process. Koopmans et al. (2006) found the inner mass density slope of lens galaxies do not significantly evolve with redshift by joint gravitational lensing and stellar-dynamical analysis of 15 massive field ETGs selected from the SLACS Survey. They found  $d\langle\gamma\rangle/dz = -0.23 \pm 0.16$ , where the  $\langle\gamma\rangle$  represents the inner density slope of total mass. Recently, Chen et al. (2019) utilized the most comprehensive compiled lens data set to measure

the evolution of the galaxy density profile. By combining the joint gravitational lensing and stellar-dynamical analysis of 161 galaxy-scale strong gravitational lensing systems from the LSD survey (Treu & Koopmans 2004), the SL2S survey (Sonnenfeld et al. 2013), the SLACS survey, the S4TM survey, the BELLS survey, and the BELLS GALLERY survey, they measured the  $d\langle\gamma\rangle/dz = -0.218^{+0.089}_{-0.087}$ . This result indicates a mild trend that the galaxies at a lower redshift have steeper slopes. Similar conclusions were also made by several other works (Koopmans et al. 2009; Ruff et al. 2011; Sonnenfeld et al. 2013; Li, Shu & Wang 2018). While the above works have claimed only weak evolution for the total mass density slope, Bolton et al. (2012a) analysed the SLACS and BELLS lenses over the redshift interval  $z \approx 0.1-0.6$ , and they found a stronger evolution trend, with magnitude  $d\langle\gamma\rangle/dz = -0.60 \pm 0.15$ . On the other hand, a recent simulation analysis based on the IllustrisTNG project gives a different picture, which claimed that the density slope remains almost constant from  $z = 1$  to 0, aside from a residual trend of becoming shallower towards  $z = 0$  (Wang et al. 2019). Our lens sample can be complimentary to previous studies in the sense that a significant fraction ( $\sim 10$  per cent) of our lenses locate at  $z > 0.7$ , which rarely appears in previous BELLS and BELLS GALLERY catalogue. These new  $z > 0.7$  lenses will allow us to extend the strong-lensing measurement of the galaxy inner density slope to higher redshift and obtain a more comprehensive understanding of the evolution of the density slope of ETGs.

Another important application of the lensed LAEs is the searching of low-mass dark matter haloes and subhaloes through the flux distortions that they induce on the lensed images. For galaxy lenses, with high-resolution imaging observations from a ground-based telescope with Laser-AO, the mass limit of detection for haloes (subhaloes) can be as low as  $\sim 10^8 M_\odot$  (Vegetti et al. 2012). Recent



**Figure 9.** An atlas that includes all the cut-out images of grade-A candidates. Images are taken from the DECaLS survey. The text label in each panel represents the ‘plate-mjd-fibre-grade.’ where the grade is evaluated according to the criterion defined in Section 3. White circles in each panel show the predicted Einstein radius calculated from the velocity dispersion under the SIE lens assumption.

theoretical investigations have shown that a sample of  $\sim 20$  with detection limit of  $\sim 10^8 M_{\odot}$  can distinguish cold dark matter model from the warm dark matter model which are made of 7 keV sterile neutrinos whose decay can explain the observed 3.5 keV X-ray emission lines in clusters and galaxies (Boyarsky et al. 2014; Bulbul et al. 2014; Li et al. 2016, 2017).

The flux perturbation induced by the dark matter clumps in lens galaxy is proportional to the product of brightness gradient of the source galaxy and the gradient of the subhalo potential perturbation. For a given detection precision of flux distortion, the higher the source gradient is, the lower the lens potential perturbation can be detected (Koopmans 2005). The LAEs are usually made of multiple compact

clumps with sharp brightness profile. Therefore, our candidates are promising places for the search of low-mass dark matter subhaloes (Shu et al. 2016a; Ritondale et al. 2019b), if high-resolution imaging follow-ups are obtained.

## ACKNOWLEDGEMENTS

We thank for Yun Chen and Zhengya Zheng for their helpful suggestions. We acknowledge the support from National Key Program for Science and Technology Research Development of China (2017YFB0203300), National Natural Science Foundation of China (nos 11773032, 118513, 11821303, 12022306, 11761131004 and

11761141012), the National Key Basic Research and Development Program of China (no. 2018YFA0404501), and RL is supported by NAOC Nebula Talents Program.

Funding for the SDSS-IV has been provided by the Alfred P. Sloan Foundation, the U.S. Department of Energy Office of Science, and the Participating Institutions. SDSS-IV acknowledges support and resources from the Center for High-Performance Computing at the University of Utah. The SDSS web site is [www.sdss.org](http://www.sdss.org).

SDSS-IV is managed by the Astrophysical Research Consortium for the Participating Institutions of the SDSS Collaboration including the Brazilian Participation Group, the Carnegie Institution for Science, Carnegie Mellon University, the Chilean Participation Group, the French Participation Group, Harvard-Smithsonian Center for Astrophysics, Instituto de Astrofísica de Canarias, The Johns Hopkins University, Kavli Institute for the Physics and Mathematics of the Universe (IPMU)/University of Tokyo, the Korean Participation Group, Lawrence Berkeley National Laboratory, Leibniz Institut für Astrophysik Potsdam (AIP), Max-Planck-Institut für Astronomie (MPIA Heidelberg), Max-Planck-Institut für Astrophysik (MPA Garching), Max-Planck-Institut für Extraterrestrische Physik (MPE), National Astronomical Observatories of China, New Mexico State University, New York University, University of Notre Dame, Observatório Nacional/MCTI, The Ohio State University, Pennsylvania State University, Shanghai Astronomical Observatory, United Kingdom Participation Group, Universidad Nacional Autónoma de México, University of Arizona, University of Colorado Boulder, University of Oxford, University of Portsmouth, University of Utah, University of Virginia, University of Washington, University of Wisconsin, Vanderbilt University, and Yale University.

## DATA AVAILABILITY

The data that support the findings of this study are available on request from the corresponding author. Most of the data are already openly available on the SDSS DR16 website, <http://skyserver.sdss.org/dr16/en/tools/explore/Summary.aspx?>

## REFERENCES

- Ahn S.-H., 2004, *ApJ*, 601, L25  
 Ahumada R. et al., 2019, *ApJS*, 249, 3  
 Alam S. et al., 2015, *ApJS*, 219, 12  
 Blanton M. R. et al., 2017, *AJ*, 154, 28  
 Bolton A. S., Burles S., Schlegel D. J., Eisenstein D. J., Brinkmann J., 2004, *AJ*, 127, 1860  
 Bolton A. S., Burles S., Koopmans L. V. E., Treu T., Gavazzi R., Moustakas L. A., Wayth R., Schlegel D. J., 2008, *ApJ*, 682, 964  
 Bolton A. S. et al., 2012a, *ApJ*, 757, 82  
 Bolton A. S. et al., 2012b, *ApJ*, 757, 82  
 Bolton A. S. et al., 2012c, *AJ*, 144, 144  
 Boyarsky A., Ruchayskiy O., Iakubovskiy D., Franse J., 2014, *Phys. Rev. Lett.*, 113, 251301  
 Brownstein J. R. et al., 2012, *ApJ*, 744, 41  
 Buchner J. et al., 2014, *A&A*, 564, A125  
 Bulbul E., Markevitch M., Foster A., Smith R. a. K., Loewenstein M., Randall S. W., 2014, *ApJ*, 789, 13  
 Chen Y., Li R., Shu Y., Cao X., 2019, *MNRAS*, 488, 3745  
 Cheng X.-L., 2020, *Res. Astron. Astrophys.*, 20, 002  
 Cornachione M. A. et al., 2018, *ApJ*, 853, 148  
 Dawson K. S. et al., 2016, *AJ*, 151, 44  
 Dey A. et al., 2019, *AJ*, 157, 168  
 Dijkstra M., Kramer R., 2012, *MNRAS*, 424, 1672  
 Dijkstra M., Haiman Z., Spaans M., 2006, *ApJ*, 649, 37  
 Feroz F., Hobson M. P., Bridges M., 2009, *MNRAS*, 398, 1601  
 Fuller S. et al., 2020, *ApJ*, 896, 156

- Gavazzi R., Treu T., Rhodes J. D., Koopmans L. V. E., Bolton A. S., Burles S., Massey R. J., Moustakas L. A., 2007, *ApJ*, 667, 176  
 Gronke M., Dijkstra M., 2016, *ApJ*, 826, 14  
 Gronke M., Bull P., Dijkstra M., 2015, *ApJ*, 812, 123  
 Gunn J. E. et al., 2006, *AJ*, 131, 2332  
 Hayashino T. et al., 2004, *AJ*, 128, 2073  
 Jiang G., Kochanek C. S., 2007, *ApJ*, 671, 1568  
 Koopmans L. V. E., 2005, *MNRAS*, 363, 1136  
 Koopmans L. V. E., Treu T., Bolton A. S., Burles S., Moustakas L. A., 2006, *ApJ*, 649, 599  
 Koopmans L. V. E. et al., 2009, *ApJ*, 703, L51  
 Kormann R., Schneider P., Bartelmann M., 1994, *A&A*, 284, 285  
 Kulas K. R., Shapley A. E., Kollmeier J. A., Zheng Z., Steidel C. C., Hainline K. N., 2012, *ApJ*, 745, 33  
 Li R., Frenk C. S., Cole S., Gao L., Bose S., Hellwing W. A., 2016, *MNRAS*, 460, 363  
 Li R., Frenk C. S., Cole S., Wang Q., Gao L., 2017, *MNRAS*, 468, 1426  
 Li R., Shu Y., Wang J., 2018, *MNRAS*, 480, 431  
 Mallery R. P. et al., 2012, *ApJ*, 760, 128  
 Neufeld D. A., 1990, *ApJ*, 350, 216  
 Orsi A., Lacey C. G., Baugh C. M., 2012, *MNRAS*, 425, 87  
 Prakash A. et al., 2016, *ApJS*, 224, 34  
 Ritondale E., Auger M. W., Vegetti S., McKean J. P., 2019a, *MNRAS*, 482, 4744  
 Ritondale E., Vegetti S., Despali G., Auger M. W., Koopmans L. V. E., McKean J. P., 2019b, *MNRAS*, 485, 2179  
 Ruff A. J., Gavazzi R., Marshall P. J., Treu T., Auger M. W., Brault F., 2011, *ApJ*, 727, 96  
 Shu Y. et al., 2015, *ApJ*, 803, 71  
 Shu Y. et al., 2016a, *ApJ*, 824, 86  
 Shu Y. et al., 2016b, *ApJ*, 833, 264  
 Shu Y. et al., 2017, *ApJ*, 851, 48  
 Sobral D. et al., 2018, *MNRAS*, 477, 2817  
 Sonnenfeld A., Treu T., Gavazzi R., Suyu S. H., Marshall P. J., Auger M. W., Nipoti C., 2013, *ApJ*, 777, 98  
 Tapken C., Appenzeller I., Noll S., Richling S., Heidt J., Meinköhn E., Mehlert D., 2007, *A&A*, 467, 63  
 Trainor R. F., Steidel C. C., Strom A. L., Rudie G. C., 2015, *ApJ*, 809, 89  
 Treu T., Koopmans L. V. E., 2004, *ApJ*, 611, 739  
 Treu T., Koopmans L. V., Bolton A. S., Burles S., Moustakas L. A., 2006, *ApJ*, 640, 662  
 U V. et al., 2015, *ApJ*, 815, 57  
 Vegetti S., Lagattuta D. J., McKean J. P., Auger M. W., Fassnacht C. D., Koopmans L. V. E., 2012, *Nature*, 481, 341  
 Verhamme A., Schaerer D., Maselli A., 2006, *A&A*, 460, 397  
 Verhamme A., Dubois Y., Blaizot J., Garel T., Bacon R., Devriendt J., Guiderdoni B., Slyz A., 2012, *A&A*, 546, A111  
 Wang Y. et al., 2019, *MNRAS*, 490, 5722  
 Warren S. J., Hewett P. C., Lewis G. F., Møller P., Iovino A., Shaver P. A., 1996, *MNRAS*, 278, 139  
 Willis J. P., Hewett P. C., Warren S. J., 2005, *MNRAS*, 363, 1369  
 Willis J. P., Hewett P. C., Warren S. J., Dye S., Maddox N., 2006, *MNRAS*, 369, 1521  
 Yamada T., Matsuda Y., Kousai K., Hayashino T., Morimoto N., Umemura M., 2012, *ApJ*, 751, 29  
 Yang H. et al., 2017, *ApJ*, 844, 171

## SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](http://mnras.oxfordjournals.org/) online.

**Table 1.** Catalogue of single-peaked candidates.

Please note: Oxford University Press is not responsible for the content or functionality of any supporting materials supplied by the authors. Any queries (other than missing material) should be directed to the corresponding author for the article.

This paper has been typeset from a  $\text{\TeX}/\text{\LaTeX}$  file prepared by the author.