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ATOMS: ALMA three-millimeter observations of massive star-forming regions – III. Catalogues of candidate hot molecular cores and hyper/ultra compact H II regions

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ABSTRACT

We have identified 453 compact dense cores in 3 mm continuum emission maps in the ALMA Three-millimetre Observations of Massive Star-forming regions survey, and compiled three catalogues of high-mass star-forming cores. One catalogue, referred to as hyper/ultra compact (H/UC)-H II catalogue, includes 89 cores that enshroud H/UC H II regions as characterized by associated compact H40 α emission. A second catalogue, referred to as pure s-CHMC, includes 32 candidate hot molecular cores (HMCs) showing rich spectra ($N \geq 20$ lines) of complex organic molecules (COMs) and not associated with H/UC-H II regions. The third catalogue, referred to as pure w-CHMC, includes 58 candidate HMCs with relatively low levels of COM richness and not associated with H/UC-H II regions. These three catalogues of dense cores provide an important foundation for future studies of the early stages of high-mass star formation across the Milky Way. We also find that nearly half of H/UC-H II cores are candidate HMCs. From the number counts of COM-containing and H/UC-H II cores, we suggest that the duration of high-mass protostellar cores showing chemically rich features is at least comparable to the lifetime of H/UC-H II regions. For cores in the H/UC-H II catalogue, the width of the H40 α line increases as the core size decreases, suggesting that the non-thermal dynamical and/or pressure line-broadening mechanisms dominate on the smaller scales of the H/UC-H II cores.

Key words: stars: formation – stars: kinematics and dynamics – ISM: clouds – H II regions.

1 INTRODUCTION

High-mass stars ($M_{\star} > 8 M_{\odot}$) play a crucial role in many astrophysical processes, for example, from the formation of solid material in the early Universe (Dunne et al. 2003) to the influence on the evolution of their host galaxies and future generations of star formation in their natal molecular clouds (e.g. Kennicutt 2005; Urquhart et al. 2014). Despite its importance on different (time and spatial) scales, high-mass star formation remains much less understood than low-mass star formation due to the observational challenges presented by increased distances, rarity, opaque surroundings, and short formation time-scales (e.g. Zinnecker & Yorke 2007; Tan et al. 2014; Motte,

Bontemps & Louvet 2018). In addition, high-mass protostars are embedded in a rich cluster environment (e.g. de Wit et al. 2005), which makes it very difficult to disentangle the physical formation processes of high-mass stars and numerous low-mass cluster members.

Despite the observational difficulties mentioned above, it has been known for the past 15 yrs that high-mass stars spend a considerable fraction of their lifetime (>10 per cent) embedded in their natal molecular clouds (e.g. van der Tak 2004). From an observational perspective, this embedded phase can be generally subdivided into four different groups of objects that represent different evolutionary stages: (1) Massive dense cores (MDCs; $M \gtrsim 30 M_{\odot}$ within ~ 0.1 pc). They are generally the nurseries of high-mass star formation (e.g. Zhang et al. 2009, 2015; Wang et al. 2011; Motte et al. 2018; Li et al. 2019; Sanhueza et al. 2019; Svoboda et al. 2019), and perhaps analogous to the prestellar cores in low-mass star formation regions

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if they are completely starless. (2) High-mass protostellar objects (HMPOs). They are very luminous, with bolometric luminosities above $10^3 L_\odot$, and deeply embedded within massive envelopes having centrally peaked temperature and density distributions. At this stage, a (pseudo) disc is formed to transfer on to the HMPOs the infalling material from the envelope (e.g. Beuther et al. 2002; McKee & Tan 2003; Cesaroni et al. 2007). (3) Hot molecular cores (HMCs). They correspond to the compact (<0.1 pc), dense ($\geq 10^5$ – 10^8 cm $^{-3}$), massive ($\sim 100 M_\odot$), and hot ($\gtrsim 100$ K) molecular material that is radiatively heated intensively (e.g. Kurtz et al. 2000; Hosokawa & Omukai 2009; Rathborne et al. 2011). Consequently, the intense radiative heating enriches the gas phase with many complex organic molecules (COMs) such as CH_3OCH_3 and CH_3OCHO . (4) Hyper/ultra compact (H/UC)-H II regions. The UC-H II regions refer to those with sizes ≤ 0.1 pc, densities $> 10^4$ cm $^{-3}$, and emission measures $\geq 10^7$ pc cm $^{-6}$ (e.g. Wood & Churchwell 1989; Hoare et al. 2007). The HC-H II regions, which are a newly identified class of compact H II regions after the discovery of UC-H II, correspond to those with sizes ≤ 0.05 pc, densities $> 10^6$ cm $^{-3}$, and emission measures $\geq 10^7$ pc cm $^{-6}$ (Kurtz & Hofner 2005; Hoare et al. 2007; Yang et al. 2019, 2021).

During the past decades, studies of the embedded phase of high-mass star formation have made much progress. After a large number of infrared dark clouds (IRDCs) were revealed by Infrared Space Observatory (*ISO*; Perault et al. 1996), Midcourse Space Experiment (*MSX*; Egan et al. 1998), and *Spitzer* (Peretto & Fuller 2009), several surveys in (sub)millimetre continuum and spectral lines followed (e.g. Ragan et al. 2006; Rathborne, Jackson & Simon 2006; Simon et al. 2006; Rathborne et al. 2010; Dirienzo et al. 2015; Sanhueza et al. 2019), and have established that massive IRDCs harbour the early, embedded stages of high-mass star and cluster formation. In particular, the recent ALMA 1 mm survey ASHES led by Sanhueza et al. (2019), which was designed to mosaic 12 massive ($> 500 M_\odot$), cold (≤ 15 K), 3.6–70 μm dark cloud clump at the resolution of ~ 4000 au and mass sensitivity better than $0.2 M_\odot$, revealed a large population of low-mass ($< 1 M_\odot$) cores but no high-mass ($> 30 M_\odot$) prestellar cores on the 0.01–0.1 pc scales. With detailed analysis, the authors also put several strong observational constraints on cloud fragmentation, the clustering mode of dense cores, and the dynamical collapse scenario at the early stages of high-mass star formation.

As for the HMPO and HMC stages, both have many observational similarities except for the abundance of COMs; for example, they both are accompanied by massive molecular outflows, infall, and even rotation (e.g. Zhang et al. 2001, 2005; Beuther et al. 2002; van der Tak 2004; Liu et al. 2017). Some important physical processes including mass accretion and disc rotation related to the HMPOs can therefore be understood from the investigation into the HMCs since the latter is generally accompanied by rich molecular line emission that bears both physical and chemical imprints of high-mass star formation (e.g. Keto & Zhang 2010; Qiu et al. 2012). Focusing on the HMCs, we find that most of the efforts so far have been towards individual sources or small samples through single-dish and/or interferometric observations in the (sub)millimetre regime (e.g. van der Tak 2004; Beuther 2007; Qin et al. 2008, 2015; Purcell et al. 2009; Cesaroni et al. 2010; Furuya, Cesaroni & Shinnaga 2011; Rathborne et al. 2011; Xu & Wang 2013; Hernández-Hernández et al. 2014; Silva et al. 2017; Csengeri et al. 2019; Belloche et al. 2020; Coletta et al. 2020; Sato et al. 2020). Although these studies have demonstrated the HMC as an ideal laboratory in which to search for infall, outflow, and rotation motions related to high-mass star formation, these important physical processes are still far from being well understood mainly due to the lack of deep, interferometric

(sub)millimetre observations towards a sufficiently large sample of HMCs. For the H/UC-H II stage, the situation improves; there are several surveys that serve to characterize the H/UC-H II regions, such as *IRAS* warm dust continuum at far-infrared wavelengths (Wood & Churchwell 1989; Kurtz, Churchwell & Wood 1994), CORNISH 5 GHz radio continuum (Hoare et al. 2012; Purcell et al. 2013), and ATLASGAL 870 μm cold dust continuum (Schuller et al. 2009; Urquhart et al. 2013) surveys. These surveys help define observationally the H/UC-H II regions reasonably well in terms of their sizes, densities, and emission measures (see above).

In addition, some H/UC-H II regions are likely still accreting mass, despite the high pressure of ionized gas (e.g. Sollins et al. 2005; Keto 2007; Keto, Zhang & Kurtz 2008). However, the role that H/UC-H II regions play in the early stages of high-mass star formation remains to be understood. Hopefully, the H/UC-H II regions can light up their immediate surroundings and allow the investigation of the properties (e.g. density distribution and velocity field) of the surrounding region in which high-mass stars are formed (e.g. Hoare et al. 2007). The immediate vicinities of H/UC-H II regions have been the subject of the (sub)millimetre observation studies, however, the lack of high angular resolution observations has impeded anything like complete understanding of H/UC-H II region physics.

In this paper, the third in a series from ATOMS¹ (e.g. Liu et al. 2020a,b, hereafter, [Paper I](#) and [Paper II](#), respectively), we take advantage of deep, high-resolution ALMA 3 mm observations (see below) of the ATOMS survey. In [Paper I](#), we introduced the main scientific goals of the survey through a case study for the G9.62+0.19 complex, including to systematically investigate the spatial distribution of various dense gas tracers in a large sample of Galactic massive star-forming clumps, to study the roles of stellar feedback in star formation, and to characterize filamentary structures inside massive clumps. In [Paper II](#), we studied the star formation scaling relations inferred from different dense gas tracers, and suggested that both the main and isotopologue lines (i.e. $\text{HCO}^+/\text{H}^{13}\text{CO}^+$, $\text{HCN}/\text{H}^{13}\text{CN}$) are good tracers of the total mass of dense gas in Galactic molecular clumps, and that the large optical depths of the main lines do not affect the interpretation of the slopes in star formation relations.

With the ATOMS survey data, in this paper we aim to establish catalogues of a large sample of both candidate HMCs and H/UC-H II regions as a crucial foundation for future studies for the early stages of high-mass star formation. The ATOMS survey targeted a large sample of 146 *IRAS* clumps (Bronfman, Nyman & May 1996; Liu et al. 2016), which have masses 5.6 to $2.5 \times 10^5 M_\odot$ with a median value of $1.4 \times 10^3 M_\odot$, radii of 0.06–4.26 pc with a median value of 0.86 pc, and bolometric luminosities of 16 – $8.1 \times 10^6 L_\odot$ with a median value of $5.7 \times 10^4 L_\odot$ (see table A1 of [Paper I](#)). All the targets except for I08076–3556 and I11590–6452 have bolometric luminosities greater than $10^3 L_\odot$, and actually more than 90 per cent of them are high-mass star-forming regions (see [Paper I](#) and [Paper II](#)). In addition, the ATOMS targets are located in the first and fourth Galactic Quadrants of the inner Galactic plane ($-80^\circ < l < 40^\circ$, $|b| < 2^\circ$) at distances of 0.4–13.0 kpc (the corresponding Galactocentric distances of 0.5–12.7 kpc, see Appendix A for the distance calculations). 27 distant ($d \geq 7$ kpc) sources are either close to the Galactic Centre or mini-starbursts (like W49N/I19078+0901), representing extreme environments for star formation. Overall, the ATOMS survey contains a diversity of objects suitable for studying the early stages of high-mass star formation, especially HMC and

¹ ATOMS: ALMA Three-millimeter Observations of Massive Star-forming regions survey.

H/UC-H II stages, with different physical conditions – including densities and luminosities – and in different environments, across a large range of Galactocentric distances.

This paper is organized as follows: Section 2 gives a brief description about the ALMA observations of the ATOMS survey, Section 3 presents the results of the extraction and search of the candidate HMC and H/UC-H II cores, Section 4 presents a discussion, and Section 5 includes the summary of our results and conclusion.

2 ALMA OBSERVATIONS

We make use of the ATOMS survey data (Project ID: 2019.1.00685.S; PI: Tie Liu, see [Paper I](#) and [Paper II](#)). The observations were conducted towards the 146 *IRAS* clumps in the single-pointing mode with both the Atacama Compact 7 m Array (ACA; Morita Array) and the 12 m array (C43-2 or C43-3 configurations) in band 3. Eight spectral windows (SPWs) were tuned to cover 11 commonly used lines including the dense gas tracers (e.g. HCO^+ , HCN , and their isotopes), HMC tracers (e.g. CH_3OH and HC_3N), shock tracers (e.g. SiO and SO), and ionized gas tracers ($\text{H}_{40\alpha}$). The basic parameters (e.g. rest frequency, transition, critical density, and upper energy temperature) of these lines can be found in Table 2 of [Paper I](#). The SPWs 1–6² at the lower side band, with spectral resolutions of $\sim 0.2\text{--}0.4\text{ km s}^{-1}$ were chosen to resolve the line profiles in order to investigate the kinematics within high-mass star forming clumps, while the SPWs 7 and 8 at the upper side band, each with a broad bandwidth of 1875 MHz at a spectral resolution of $\sim 1.6\text{ km s}^{-1}$, were chosen for continuum emission and line scan observations. Note that the original spectral resolution of the SPWs 7 and 8 is $\sim 1.6\text{ km s}^{-1}$ instead of $\sim 3\text{ km s}^{-1}$ as given in [Paper I](#).

The ACA and 12 m array data were calibrated and imaged separately with the CASA software package version 5.6 (McMullin et al. 2007), where the continuum image centred at $\sim 93.8\text{ GHz}$ was cleaned in an aggregated $\sim 4\text{ GHz}$ frequency bandwidth free of the line emission. More details on the data reduction can be found in [Paper I](#) and [Paper II](#). As we focus in this work on the very dense, compact HMCs and H/UC-H II regions, in what follows we will only consider the 12 m array data and the analysis is focused on the primary-beam corrected data. The reduced 12 m array continuum image and line cubes for the 146 target clumps have angular resolutions of $\sim 1''.2\text{--}1''.9$ (for reference, 2 arcsec corresponds to 0.1 pc at a distance of 10 kpc), and maximum recoverable angular scales $\sim 14''.5\text{--}20''.3$. The sensitivity of the 12 m array data is better than 10 mJy beam^{-1} per 0.122 MHz channel (see Table 1 of [Paper I](#)) for lines, and $\sim 0.4\text{ mJy beam}^{-1}$ for continuum.

3 ANALYSIS AND RESULTS

3.1 Core extraction

To extract compact cores from the 3 mm continuum maps, we make use of both the *Dendrogram* algorithm³ and CASA *imfit* function. *Dendrogram* can be easily used to extract the dense leaf-like structures (hereafter referred to as cores), which are the smallest

structures without substructures in the terminology of the algorithm, and thus customarily taken to be the candidates for cores. This technique does not always provide good measurements of the core parameters such as size and position angle, while CASA *imfit* performs better in this regard through a two-dimensional Gaussian fit to the emission. A visual inspection of the continuum images of the 146 target sources shows that their outer parts, i.e. the area beyond a radius of 0.01° from the image centre, are much noisier than the inner parts, where the 0.01° radius corresponds to the $\text{pblimit} = 0.4$. Therefore, we masked out the outer part of the images before identifying the core structures. To identify cores we followed two steps. First, using *Dendrogram* we find cores and determine their parameters (i.e. centre position, peak flux density, minor and major axis sizes, and position angle), towards each target clump. Then we made a more accurate measurement of the core parameters using the CASA task *imfit* adopting as initial guesses the parameters determined in the previous step. This approach has been employed in Li et al. (2020), and demonstrated to work very well.

In the first step, as inputs of the three key parameters to the *Dendrogram* algorithm we used: (i) $\text{min_value} = 2\sigma_{3\text{mm}}$, where $\sigma_{3\text{mm}}$ is the noise level. In practice, we first run the *Dendrogram* algorithm to identify the compact cores in the 3 mm continuum image, then subtract them to obtain a nearly flat residual 3 mm image, and finally estimate the rms level of the residual image as the noise ($\sigma_{3\text{mm}}$) of the original 3 mm continuum image as the starting level to construct the dendrogram. In addition, we used (ii) $\text{min_delta} = \sigma_{3\text{mm}}$, the default value in the algorithm to recognize an independent leaf (i.e. the core structure); and (iii) $\text{min_npix} = N$ pixels, where N was chosen to be equivalent to the area of one beam (depending on the pixel and beam sizes of the image of each target clump). This was imposed to ensure a core to be resolvable. Given the above three parameters, the *Dendrogram* algorithm automatically locates the cores, and measures their parameters (see above). To avoid spurious cores, we performed a post-selection to ensure the selected cores must have peak flux greater than $5\sigma_{3\text{mm}}$.

In the second step, we used the parameters of the post-selected cores determined in the first step as inputs to run the CASA *imfit* task to accurately identify the cores, and measure their parameters, including peak position, major and minor sizes (FWHM_{maj} , FWHM_{min}), position angle (PA), peak flux ($F_{3\text{mm}}^{\text{p}}$), and integrated flux ($F_{3\text{mm}}^{\text{int}}$). The measurement uncertainties on the flux are given by CASA *imfit*. The uncertainties related to the missing flux effect are not taken into account in this work since it is found not significant for compact cores by comparing the 12M data with the combined 12M + 7M data ([Paper I](#)). Through careful visual inspection of the 3 mm continuum map overlaid with the identified cores, we removed a few fake cores with poorly fitted shapes. Such cores tend to have a large aspect ratio (>3) between the major and minor axis sizes, appearing either as a filamentary structure or as a diffuse emission feature. As a result, we finally obtained 453 cores from the 146 ATOMS target clumps.

The entire extraction procedure is illustrated in Fig. 1, where we take the target clump 118032–2032 as an example of identifying the cores that are contained. The procedure starts from the compact core extraction on the 3 mm continuum map with *Dendrogram* in the left-hand panel, and then moves to the finer core fitting with CASA *imfit* in the right-hand panel. Due to the marginal angular resolution (~ 2 arcsec) and small intensity contrast in the 3 mm continuum emission maps, some cores may be missed in the core extraction. For example, comparing with the higher intensity-contrast range HC_3N (11–10) emission map, one can see that one molecular core between G009.6197+00.1940 and G009.6207+00.1950 revealed by HC_3N (11–10) is missing in the core extraction procedure. This is

²The frequency ranges of the eight SPWs are [86.311, 86.369] GHz for SPW 1, [86.725, 86.784] GHz for SPW 2, [86.818, 86.876] GHz for SPW 3, [87.288, 87.346] GHz for SPW 4, [88.603, 88.661] GHz for SPW 5, [89.159, 89.218] GHz for SPW 6, [97.530, 99.404] GHz for SPW 7, and [99.468, 101.341] GHz for SPW 8.

³<https://dendrograms.readthedocs.io/en/stable/>

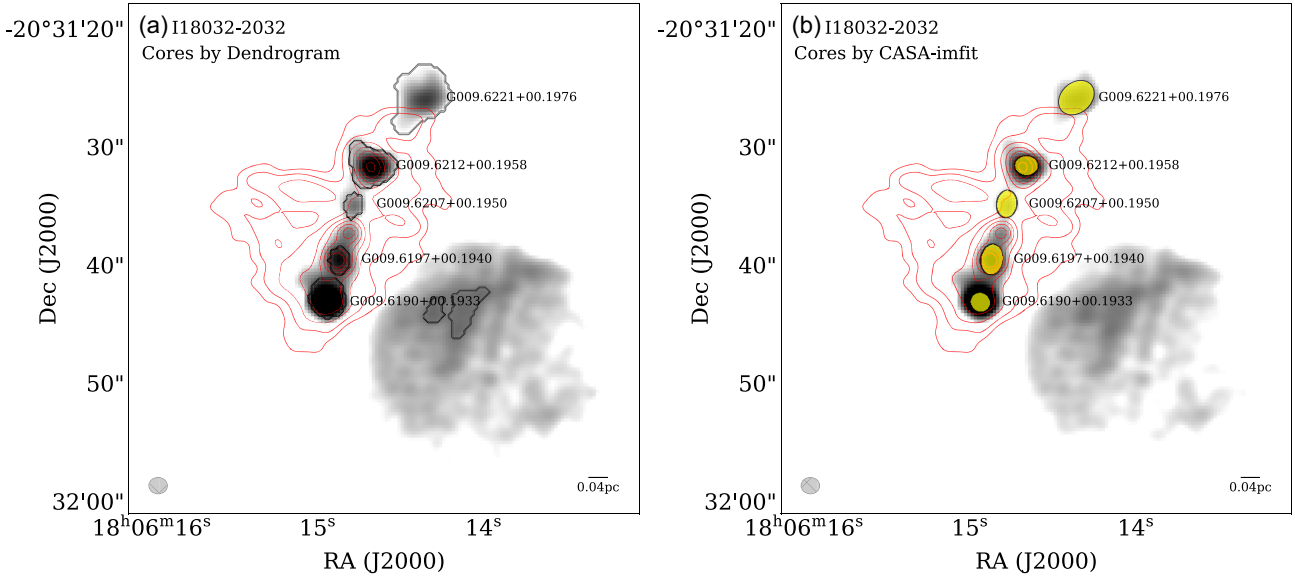


Figure 1. Illustration of the core extraction. (a) The preliminary extraction of seven core structures (in black contour) in the target clump I18032–2302 from 3 mm continuum emission by *Dendrogram*. (b) The final extraction of five finer cores (in yellow ellipses) by *CASA-imfit*. In both panels, 3 mm continuum (in grey-scale) is overlaid with the velocity-integrated intensity of HC_3N (in red contours) over $[-35, 45] \text{ km s}^{-1}$. The cores labelled with *Galactic* names are the cores remaining after post-selection (see text). One core, between G009.6197+00.1940 and G009.6207+00.1950, is missing in the second extraction procedure due to either insufficient angular resolution or low contrast against its background. Note that the contour levels start at 3 rms with the steps set by a dynamically determined power law of the form $D = 3 \times N^p + 2$, where D is the dynamic range of the intensity map (i.e. the ratio between the peak and the rms noise), N is the number of the contours used (eight in this case), and p the contour power-law index. The beam of the 3 mm continuum map and a scale bar of 0.04 pc are indicated at the bottom left and bottom right of each panel, respectively.

because that core cannot be separated from its adjacent cores or background emission in the 3 mm continuum emission map but emerges in higher intensity-contrast HC_3N (11–10) emission map. This issue very rarely occurs in the full sample and predominantly affects low flux (so presumably low-mass) core (specifically, those in the ‘unknown core’ category that contains the cores with unknown nature of star formation so far, see Section 3.3). Hence, the analysis in what follows is largely unaffected by this issue.

In addition, with both *Dendrogram* and *CASA-imfit* used in the core extraction, several major parameters measured from both methods can be compared. Actually, we find that the flux and position measurements are consistent in both methods. However, the size measured by *CASA-imfit* is larger than that in *Dendrogram* mainly for the low-flux cores (i.e. in the ‘unknown cores’ category, see Section 3.3). The position angle measurement in *CASA-imfit* generally matches the 3 mm continuum dense structure orientation better than that in *Dendrogram*. These differences mainly come from the different methods of parameter determination used in *Dendrogram* and *CASA-imfit*. In the former, the size and position angle parameters are derived from the intensity-weighted second moments along the two spatial dimensions (Rosolowsky et al. 2008) while in the latter the two parameters are directly derived from Gaussian fitting to the unweighted intensity spatial distribution. Therefore, to better reflect the core geometry as seen in 3 mm continuum, we prefer to use the core parameters measured by *CASA-imfit*.

3.2 Search for candidate HMCs and H/UC-H II cores

One important step towards understanding high-mass star formation is to search for the massive cores that harbour the early stages of high-mass star and cluster formation. Clues to high-mass star formation can be inferred from the properties of massive starless

and/or prestellar cores (e.g. Yuan et al. 2017), the spectral energy distribution (SED) of HMPOs (e.g. Rathborne et al. 2005), HMCs (e.g. Kurtz et al. 2000; Rathborne et al. 2011), H/UC-H II regions (e.g. Churchwell 2002; Hoare et al. 2007), and some masers (e.g. Class II CH_3OH masers), which exclusively trace high-mass star formation (e.g. Wang et al. 2014). Here, we focus on the search of massive cores associated with HMCs and H/UC-H II regions. The advantage of the ATOMS survey strategy is that the two wide SPWs 7 and 8 cover many transitions of the COMs (i.e. HMC tracers; see Fig. 12 of Paper I) and the $\text{H}40\alpha$ transition (i.e. the ionized gas tracer). In what follows, we discuss the search for H/UC-H II cores in Section 3.2.1 and COM-containing cores in Section 3.2.2 that could correspond to candidate HMCs.

3.2.1 Search for H/UC-H II cores

The entire procedure of the H/UC-H II core search is illustrated in Fig. 2, where the core G009.6190+00.1933 from the clump I18032–2302 is taken as an example. In Fig. 2(a), the velocity-integrated intensity of $\text{H}40\alpha$ emission over a velocity range of 150 km s^{-1} centred at the systemic velocity of the core (in both the grey-scale and red contours) is overlaid with the 3 mm continuum (cyan contours). The core in question is indicated by the white ellipse. Fig. 2(b) presents the beam-averaged spectrum of $\text{H}40\alpha$ at the core centre. If the $\text{H}40\alpha$ emission appears compact and is spatially associated with the compact continuum core, the core will be classified as an H/UC-H II core. Following this rule, we finally obtain 89 H/UC-H II cores. Their parameters determined by *CASA-imfit* are given in Table 1, including the clump and core names (i.e. columns 1 and 2), the measured sizes of the major and minor axes (i.e. columns 4 and 5), the deconvolved sizes of both axes (i.e. column 6), the core-integrated 3 mm flux (i.e. column 7),

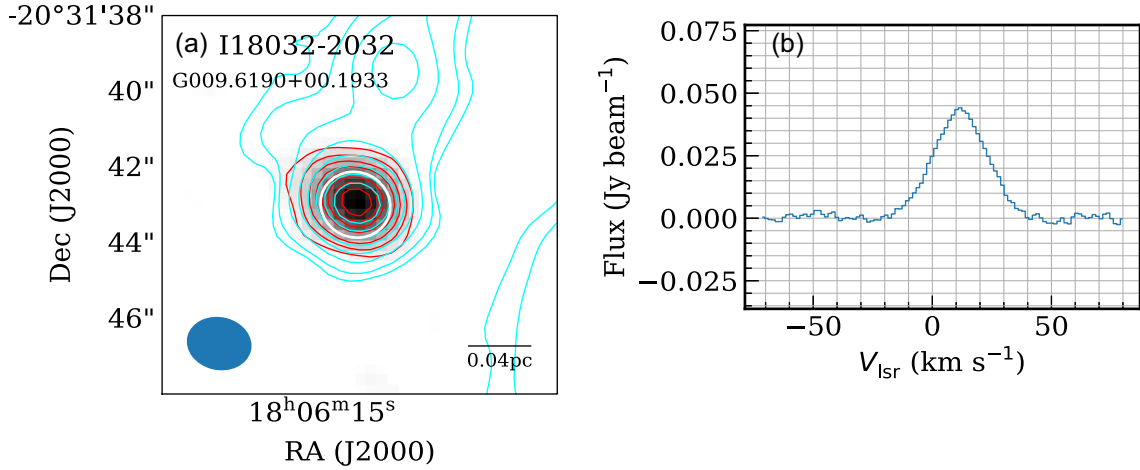


Figure 2. Illustration of the H/UC-H II core search. (a) Velocity-integrated intensity map of H40α (in both grey-scale and red contours) over $[-70, 80]$ km s⁻¹ for the core G009.6190+00.1933 (in white ellipse) in the target clump I18032–2032. The map is centred at the core peak with a size of 10 arcsec. For comparison, the 3 mm continuum is also superimposed in cyan contours. The synthesized beam and a scale bar of 0.04 pc are drawn at the bottom of the panel. (b) Beam (2 arcsec)-averaged spectrum of H40α at the core centre. G009.6190+00.1933 is an H/UC-H II core as characterized by strong and compact H40α emission.

the core-peak flux (i.e. column 8). The radii (i.e. column 9) of the cores were calculated assuming the same distance as that of their natal clumps (see Paper I). In addition, we measured the line peak intensity ($F_{H40\alpha}^p$), the velocity centroid ($V_{lsr, H40\alpha}$), and the line width ($\Delta V_{H40\alpha}$) of the H40α line for each H/UC-H II core by fitting a single-Gaussian component to the beam-averaged spectrum of H40α. For comparison (see Section 4.5), the deconvoluted size of the H/UC-H II cores from H40α emission was also measured with CASA-*imfit* and the corresponding radius ($\Delta R_{core}^{H40\alpha}$) was computed as above. These parameters can also be found in last five columns of Table 1. The $\Delta R_{core}^{H40\alpha}$ and $\Delta V_{H40\alpha}$ together suggest that the cores with associated compact H40α emission are indeed H/UC-H II sources (see Section 4.5).

3.2.2 Search for COM-containing cores

To quantitatively describe the chemical richness of a core, we use the number of lines (N_{line}) detected in the two line-scan windows SPWs 7 and 8. These spectral windows cover tens of transitions from COMs (e.g. CH₃OCHO, CH₃CHO, CH₃OH, C₂H₅CN) that are usually detectable in an HMC, and only two or three transitions of non-HMC-tracer molecules (i.e. CS (2–1), SO (3–2) or H₂CO, see fig. 12 of Paper I). Counting the emission lines is therefore useful in a search for the COM-containing cores.

The number of emission lines was determined by searching for emission peaks above the 3σ level across the entire spectrum of the two SPWs, which was also adopted in Sánchez-Monge et al. (2017). In practice, the 1σ level was calculated individually for each source from the standard deviation of the amplitudes in the line-free emission channels usually located on either or both ends of the wide SPWs. This method works very well for the sources without a forest of detectable lines. However, it is generally conservative for the sources (e.g. the one in Fig. 3a) with a forest of detectable lines since it is difficult to find completely line-free channels, leading to an overestimate of the sigma level. To avoid counting spurious peaks, for example, from the broad-line wings of the spectrum of the outflow sources, the separation between the nearest neighbours was required to be at least five velocity channels (i.e. ~ 8 km s⁻¹). In addition,

a valid peak detection was required to have at least three channels above the 3σ level.

Fig. 3 displays an example of the search for the emission lines from the beam-averaged spectrum of the two SPWs at the peak position of two cores, demonstrating that our method works well in counting emission lines. Note that the two sources given in the example were in different ALMA observing clusters with the centre frequencies slightly different, so that the frequency ranges of the sources are different accordingly. The final COM-containing cores are defined by the following two criteria, $N_{line} \geq 5$ and a compact feature of HC₃N emission. The former ensures at least one transition of the COMs for each candidate (see above, and Fig. 3), while the latter requires each candidate to be a compact source in molecular line emission. Fig. 4 presents an example of the compact COM-containing core in HC₃N emission. Fig. 4(a) shows the velocity-integrated intensity of HC₃N (in both grey-scale and red contours) over a velocity extent of 80 km s⁻¹ centred at the systemic velocity of the core (in white ellipse) and 3 mm continuum (in cyan contours), while Fig. 4(b) displays the beam-averaged spectrum of HC₃N at the core centre.

Following the above two criteria, 138 COM-containing cores were selected. We find that 48 out of these 138 COM-containing cores are spatially associated with H/UC-H II cores (positions matching within a radius of 2 arcsec, comparable to the beam size), and thus are H/UC-H II-associated candidate HMCs (cHMCs, hereafter). Note that the HMC is a definition from the molecular chemistry study, and has been observationally found to overlap the H/UC-H II regions. Given their masses and densities (see Section 4.2), most of the remaining 90 COM-containing cores are cHMCs and they all are thus treated as HMC candidates for simplicity.

To make a robust sample of cHMCs, the 138 COM-containing cores are divided into two groups of cHMCs using the values of N_{line} . As shown in Fig. 5, one group, regarded as strong-cHMCs (s-cHMCs), contains 56 high-level COM-containing cores with $N_{line} \geq 20$, and the other group, regarded as weak-cHMCs (w-cHMCs), contains 82 low-level COM-containing cores with $5 \leq N_{line} < 20$. In each group, there are 24 cHMCs associated with H/UC-H II cores. The threshold $N_{line} = 20$ is somewhat arbitrary, based on our visual inspection of the spectra of all cHMCs, but does not matter too much in the following statistical analysis. Moreover, to distinguish

Table 1. Parameters of the H/UC-H II cores.

Name	Gal. name	D (kpc)	FWHM maj (arcsec) \times min (arcsec)	PA (deg)	FWHM _{dec} maj (arcsec) \times min (arcsec)	$F_{3\text{mm}}^{\text{int.}}$ (mJy)	$F_{3\text{mm}}^p$ (mJy beam ⁻¹)	R_{core} (10 ⁻² pc)	$F_{\text{H40}\alpha}^p$ (mJy beam ⁻¹)	$V_{\text{lsr, H40}\alpha}$ (km s ⁻¹)	$\Delta V_{\text{H40}\alpha}$ (km s ⁻¹)	FWHM _{dec} maj (arcsec) \times min (arcsec)	$R_{\text{H40}\alpha}^{\text{dec}}$ (10 ⁻² pc)
clump	core												
I09002–4732	G268.4225–00.8491	2.0 ^{+0.7} _{-0.8}	4.6 \times 4.0	15.8	4.5 \times 3.6	3058.0 \pm 29.0	329.7 \pm 2.9	2.0 ^{+0.6} _{-0.8}	242.7 \pm 1.5	-2.0 \pm 1.6	34.5 \pm 1.6	4.4 \times 3.3	1.9 ^{+0.6} _{-0.7}
I12320–6122	G300.9688+01.1475 ^a	3.1 ^{+0.6} _{-0.5}	2.6 \times 2.4	11.2	1.6 \times 0.9	477.4 \pm 9.5	333.5 \pm 4.3	0.9 ^{+0.2} _{-0.1}	191.6 \pm 2.0	-41.4 \pm 1.6	37.3 \pm 1.6	1.5 \times 0.7	0.8 ^{+0.2} _{-0.1}
I12326–6245	G301.1364–00.2257 ^a	5.6 ^{+0.5} _{-0.8}	2.9 \times 2.2	172.0	1.8 \times 1.0	2003.0 \pm 49.0	1351.0 \pm 21.0	1.8 ^{+0.1} _{-0.3}	526.5 \pm 3.2	-63.8 \pm 1.6	55.0 \pm 1.6	1.6 \times 0.8	1.6 ^{+0.1} _{-0.2}
I12383–6128	G301.7314+01.1033	2.8 ^{+0.7} _{-0.6}	4.1 \times 3.7	17.0	3.3 \times 3.2	96.7 \pm 2.2	27.4 \pm 0.5	2.2 ^{+0.5} _{-0.5}	32.8 \pm 0.9	-38.3 \pm 1.6	24.7 \pm 1.6	3.2 \times 2.3	1.9 ^{+0.4} _{-0.4}
I12572–6316	G303.9302–00.6879 ^a	11.5 ^{+0.3} _{-0.3}	2.8 \times 2.3	7.3	1.5 \times 1.3	48.2 \pm 1.0	32.3 \pm 0.5	3.9 ^{+0.1} _{-0.1}	17.9 \pm 0.8	24.0 \pm 1.6	24.5 \pm 1.6	1.4 \times 1.0	3.2 ^{+0.1} _{-0.1}
I13080–6229	G305.1961+00.0331 ^a	2.3 ^{+0.6} _{-0.6}	12.6 \times 8.9	78.0	12.5 \times 8.6	2014.0 \pm 46.0	76.6 \pm 1.7	5.8 ^{+1.5} _{-1.4}	70.0 \pm 1.8	-35.8 \pm 1.6	28.0 \pm 1.6	9.3 \times 5.2	3.9 ^{+1.0} _{-1.0}
I13291–6249	G307.5605–00.5870	7.8 ^{+0.5} _{-0.5}	7.6 \times 5.9	24.2	7.3 \times 5.6	562.4 \pm 9.3	53.1 \pm 0.8	12.1 ^{+0.8} _{-0.7}	51.0 \pm 1.5	-38.2 \pm 1.6	28.6 \pm 1.6	4.4 \times 3.2	7.1 ^{+0.4} _{-0.4}
I13471–6120	G309.9206+00.4788 ^a	6.6 ^{+0.6} _{-0.6}	2.4 \times 1.9	71.9	1.0 \times 1.0	848.0 \pm 21.0	648.0 \pm 10.0	1.6 ^{+0.1} _{-0.1}	347.8 \pm 1.8	-57.0 \pm 1.6	34.1 \pm 1.6	1.1 \times 1.0	1.6 ^{+0.1} _{-0.1}
I14013–6105	G311.6273+00.2902	3.5 ^{+0.7} _{-0.5}	5.0 \times 4.8	155.0	4.8 \times 4.3	1198.0 \pm 34.0	171.0 \pm 4.3	3.9 ^{+0.8} _{-0.5}	133.1 \pm 1.5	-58.3 \pm 1.6	37.6 \pm 1.6	4.1 \times 3.0	3.0 ^{+0.6} _{-0.4}
I14050–6056	G312.1084+00.3093	3.1 ^{+0.5} _{-0.5}	5.4 \times 4.8	78.3	5.0 \times 4.5	487.0 \pm 16.0	63.3 \pm 1.9	3.5 ^{+0.6} _{-0.6}	50.7 \pm 2.0	-46.5 \pm 1.6	41.0 \pm 1.6	3.5 \times 2.6	2.2 ^{+0.4} _{-0.4}

Note. ^aThis core is associated with a COM-containing core or COM-containing cores. FWHM is the measured core size while FWHM_{dec} is the corresponding deconvolved size. R_{core} follows the form of $R_{\text{eff}}/3600 \times \pi/180 \times D$, where R_{eff} is the effective radius equal to $\sqrt{\text{FWHM}_{\text{dec}}/2}$, and D is the distance of the core. The errors on R_{core} and $R_{\text{H40}\alpha}^{\text{dec}}$ mainly result from the distance uncertainties and those on the fluxes arise from the 2D Gaussian fitting in the core extraction. Only a small portion of the data are provided here and the full table will be available as supplementary material.

the cHMCs between with H/UC-H II and without H/UC-H II cores, 32 pure s-cHMCs and 58 pure w-cHMCs are separated from the entire sample of the 138 cHMCs (see Fig. 5), and compiled into the two catalogues, one called pure s-cHMCs and the other one pure w-cHMCs. That is, pure cores refer to those without H/UC-H II signatures. The parameters of the two catalogues are tabulated in Table 2 for pure s-cHMCs, and in Table 3 for pure w-cHMCs, including the clump and core names (columns 1 and 2), the measured sizes of the major and minor axes and the position angle (columns 4 and 5), the deconvolved sizes of both axes (column 6), the core-integrated 3 mm flux (column 7), the core-peak flux (column 8), the radius (column 9), the number of lines (column 10), the mass (column 11), the number density (column 12), and the mass surface density (column 13). The last three parameters were calculated following the equations in Appendix B for a fixed temperature of 100 K typical of a hot core. The median radius is 1.6×10^{-2} pc in an interquartile range (IQR)⁴ of $[1.0, 2.3] \times 10^{-2}$ pc for the catalogue of 32 pure s-cHMCs, and 2.6×10^{-2} pc in an IQR of $[1.7, 5.2] \times 10^{-2}$ pc for the catalogue of 58 pure w-cHMCs. These results suggest that most of the 90 COM-containing cores are indeed compact.

Note that contamination from free-free radiation to continuum may not be trivial for the cores associated with H/UC-H II. In practice, we calculated at a temperature of 100 K the masses for two types of cores: pure cHMCs and cHMCs associated with H/UC-H II sources (i.e. compact H40 α emission). In comparison, we found that the mass of the latter cores ($\sim 10^4 M_{\odot}$) is on average about two magnitude higher than that of the former cores ($\sim 10^2 M_{\odot}$). This result suggests that contamination by free-free radiation is non-negligible in calculating the flux-related parameters. Since this contamination cannot be accurately determined, we do not give the estimate of the continuum flux-related parameters for the H/UC-H II cores such as the mass.

3.3 Compact dense core list not associated with either cHMC or H/UC-H II

As described above, we extract 453 cores in 3 mm continuum. Of these, 89 are classified as H/UC-H II, 32 as s-cHMC, and 58 as w-cHMC. The remaining 274 cores lack observationally quantifiable metrics for association into any of the three categories above. Hence, we refer to these as ‘unknown cores’ (see Table 4). These ‘unknown cores’ could be simple continuum flux condensations in an earlier evolutionary stage. Indeed, chemically poor HMPOs and candidate massive prestellar cores have been detected in some massive protoclusters (e.g. Liu et al. 2017). However, it is also possible that some (to date) unknown but potentially significant fraction of these ‘unknown cores’ may actually be cHMC or H/UC-H II cores that fall below our sensitivity limit due to spatial resolution, distance, and/or flux limitations. These will be investigated in more detail in future work.

For reference, the mass, number density, and mass surface density of the ‘unknown’ cores (see Table 4) were calculated (see Appendix B) assuming a dust temperature of 25 K. This temperature is comparable to the median value of the clumps that do not have cHMC or H/UC-H II signatures. As a result, these cores have a median radius of 2.4×10^{-2} pc in an IQR of $[1.3, 4.6] \times 10^{-2}$ pc, a median mass of $19 M_{\odot}$ in an IQR of $[5, 99] M_{\odot}$, a median number density of $0.6 \times 10^6 \text{ cm}^{-3}$ in an IQR of $[0.3, 1.3] \times 10^6 \text{ cm}^{-3}$,

⁴The IQR is a measure of statistical dispersion, being equal to the difference between 75th and 25th percentiles, or between upper and lower quartiles.

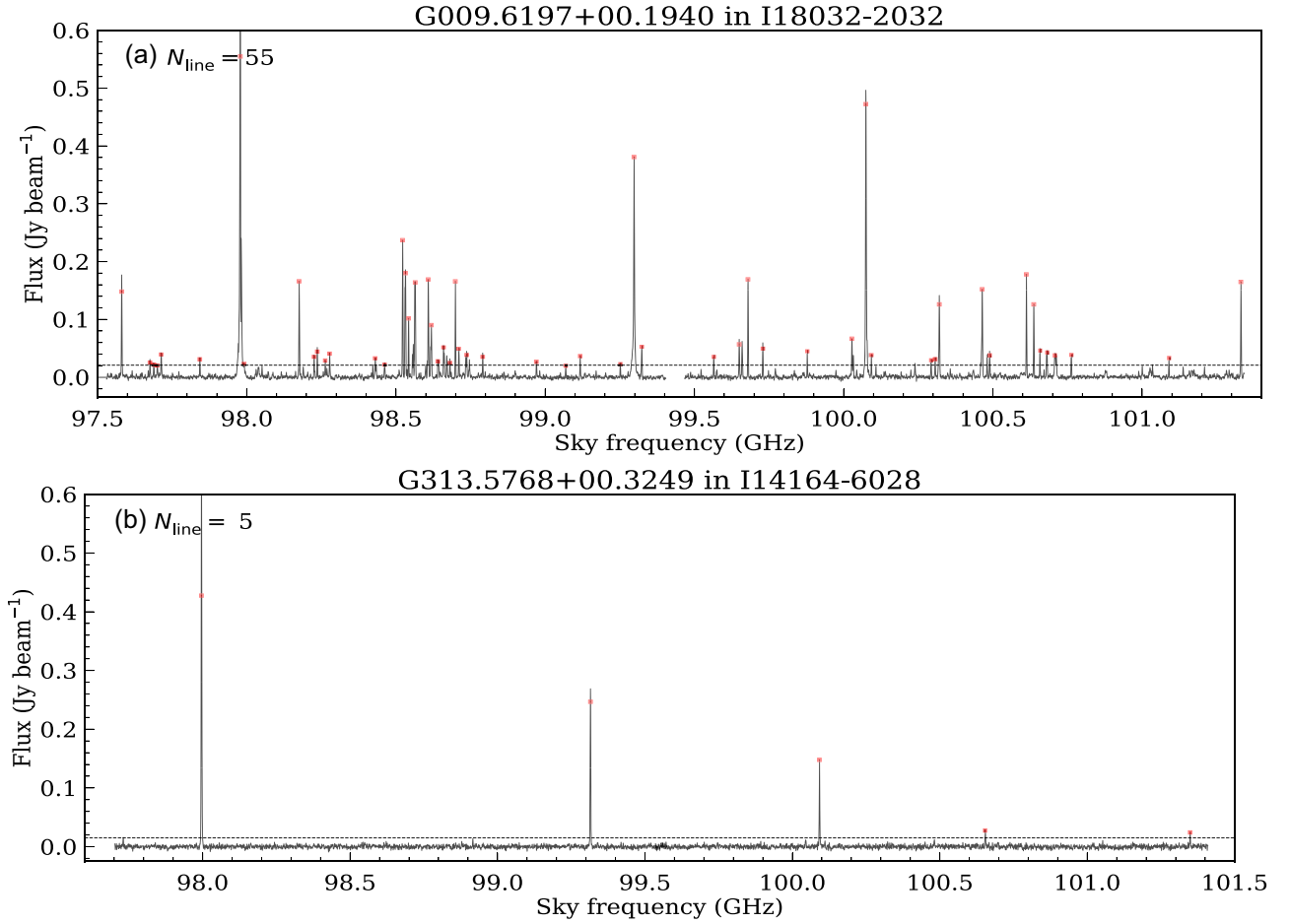


Figure 3. Beam (2 arcsec)-averaged spectrum of SPWs 7 and 8 at the emission peak of the two COM-containing cores (see definition in text) from two ATOMS clumps. The core and clump names are labelled on top of each panel. The red squares indicate the emission lines above 3 rms and their total number is given to the parameter N_{line} . The horizontal dashed line indicates the 3 rms noise level across the two spectral windows. For reference, the five lines of panel (b) in an order of decreasing intensity are CS (2–1), SO (3–2), HC₃N (11–10), CH₃OH, and H₂CO, respectively.

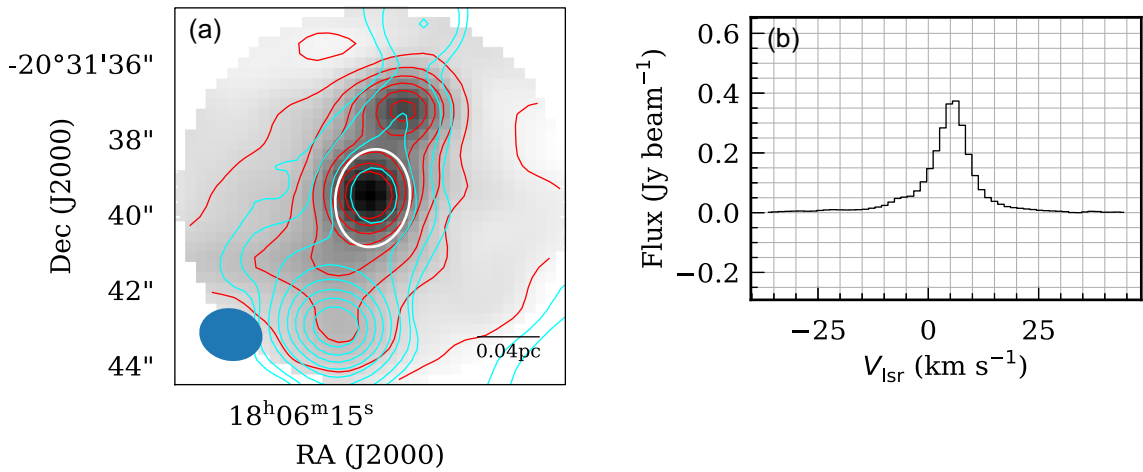


Figure 4. (a) Velocity-integrated intensity map of HC₃N over $[-35, 45]$ km s⁻¹ (in both grey-scale and red contours) for the G009.6197+00.1940 core (in white ellipse) and its surroundings. The map is centred at the core peak with a size of 10 arcsec. For comparison, the 3 mm continuum is superimposed in cyan contours. The synthesized beam and a scale bar of 0.04 pc are drawn at the bottom left of the panel. (b) Beam (2 arcsec)-averaged spectrum of HC₃N at the core centre (i.e. the white ellipse in panel a).

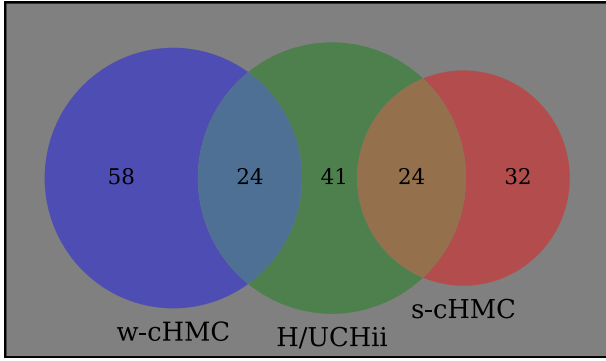


Figure 5. Venn diagram showing the relation between 138 COM-containing and 89 H/UC-H II cores. The 138 COM-containing cores consist of 58 pure and 24 H/UC-H II-associated w-cHMCs and 32 pure and 24 H/UC-H II-associated s-cHMCs.

and a median mass surface density of 2.2 g cm^{-2} in an IQR of $[1.1, 4.3] \text{ g cm}^{-2}$. The properties of the ‘unknown’ cores will be more thoroughly investigated in a forthcoming paper.

4 DISCUSSION

4.1 Limitations in dense core extraction and classification

Our ATOMS survey consists of a uniform sensitivity and angular resolution survey towards regions located over a large range of distances. Consequently, the dense core extraction, as discussed above, is affected by variations of factors of order of ~ 10 in distance across the sample. Here, we focus on the most robust (brightest) portion of our sample (i.e. w-cHMC, s-cHMC, and H/UC-H II cores, see Section 3.2).

Fig. 6 shows how the radii (R) and masses (M) depend on distances (D) for the three groups (w-cHMC, s-cHMC, and H/UC-H II) of dense cores. There are clear increasing trends in R versus D and M versus D relations, further suggesting that distant sources were not well resolved. With future higher resolution observations, we expect that these distant cores with sizes of $\sim 0.1 \text{ pc}$ would break down to smaller sub-cores or condensations with sizes of $\sim 0.01 \text{ pc}$. Moreover, the masses of the w-cHMC and s-cHMC cores can be found in Fig. 6(b) far above the detection limit, which is in agreement with the nature of bright emission in continuum for those two types of cores.

Moreover, beam dilution impairs detection of hot and H/UC-H II cores at large distances. Therefore, we assume a smaller detection rate of such cores at larger distances. We tested this assumption by re-examining the detectability of the COM and H40 α lines after pushing the observations of the relatively nearby sources to a large distance. In practice, we smoothed the cubes of SPWs 7 and 8 for 10 cHMCs and 5 H/UC-H II cores at distances $\leq 2 \text{ kpc}$ to the larger beam size that was transferred from the ATOMS typical angular resolution ($\sim 1.6 \text{ arcsec}$) at a fixed far distance of 5 kpc . In consequence, only 30 per cent of the selected cHMC sources were re-identified using the same HMC identification method as in Section 3.2.2, while 80 per cent of the selected H/UC-H II sources have been re-identified. This result suggests that beam dilution affects the detectability for cHMCs more than it does for H/UC-H II cores at the ATOMS sensitivity.

Furthermore, Fig. 7 shows the distance distributions of three groups of clumps, which contain w-cHMC, s-cHMC, and H/UC-H II. For comparison, we also plot the distance distribution of the full sample. Indeed, we find that the candidate hot cores (w-cHMC and s-cHMC) tend to be located at nearer distances (Figs 7b and c), with

their cumulative distribution curves (in red and blue colours in Fig. 7a) above the curve (in grey colour) for the full sample.

We test whether the distance distributions of clumps containing different kinds of dense cores follow the same distribution as the full sample with Kolmogorov–Smirnov test. The unknown hypothesis is that two distributions follow the same distribution. The tests are consistent with the unknown hypothesis for the w-cHMCs group (P -value = 0.37) and H/UC-H II group (P -value = 0.67) due to their P -values being greater than 0.05, suggesting that the detection rates of w-cHMCs and H/UC-H II are not greatly affected by the distance issue. In contrast, the P -value for s-cHMCs is 0.06, indicating that the detection of s-cHMCs is significantly limited by distances. It may imply that the detection rate of s-cHMCs is significantly underestimated for distant sources. This is in line with the cumulative distribution function of distances (Fig. 7a). In the figure, clumps with s-cHMC appear to show smaller distances ($\sim 2.6 \text{ kpc}$ on average) than those ($\sim 3.2 \text{ kpc}$ on average) with w-cHMC, indicating that the detection of hot cores at large distances is incomplete.

4.2 Nature of the COM-containing and not H/UC-H II-associated cores

Since emission in the COM transitions covered by the ATOMS observations can be detected towards several distinct objects, including low/high-mass prestellar cores, HMPO cores, HMCs, and even hot corinos associated with low-mass protostars (e.g. Vastel et al. 2014; Jiménez-Serra et al. 2016; Soma et al. 2018; Csengeri et al. 2019; Molet et al. 2019; Hsu et al. 2020; Jorgensen, Bellocche & Garrod 2020), it is necessary to constrain the nature of the 90 COM-containing cores that are not associated with H/UC-H II (see Section 3.2.2). To this end, we investigate the distribution of both the mass (M_{core}) and the mass surface density (Σ_{core}) of the 90 COM-containing cores, as shown in Fig. 8. They have a median value of mass equal to $\sim 57 M_{\odot}$ in an IQR of $19\text{--}138 M_{\odot}$, and surface densities with a median value of $\sim 6.5 \text{ g cm}^{-2}$ in an IQR of $3.3\text{--}15.6 \text{ g cm}^{-2}$.

Theoretically, cores (on scale $< 0.1 \text{ pc}$) with $\Sigma_{\text{crit}} \geq 1 \text{ g cm}^{-2}$ will most likely form high-mass stars (Krumholz & McKee 2008). Searching from existing maser catalogue⁵ (Ladeyschikov, Bayandina & Sobolev 2019), we find that ~ 34 per cent of the 90 COM-containing cores are associated with Class II CH₃OH masers located within the core size (red dots in Fig. 8). Class II CH₃OH masers have been found to be exclusively associated with high-mass star-forming regions (e.g. Cyganowski et al. 2009). The spread of Σ_{core} for the cores associated with the CH₃OH masers is consistent with the theoretical prediction (i.e. the red dots versus the vertical dashed line in the figure). Following the threshold $\Sigma_{\text{crit}} \geq 1 \text{ g cm}^{-2}$, we find that ~ 96 per cent of the 90 COM-containing cores above it will most likely form high-mass stars. We note that there are six COM-containing cores, three of which are associated with Class II masers, above the threshold but with $M_{\text{core}} < 10 M_{\odot}$. We suggest that these cores could have been in the process of forming high-mass stars for a significant fraction of their lifetime, so that a large fraction of their mass has been accreted on to the central HMPOs. We also note that if the actual temperature of the 90 COM-containing cores is lower than 100 K , the possibility that they form high-mass stars will be higher. For reference, if the temperature is assumed to be 50 K , the mass and surface density will be about two times higher.

Overall, the probability that the identified COM-containing cores are associated with low-mass star forming objects (i.e. low-mass

⁵<http://maserdb.net/download.pl>

Table 2. Parameters of the COM-containing but not H/UC-H II-associated cores with $N_{\text{line}} \geq 20$ (or pure s-cHMC).

Name	Gal. name	D	FWHM	PA	FWHM _{dec}	$F_{3\text{mm}}^{\text{int}}$	$F_{3\text{mm}}^p$	R_{core}	N_{line}	$\log(M_{\text{core}})$	$\log(n_{\text{core}})$	Σ_{core}
clump	core	(kpc)	maj (arcsec) \times min (arcsec)	(deg)	maj (arcsec) \times min (arcsec)	(mJy)	(mJy beam ⁻¹)	(10 ⁻² pc)		(M _⊙)	(cm ⁻³)	(g cm ⁻²)
I08303–4303	G261.6444–02.0890	2.4 ^{+0.5} _{-0.6}	2.3 \times 2.1	58.0	1.9 \times 1.5	18.2 \pm 0.8	7.0 \pm 0.2	1.0 ^{+0.2} _{-0.2}	28	1.1 ^{+0.7} _{-0.7}	6.6 ^{+3.9} _{-3.9}	8.2
I08470–4243	G263.2500+00.5137	2.4 ^{+0.6} _{-0.6}	2.6 \times 2.1	147.6	2.1 \times 1.5	38.0 \pm 0.9	13.5 \pm 0.2	1.0 ^{+0.2} _{-0.3}	61	1.4 ^{+1.1} _{-1.1}	6.8 ^{+4.2} _{-4.2}	15.6
I13134–6242	G305.7992–00.2447	2.0 ^{+0.6} _{-0.4}	3.4 \times 2.7	171.5	2.4 \times 2.1	103.8 \pm 2.4	47.2 \pm 0.8	1.1 ^{+0.3} _{-0.2}	48	1.7 ^{+1.5} _{-1.3}	7.0 ^{+4.2} _{-4.1}	26.6
I13484–6100	G310.1436+00.7598	6.9 ^{+0.5} _{-0.7}	3.5 \times 3.2	77.0	2.8 \times 2.7	41.1 \pm 1.7	13.0 \pm 0.4	4.6 ^{+0.3} _{-0.5}	36	2.4 ^{+1.5} _{-1.7}	5.9 ^{+3.9} _{-4.0}	7.0
I14498–5856	G318.0482+00.0863	2.9 ^{+0.4} _{-0.3}	5.2 \times 2.6	48.3	4.0 \times 2.0	9.1 \pm 0.0	4.2 \pm 0.2	2.0 ^{+0.3} _{-0.2}	33	0.9 ^{+0.4} _{-0.3}	5.8 ^{+3.1} _{-3.1}	1.5
I14498–5856	G318.0497+00.0868	2.9 ^{+0.4} _{-0.3}	3.9 \times 3.0	63.4	3.5 \times 2.4	60.7 \pm 1.4	15.9 \pm 0.3	2.0 ^{+0.3} _{-0.2}	39	1.8 ^{+1.2} _{-1.1}	6.4 ^{+3.7} _{-3.6}	9.3
I15394–5358	G326.4747+00.7030	2.4 ^{+0.3} _{-0.2}	5.2 \times 4.4	140.9	4.9 \times 4.0	226.7 \pm 5.4	28.4 \pm 0.6	2.5 ^{+0.3} _{-0.2}	33	2.2 ^{+1.6} _{-1.4}	6.6 ^{+3.6} _{-3.5}	14.9

Note. The parameters FWHM, FWHM_{dec}, and R_{core} are derived in the same way as in Table 1. The errors on R_{core} , M_{core} , and n_{core} mainly result from the distance uncertainties, and those on the fluxes from the 2D Gaussian fitting in the core extraction. Only a small portion of the data are provided here and the full table will be available as supplementary material.

Table 3. Parameters of the COM-containing but not H/UC-H II-associated cores with $5 \leq N_{\text{line}} < 20$ (or pure w-cHMC).

Name	Gal. name	D	FWHM	PA	FWHM _{dec}	$F_{3\text{mm}}^{\text{int}}$	$F_{3\text{mm}}^p$	R_{core}	N_{line}	$\log(M_{\text{core}})$	$\log(n_{\text{core}})$	Σ_{core}
clump	core	(kpc)	maj (arcsec) \times min (arcsec)	(deg)	maj (arcsec) \times min (arcsec)	(mJy)	(mJy beam ⁻¹)	(10 ⁻² pc)		(M _⊙)	(cm ⁻³)	(g cm ⁻²)
I08448–4343	G263.7737–00.4326	1.4 ^{+0.6} _{-0.8}	3.5 \times 2.3	113.0	3.1 \times 2.0	5.6 \pm 1.2	1.3 \pm 0.2	0.9 ^{+0.4} _{-0.5}	9	0.1 ^{+0.1} _{-0.1}	5.9 ^{+2.9} _{-3.0}	1.2
I09018–4816	G269.1523–01.1279	2.8 ^{+0.6} _{-0.7}	4.3 \times 2.7	178.5	4.1 \times 2.2	51.9 \pm 2.1	8.6 \pm 0.3	2.1 ^{+0.4} _{-0.5}	12	1.7 ^{+1.3} _{-1.4}	6.3 ^{+3.8} _{-3.8}	7.4
I10365–5803	G286.2074+00.1698	3.4 ^{+0.6} _{-0.6}	3.6 \times 3.0	81.9	3.1 \times 2.5	31.7 \pm 1.0	9.2 \pm 0.2	2.3 ^{+0.4} _{-0.4}	10	1.6 ^{+1.1} _{-1.2}	6.1 ^{+3.6} _{-3.7}	5.3
I11298–6155	G293.8274–00.7461	9.4 ^{+0.4} _{-0.3}	4.0 \times 3.0	92.8	3.6 \times 2.3	41.8 \pm 1.9	11.8 \pm 0.4	6.6 ^{+0.3} _{-0.2}	10	2.6 ^{+1.6} _{-1.5}	5.7 ^{+3.8} _{-3.7}	6.5
I11332–6258	G294.5115–01.6213	1.1 ^{+0.7} _{-0.6}	5.9 \times 4.7	81.2	5.6 \times 4.2	30.5 \pm 2.0	3.8 \pm 0.2	1.3 ^{+0.8} _{-0.7}	6	0.6 ^{+0.7} _{-0.7}	5.9 ^{+2.8} _{-2.8}	1.7
I12326–6245	G301.1363–00.2235	5.6 ^{+0.5} _{-0.8}	3.4 \times 2.1	178.0	2.3 \times 1.1	212.0 \pm 52.0	130.0 \pm 21.0	2.2 ^{+0.2} _{-0.3}	7	2.9 ^{+2.1} _{-2.3}	7.1 ^{+4.9} _{-5.2}	108.1
I13079–6218	G305.2084+00.2060	3.0 ^{+0.6} _{-0.7}	3.5 \times 3.0	22.8	2.8 \times 2.2	97.0 \pm 2.6	39.1 \pm 0.8	1.8 ^{+0.4} _{-0.4}	11	2.0 ^{+1.6} _{-1.7}	6.7 ^{+4.2} _{-4.3}	20.3
I13079–6218	G305.2022+00.2072	3.0 ^{+0.6} _{-0.7}	3.5 \times 3.2	105.0	3.0 \times 2.0	20.5 \pm 2.8	7.8 \pm 0.8	1.8 ^{+0.4} _{-0.4}	11	1.3 ^{+0.9} _{-1.0}	6.0 ^{+3.5} _{-3.6}	4.4

Note. The parameters FWHM, FWHM_{dec}, and R_{core} are derived in the same way as in Table 1. The errors on R_{core} , M_{core} , and n_{core} mainly result from the distance uncertainties, and those on the fluxes from the 2D Gaussian fitting in the core extraction. Only a small portion of the data are provided here and the full table will be available as supplementary material.

Table 4. Parameters of the ‘unknown’ cores (i.e. no signature of H/UC-H II nor cHMC).

Name	Gal. name	D	FWHM	PA	FWHM _{dec}	$F_{3\text{mm}}^{\text{int}}$	$F_{3\text{mm}}^p$	R_{core}	$\log(M_{\text{core}})$	$\log(n_{\text{core}})$	Σ_{core}
clump	core	(kpc)	maj (arcsec) \times min (arcsec)	(deg)	maj (arcsec) \times min (arcsec)	(mJy)	(mJy beam ⁻¹)	(10 ⁻² pc)	(M _⊙)	(cm ⁻³)	(g cm ⁻²)
I08076–3556	G253.2971–01.6156	1.0 ^{+0.5} _{-0.5}	5.1 \times 3.3	86.4	4.9 \times 3.0	11.6 \pm 1.2	1.3 \pm 0.1	0.9 ^{+0.4} _{-0.5}	0.1 ^{+0.1} _{-0.1}	5.9 ^{+2.5} _{-2.6}	0.9
I08076–3556	G253.2930–01.6113	1.0 ^{+0.5} _{-0.5}	2.2 \times 1.6	99.3	1.6 \times 1.1	18.8 \pm 0.5	9.8 \pm 0.2	0.3 ^{+0.1} _{-0.2}	0.3 ^{+0.3} _{-0.3}	7.2 ^{+3.9} _{-4.0}	13.0
I08303–4303	G261.6446–02.0899	2.4 ^{+0.5} _{-0.6}	4.9 \times 2.5	6.9	4.7 \times 2.0	23.7 \pm 1.4	3.6 \pm 0.2	1.8 ^{+0.4} _{-0.4}	1.1 ^{+0.8} _{-0.8}	6.0 ^{+3.3} _{-3.4}	2.9
I08303–4303	G261.6444–02.0876	2.4 ^{+0.5} _{-0.6}	4.0 \times 3.8	40.0	3.8 \times 3.5	15.2 \pm 1.6	1.8 \pm 0.2	2.1 ^{+0.5} _{-0.5}	1.0 ^{+0.6} _{-0.6}	5.7 ^{+3.0} _{-3.0}	1.3
I08448–4343	G263.7776–00.4332	1.4 ^{+0.6} _{-0.8}	5.2 \times 3.0	116.0	5.0 \times 2.8	8.8 \pm 1.7	1.1 \pm 0.2	1.3 ^{+0.6} _{-0.7}	0.3 ^{+0.2} _{-0.3}	5.6 ^{+2.7} _{-2.8}	0.7
I08448–4343	G263.7729–00.4364	1.4 ^{+0.6} _{-0.8}	2.5 \times 1.8	72.0	2.1 \times 1.1	2.1 \pm 0.8	0.9 \pm 0.2	0.5 ^{+0.2} _{-0.3}	–0.3 ^{+–0.4} _{–0.3}	6.0 ^{+3.0} _{-3.1}	1.1
I08448–4343	G263.7724–00.4366	1.4 ^{+0.6} _{-0.8}	2.2 \times 1.8	103.0	1.5 \times 1.3	1.9 \pm 0.7	0.9 \pm 0.2	0.5 ^{+0.2} _{-0.3}	–0.4 ^{+–0.5} _{–0.4}	6.0 ^{+3.0} _{-3.1}	1.1
I08448–4343	G263.7723–00.4350	1.4 ^{+0.6} _{-0.8}	4.4 \times 2.2	80.6	3.6 \times 1.8	7.3 \pm 0.0	5.8 \pm 0.2	0.9 ^{+0.4} _{-0.5}	0.2 ^{+0.1} _{-0.2}	6.5 ^{+3.6} _{-3.7}	1.3
I08448–4343	G263.7744–00.4328	1.4 ^{+0.6} _{-0.8}	2.5 \times 2.2	18.0	2.1 \times 1.6	11.5 \pm 0.9	3.9 \pm 0.2	0.6 ^{+0.3} _{-0.3}	0.4 ^{+0.3} _{-0.4}	6.5 ^{+3.5} _{-3.6}	3.9
I08448–4343	G263.7712–00.4363	1.4 ^{+0.6} _{-0.8}	3.2 \times 1.8	157.0	2.9 \times 1.1	2.4 \pm 0.9	0.8 \pm 0.2	0.6 ^{+0.3} _{-0.3}	–0.3 ^{+–0.4} _{–0.3}	5.8 ^{+2.9} _{-2.9}	0.9

Note. The parameters FWHM, FWHM_{dec}, and R_{core} are derived in the same way as in Table 1. The errors on R_{core} , M_{core} , and n_{core} mainly come from the distance uncertainties, those on the fluxes from the 2D Gaussian fitting in the core extraction. Only a small portion of the data are provided here, and the full table will be available as supplementary material.

prestellar cores and hot corinos) is rather low. Given that emission in the COM transitions available in the ATOMS data has been already observed towards other HMPO and HMC regions (Csengeri et al. 2019; Jorgensen, Belloche & Garrod 2020), we conclude that most (if not all) of the 90 COM-containing cores should be either HMPO cores or HMCs. Note that HMPOs are simply assumed here to be younger than HMCs although this time-scale relation has yet to be determined.

To distinguish definitely between the two types of high-mass star-forming cores requires knowledge of their temperatures. The temperature is usually computed from a rotational diagram analysis

of multiple transitions of one or more COM species (e.g. CH₃OCHO and CH₃OH). However, the range of upper level energy of the COM lines covered in the ATOMS survey is too narrow to permit a reliable temperature determination. This thus requires follow-up observations of higher frequency transitions that arise from higher energy levels. For simplicity, hereafter, we treat the 90 COM-containing cores as cHMCs (see Section 3.2.2) since the detection rate of COMs towards HMCs is presumably higher than towards HMPOs.

In short, most of the 90 COM-containing cores are presumably in the very early stages of high-mass star formation. We include them

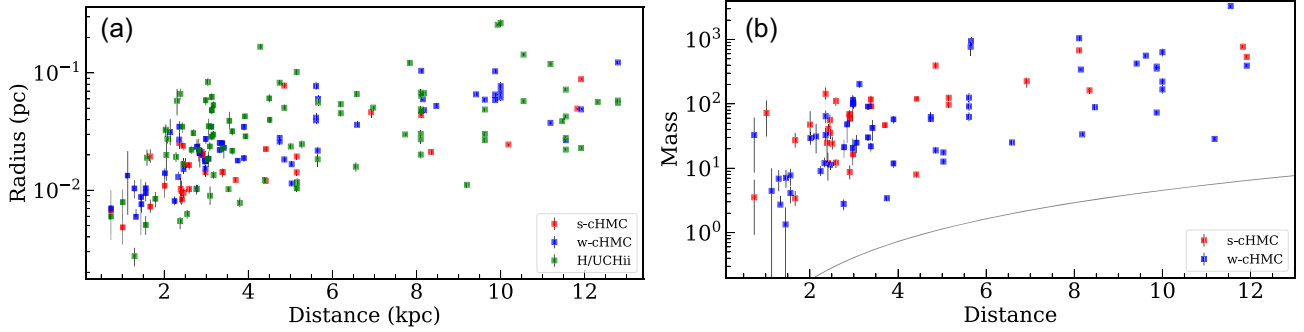


Figure 6. (a) Radius versus distance for w-cHMC, s-cHMC, and H/UC-H II cores. (b) Mass versus distance for w-cHMC and s-cHMC cores. H/UC-H II cores are not shown here since their masses cannot be estimated from 3 mm continuum due to free-free contamination (see Section 3.2.2). The curve is the mass sensitivity distribution as a function of the distance.

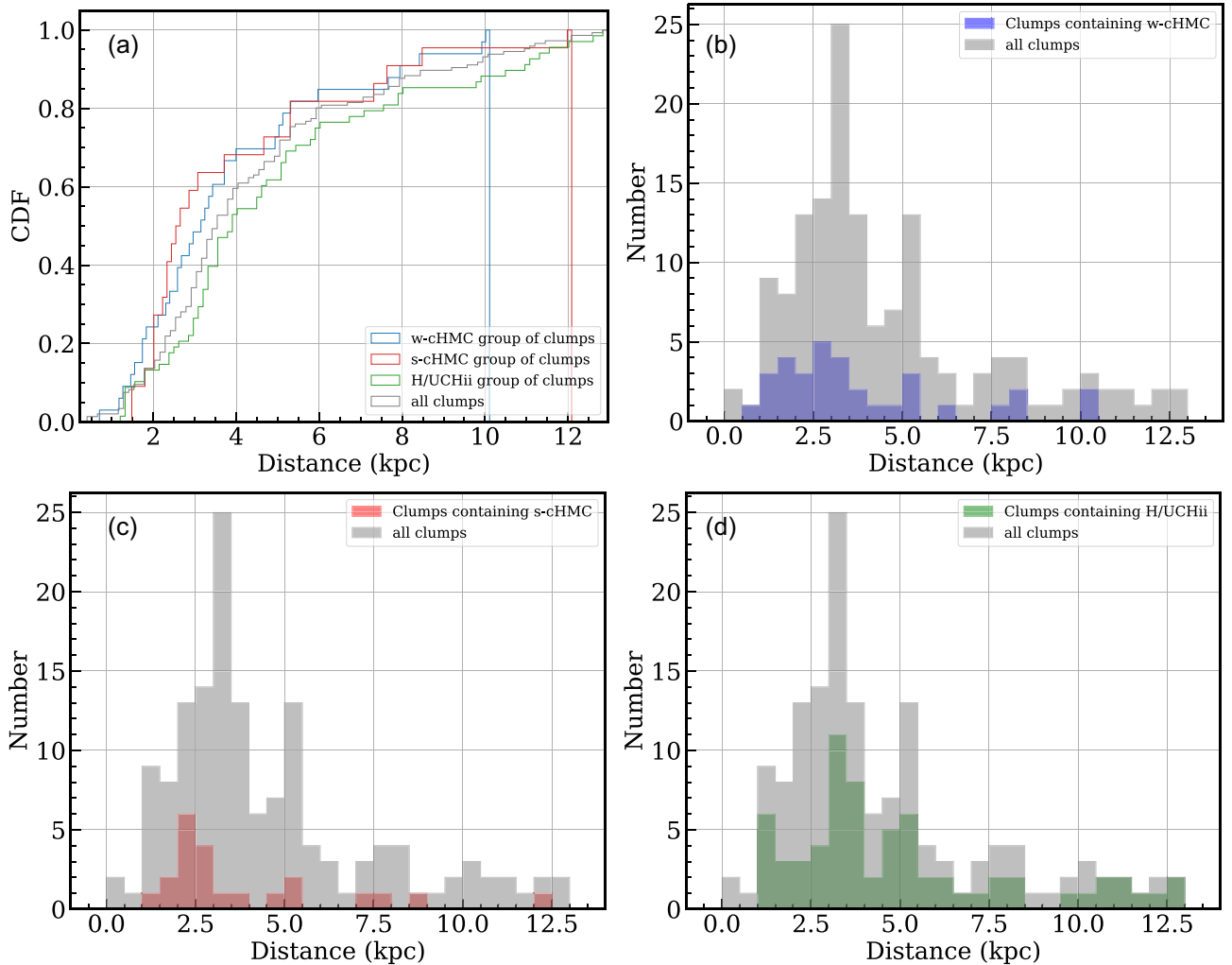


Figure 7. Distance distribution of the four groups of clumps. (a) Cumulative distribution function. (b)–(d) Histograms of all clumps against w-cHMC group, s-cHMC group, and H/UC-H II group of clumps, respectively.

in two separate catalogues: one including 32 s-cHMCs and the other including 58 w-cHMCs, with the former catalogue presumably containing more robust HMC candidates than the latter. These catalogues constitute a unique data base of objects for future follow-up higher resolution observations to investigate the dynamical processes related to high-mass star formation, such as infall, outflow, and rotation.

4.3 Constraints on the duration of the HMC phase

Because heating by the forming star precedes ionization, the general picture is that HMCs precede H/UC-H II in the evolution of a core to a massive star (Garay & Lizano 1999; Kurtz et al. 2000; van der Tak 2004; Cesaroni 2005; Rathborne et al. 2011). Our large sample, unbiased towards either HMCs or H/UC-H II cores, enables

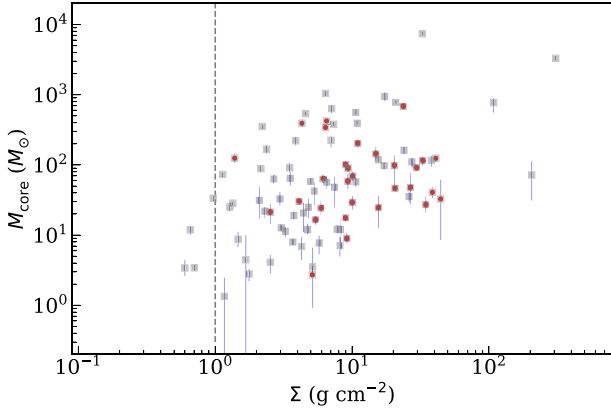


Figure 8. Distributions of the mass and mass surface density of the 90 COM-containing cores that are not associated with the H/UC-H II signature. The two parameters were calculated assuming a temperature of 100 K typical of HMCs (see Appendix B). The dashed vertical line indicates a stringent theoretical threshold of 1 g cm^{-2} above which the cores most likely form high-mass stars. The red dots highlight the cores associated with the Class II CH_3OH masers, showing that they are widely distributed, but above the 1 g cm^{-2} line.

further study. First, the two phenomena are not exclusive, as there is a substantial overlap between HMCs and H/UC-H II (Fig. 5). Nearly one third (48/138) of H/UC-H II in the sample show COMs, indicating that hot cores may persist during nearly one third of H/UC-H II lifetime even when the central high-mass protostars start to ionize the surroundings.

Second, with certain assumptions, we can use the numbers of candidate HMCs and cores with H/UC-H II to estimate the relative duration of the HMC phase. For a rigorous result, the survey must be complete, the evolution through the phases must be a continuous process proceeding at a steady rate for a period longer than the last phase considered, and other variables, such as the mass of the star, cannot be important. Further study of this sample is needed to check these assumptions, so we offer here only some preliminary considerations.

As discussed in Section 4.1, the ATOMS survey is more complete in the detection of H/UC-H II regions than HMCs, so we have a lower limit on the number counts of HMCs. The total number (138) of COM-containing dense cores (w-CHMC and s-CHMC) is already larger than the number (89) of H/UC-H II cores. The number of strong candidate HMCs (56) is smaller than the number (89) of H/UC-H II cores but may be greatly underestimated due to the distance effect. The most important unknown is the mass of the forming star because the number of ionizing photons is a very strong function of stellar mass. If all these cores will form stars capable of producing an H/UC-H II, the current lack of one is due to evolution. Further study of the luminosity of the individual cores with *JWST* observations could test this assumption.

With all these caveats, we suggest that the duration of candidate hot cores is at least comparable to the lifetime of H/UC-H II and is probably greater.

4.4 The origin of chemical differentiation among high-mass star-forming cores

Some high-mass star-forming cores are found to be chemically rich while others are not. Even cHMCs show differing richness in molecular lines especially in COMs. The median N_{line} values are 8

within an IQR of [5, 16] for the 58 w-CHMCs, and 57 within an IQR of [28, 123] for the 32 s-CHMCs. Would this different chemical richness (i.e. N_{line}) be related to an evolutionary sequence from w-CHMCs \rightarrow s-CHMCs?

The chemical differentiation among high-mass star-forming cores could be partially attributed to an evolutionary effect. More evolved cores with higher luminosity can heat a larger volume of their surroundings, evaporating molecules from ices and thus increasing the gaseous molecular abundances.

The 146 analysed ATOMS clumps can be classified into four main groups according to the types of cores they harbour: (i) the w-CHMC group with 28 clumps hosting w-CHMC cores, (ii) the s-CHMC group with 17 clumps hosting s-CHMC cores, (iii) the H/UC-H II group with 68 clumps containing H/UC-H II cores, and (iv) the ‘unknown’ group with 28 clumps without cHMC or H/UC-H II cores. There are five clumps with both w-CHMC and s-CHMC cores. Note that if a clump hosts both H/UC-H II core and cHMC, it will be placed into the H/UC-H II group. In addition, the w-CHMC group could contain s-CHMC cores that are not detectable yet or appear as w-CHMC cores due to the limited angular resolution and sensitivity. That is, there could be an underlying overlap between w-CHMC and s-CHMC groups of clumps, so the following analysis should be treated with caution. To explore the evolutionary stage of the clumps in the different groups, we resort to two parameters available for the ATOMS clumps (see Paper I): the bolometric luminosity to mass ratio $L_{\text{bol}}/M_{\text{clump}}$ and the dust temperature T_d . These parameters, which cannot be directly determined for the cores with the ATOMS data alone, are distance independent and have been proposed as indicator of the evolutionary stage of clumps (e.g. Guzmán et al. 2015).

Fig. 9 plots the distribution of the two parameters for the four groups of clumps. The ‘unknown’, w-CHMC, s-CHMC, and H/UC-H II groups have respective median L/M values of 28.1, 26.9, 29.7, and $32.7 L_{\odot}/M_{\odot}$. These values fall into the range of $2 L_{\odot}/M_{\odot} \leq L/M \leq 40 L_{\odot}/M_{\odot}$ observationally defined for the high-mass star-forming objects (Elia et al. 2017; Giannetti et al. 2017). This result suggests that the four groups of clumps harbour embedded phases of high-mass star formation. As such, many of the clumps in the ‘unknown’ group could not be really devoid of high-mass star-forming cores, and instead harbour HMPO objects or even more evolved HMC that could not be detected with the sensitivity of the ATOMS data. If we only consider the w-CHMC, s-CHMC, and H/UC-H II groups, there is a weak increasing trend in the median values of the $L_{\text{bol}}/M_{\text{clump}}$ ratio. A similar, but stronger, trend also appears in the temperature distribution among these three groups of clumps.

The differences in $L_{\text{bol}}/M_{\text{clump}}$ and T_d of the natal clumps are not very significant. If the evolution from HMCs to H/UC-H II is correct for individual cores, the explanation for the poor correlation with clump-scale properties may be that the evolution of individual cores to stars within a clump differs, and the clump-scale evolution of L_{bol} is driven by the most massive/evolved core. To study evolution at the core level, higher angular resolution and sensitivity in the mid-infrared to far-infrared is needed to pinpoint the embedded protostars, which cannot be well resolved at the current resolution of the infrared data, through the spectral energy distribution over a large dynamical frequency range.

In addition, recent studies of low-mass cores indicate that chemical richness of star-forming cores may not be related to evolutionary sequence. Belloche et al. (2020) suggested a luminosity threshold of $4 L_{\odot}$, above which low-mass cores should exhibit spectral features from at least one COM and may contain hot corinos. However, in a recent study of a large sample of 49 low-mass star-forming cores located in the Orion complex by Hsu et al. (2020), many protostellar

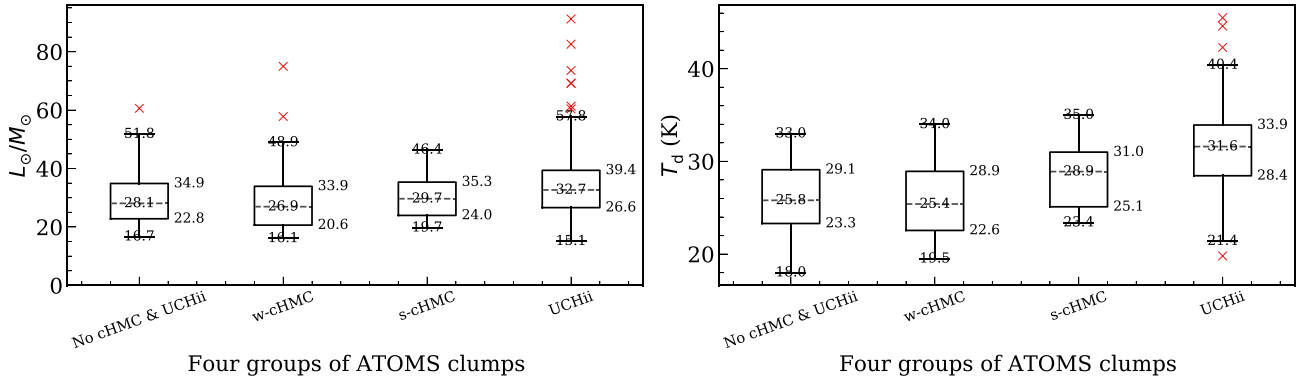


Figure 9. Box and whisker plot showing four groups of clumps the distribution of the L/M ratio (panel a), dust temperature T_d (panel b). The numbers associated with the boxes from the top to bottom represent the upper quartile, median (inside the box), and lower quartile, respectively. The red crosses indicate the outliers outside 1.5 times the interquartile range either above the upper quartile or below the lower quartile.

objects with luminosities much greater than $4L_{\odot}$ do not show any clear sign of COMs or hot corinos. This indicates that the chemical differentiation among star-forming cores depends on their chemical history, as well as their luminosity.

An alternative scenario could be that some star-forming cores in their ‘cold’ core phase form abundant molecules such as COMs on ice mantles through grain-surface chemistry while other cores do not. These different behaviours are probably related to varying physical and chemical properties as well as initial conditions of different sources. When molecules evaporate from ice in the ‘hot’ core phase, their gaseous molecular abundances should also vary greatly. This scenario could be tested by detailed chemical studies towards a sample of ‘cold clumps’ like in infrared dark clouds that lack active star formation and may represent the very initial conditions of high-mass star formation. Detailed chemical investigations are beyond the scope of this work. More detailed chemical studies of molecular abundances of dense cores will be presented in forthcoming papers.

4.5 Size versus $H40\alpha$ linewidth relationship for H/UC-H II cores

The 89 H/UC-H II cores could be associated with either HC-H II or UC-H II regions, and we attempt to distinguish them by their observable properties (see Section 1). To this end, we focus on the size and line width of the radio recombination line ($\Delta V_{H40\alpha}$) of the H/UC-H II cores, as shown in Fig. 10. These cores have two types of sizes, one measured from 3 mm continuum and the other one from $H40\alpha$ emission. As shown in Fig. 10(a), the two types of sizes match each other well until the size is greater than 0.1 pc. This result suggests that $H40\alpha$ emission mainly contributes to 3 mm continuum when the H/UC-H II core size is less than 0.1 pc. From distance information of the cores as indicated in the colourscale in Fig. 10(a), it can be seen that the deviation of the two types of core sizes does not depend on the distance. We suspect that when the H/UC-H II core size is greater than 0.1 pc, more cold dust emission could be included within it, leading to an overestimation of the size of the real ionized gas region. In view of this, the size measured from $H40\alpha$ emission for the 89 H/UC-H II cores will be discussed in the following. These cores have sizes from $\sim 5.3 \times 10^{-3}$ to 3.7×10^{-1} pc with a median value of 5.7×10^{-2} pc.

The $\Delta V_{H40\alpha}$ of the 89 H/UC-H II cores range from ~ 15 to 55 km s^{-1} with a median value of 28 km s^{-1} . The measurement of $\Delta V_{H40\alpha}$ is not found to be distance-dependent, as shown in Fig. 10(b). Actually, the $\Delta V_{H40\alpha}$ values are typical of both UC-H II and

HC-H II regions, namely in a range of ~ 10 – 40 km s^{-1} for the former and 40 – 60 km s^{-1} or even higher for the latter (e.g. Hoare et al. 2007; Rivera-Soto et al. 2020). Following the typical size and line width of radio recombination lines of H/UC-H II regions (see above), the 89 H/UC-H II cores can be divided into 79 UC-H II and 10 HC-H II cores (see Fig. 10c). That is, most (~ 90 per cent) of the compact-H II cores are UC-H II cores while only a small fraction (~ 10 per cent) of them are HC-H II cores. Therefore, if the lifetime of UC-H II is of the order of 10^5 yr , the lifetime of HC-H II will be $\sim 10^4 \text{ yr}$ (e.g. Wood & Churchwell 1989; Davies et al. 2011).

In general, the widths of radio recombination lines are due to a combination of thermal, dynamical, and/or pressure broadening mechanisms (e.g. Keto et al. 2008; Nguyen-Luong et al. 2017; Rivera-Soto et al. 2020). The thermal broadening, ΔV_{th} , is given by $\Delta V_{th} = (8 \ln 2 k_B T_e / m_H)^{1/2}$ where k_B is the Boltzmann constant and m_H is the mass of the atom hydrogen. For electron temperatures of $T_e = 6000, 8000$, and 10000 K , the respective thermal widths are $16.6, 19.1$, and 21.4 km s^{-1} (see the three vertical colour-gradient lines in Fig. 10c). Most (~ 96 per cent) of the H/UC-H II cores are found to have $\Delta V_{H40\alpha}$ widths larger than $\sim 19 \text{ km s}^{-1}$, which corresponds to the thermal broadening for an electron temperature of 8000 K typical of an ionized plasma (e.g. Osterbrock & Ferland 2006). This result indicates that in addition to the thermal broadening, non-thermal (e.g. dynamical and/or pressure) broadening mechanisms are responsible for the observed large $\Delta V_{H40\alpha}$ in most of the 89 H/UC-H II cores.

Moreover, we find that the $H40\alpha$ line width decreases as the size of the core increases. To investigate whether the trend is distance-dependent or not, we present the distance information of each H/UC-H II core in colourscale in Fig. 10. The figure shows that the decreasing trend persists rather than breaks at different distances, and thus it is physically meaningful. It indicates that non-thermal (e.g. dynamical and/or pressure) mechanisms dominate the broadening on smaller scales of the H/UC-H II cores (Garay & Lizano 1999).

5 SUMMARY AND CONCLUSION

Using the 3 mm ATOMS survey data, we have searched for the presence of compact cores towards 146 IRAS clumps located within the inner Galactic disc in the range $-80^\circ < l < 40^\circ$ and $|b| < 2^\circ$. The compact cores were extracted from 3 mm continuum emission maps through a combination of the *Dendrogram* algorithm and CASA *imfit* task. The main results of this work are summarized as follows:

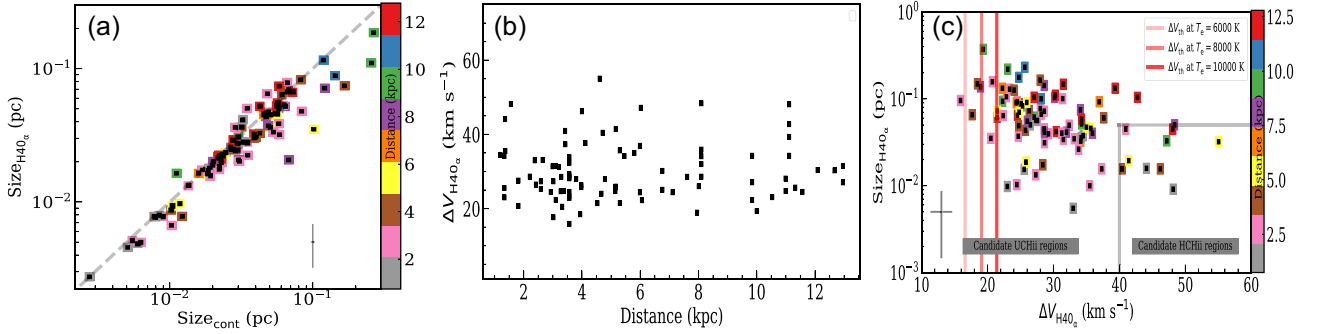


Figure 10. Size and H40 α line width of the 89 H/UC-H II cores. (a) Comparison between the two types of core sizes, one measured from 3 mm continuum (in x-axis) and the other from H40 α emission (in y-axis). The colour scheme encodes distance information of the cores. The typical error bars for the two measurements are presented at the bottom right. (b) Distribution of the H40 α line width as a function of distance. (c) Relation between the size of the H/UC-H II cores measured from H40 α emission and the line width of H40 α ($\Delta V_{\text{H40}\alpha}$). The colour scheme is the same as in panel (a). The H/UC-H II cores are roughly divided into 79 UC-H II and 10 HC-H II classes. For reference, the thermal-broadening line widths are given in three colour-gradient vertical lines at 16.6, 19.1, and 21.4 km s $^{-1}$ for the electronic temperatures of 6000, 8000, and 10000 K, respectively. The typical error bars for the two parameters are presented at the bottom left.

(i) We compiled three catalogues of compact cores with different characteristics. One catalogue, referred to as the H/UC-H II catalogue, includes 89 cores that enshroud H/UC-H II regions as indicated by the presence of compact H40 α emission; a second catalogue, referred to as the pure s-cHMC catalogue, includes 32 s-cHMCs associated with a rich spectrum ($N \geq 20$ lines) of COMs; and the third catalogue, referred to as the pure w-cHMC catalogue, includes 58 w-cHMCs with relatively low levels of COM richness (i.e. $5 \leq N_{\text{line}} < 20$). These three catalogues of compact cores provide a crucial foundation for future studies of the early stages of high-mass star formation across the Milky Way. We find that nearly half of H/UC-H II cores are candidate HMCs.

(ii) For completeness, another catalogue that includes the remaining 274 cores without the signature of either cHMC or H/UC-H II, referred to as the ‘unknown’ catalogue, is also provided. These cores deserve future detailed investigation, since some of them could be in the evolutionary stages of high-mass star formation prior to formation of an HMC, and thus could be good targets for studying initial conditions of high-mass star formation.

(iii) The detection rate of candidate distant hot cores is greatly underestimated due to relatively poorer linear resolution and line sensitivity than for nearer sources. However, the total number (138) of the detected COM-containing dense cores is still larger than the number (89) of H/UC-H II cores, indicating that the duration of high-mass protostellar cores (or cores) showing line-rich features could be at least comparable to the lifetime of H/UC-H II phase.

(iv) We find weak increasing trends in L/M ratio and dust temperature (T_d) for the natal clumps that contain different kinds of dense cores (w-cHMC \rightarrow s-cHMC \rightarrow H/UC-H II). However, the chemical classes indicating core evolution are not well traced by the $L_{\text{bol}}/M_{\text{clump}}$, suggesting that the evolution of individual cores is not captured by that evolutionary indicator used for clumps.

(v) Based on the relationship between size and H40 α line width, we conclude that of the 89 H/UC-H II cores, 79 are associated with UC-H II and 10 with HC-H II regions. We also find that the width of the H40 α line increases as the core size decreases, suggesting that non-thermal (e.g. dynamical and/or pressure) line-broadening mechanisms are dominant on the smaller scales of the H/UC-H II cores.

The 3 mm interferometric catalogues of high-mass star-forming cores provided here can be used to study the kinematics and dynamics

of molecular gas in the vicinity of newly formed (proto)OB stars and somewhat-evolved H/UC-H II regions. Such studies will definitely advance our understanding of the dynamical processes in the early stages of high-mass star formation, such as infall, outflow, and rotational motions, and our knowledge of the roles of feedback (e.g. outflows, stellar winds, and H/UC-H II regions) in the process of high-mass star formation.

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a PYTHON package to compute dendrograms of Astronomical data (<http://www.dendrograms.org/>). This research made use of ASTROPY, a community-developed core PYTHON package for Astronomy (Astropy Collaboration 2018).

DATA AVAILABILITY

The data underlying this article are available in the article and in its online supplementary material.

REFERENCES

- Astropy Collaboration et al., 2018, *AJ*, 156, 123
- Belloche A. et al., 2020, *A&A*, 635, A198
- Beuther H., 2007, in Elmegreen B. G., Palous J., eds, *Proc. IAU Symp. 237, Triggered Star Formation in a Turbulent ISM*. Kluwer, Dordrecht, p. 148
- Beuther H., Schilke P., Sridharan T. K., Menten K. M., Walmsley C. M., Wyrowski F., 2002, *A&A*, 383, 892
- Bronfman L., Nyman L.-Å., May J., 1996, *A&AS*, 115, 81
- Cesaroni R., 2005, in Cesaroni R., Felli M., Churchwell E., Walmsley M., eds, *Proc. IAU Symp. 227, Massive Star Birth: A Crossroads of Astrophysics*. Kluwer, Dordrecht, p. 59
- Cesaroni R., Galli D., Lodato G., Walmsley C. M., Zhang Q., 2007, in Reipurth B., Jewitt D., Keil K., eds, *Protostars and Planets V*. University of Arizona Press, Tucson, p. 197
- Cesaroni R., Hofner P., Araya E., Kurtz S., 2010, *A&A*, 509, A50
- Churchwell E., 2002, *ARA&A*, 40, 27
- Coletta A., Fontani F., Rivilla V. M., Mininni C., Colzi L., Sánchez-Monge Á., Beltrán M. T., 2020, *A&A*, 641, A54
- Csengeri T., Belloche A., Bontemps S., Wyrowski F., Menten K. M., Bouscasse L., 2019, *A&A*, 632, A57
- Cyganowski C. J., Brogan C. L., Hunter T. R., Churchwell E., 2009, *ApJ*, 702, 1615
- Davies B., Hoare M. G., Lumsden S. L., Hosokawa T., Oudmaijer R. D., Urquhart J. S., Mottram J. C., Stead J., 2011, *MNRAS*, 416, 972
- de Wit W. J., Testi L., Palla F., Zinnecker H., 2005, *A&A*, 437, 247
- Dirienzo W. J., Brogan C., Indebetouw R., Chandler C. J., Friesen R. K., Devine K. E., 2015, *AJ*, 150, 159
- Dunne L., Eales S., Ivison R., Morgan H., Edmunds M., 2003, *Nature*, 424, 285
- Egan M. P., Shipman R. F., Price S. D., Carey S. J., Clark F. O., Cohen M., 1998, *ApJ*, 494, L199
- Elia D. et al., 2017, *MNRAS*, 471, 100
- Faúndez S., Bronfman L., Garay G., Chini R., Nyman L.-Å., May J., 2004, *A&A*, 426, 97
- Furuya R. S., Cesaroni R., Shinnaga H., 2011, *A&A*, 525, A72
- Garay G., Lizano S., 1999, *PASP*, 111, 1049
- Giannetti A., Leurini S., Wyrowski F., Urquhart J., Csengeri T., Menten K. M., König C., Güsten R., 2017, *A&A*, 603, A33
- Guzmán A. E., Sanhueza P., Contreras Y., Smith H. A., Jackson J. M., Hoq S., Rathborne J. M., 2015, *ApJ*, 815, 130
- Hernández-Hernández V., Zapata L., Kurtz S., Garay G., 2014, *ApJ*, 786, 38
- Hoare M. G., Kurtz S. E., Lizano S., Keto E., Hofner P., 2007, in Reipurth B., Jewitt D., Keil K., eds, *Protostars and Planets V*. University of Arizona Press, Tucson, p. 181
- Hoare M. G. et al., 2012, *PASP*, 124, 939
- Hosokawa T., Omukai K., 2009, *ApJ*, 691, 823
- Hsu S.-Y. et al., 2020, *ApJ*, 898, 107
- Jiménez-Serra I. et al., 2016, *ApJ*, 830, L6
- Jorgensen J. K., Belloche A., Garrod R. T., 2020, *ARA&A*, 58, 727
- Kennicutt R. C., 2005, in Cesaroni R., Felli M., Churchwell E., Walmsley M., eds, *Proc. IAU Symp. 227, Massive Star Birth: A Crossroads of Astrophysics*. Kluwer, Dordrecht, p. 3
- Keto E., 2007, *ApJ*, 666, 976
- Keto E., Zhang Q., 2010, *MNRAS*, 406, 102
- Keto E., Zhang Q., Kurtz S., 2008, *ApJ*, 672, 423
- Krumholz M. R., McKee C. F., 2008, *Nature*, 451, 1082
- Kurtz S., Hofner P., 2005, *AJ*, 130, 711
- Kurtz S., Churchwell E., Wood D. O. S., 1994, *ApJS*, 91, 659
- Kurtz S., Cesaroni R., Churchwell E., Hofner P., Walmsley C. M., 2000, in Mannings V., Boss A. P., Russell S. S., eds, *Protostars and Planets IV*. University of Arizona Press, Tucson, p. 299
- Ladeyschikov D. A., Bayandina O. S., Sobolev A. M., 2019, *AJ*, 158, 233
- Li S., Zhang Q., Pillai T., Stephens I. W., Wang J., Li F., 2019, *ApJ*, 886, 130
- Li S. et al., 2020, *ApJ*, 896, 110
- Liu T. et al., 2016, *ApJ*, 829, 59
- Liu T. et al., 2017, *ApJ*, 849, 25
- Liu T. et al., 2020a, *MNRAS*, 496, 2790 (Paper I)
- Liu T. et al., 2020b, *MNRAS*, 496, 2821 (Paper II)
- McKee C. F., Tan J. C., 2003, *ApJ*, 585, 850
- McMullin J. P., Waters B., Schiebel D., Young W., Golap K., 2007, in Shaw R. A., Hill F., Bell D. J., eds, *ASP Conf. Ser. Vol. 376, Astronomical Data Analysis Software and Systems XVI*. Astron. Soc. Pac., San Francisco, p. 127
- Molet J. et al., 2019, *A&A*, 626, A132
- Motte F., Bontemps S., Louvet F., 2018, *ARA&A*, 56, 41
- Nguyen-Luong Q. et al., 2017, *ApJ*, 844, L25
- Osorio M., Lizano S., D'Alessio P., 1999, *ApJ*, 525, 808
- Ossenkopf V., Henning T., 1994, *A&A*, 291, 943
- Osterbrock D. E., Ferland G. J., 2006, *Astrophysics of Gaseous Nebulae and Active Galactic Nuclei*. University Science Books, California
- Perault M. et al., 1996, *A&A*, 315, L165
- Peretto N., Fuller G. A., 2009, *A&A*, 505, 405
- Pollack J. B., Hollenbach D., Beckwith S., Simonelli D. P., Roush T., Fong W., 1994, *ApJ*, 421, 615
- Purcell C. R., Longmore S. N., Burton M. G., Walsh A. J., Minier V., Cunningham M. R., Balasubramanyam R., 2009, *MNRAS*, 394, 323
- Purcell C. R. et al., 2013, *ApJS*, 205, 1
- Qin S.-L., Huang M., Wu Y., Xue R., Chen S., 2008, *ApJ*, 686, L21
- Qin S.-L., Schilke P., Wu J., Wu Y., Liu T., Liu Y., Sánchez-Monge Á., 2015, *ApJ*, 803, 39
- Qiu K., Zhang Q., Beuther H., Fallscheer C., 2012, *ApJ*, 756, 170
- Ragan S. E., Bergin E. A., Plume R., Gibson D. L., Wilner D. J., O'Brien S., Hails E., 2006, *ApJS*, 166, 567
- Rathborne J. M., Jackson J. M., Chambers E. T., Simon R., Shipman R., Frieswijk W., 2005, *ApJ*, 630, L181
- Rathborne J. M., Jackson J. M., Simon R., 2006, *ApJ*, 641, 389
- Rathborne J. M., Jackson J. M., Chambers E. T., Stojimirovic I., Simon R., Shipman R., Frieswijk W., 2010, *ApJ*, 715, 310
- Rathborne J. M., Garay G., Jackson J. M., Longmore S., Zhang Q., Simon R., 2011, *ApJ*, 741, 120
- Reid M. J. et al., 2014, *ApJ*, 783, 130
- Reid M. J. et al., 2019, *ApJ*, 885, 131
- Rivera-Soto R., Galván-Madrid R., Ginsburg A., Kurtz S., 2020, *ApJ*, 899, 94
- Rosolowsky E. W., Pineda J. E., Kauffmann J., Goodman A. A., 2008, *ApJ*, 679, 1338
- Sánchez-Monge Á. et al., 2017, *A&A*, 604, A6
- Sanhueza P. et al., 2019, *ApJ*, 886, 102
- Sato K., Hasegawa T., Umegaki T., Saito H., Kuno N., Seta M., Sakamoto S., 2020, preprint ([arXiv:2006.13191](https://arxiv.org/abs/2006.13191))
- Schuller F. et al., 2009, *A&A*, 504, 415
- Silva A. et al., 2017, *ApJ*, 847, 87
- Simon R., Jackson J. M., Rathborne J. M., Chambers E. T., 2006, *ApJ*, 639, 227
- Sollins P. K., Zhang Q., Keto E., Ho P. T. P., 2005, *ApJ*, 624, L49
- Soma T., Sakai N., Watanabe Y., Yamamoto S., 2018, *ApJ*, 854, 116
- Svoboda B. E. et al., 2019, *ApJ*, 886, 36
- Tan J. C., Beltrán M. T., Caselli P., Fontani F., Fuente A., Krumholz M. R., McKee C. F., Stolte A., 2014, in Beuther H., Klessen R. S., Dullemond C. P., Henning T., eds, *Protostars and Planets VI*. University of Arizona Press, Tucson, p. 149
- Urquhart J. S. et al., 2013, *MNRAS*, 435, 400
- Urquhart J. S. et al., 2014, *MNRAS*, 443, 1555
- Urquhart J. S. et al., 2018, *MNRAS*, 473, 1059

- van der Tak F. F. S., 2004, in Burton M., Jayawardhana R., Bourke T., eds, *Proc. IAU Symp. 221, Star Formation at High Angular Resolution*. Kluwer, Dordrecht, p. 59
- Vastel C., Ceccarelli C., Lefloch B., Bachiller R., 2014, *ApJ*, 795, L2
- Wang K., Zhang Q., Wu Y., Zhang H., 2011, *ApJ*, 735, 64
- Wang K. et al., 2014, *MNRAS*, 439, 3275
- Wenger T. V., Balser D. S., Anderson L. D., Bania T. M., 2018, *ApJ*, 856, 52
- Wood D. O. S., Churchwell E., 1989, *ApJS*, 69, 831
- Xu J.-L., Wang J.-J., 2013, *MNRAS*, 431, 2385
- Yang A. Y., Thompson M. A., Tian W. W., Bihr S., Beuther H., Hindson L., 2019, *MNRAS*, 482, 2681
- Yang A. Y. et al., 2021, *A&A*, 645, A110
- Young C. H., Evans N. J., 2005, *ApJ*, 627, 293
- Yuan J. et al., 2017, *ApJS*, 231, 11
- Zhang Q., Hunter T. R., Brand J., Sridharan T. K., Molinari S., Kramer M. A., Cesaroni R., 2001, *ApJ*, 552, L167
- Zhang Q., Hunter T. R., Brand J., Sridharan T. K., Cesaroni R., Molinari S., Wang J., Kramer M., 2005, *ApJ*, 625, 864
- Zhang Q., Wang Y., Pillai T., Rathborne J., 2009, *ApJ*, 696, 268
- Zhang Q., Wang K., Lu X., Jiménez-Serra I., 2015, *ApJ*, 804, 141
- Zinnecker H., Yorke H. W., 2007, *ARA&A*, 45, 481

SUPPORTING INFORMATION

Supplementary data are available at [MNRAS](https://www.mnras.org/) online.

Table S1. Parameters of 89 H/CH II cores .

Table S2. Parameters of 32 COM-containing cores with $N_{\text{line}} \geq 20$ (or pure s-cHMC).

Table S3. Parameters of 58 COM-containing cores with $5 \leq N_{\text{line}} < 20$ (or pure w-cHMC).

Table S4. Parameters of 274 ‘unknown’ cores (i.e. no signature of H/UC-HII nor cHMC).

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APPENDIX A: CALCULATION OF DISTANCES

Due to the lack of distance uncertainty information from the literature (Faúndez et al. 2004; Urquhart et al. 2018) for the 146 target clumps, their kinematic distances were recalculated using the Monte Carlo kinematic distance tool developed by Wenger et al. (2018). With the Galactic coordinates and radial velocity parameters of a source inputted, the tool can return the distances and the uncertainties for the source leveraging the rotation curve of Reid et al. (2014) with the updated solar motion parameters (Reid et al. 2019) and the

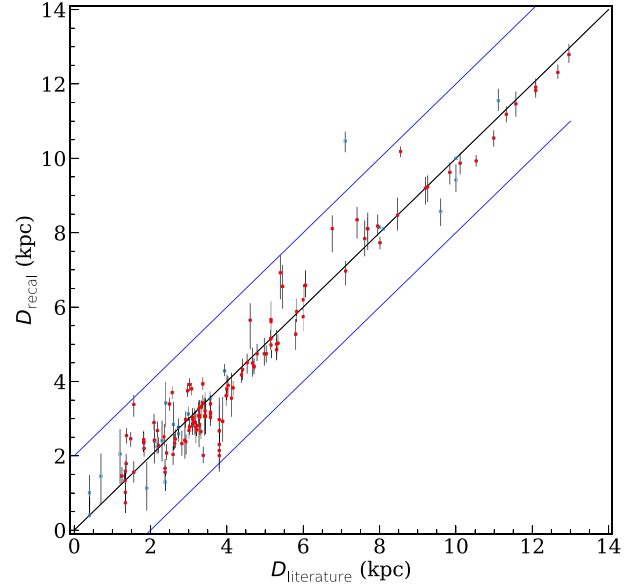


Figure A1. Comparison between the distances calculated in this work and those from the literature for the 146 ATOMS clumps. The black and blue lines indicate unity and 2.0 kpc offsets, respectively.

Monte Carlo technique with 10 000 samplings. The radial velocity measurements by Liu et al. (2016) rather than by Bronfman et al. (1996) were adopted in this work since the velocity resolution (0.42 km s^{-1}) in Liu et al. (2016) at which the measurements were made better than that (0.76 km s^{-1}) in Bronfman et al. (1996).

Apart from the two sources located near the Galactic Centre, kinematic distance solutions have been calculated for 144 ATOMS clumps. The distance ambiguities (if any) have been resolved following Faúndez et al. (2004) and Urquhart et al. (2018). Moreover, the kinematic distance solutions have been improved for the three sources using maser parallax distances available in the literature (Reid et al. 2014; Urquhart et al. 2018). Overall, the typical relative distance uncertainty for the 144 ATOMS clumps is ~ 7 per cent. For the remaining two sources located near the Galactic Centre, their kinematic distances could not be calculated from the rotation curve and thus the values from the literature along with a typical relative distance uncertainty of ~ 7 per cent have been adopted. The recalculated distances are listed in Table A1.

Comparing the recalculated distances with those from Faúndez et al. (2004) and Urquhart et al. (2018), we find a good agreement (Pearson correlation coefficient ρ is 0.97) as shown in Fig. A1. In addition, the relative distance uncertainties are found to have an IQR of $\sim [5 \text{ per cent}, 12 \text{ per cent}]$ with a median value of ~ 7 per cent.

Table A1. Distances of 146 ATOMS clumps.

ID	Name	l (deg)	b (deg)	V_{lsr} (km s ⁻¹)	D (kpc)	ID	Name	l (deg)	b (deg)	V_{lsr} (km s ⁻¹)	D (kpc)
1	I08076–3556	253.2952	−1.6159	5.9 ± 0.4	1.0 ^{+0.5} _{−0.5}	74	I16385–4619	338.5690	−0.1441	−117.8 ± 0.4	7.0 ^{+0.3} _{−0.4}
2	I08303–4303	261.6457	−2.0911	15.2 ± 0.4	2.4 ^{+0.5} _{−0.6}	75	I16424–4531	339.6226	−0.1225	−31.1 ± 0.4	2.3 ^{+0.3} _{−0.2}
3	I08448–4343	263.7731	−0.4337	3.4 ± 0.4	1.4 ^{+0.6} _{−0.8}	76	I16445–4459	340.2492	−0.0458	−121.4 ± 0.4	8.2 ^{+0.3} _{−0.3}
4	I08470–4243	263.2493	0.5114	13.5 ± 0.4	2.4 ^{+0.6} _{−0.6}	77	I16458–4512	340.2475	−0.3741	−49.9 ± 0.4	3.5 ^{+0.2} _{−0.2}
5	I09002–4732	268.4203	−0.8476	3.9 ± 0.4	2.0 ^{+0.7} _{−0.8}	78	I16484–4603	339.8842	−1.2572	−31.6 ± 0.4	2.4 ^{+0.2} _{−0.2}
6	I09018–4816	269.1529	−1.1308	10.8 ± 0.4	2.8 ^{+0.6} _{−0.7}	79	I16487–4423	341.2158	−0.2358	−43.2 ± 0.4	3.3 ^{+0.2} _{−0.2}
7	I09094–4803	269.8498	−0.0633	74.6 ± 0.5	8.6 ^{+0.3} _{−0.4}	80	I16489–4431	341.1258	−0.3475	−40.4 ± 0.4	3.1 ^{+0.2} _{−0.2}
8	I10365–5803	286.2076	0.1708	−18.9 ± 0.4	3.4 ^{+0.6} _{−0.6}	81	I16506–4512	340.7842	−1.0157	−25.8 ± 0.4	2.1 ^{+0.2} _{−0.2}
9	I11298–6155	293.8275	−0.7457	33.0 ± 0.4	9.4 ^{+0.4} _{−0.3}	82	I16524–4300	342.7076	0.1258	−40.6 ± 0.4	3.2 ^{+0.2} _{−0.2}
10	I11332–6258	294.5109	−1.6218	−16.1 ± 0.5	1.1 ^{+0.7} _{−0.6}	83	I16547–4247	343.1275	−0.0625	−29.7 ± 0.4	2.6 ^{+0.2} _{−0.2}
11	I11590–6452	297.7209	−2.7820	−3.9 ± 0.4	0.4 ^{+0.5} _{−0.1}	84	I16562–3959	345.4926	1.4686	−11.9 ± 0.4	1.3 ^{+0.2} _{−0.2}
12	I12320–6122	300.9692	1.1456	−42.5 ± 0.4	3.1 ^{+0.6} _{−0.7}	85	I16571–4029	345.1927	1.0288	−14.7 ± 0.4	1.6 ^{+0.2} _{−0.2}
13	I12326–6245	301.1358	−0.2258	−39.1 ± 0.4	5.6 ^{+0.5} _{−0.8}	86	I17006–4215	344.2208	−0.5941	−24.2 ± 0.4	2.3 ^{+0.2} _{−0.2}
14	I12383–6128	301.7308	1.1040	−37.9 ± 0.4	2.8 ^{+0.7} _{−0.6}	87	I17008–4040	345.5043	0.3475	−15.9 ± 0.4	1.7 ^{+0.2} _{−0.2}
15	I12572–6316	303.9308	−0.6874	29.5 ± 0.4	11.5 ^{+0.3} _{−0.3}	88	I17016–4124	345.0029	−0.2241	−26.5 ± 0.4	2.5 ^{+0.2} _{−0.2}
16	I13079–6218	305.2091	0.2058	−40.5 ± 0.4	3.0 ^{+0.6} _{−0.7}	89	I17136–3617	350.5074	0.9590	−10.3 ± 0.4	1.6 ^{+0.3} _{−0.2}
17	I13080–6229	305.1958	0.0342	−34.0 ± 0.4	2.3 ^{+0.6} _{−0.6}	90	I17143–3700	350.0158	0.4325	−31.3 ± 0.4	12.3 ^{+0.2} _{−0.2}
18	I13111–6228	305.5624	0.0142	−38.7 ± 0.4	2.7 ^{+0.7} _{−0.6}	91	I17158–3901	348.5492	−0.9789	−15.8 ± 0.4	2.0 ^{+0.2} _{−0.2}
19	I13134–6242	305.7989	−0.2441	−31.6 ± 0.4	2.0 ^{+0.6} _{−0.4}	92	I17160–3707	350.1025	0.0825	−69.7 ± 0.4	9.9 ^{+0.2} _{−0.2}
20	I13140–6226	305.8872	0.0158	−32.7 ± 0.4	2.1 ^{+0.6} _{−0.5}	93	I17175–3544	351.4159	0.6457	−5.7 ± 0.4	1.0 ^{+0.3} _{−0.3}
21	I13291–6229	307.6158	−0.2592	−37.1 ± 0.4	2.4 ^{+0.6} _{−0.5}	94	I17204–3636	351.0411	−0.3357	−17.8 ± 0.4	2.7 ^{+0.2} _{−0.2}
22	I13291–6249	307.5608	−0.5875	−32.4 ± 0.4	7.8 ^{+0.5} _{−0.5}	95	I17220–3609	351.5809	−0.3524	−95.0 ± 0.4	7.7 ^{+0.2} _{−0.2}
23	I13295–6152	307.7558	0.3525	−43.5 ± 0.5	2.9 ^{+0.7} _{−0.6}	96	I17233–3606	351.7742	−0.5374	−4.2 ± 0.4	0.7 ^{+0.3} _{−0.3}
24	I13471–6120	309.9208	0.4775	−57.3 ± 0.4	6.6 ^{+0.6} _{−0.6}	97	I17244–3536	352.3158	−0.4425	−9.4 ± 0.4	1.8 ^{+0.3} _{−0.3}
25	I13484–6100	310.1442	0.7591	−54.1 ± 0.4	6.9 ^{+0.5} _{−0.7}	98	I17258–3637	351.6326	−1.2538	−11.7 ± 0.4	2.0 ^{+0.2} _{−0.3}
26	I14013–6105	311.6257	0.2891	−54.6 ± 0.4	3.5 ^{+0.7} _{−0.5}	99	I17269–3312	354.5959	0.4691	−20.7 ± 0.5	4.2 ^{+0.2} _{−0.2}
27	I14050–6056	312.1078	0.3091	−48.1 ± 0.4	3.1 ^{+0.5} _{−0.5}	100	I17271–3439	353.4091	−0.3607	−15.9 ± 0.6	3.0 ^{+0.2} _{−0.2}
28	I14164–6028	313.5758	0.3242	−45.2 ± 0.4	2.8 ^{+0.4} _{−0.4}	101	I17278–3541	352.6308	−1.0674	0.7 ± 0.4	1.3 ^{+0.2} _{−0.0}
29	I14206–6151	313.5758	−1.1540	−48.8 ± 0.4	3.0 ^{+0.5} _{−0.5}	102	I17439–2845	0.3160	−0.2008	18.8 ± 0.4	8.1 ^{+0.0} _{−0.0}
30	I14212–6131	313.7658	−0.8624	−49.2 ± 0.4	3.1 ^{+0.5} _{−0.5}	103	I17441–2822	0.6659	−0.0341	51.5 ± 1.6	8.1 ^{+0.6} _{−0.6}
31	I14382–6017	316.1392	−0.5058	−60.3 ± 0.4	8.1 ^{+0.4} _{−0.6}	104	I17455–2800	1.1258	−0.1075	−14.2 ± 0.4	10.0 ^{+0.7} _{−0.7}
32	I14453–5912	317.4074	0.1092	−40.2 ± 0.4	2.3 ^{+0.3} _{−0.3}	105	I17545–2357	5.6374	0.2375	9.2 ± 0.4	2.4 ^{+0.2} _{−0.3}
33	I14498–5856	318.0495	0.0858	−49.7 ± 0.4	2.9 ^{+0.4} _{−0.3}	106	I17589–2312	6.7959	−0.2575	21.6 ± 0.5	3.7 ^{+0.2} _{−0.2}
34	I15122–5801	321.0545	−0.5073	−60.3 ± 0.4	9.2 ^{+0.3} _{−0.4}	107	I17599–2148	8.1408	0.2225	19.8 ± 0.4	3.1 ^{+0.2} _{−0.2}
35	I15254–5621	323.4591	−0.0791	−66.9 ± 0.4	3.8 ^{+0.4} _{−0.3}	108	I18032–2032	9.6209	0.1942	5.9 ± 0.4	5.2 ^{+0.2} _{−0.3}
36	I15290–5546	324.2011	0.1208	−88.5 ± 0.4	8.1 ^{+0.4} _{−0.6}	109	I18056–1952	10.4725	0.0275	66.7 ± 0.4	10.2 ^{+0.1} _{−0.2}
37	I15384–5348	326.4457	0.9073	−40.6 ± 0.4	2.4 ^{+0.3} _{−0.3}	110	I18075–2040	10.0008	−0.7541	31.5 ± 0.4	3.8 ^{+0.2} _{−0.2}
38	I15394–5358	326.4741	0.7024	−40.7 ± 0.4	2.4 ^{+0.3} _{−0.2}	111	I18079–1756	12.4176	0.5058	18.4 ± 0.4	2.2 ^{+0.2} _{−0.2}
39	I15408–5356	326.6573	0.5940	−40.1 ± 0.4	2.4 ^{+0.3} _{−0.3}	112	I18089–1732	12.8875	0.4891	33.5 ± 0.4	3.4 ^{+0.2} _{−0.2}
40	I15411–5352	326.7240	0.6140	−41.0 ± 0.4	2.4 ^{+0.2} _{−0.3}	113	I18110–1854	11.9355	−0.6156	39.0 ± 0.4	3.9 ^{+0.2} _{−0.2}
41	I15437–5343	327.1194	0.5090	−83.8 ± 0.5	4.7 ^{+0.4} _{−0.3}	114	I18116–1646	13.8725	0.2808	50.2 ± 0.4	4.3 ^{+0.2} _{−0.1}
42	I15439–5449	326.4724	−0.3774	−53.4 ± 0.4	3.1 ^{+0.3} _{−0.3}	115	I18117–1753	12.9075	−0.2591	38.0 ± 0.4	3.7 ^{+0.2} _{−0.2}
43	I15502–5302	328.3075	0.4309	−91.6 ± 0.4	5.3 ^{+0.4} _{−0.4}	116	I18134–1942	11.4974	−1.4853	10.7 ± 0.4	1.5 ^{+0.2} _{−0.2}
44	I15520–5234	328.8092	0.6325	−41.2 ± 0.4	2.5 ^{+0.3} _{−0.2}	117	I18139–1842	12.4310	−1.1138	40.0 ± 0.4	3.9 ^{+0.2} _{−0.2}
45	I15522–5411	327.8092	−0.6358	−46.8 ± 0.4	2.8 ^{+0.2} _{−0.3}	118	I18159–1648	14.3308	−0.6440	23.9 ± 0.4	2.5 ^{+0.2} _{−0.2}
46	I15557–5215	329.4691	0.5024	−67.6 ± 0.4	3.9 ^{+0.3} _{−0.3}	119	I18182–1433	16.5858	−0.0508	60.1 ± 0.4	4.4 ^{+0.2} _{−0.2}
47	I15567–5236	329.3374	0.1475	−107.3 ± 0.4	6.2 ^{+0.4} _{−0.4}	120	I18223–1243	18.6542	−0.0591	46.0 ± 0.4	3.4 ^{+0.2} _{−0.2}
48	I15570–5227	329.4724	0.2158	−100.2 ± 0.4	5.7 ^{+0.5} _{−0.4}	121	I18228–1312	18.3010	−0.3891	33.6 ± 0.4	2.7 ^{+0.2} _{−0.2}
49	I15584–5247	329.4241	−0.1624	−76.2 ± 0.4	4.3 ^{+0.3} _{−0.3}	122	I18236–1205	19.3625	−0.0308	27.8 ± 0.4	2.2 ^{+0.2} _{−0.2}
50	I15596–5301	329.4058	−0.4591	−73.9 ± 0.4	9.9 ^{+0.3} _{−0.3}	123	I18264–1152	19.8825	−0.5341	45.4 ± 0.4	3.3 ^{+0.2} _{−0.2}

Table A1 – continued

ID	Name	l (deg)	b (deg)	V_{lsr} (km s ⁻¹)	D (kpc)	ID	Name	l (deg)	b (deg)	V_{lsr} (km s ⁻¹)	D (kpc)
51	I16026–5035	331.3608	1.0657	-78.5 ± 0.4	$4.5^{+0.2}_{-0.3}$	124	I18290–0924	22.3508	0.0675	84.3 ± 0.4	$5.0^{+0.2}_{-0.2}$
52	I16037–5223	330.2943	-0.3941	-81.0 ± 0.4	$9.6^{+0.3}_{-0.3}$	125	I18308–0503	26.4212	1.6902	43.6 ± 0.4	$2.8^{+0.2}_{-0.2}$
53	I16060–5146	330.9542	-0.1825	-88.8 ± 0.4	$5.0^{+0.4}_{-0.3}$	126	I18311–0809	23.7107	0.1708	113.5 ± 0.4	$6.6^{+0.4}_{-0.3}$
54	I16065–5158	330.8792	-0.3675	-61.5 ± 0.4	$3.6^{+0.2}_{-0.3}$	127	I18314–0720	24.4709	0.4874	102.1 ± 0.5	$5.9^{+0.4}_{-0.3}$
55	I16071–5142	331.1325	-0.2442	-85.9 ± 0.5	$4.9^{+0.3}_{-0.3}$	128	I18316–0602	25.6492	1.0507	45.3 ± 0.4	$2.9^{+0.2}_{-0.2}$
56	I16076–5134	331.2792	-0.1891	-86.0 ± 0.4	$4.9^{+0.3}_{-0.3}$	129	I18317–0513	26.3808	1.4071	42.2 ± 0.4	$2.7^{+0.2}_{-0.2}$
57	I16119–5048	332.2958	-0.0941	-48.4 ± 0.4	$3.0^{+0.2}_{-0.3}$	130	I18317–0757	23.9539	0.1508	80.5 ± 0.4	$4.7^{+0.3}_{-0.2}$
58	I16132–5039	332.5440	-0.1241	-47.0 ± 0.4	$2.9^{+0.3}_{-0.2}$	131	I18341–0727	24.6725	-0.1508	113.5 ± 0.4	$6.6^{+0.4}_{-0.3}$
59	I16158–5055	332.6474	-0.6090	-49.4 ± 0.4	$3.0^{+0.2}_{-0.2}$	132	I18411–0338	28.8608	0.0659	103.7 ± 0.4	$8.3^{+0.3}_{-0.5}$
60	I16164–5046	332.8256	-0.5490	-56.2 ± 0.4	$3.4^{+0.2}_{-0.2}$	133	I18434–0242	29.9538	-0.0158	98.1 ± 0.4	$5.7^{+0.4}_{-0.4}$
61	I16172–5028	333.1344	-0.4307	-52.1 ± 0.4	$3.2^{+0.2}_{-0.2}$	134	I18440–0148	30.8175	0.2742	98.3 ± 0.4	$5.7^{+0.5}_{-0.4}$
62	I16177–5018	333.3077	-0.3657	-50.9 ± 0.4	$3.1^{+0.2}_{-0.2}$	135	I18445–0222	30.3860	-0.1041	87.2 ± 0.4	$5.0^{+0.3}_{-0.3}$
63	I16272–4837	335.5857	-0.2908	-45.8 ± 0.4	$3.0^{+0.2}_{-0.2}$	136	I18461–0113	31.5808	0.0775	97.1 ± 0.4	$5.6^{+0.5}_{-0.4}$
64	I16297–4757	336.3593	-0.1374	-79.9 ± 0.4	$4.7^{+0.2}_{-0.2}$	137	I18469–0132	31.3958	-0.2575	86.9 ± 0.4	$5.2^{+0.4}_{-0.3}$
65	I16304–4710	337.0042	0.3225	-61.0 ± 0.4	$11.2^{+0.2}_{-0.2}$	138	I18479–0005	32.7973	0.1908	15.1 ± 0.4	$12.8^{+0.3}_{-0.2}$
66	I16313–4729	336.8658	0.0042	-72.3 ± 0.4	$4.4^{+0.2}_{-0.2}$	139	I18502+0051	34.2601	-0.4259	53.0 ± 0.4	$10.5^{+0.3}_{-0.3}$
67	I16318–4724	336.9942	-0.0275	-120.5 ± 0.4	$8.1^{+0.3}_{-0.3}$	140	I18507+0110	34.2575	0.1542	58.5 ± 0.4	$1.6^{+0.3}_{-0.3}$
68	I16330–4725	337.1208	-0.1741	-74.1 ± 0.4	$10.5^{+0.2}_{-0.2}$	141	I18507+0121	34.4108	0.2342	58.2 ± 0.4	$3.4^{+0.3}_{-0.3}$
69	I16344–4658	337.6125	-0.0592	-48.9 ± 0.4	$11.8^{+0.2}_{-0.2}$	142	I18517+0437	37.4308	1.5188	44.1 ± 0.4	$2.5^{+0.4}_{-0.2}$
70	I16348–4654	337.7042	-0.0541	-47.2 ± 0.4	$11.9^{+0.2}_{-0.2}$	143	I18530+0215	35.4657	0.1408	77.8 ± 0.4	$4.5^{+0.4}_{-0.4}$
71	I16351–4722	337.4058	-0.4025	-40.3 ± 0.4	$2.8^{+0.2}_{-0.2}$	144	I19078+0901	43.1658	0.0109	7.2 ± 0.4	$11.5^{+0.3}_{-0.3}$
72	I16362–4639	338.0642	-0.0675	-37.4 ± 0.4	$2.7^{+0.2}_{-0.2}$	145	I19095+0930	43.7941	-0.1275	44.7 ± 0.4	$9.2^{+0.3}_{-0.4}$
73	I16372–4545	338.8506	0.4091	-57.5 ± 0.4	$3.8^{+0.2}_{-0.2}$	146	I19097+0847	43.1791	-0.5191	57.3 ± 0.4	$8.5^{+0.5}_{-0.4}$

Note. I17441–2822 and 17455–2800 are close to the Galactic Centre, whose distance values are adopted from Faúndez et al. (2004); I18469–0132, I18507+0110, and I19097+0847 have the maser parallax distance measurements (Reid et al. 2014).

APPENDIX B: CALCULATION OF MASS AND MASS SURFACE DENSITY

To understand the nature of the 90 COM-containing cores not associated with the H/UC-H II signature (see Section 3.2.2) and 274 ‘unknown’ cores, we calculated their mass M_{core} in the following form:

$$M_{\text{core}} = \frac{F_{\nu}^{\text{int}} R_{\text{gd}} D^2}{B_{\nu}(T_{\text{dust}}) \kappa_{\nu}}, \quad (\text{B1})$$

where F_{ν}^{int} is the measured integrated 3 mm flux over the core size, R_{gd} is the gas-to-dust mass ratio (assumed to be 100), D is the distance to the core, B_{ν} is the Planck function for a dust temperature T_{dust} , and κ_{ν} is the opacity assumed to be $0.18 \text{ cm}^2 \text{ g}^{-1}$ of dust at $\nu \sim 94 \text{ GHz}$, taken directly from the ‘OH5’ dust model, a combination of dust from Ossenkopf & Henning (1994) and Pollack et al. (1994) and extended to longer wavelengths by Young & Evans (2005). Moreover, the mass surface density Σ_{core} was derived from $\Sigma_{\text{core}} = M_{\text{core}}/(\pi R_{\text{core}}^2)$, where R_{core} is the core radius equal to the geometric mean of the $\text{FWHM}_{\text{maj}}^{\text{dec}}$ and $\text{FWHM}_{\text{min}}^{\text{dec}}$ at the respective core distance. Given the median radius of the 90 COM-containing cores $\sim 0.02 \text{ pc}$, the $T_{\text{dust}} \sim 100 \text{ K}$ was assumed for a first-order approximation in the calculation, which is typical of HMCs at a radius of $\sim 0.01 \text{ pc}$ (Osorio, Lizano & D’Alessio 1999). For the 274 ‘unknown’ cores, the temperature 25 K is assumed, which is comparable to the mean value over the clumps that do not have cHMC or H/UC-H II signatures.

The peak column densities are estimated from the peak flux density of the cores using the following equation:

$$N_{\text{H}_2} = \frac{F_{\nu}^{\text{p}} R_{\text{gd}}}{B_{\nu}(T_{\text{dust}}) \Omega \kappa_{\nu} \mu m_{\text{H}}}, \quad (\text{B2})$$

where Ω is the beam solid angle, μ is the mean molecular weight of the gas, assumed to be 2.8, and m_{H} is the mass of the hydrogen atom. In turn, the number densities can be estimated via $n_{\text{H}_2} = N_{\text{H}_2}/2R$, assuming a spherical geometry for cores.

The three derived parameters M_{core} , n_{H_2} , and Σ_{core} are arranged in last three columns of Table 2 for the 32 s-cHMCs, of Table 3 for 58 w-cHMCs, and of Table 4 for ‘unknown’ cores.

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