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# Gas kinematics of key prebiotic molecules in GV Tau N revealed with an ALMA, PdBI, and Herschel synergy

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## ABSTRACT

A large effort has been made to detect warm gas in the planet formation zone of circumstellar discs using space and ground-based near-infrared facilities. GV Tau N, the most obscured component of the GV Tau system, is an outstanding source, being one of the first targets detected in HCN and the only one detected in CH<sub>4</sub> so far. Although near-infrared observations have shed light on its chemical content, the physical structure and kinematics of the circumstellar matter remained unknown. We use interferometric images of the HCN 3→2 and <sup>13</sup>CO 3→2 lines, and far-IR observations of <sup>13</sup>CO, HCN, CN, and H<sub>2</sub>O transitions to discern the morphology, kinematics, and chemistry of the dense gas close to the star. These observations constitute the first detection of H<sub>2</sub>O towards GV Tau N. Moreover, ALMA high spatial resolution (~ 7 au) images of the continuum at 1.1 mm and the HCN 3→2 line resolve different gas components towards GV Tau N, a gaseous disc with R~25 au, an ionized jet, and one (or two) molecular outflows. The asymmetric morphology of the gaseous disc shows that it has been eroded by the jet. All observations can be explained if GV Tau N is binary, and the primary component has a highly inclined individual disc relative to the circumbinary disc. We discuss the origin of the water and the other molecules emission according to this scenario. In particular, we propose that the water emission would come from the disrupted gaseous disc and the molecular outflows.

**Key words:** astrochemistry – protoplanetary discs – stars: formation – stars: individual (GV Tau N) – ISM: jets and outflows – stars: variables: T Tauri, Herbig Ae/Be.

## 1 INTRODUCTION

A fraction of the gas and dust in protoplanetary discs will end up in planets and may constitute the basis to form prebiotic species. A large effort has been made to detect the warm gas in the planet formation zone using Spitzer and ground-based facilities such as the Very Large Telescope (VLT) and the Keck Observatory telescopes. In a pioneering work, Lahuis et al. (2006) detected strong HCN absorption towards one source, IRS 46, from a sample of more than 100 Class I and II sources located in nearby star-forming regions. Later, Gibb et al. (2007) detected the HCN absorption in GV Tau. Since then, the observation of absorption and emission lines have provided a view of the molecular content of the very inner regions of protoplanetary discs with several detections of molecules such as CO, CO<sub>2</sub>, H<sub>2</sub>O, OH, HCN, and C<sub>2</sub>H<sub>2</sub> (Salyk et al. 2007, 2008, 2011, 2019; Carr & Najita 2008, 2011, 2014; Doppmann, Najita & Carr 2008; Najita et al. 2010, 2013; Pontoppidan et al. 2010; Doppmann et al. 2011; Fedele et al. 2011; Kruger et al. 2011; Mandell et al. 2012; Bast et al. 2013; Gibb & Horne 2013; Sargent et al. 2014; Banzatti et al. 2017). In addition, the organic species CH<sub>4</sub> has been detected towards GV Tau N (Gibb & Horne 2013). In many cases, the

profiles of the detected lines are broad and centrally peaked (Mandell et al. 2012; Salyk et al. 2019). This kind of centrally peaked profile is observed even for discs with previously determined high inclination angle and has been interpreted as emission from a disc wind. By now, direct imaging of the circumstellar material remains difficult with current telescopes, which challenges the interpretation.

The launch of the Herschel Space Observatory allowed to investigate the chemical content of discs at far-IR wavelengths (Thi et al. 2011; Salinas et al. 2016), providing valuable information on the abundance of H<sub>2</sub>O in large disc samples (Hogerheijde et al. 2011; Fedele et al. 2012, 2013; Meeus et al. 2012; Riviere-Marichalar et al. 2012; Podio et al. 2013; Salinas et al. 2016; Alonso-Martínez et al. 2017). Warm water coming from the inner R ~ 2–3 au disc region was detected through the 63.3 μm line observed with PACS (Meeus et al. 2012; Riviere-Marichalar et al. 2012). The emission of the ground ortho and para water lines is thought to originate in the disc surfaces where UV radiation desorbs the water from the icy mantles (Hogerheijde et al. 2011; Podio et al. 2013; Salinas et al. 2016). In Class I sources, a fraction of the warm water can be emitted in shocks along the outflow cavity as well (Kristensen et al. 2012; Karska et al. 2013; Mottram et al. 2014). Thus, the study of far-IR H<sub>2</sub>O lines provides essential information to characterize the H<sub>2</sub>O abundance from the inner to the cold outer regions of discs. Moreover, several guaranteed and open time Herschel Key Programs:

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‘Water In star-forming regions with Herschel’ (van Dishoeck et al. 2011), ‘CHEMical Survey of Star-forming regions’ (Ceccarelli et al. 2010), ‘Dust, Ice and Gas in Time’ (Green et al. 2013), and ‘GAS in protoplanetary systems’ (Mathews et al. 2010; Dent et al. 2013), have provided information on water in a large number of low-mass protostars at different evolutionary stages, allowing us to figure out how water is transported from the molecular cloud to planet-forming discs.

The chemical composition of discs in Class I protostars is receiving increasing attention. The presence of discs in all Class 0 protostars (ages  $\sim 0.2$  Myr) is still debated (see e.g. Maret et al. 2020). If indeed discs are present, their masses have been suggested to be high ( $\sim 0.05$ – $0.1 M_{\odot}$ ; Jørgensen et al. 2009) and they might be gravitationally unstable. Conversely, the masses of Class II discs in Taurus and Orion (ages  $\sim 1$ – $5$  Myr) have been well studied and are found to have a median mass of  $0.001 M_{\odot}$  (e.g. Eisner et al. 2008; Andrews et al. 2013; Tobin et al. 2020). These median masses are low compared with the amount of matter needed to form giant planets, estimated to be  $0.01$ – $0.1 M_{\odot}$  (e.g. Weidenschilling 1977; Desch 2007). Class I discs would then be in a transitional stage between massive, highly unstable protoplanetary discs to stable discs in which planet formation is progressing. The discs around Class I objects may therefore accurately represent the initial mass budget for forming planets.

GV Tau (Haro 6-10) is a T Tauri binary system embedded in the L1524 molecular cloud ( $d = 180 \pm 17$  pc; Gaia Collaboration 2018). It is one of the few young binary systems with the primary source (GV Tau S) optically visible, and the companion (GV Tau N), located  $1.2$  arcsec to the North, obscured. It is associated with a parsec-scale Herbig–Haro flow that extends for  $1.6$  pc to the North and  $\sim 1$  pc to the South (Devine et al. 1999; Movsessian & Magakian 1999). Both stars have been classified as Class I objects with masses of  $M \sim 0.8 M_{\odot}$  and ages of  $\sim 3$  Myr (Doppmann et al. 2008). Only the GV Tau N disc has been detected in the HCN and  $C_2H_2$  absorption ro-vibrational lines (Gibb et al. 2007, 2008). Indeed, GV Tau N is an outstanding source, being one of the first targets detected in HCN and the only one detected in  $CH_4$  (Gibb & Horne 2013). Moreover, the mass of its protoplanetary disc falls near the Minimum Mass Solar Nebula, and it may have just enough mass to form giant planets (Sheehan & Eisner 2014). Being a Class I disc with a chemistry rich in organic species, GV Tau N is a priority target to investigate how organics are delivered from the gas-rich disc to planets. Fuente, Cernicharo & Agúndez (2012) reported interferometric images of the HCN  $3 \rightarrow 2$  and  $HCO^+ 3 \rightarrow 2$  lines at an angular resolution of  $\sim 50$  au. Only the HCN  $3 \rightarrow 2$  emission was detected towards GV Tau N, and was interpreted as coming from the unresolved planet-forming region.

This paper explores the physical structure, kinematics, and chemistry of the gas in the planet-forming region of GV Tau N based on high spatial resolution images of the HCN  $3 \rightarrow 2$  and the  $^{13}CO 3 \rightarrow 2$  lines, and far-IR observations of the high excitation lines of  $^{13}CO$ ,  $H_2O$ , HCN, and CN. In particular, we present the first water detection in this system.

## 2 OBSERVATIONS

GV Tau was observed with the IRAM<sup>1</sup> Plateau de Bure Interferometer (PdBI), the Atacama Large Millimeter/sub-millimeter Array

(ALMA), and the HIFI instrument (de Graauw et al. 2010) on board the Herschel Space Telescope (Pilbratt et al. 2010). Observations of the  $^{13}CO 3 \rightarrow 2$  and  $C^{18}O 3 \rightarrow 2$  lines were carried out using the PdBI in its A configuration during 2014 October–November. This configuration provided a beam of  $0.35$  arcsec  $\times$   $0.18$  arcsec PA  $25^\circ$  ( $\sim 63$  au  $\times$   $32$  au at the Taurus distance). During the observations two 40 MHz correlator units were placed at the frequencies of the  $^{13}CO$  (330.588 GHz) and  $C^{18}O 3 \rightarrow 2$  (329.330 GHz) lines providing a spectral resolution of 78 kHz. The  $C^{18}O 3 \rightarrow 2$  line was not detected with an rms of  $46$  mJy beam<sup>-1</sup>. Calibration and imaging were performed using the Grenoble Image and Line Data Analysis Software (GILDAS; Pety 2005).

Our study uses the HCN  $3 \rightarrow 2$  line observations obtained from the ALMA archive. These observations were performed during Cycle 4 within project number 2016.1.00813.S. The line was observed with a spectral resolution of  $\sim 488.25$  kHz and the synthesized beam is  $0.07$  arcsec  $\times$   $0.04$  arcsec with PA  $27^\circ$ . This provides an unprecedented spatial resolution of  $\sim 12$  au  $\times$   $7$  au and the achieved rms is  $2.4$  mJy beam<sup>-1</sup>. We produced a  $1.1$  mm continuum image with the channels and spectral windows empty of line emission. The continuum image has a synthesized beam of  $0.06$  arcsec  $\times$   $0.03$  arcsec with PA  $27^\circ$ , and an rms of  $0.13$  mJy beam<sup>-1</sup>. The data were manually calibrated using the 4.7.1 version of the Common Astronomy Software Applications (CASA; McMullin et al. 2007), while the imaging process was performed using CASA version 5.4.1 following the National Radio Astronomy Observatory (NRAO) imaging guide.

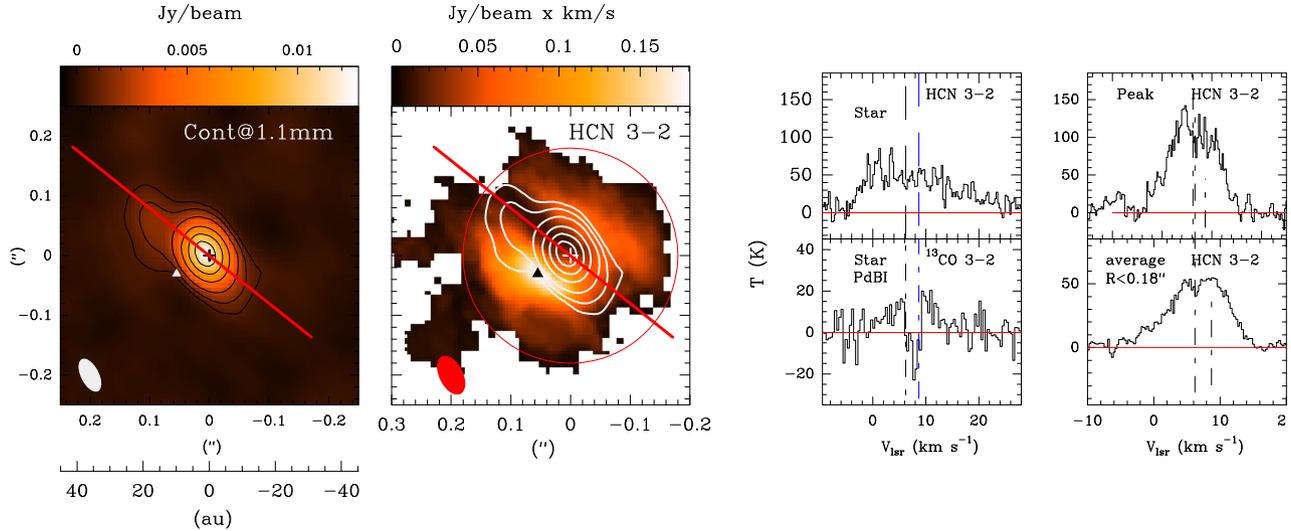
The Herschel Space Observatory (Pilbratt et al. 2010) observations presented here were obtained with the heterodyne instrument for the far-infrared HIFI (de Graauw et al. 2010; OT1\_afuente.2). The basic data reduction was performed using the standard pipeline provided with the version 7.0 of HIPE (Ott 2010) and then exported to GILDAS/CLASS (Pety 2005) for a more detailed analysis.

## 3 (SUB-)MILLIMETER INTERFEROMETRIC IMAGES

Fig. 1 shows the high angular resolution  $1.1$  mm continuum image obtained from ALMA data. Based on a 2D Gaussian fit, we derive the flux, position, and size of the compact continuum source. The estimated flux is  $\sim 48$  mJy, which is in good agreement with the previous measurement by Fuente et al. (2012). The size of the compact source is  $94 \pm 1$  mas  $\times$   $66 \pm 1$  mas, which constrains the radius of the dusty circumstellar disc to  $R \sim 8$  au. In addition to the compact source, elongated emission is detected in the direction of the outflow at  $>10\sigma$  level (see Fig. 1). It should be noted that the coordinates of the compact continuum source are  $\alpha(J2000) = 04^h 29^m 23^s.731$ ,  $\delta(J2000) = 24^\circ 33' 01''.11$ . This position is offset  $0.19$  arcsec to the South compared to the position measured by Fuente et al. (2012), which would imply proper motion of  $\sim (0, -34)$  mas yr<sup>-1</sup>. The most recent estimation of the proper motions of GV Tau N are  $(+8.695 \pm 0.853)$  mas yr<sup>-1</sup> in right ascension, and  $(-25.085 \pm 0.672)$  mas yr<sup>-1</sup> in declination (Gaia Collaboration 2018). Taking into account the lower angular resolution of Fuente et al. (2012), beam  $\sim 0.5$  arcsec, we consider that the observed displacement can be due to the proper motions of GV Tau N. The position of the continuum source is adopted as offset (0 arcsec, 0 arcsec) through this paper.

Reipurth et al. (2004) measured a  $3.6$  cm flux of  $0.10 \pm 0.01$  mJy towards GV Tau N. We estimate a flux of  $4.6 \pm 0.5$  mJy for the extended component, which would imply a  $3.6$  cm to  $1.1$  mm spectral index,  $\alpha_{1.1mm}^{3.6cm} \sim +1.1$ . This spectral index is higher than the one

<sup>1</sup>IRAM is supported by INSU/CNRS (France), MPG (Germany), and IGN (Spain).



**Figure 1.** *Left:* Continuum image at 1.1 mm as observed with ALMA. Levels are  $5\sigma$ ,  $10\sigma$ ,  $20\sigma$ – $100\sigma$  by  $20\sigma$ , where  $\sigma = 0.13$  mJy beam<sup>-1</sup>. *Centre:* Zero moment map of the HCN 3→2 line in units of Jy beam<sup>-1</sup> × km s<sup>-1</sup>. Contour levels are 0.01, 0.025, 0.05, 0.1, 0.15, and 0.175 Jy beam<sup>-1</sup> × km/s. The red line indicates the jet direction derived from the 1.1 mm continuum observations (*Jet 1*), and the circle marks the radius of the gaseous disc,  $R = 0.18$  arcsec. In white, the contours of the 1.1 mm continuum emission. *Right:* Interferometric spectra of the HCN 3→2 and <sup>13</sup>CO 3→2 lines towards the star position (the black and white crosses in left-hand and centre panels, respectively) as observed with ALMA and PdBI. The HCN 3→2 spectrum at the emission peak in the gaseous disc (the white and black triangles in the left-hand and centre panels) and the averaged HCN 3→2 spectrum in the region  $R < 0.18$  arcsec are shown in the right-hand panel. The intensity scale in the spectra is main beam brightness temperature. Vertical lines mark  $v_{\text{lsr}} = 6.3$  km s<sup>-1</sup> and  $8.7$  km s<sup>-1</sup>.

expected for an ionized spherical wind,  $\alpha \sim +0.6$ , and usually found in T Tauri and Herbig Ae/Be stars (Skinner, Brown & Stewart 1993; Alonso-Albi et al. 2009). This higher spectral index is still consistent with an ionized wind but with unconventional assumptions, i.e. a velocity power law  $v(R) \propto R$  or a density law,  $n_e(R) \propto R^{-3}$  (Panagia 1988). We should also consider the possibility that radio emission has varied on time-scales of tens of years. Variability is a common phenomenon in young T Tauri systems (Ubach et al. 2017; Liu et al. 2017). It would be also possible that we have overestimated the flux coming from the jet because a fraction of the extended emission is coming from dust thermal emission. Simultaneous measurements of the flux at several wavelengths would be desirable to further constrain the cm–mm spectral index.

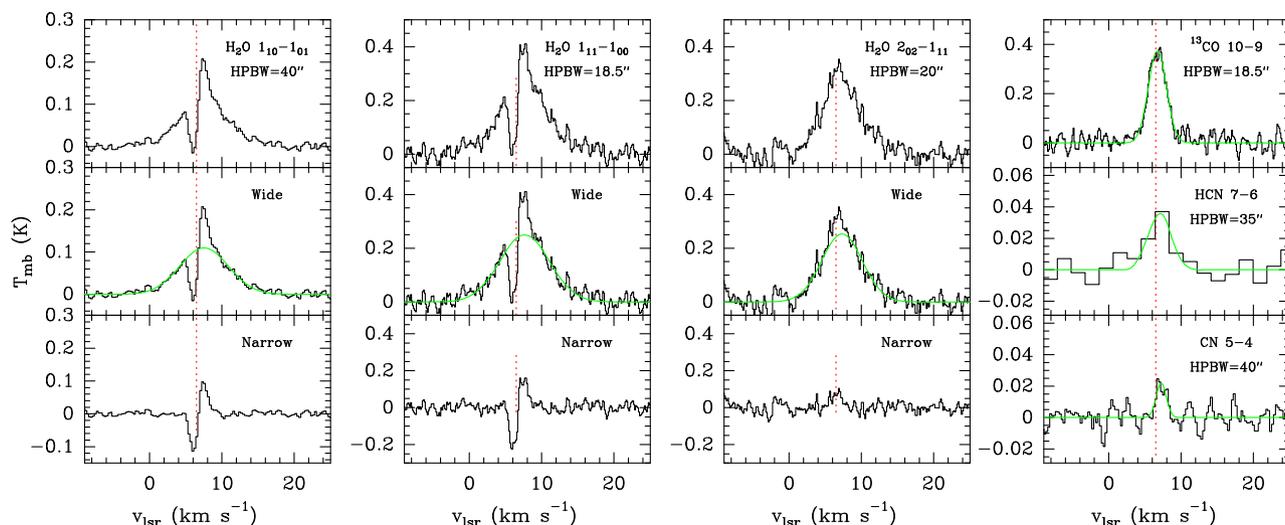
Mid-infrared interferometry (VLTI/MIDI) observations to probe the circumstellar material around GV Tau N and S were carried out by Roccatagliata et al. (2011). They modelled the emission assuming the existence of two different blackbody components, which accounts for the different disc regions. They obtained that the emission within 1.5 au has a temperature of 900 K, while a second colder emission at 150 K originates within 10 au from the central star. The sizes of these two components are consistent with the size of the 1.1 mm compact source as observed with ALMA for which we derive a radius  $R \sim 8$  au. Based on the measured visibilities, they estimate an inclination,  $i = 80^\circ$  for the GV Tau N disc. This value is consistent with GV Tau N being a highly obscured star.

Intense emission of the HCN 3→2 line is detected towards GV Tau N. We have produced the zero-moment map using only the channels with emission  $>3\sigma$ . Based on this map, we estimate that the size of the HCN emission is  $R \sim 0.18$  arcsec, significantly larger than the small compact continuum clump (see Fig. 1). The total velocity integrated flux within  $R \sim 0.18$  arcsec is  $\sim 1.4$  Jy beam<sup>-1</sup> × km s<sup>-1</sup>. To check for possible filtering in the HCN 3→2 line map, we have compared ALMA fluxes with those measured by Fuente et al. (2012) using the PdBI. Fuente et al. (2012) derived a total flux

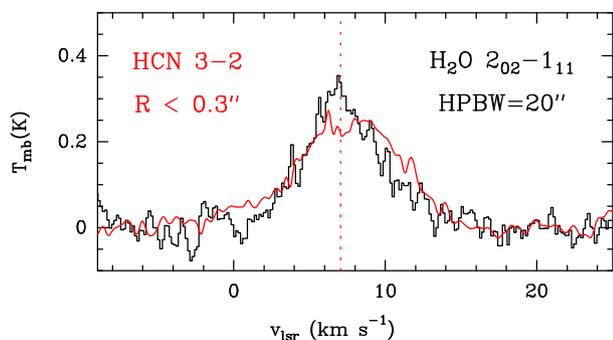
of  $3.64 \pm 0.21$  Jy × km s<sup>-1</sup> in a region of  $0.32$  arcsec ×  $0.25$  arcsec. Therefore, the high spatial resolution ALMA observations recover  $\sim 40$  percent of the flux detected in PdBI observations. This is consistent with the zero-moment map shown in Fig. 1, which presents evidence of a resolved component that extends towards the East.

An intense and asymmetric molecular ring is observed in the HCN 3→2 image shown in Fig. 1. The inner and outer radii of this ring are  $R_{\text{in}} \sim 0.04$  arcsec (6 au) and  $R_{\text{out}} \sim 0.18$  arcsec (17 au), respectively. In the azimuthal direction, the minimum of emission is located in the outflow direction, suggesting that the jet has eroded the circumstellar material, sweeping out the molecular gas. In Fig. 1, we also show the HCN 3→2 spectra towards the star and the emission peak in the molecular ring. The velocity profile towards the star is wide, with Full Width Half Power (FWHP) of  $\sim 12$  km s<sup>-1</sup>, and asymmetric, with the emission at highest velocities in the redshifted range. The spectrum towards the emission peak is narrower with FWHP  $\sim 10$  km s<sup>-1</sup>, and presents two peaks located at blue-shifted and red-shifted velocities with respect to the systemic velocity, which is not expected in a pure rotation scenario. This wide line is not surprising and it is most likely due to the strong interaction between the jet and the surrounding molecular gas. Also, self-absorption by the lower density envelope at the systemic velocity would produce the two-peaks feature.

The <sup>13</sup>CO 3→2 line is detected towards GV Tau N with a complex profile (see Fig. 1). The emission presents high velocity wings at the same velocities as the HCN 3→2 spectrum. In addition, a narrow, FWHP  $\sim 2$  km s<sup>-1</sup>, absorption feature is observed at  $v_{\text{lsr}} = 9.0$  km s<sup>-1</sup>, i.e. redshifted from the systemic velocity of  $v_{\text{lsr}} = 6.3$  km s<sup>-1</sup> and consistent with the emission coming from infalling gas. It is interesting to note that Doppmann et al. (2008) detected HCN and C<sub>2</sub>H<sub>2</sub> absorbing gas at  $v_{\text{lsr}} = 8.7 \pm 0.3$  km s<sup>-1</sup> (the blue-dashed line in Fig. 1), which is close to the velocity of the absorption detected in <sup>13</sup>CO. The signal-to-noise ratio of our <sup>13</sup>CO 3→2 observations does not allow to obtain a good quality image.



**Figure 2.** Spectra of the  $^{13}\text{CO } 10 \rightarrow 9$ ,  $\text{H}_2\text{O } 1_{11} \rightarrow 0_{00}$ ,  $1_{10} \rightarrow 1_{01}$ , and  $2_{02} \rightarrow 1_{11}$  lines obtained with the HIFI instrument on board the Herschel Space Observatory towards GV Tau N. The systemic velocity,  $v_{\text{lsr}} = 6.3 \text{ km s}^{-1}$ , is indicated with the red-dashed line.



**Figure 3.** Spectra of the  $\text{H}_2\text{O } 2_{02} \rightarrow 1_{11}$  line obtained with the HIFI instrument on board the Herschel Space Observatory towards GV Tau N. For comparison, we have plotted the averaged ALMA spectrum of  $\text{HCN } 3 \rightarrow 2$  in a circle of  $R < 0.3$  arcsec around GV Tau N scaled to match the intensity of the  $\text{H}_2\text{O } 2_{02} \rightarrow 1_{11}$  line. The systemic velocity,  $v_{\text{lsr}} = 6.3 \text{ km s}^{-1}$ , is indicated with the red-dashed line.

#### 4 FAR-INFRARED LINES

High excitation lines of  $^{13}\text{CO}$ ,  $\text{H}_2\text{O}$ ,  $\text{HCN}$ , and  $\text{CN}$  were observed with high spectral resolution using the instrument HIFI onboard Herschel (see Fig. 2). High signal-to-noise detections were obtained for the  $^{13}\text{CO } 10 \rightarrow 9$ , and  $\text{H}_2\text{O } 1_{11} \rightarrow 0_{00}$ ,  $1_{10} \rightarrow 1_{01}$ , and  $2_{02} \rightarrow 1_{11}$  lines. The  $\text{HCN } 7 \rightarrow 6$  and  $\text{CN } 6 \rightarrow 5$  lines were detected at  $4\sigma$  level (see Fig. 2). Unfortunately, the Herschel angular resolution does not allow to resolve the two components of the binary, GV Tau N, and GV Tau S, and we need to use the kinematic information to discern the origin of the observed lines.

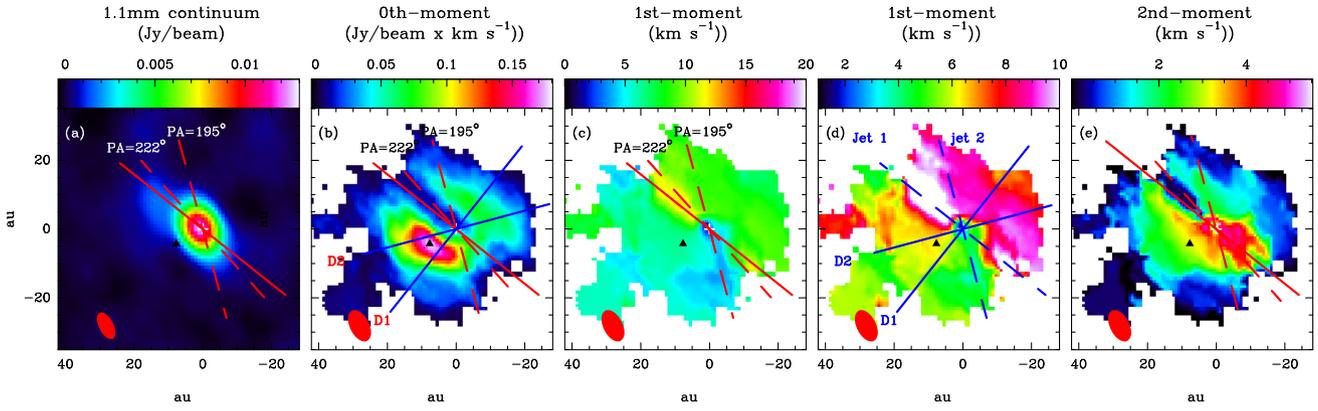
It is remarkable the difference between the profile of the  $^{13}\text{CO } 10 \rightarrow 9$  line and those of the  $\text{H}_2\text{O}$  lines. The  $^{13}\text{CO } 10 \rightarrow 9$  line presents a narrow profile, with FWHP  $\sim 2.7 \text{ km s}^{-1}$ , and is centred at  $v_{\text{lsr}} = 6.74 \pm 0.05 \text{ km s}^{-1}$ . The profiles of the ground ortho- and para-water lines are significantly wider, and show a strong absorption at  $v_{\text{lsr}} \sim 6.0 \text{ km s}^{-1}$ . The absorption is not observed in the higher excitation p- $\text{H}_2\text{O } 2_{02} \rightarrow 1_{11}$  line ( $E_u/k \sim 100 \text{ K}$ ). Indeed, the profile of the water lines can be fitted using two components: (i) a wide Gaussian with FWHP  $\sim 8 \text{ km s}^{-1}$ , and (ii) a narrower component with FWHP  $\sim 2.7 \text{ km s}^{-1}$ , similar to  $^{13}\text{CO}$ . Fig. 2

shows the decomposition of the water line profiles in these two components that we hereafter name wide and narrow components. The Gaussian fits to these components are shown in Table A1. It is specially interesting to discern whether the emission of the wide component arises in the compact gaseous disc detected with the ALMA maps or it is coming from a large scale outflow(s). To explore the origin of this component, we compare its profile with the mean  $\text{HCN } 3 \rightarrow 2$  spectrum in the region  $R < 0.3$  arcsec (see Fig. 3). The wings of the p- $\text{H}_2\text{O } 2_{02} \rightarrow 1_{11}$  line are similar to those of the interferometric  $\text{HCN } 3 \rightarrow 2$  line, suggesting that the origin might be similar. We further discuss the origin of this component in Section 6.

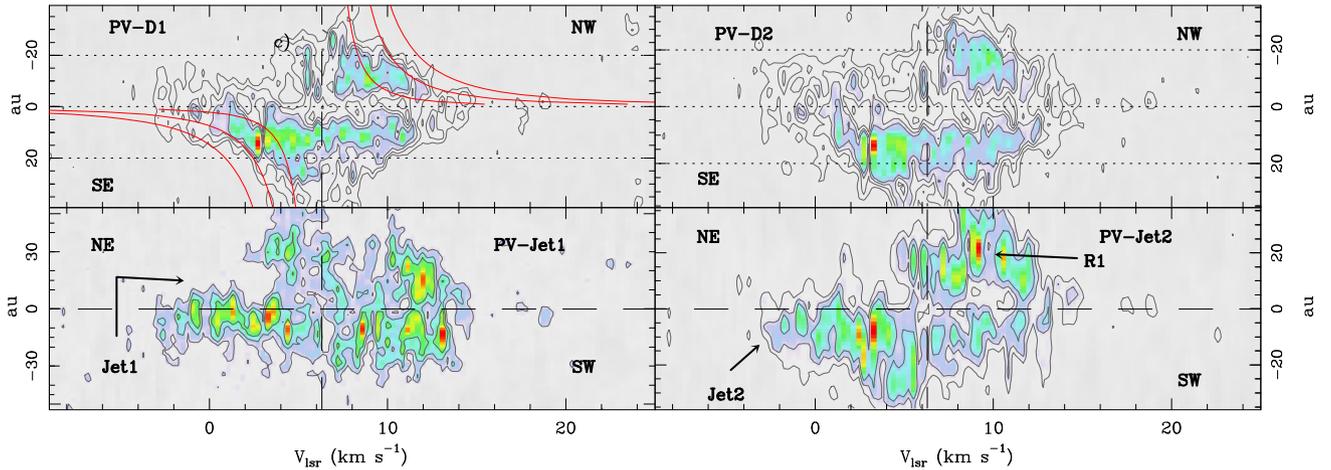
In the bottom panels of Fig. 2, we show the profiles of the water narrow component. This component presents absorption at blue-shifted velocities. We recall that the absorptions observed in the near-infrared (NIR) observations lie at red-shifted velocities,  $\sim 8.7 \text{ km s}^{-1}$ . An absorption at similar velocity,  $\sim 9 \text{ km s}^{-1}$ , has been observed in our PdBI spectrum of the  $^{13}\text{CO } 3 \rightarrow 2$  line. Thus, the blue absorption features detected in the FIR water lines are more likely caused by a colder layer in the outer envelope. Fig. 2 also shows the spectra of the  $\text{HCN } 7 \rightarrow 6$  and  $\text{CN } 5 \rightarrow 4$  lines. These spectra only present the narrow component. We cannot discard, however, the existence of a weaker wide component that it is not detected because of the low S/N of the detections. This should be the case of  $\text{HCN } 7 \rightarrow 6$ , for which emission presents a slightly wider spectrum. It should be noticed that the peak of the weak and narrow  $\text{HCN } 7 \rightarrow 6$  lines is at the velocity of the dip observed in the ALMA  $\text{HCN } 3 \rightarrow 2$  profile, suggesting that the narrow line observed with Herschel is probing the lower density envelope.

#### 5 KINEMATICS OF THE MOLECULA GAS

In Fig. 4, we present the first and second moment maps of the  $\text{HCN } 3 \rightarrow 2$  emission as observed with ALMA. The emission at high red- and blue-shifted velocities,  $|v - v_0| > 4 \text{ km s}^{-1}$ , shows the presence of two misaligned outflow lobes: the red lobe has the same direction as the ionized jet traced by the 1mm continuum emission (hereafter, *Jet 1*), which is close to the direction defined by the giant Herbig-Haro flow,  $\text{PA} = 222^\circ$  (Devine et al. 1999); and the blue lobe is placed at  $\text{PA} = 195^\circ$ , the direction of the Herbig-Haro jet detected by



**Figure 4.** (a) to (d) Continuum image at 1.1 mm and moment maps of the HCN 3→2 emission as observed with ALMA. The dashed lines indicate the direction of the giant Herbig–Haro flow at PA = 222° (Devine et al. 1999) and of the optical jet detected at PA = 195° (Movsessian & Magakian 1999). We also indicate the directions perpendicular to *Jet 1*, named *D1*, and to *Jet 2*, named *D2*. Panels (c) and (d) show the same map with different colour scales to highlight the bipolar outflow(s) (c) or the rotation motion (d). The directions *Jet 1*, *Jet 2*, *D1*, and *D2* are selected to perform the position–velocity diagrams.



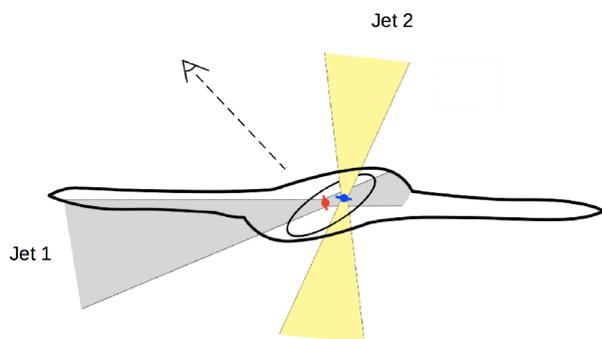
**Figure 5.** Position–velocity diagrams of the HCN 3→2 line along the strips *Jet 1*, *Jet 2*, *D1*, and *D2*. In panel *D1*, we show the Keplerian velocity curves assuming  $M_* = 0.8 M_\odot$  and  $i = 20^\circ, 40^\circ, \text{ and } 65^\circ$  ( $i = 0^\circ$  would be face-on) for the circumbinary disc. Observational data fits better the  $i = 40^\circ$  curve. Contour levels are  $2\sigma$ – $5\sigma$  in steps of  $1\sigma$  ( $\sigma = 2.4 \text{ mJy beam}^{-1}$ ).

Movsessian & Magakian (1999; hereafter, *Jet 2*). This misalignment can be interpreted as the consequence of a precessing bipolar outflow. Alternatively, these lobes might belong to two different bipolar outflows.

To discern between these two scenarios, we inspect the channel maps shown in Figs A1 and A2. The directions of *Jet 1* and *Jet 2* are indicated in each panel. The red lobe of *Jet 1* is clearly seen in panels  $v_{\text{lsr}} = +10.28$  to  $+13.36 \text{ km s}^{-1}$ , located SW from the star. The blue counterpart is missing in the HCN 3→2 channels maps. The molecular outflow is formed by the circumstellar matter that is being swept up by the ionized jet. The lack of a blue lobe is due to the absence of the molecular gas in this direction, where a wide cavity has already been excavated by the jet, and the molecular gas has been cleaned. The existence of high velocity gas along *Jet 2* is not such clear since the kinematic signatures of the blue and red lobes are mixed with the disc rotation. High velocity gas along *Jet 2* is detected in panels  $v_{\text{lsr}} = -1.48$  to  $+2.16 \text{ km s}^{-1}$ . Fig. A2 shows that high velocity gas is detected in panels  $v_{\text{lsr}} = +10.00$  to  $+14.76 \text{ km s}^{-1}$ , which is located in the *Jet 2* direction and can be interpreted

as the red lobe of this outflow. Given the angular resolution of the observations, it might also be associated to the walls of the cavity excavated by *Jet 1*.

Fig. 5 shows the position–velocity (P–V) diagrams along *Jet 1* and *Jet 2*. The *Jet 1* P–V diagram presents clumpy emission with small changes in the velocity between one small bright clump and the following. These bright spots are also observed as narrow features in the HCN 3→2 spectrum shown in Fig. 1. This kind of ‘saw-like’ features have been detected in high-velocity jets and were interpreted as oblique shocks formed in the interphase between the high-velocity jet and the circumstellar matter (Santiago-García et al. 2009). In this scenario, the clumps at red-shifted and blue-shifted velocities correspond to the back and front walls of the cavity excavated by the jet in the gaseous disc. This interpretation is also consistent with the morphology observed in the zero moment map of the HCN 3→2 emission (see Fig. 1). The outflow *Jet 1* crosses the molecular ring through the azimuthal minimum of the HCN 3→2 emission, probing that the gas has been swept up.



**Figure 6.** Sketch of the geometry proposed for GV Tau N to explain ALMA spectroscopic observations.

The P–V diagram along *Jet 2* is also shown in Fig. 5. The outflow *Jet 2* presents bipolar morphology with the highest velocities close to the star. In this case, the blue lobe is more intense than the red lobe, contrary to the case of *Jet 1*. An intense clump is detected at  $10 \text{ km s}^{-1}$  in the direction of *Jet 2* (R1 in Fig. 5) that might be disc gas accelerated by the *Jet 2*. However, we need to be cautious in our conclusions because the sensitivity and angular resolution of the HCN 3→2 observations do not allow to fully disentangle *Jet 2* from the gas that is accelerated in the walls of the cavity excavated by *Jet 1*.

To explore the kinematics of the circumbinary disc, we performed the P–V diagrams along the directions perpendicular to *Jet 1* and *Jet 2*, named *D1* and *D2*, respectively (Fig. 5). The P–V diagrams along *D1* and *D2* present emission in the four quadrants, which is not compatible with a pure rotating disc. A combination of rotation and outflow is necessary to explain this pattern. Furthermore, there is a gap around the position of the stars with a size,  $R \sim 8 \text{ au}$ , in both diagrams. In spite of the highly disturbed kinematics, we can see a hint for Keplerian rotation in the P–V diagram along *D1*. For comparison, we compare the observations with the Keplerian curves for a star with  $M_* = 0.8 M_\odot$  and inclination angles  $i = 20^\circ$ ,  $40^\circ$ , and  $65^\circ$ . We find reasonable agreement between the observations in the inner  $R < 20 \text{ au}$  region and the Keplerian curves assuming  $i = 40^\circ$  curve. This suggests that the HCN 3→2 emission comes from a rotating disc in which kinematics is highly perturbed by the interaction with the bipolar outflow(s). We recall that there is a degeneracy between the stellar mass and the inclination angle. If GV Tau N is a binary, the stellar mass content would be larger, and a lower inclination angle would be required to fit the observations. This would be consistent with the observed morphology, which is that expected for an almost face-on disc.

Doppmann et al. (2008) proposed that GV Tau N could be a binary to explain the large difference between the velocity of the lines coming from the stellar photosphere and the molecular gas. The existence of two bipolar outflows would also explain the kinematics observed in the ALMA HCN 3→2 cube. The most energetic outflow, *Jet 1*, should be in (or close to) the equatorial plane of the circumbinary disc to explain the cavity cleaned of gas observed in the HCN maps. The second outflow would be perpendicular, or close to the perpendicular direction, to the circumbinary disc, hence flowing away with little interaction with the molecular gas (see a sketch in Fig. 6). This geometry would imply that the two components of the binary are highly misaligned.

Binarity could also explain the different morphology of the gas and dust around GV Tau N, the size of the dust continuum emission being smaller. Stellar binaries exert gravitational torques on inclined

**Table 1.** Summary of molecular observations.

Species	$T_{\text{rot}}(\text{K})$	$N(\text{cm}^{-2})$	Ref
NIR molecular absorption			
$^{12}\text{CO}$	$200 \pm 40$	$\sim 1.2 \times 10^{19}$	(1)
$^{13}\text{CO}$	$260 \pm 20$	$\sim 1.1 \times 10^{17}$	(1)
$\text{C}^{18}\text{O}$	$140 \pm 50$	$(1.4 \pm 0.5) \times 10^{16}$	(1)
$\text{C}_2\text{H}_2$	$170 \pm 20$	$(7.3 \pm 0.2) \times 10^{16}$	(1)
	$720 \pm 60$	$(1.4 \pm 0.3) \times 10^{16}$	(3)
$\text{CH}_4$	$750 \pm 50$	$(2.8 \pm 0.2) \times 10^{17}$	(4)
$\text{CO}_2$	$250 \pm 25$	$(5.1 \pm 0.7) \times 10^{16}$	(3)
$\text{HCN}$	$115 \pm 11$	$(3.7 \pm 0.3) \times 10^{16}$	(1)
	$440 \pm 40$	$(1.8 \pm 0.4) \times 10^{16}$	(3)
	550	$1.5 \times 10^{17}$	(2)
FIR molecular emission			
Species	$T_{\text{rot}}(\text{K})$	$N(\text{cm}^{-2})$	Comments
o- $\text{H}_2\text{O}$	500	$> 10^{17}$	Wide
p- $\text{H}_2\text{O}$	500	$> 10^{17}$	Wide
$^{13}\text{CO}$	250	$5.0 \times 10^{14}$	Narrow <sup>a</sup>
$\text{HCN}$	250	$4.0 \times 10^{11}$	Narrow <sup>a</sup>
$\text{CN}$	250	$6.0 \times 10^{11}$	Narrow <sup>a</sup>
mm interferometric emission			
Species	$T_{\text{rot}}(\text{K})$	$N(\text{cm}^{-2})$	Comments
$^{13}\text{CO}$	250 – 500	$(2.6 – 5.0) \times 10^{17}$	Wide
$\text{HCN}$	250 – 500	$(2.0 – 3.5) \times 10^{15}$	Wide

Notes. References: (1) Gibb et al. 2007, 2008; (2) Doppmann et al. 2008; (3) Bast et al. 2013; (4) Gibb & Horne 2013

<sup>a</sup>Beam averaged column densities (beam = 18.5 arcsec).

circumbinary discs causing radially differential precession and, under some circumstances, warps in the disc. Large dust grains in thick discs pile up at the warp location, forming narrow dust rings within an extended gaseous disc (Aly & Lodato 2020). This warped narrow dust ring could be the emitting region of the  $T = 150 \text{ K}$  component detected in the mid-IR interferometric observations by Roccatagliata et al. (2011).

Misaligned binaries are common in young stellar systems (see e.g. Brinch et al. 2016; Takakuwa et al. 2017). GV Tau S was also proposed to be a spectroscopic binary by Doppmann et al. (2008). If confirmed, GV Tau would be a quadruple star system such as GG Tau (Andrews et al. 2014) and HD98800B (Kennedy et al. 2019).

## 6 MOLECULAR CHEMISTRY

In this section, we compile the column densities and rotation temperatures that have been derived using observations from NIR to mm wavelengths (see Table 1). Even when comparing lines of the same species, ro-vibrational and rotational lines are emitted from different regions of this complex system and the comparison is not straightforward. Our goal is to put together all the pieces of information to figure out the overall scenario.

The physical and chemical conditions in the inner  $R < 1–10 \text{ au}$  are better probed by NIR observations. Different rotation temperatures and molecular column densities have been measured with NIR molecular absorption observations towards GV Tau N (see Table 1). In particular, Gibb et al. (2007) derived  $T_{\text{rot}} = 115 \pm 11 \text{ K}$  and  $N(\text{HCN}) = (3.7 \pm 0.3) \times 10^{16} \text{ cm}^{-2}$  in the first detection of the HCN absorption lines towards this star. Later, Doppmann et al. (2008) and Bast et al. (2013) estimated a higher rotation temperature,  $T_{\text{rot}} > 500 \text{ K}$  for HCN (see Table 1). Doppmann et al. (2008) argued that this difference is due to the different set of lines used in the calculations. Gibb et al. (2007) calculations were based on lines at

lower energy and obtained a lower rotation temperature, as expected in the case of a steep gas kinetic temperature gradient along the line of sight. The velocity centroid of the  $\text{C}_2\text{H}_2$  and HCN absorption lines is  $v_{\text{lsr}} = 8.7 \pm 0.3 \text{ km s}^{-1}$ , similar to that of the absorption feature that we have detected in the  $^{13}\text{CO}$  3→2 line on the top of high-velocity wings, thus consistent with the scenario of these molecules coming from infalling gas towards the star. Gibb et al. (2007, 2008) also derived  $^{12}\text{CO}$ ,  $^{13}\text{CO}$ , and  $\text{C}^{18}\text{O}$  rotation temperatures based on molecular absorption (see Table 1). They obtained  $T_{\text{rot}} = 140\text{--}260 \text{ K}$ , suggesting that the observed lines are probing molecular gas located further from the star than HCN and  $\text{C}_2\text{H}_2$ , although along the line of sight.

We present the first detection of water towards GV Tau (see Fig. 2). In particular, three water lines, p- $\text{H}_2\text{O}$   $1_{11} \rightarrow 0_{00}$ , o- $\text{H}_2\text{O}$   $1_{10} \rightarrow 1_{01}$ , and p- $\text{H}_2\text{O}$   $2_{02} \rightarrow 1_{11}$  have been detected. The emission of the ground ortho and para water lines is thought to originate in the disc surfaces and in shocks along the outflow cavity in Class I sources (Hogerheijde et al. 2011; Kristensen et al. 2012; Karska et al. 2013; Podio et al. 2013; Mottram et al. 2014; Salinas et al. 2016). The FIR water lines show two well differentiated components: one with a wide profile, very similar to that derived from the interferometric HCN 3→2 observations; the second one is characterized by a narrow profile and was interpreted as coming from the outer envelope. In the following, we explore the possibility that the emission of the wide component comes from the GV Tau N disrupted gaseous disc detected with ALMA observations of the HCN 3→2 line. First, we would assume that the bulk of the water emission is coming from GV Tau N. This is a reasonable assumption taking into account that previous NIR and millimeter observations showed that molecular absorption/emission is mainly coming from the North component (Gibb et al. 2007, 2008; Doppmann et al. 2008; Fuente et al. 2012; Bast et al. 2013). Secondly, we use the two observed p- $\text{H}_2\text{O}$  line to investigate the physical conditions of the emitting region. Based on our Gaussian fits to the wide component (Table A1), and after correcting for the small difference in the beam size, we derive an integrated intensity ratio,  $R_{\text{obs}}(\text{H}_2\text{O}@987 \text{ GHz}/1113 \text{ GHz}) = 1.0 \pm 0.1$  for the two p- $\text{H}_2\text{O}$  lines observed. This ratio is not consistent with optically thin emission. In the optically thin limit, and assuming Local Thermodynamic Equilibrium (LTE), we estimate  $R_{\text{LTE}}(\text{H}_2\text{O}@987 \text{ GHz}/1113 \text{ GHz}) = 0.5$  at 200 K, and increases with temperature until  $\sim 0.63$  for  $T_k > 500 \text{ K}$ . The observed ratio can only be explained if the emission of p- $\text{H}_2\text{O}$  is optically thick. This would imply  $N(\text{p-}\text{H}_2\text{O}) > \text{a few } 10^{17} \text{ cm}^{-2}$ . This high value of the water column density is within the range of water column densities derived towards T Tauri stars using Spitzer observations (Salyk et al. 2007). Thus, it is not unreasonable that the emission comes from the compact disrupted gaseous disc. To further check the consistency of our interpretation, we perform the following calculation: assuming  $T_{\text{rot}} \sim 500 \text{ K}$  as derived from HCN absorption observations, optically thick emission, and a source size of 0.5 arcsec (see Fig. 1), we predict  $T_{\text{mb}} \sim 0.23 \text{ K}$  for the p- $\text{H}_2\text{O}$  lines after diluting to the Herschel beam, which is fully consistent with Herschel observations.

It is also interesting to compare the intensities of the ground p- $\text{H}_2\text{O}$  line with the o- $\text{H}_2\text{O}$   $1_{10} \rightarrow 1_{01}$  line at 557 GHz. In the optically thick case, we would expect  $R_{\text{thick}}(\text{H}_2\text{O}@1113 \text{ GHz}/557 \text{ GHz}) \sim 4.7$  because of the different observational beams. Instead, the observed value is  $R_{\text{obs}}(\text{H}_2\text{O}@1113 \text{ GHz}/557 \text{ GHz}) \sim 2.4 \pm 0.1$ . This can be explained if the size of the emitting region is larger for ortho- $\text{H}_2\text{O}$  than for the para- $\text{H}_2\text{O}$ . Indeed, this is expected according to the critical densities of these line ( $\sim 2 \times 10^7$  and  $\sim 2 \times 10^8 \text{ cm}^{-3}$  at 50 K for the 557 and 1113 GHz lines, respectively), and the higher abun-

dance of o- $\text{H}_2\text{O}$  (the ortho-to-para ratio is  $\sim 3$  at high temperatures). Direct imaging of these lines at high angular resolution is required to confirm this scenario.

We have also derived  $^{13}\text{CO}$ , HCN, and CN column densities based on FIR observations. In these cases, the lines show a narrow profile and therefore we assume a lower rotation temperature  $T_{\text{rot}} = 250 \text{ K}$ , similar to those measured from NIR absorption lines of  $^{13}\text{CO}$  (Gibb et al. 2007, 2008). Again, the unknown source size introduces a large uncertainty in our calculations. To make a fair comparison between molecules, we assumed a source size of 18.5 arcsec. With these assumptions, we derive the following source-averaged column densities:  $N(^{13}\text{CO}) = 5.0 \times 10^{14} \text{ cm}^{-2}$ ,  $N(\text{HCN}) = 4.0 \times 10^{11} \text{ cm}^{-2}$ , and  $N(\text{CN}) = 6.0 \times 10^{11} \text{ cm}^{-2}$ .

Finally, we explore the chemistry of the compact emission around GV Tau N using the millimeter interferometric images. Fuente et al. (2012) concluded that the HCN/ $\text{HCO}^+$  abundance ratio should be  $> 300$  to be compatible with the non-detection of  $\text{HCO}^+$  3→2 with PdBI observations. Here, we will compare the HCN 3→2 observations obtained with ALMA with the PdBI observations of the  $^{13}\text{CO}$  →2 line. For this comparison, we use the averaged HCN 3→2 spectrum in  $R < 0.18$  arcsec, since the beam of  $^{13}\text{CO}$  3→2 observations is  $\sim 0.3$  arcsec. We can fit the PdBI  $^{13}\text{CO}$  3→2 line profile, assuming LTE,  $T_{\text{rot}} = 250 - 500 \text{ K}$  and  $N(^{13}\text{CO}) = (2.6 - 5.0) \times 10^{17} \text{ cm}^{-2}$ . At these high temperatures, the emission of the  $^{13}\text{CO}$  3→2 is optically thin and it is not very dependent on the assumed gas kinetic temperature. This estimate is in reasonable agreement with the value derived from the NIR absorption lines, suggesting that we are probing the same gas component. The  $R < 0.18$  arcsec averaged profile of the HCN 3→2 emission can be fitted with  $T_{\text{rot}}=250\text{--}500 \text{ K}$  and  $N(\text{HCN}) = (2.0\text{--}3.5) \times 10^{15} \text{ cm}^{-2}$ , i.e. two orders of magnitude lower than the values derived from the NIR molecular absorptions. This suggests that the mm emission is probing a different gas component than the molecular absorption in NIR. Molecular absorptions arise in a very small and hot region where the HCN lines are optically thick, and  $T_b(\text{HCN } 3\rightarrow 2) \sim 500 \text{ K}$ . To be compatible with ALMA observations, the size of this region should be  $R \sim 2\text{--}3 \text{ au}$  around the star, confirming its origin in the GV Tau N individual disc.

The observation of  $^{13}\text{CO}$  and HCN with different techniques provides valuable information on the chemical changes occurring in this complex system. As shown in Table 1, we obtain that the HCN/ $^{13}\text{CO}$  abundance ratio is  $\sim 10^{-3}$  based on the FIR observations,  $\sim 10^{-2}$  in the gaseous disc traced with ALMA and PdBI observations, and  $\sim 1$  in the close vicinity of the star where NIR molecular absorptions arise. These different values of the HCN/ $^{13}\text{CO}$  ratio testify for chemical gradients in this system, with the HCN/ $^{13}\text{CO}$  ratio decreasing with the distance from the star. This behaviour of the HCN abundance, peaking at  $R \sim 2\text{--}3 \text{ au}$  from the star is consistent with chemical models predictions (see e.g. Walsh, Millar & Nomura 2010; Agúndez et al. 2018). These models also predict that a high abundance of  $\text{H}_2\text{O}$  is expected in this inner region, coeval with HCN. The narrow component appearing in the  $\text{H}_2\text{O}$  line profiles is thought to come from a more extended envelope component.

## 7 CONCLUDING REMARKS

The high spatial resolution of ALMA observations allows us, for the first time, to discern the different gas components in the vicinity of GV Tau N. The young star GV Tau N is located in the inner gap of a gaseous Keplerian disc of  $R \sim 20 \text{ au}$ . An ionized jet is detected in the continuum at 1.1 mm. Furthermore, ALMA observations present evidence of the existence of two molecular outflows, one

of them associated with the continuum jet. We propose that all these observations can be easily explained if GV Tau N is a binary, as already suggested by Doppmann et al. (2008). The individual disc of the primary is highly misaligned relative to the circumbinary disc, and its jet is disrupting the molecular ring (see Fig. 6).

The present detection of the o-H<sub>2</sub>O and p-H<sub>2</sub>O line in GV Tau N is of paramount importance. Determining the water abundance in cold proto-planetary discs has been remarkably difficult. Water vapor detections using the ground state H<sub>2</sub>O lines have been reported for just two discs, the well-known transition disc TW Hya (Hogerheijde et al. 2011; Salinas et al. 2016), and the Class I disc DG Tau (Podio et al. 2013). Stringent upper limits to the water abundance have been reported for a dozen other sources at an order of magnitude lower than expected (Bergin et al. 2010; Du et al. 2017). Recently, Harsono et al. (2020) unsuccessfully searched for water in five Class I discs using ALMA observations of millimeter lines of the rarer isotopologue H<sub>2</sub><sup>18</sup>O. The detection of cold water towards Class I protostar GV Tau stands as a major result that deserves further study.

We also present observations of high excitation lines of <sup>13</sup>CO, HCN, and CN as observed with the HIFI/Herschel instrument, which provide information on the column densities of these species at spatial scales of  $\sim$  a few 10 arcsec. As shown in Table 1, the HCN/<sup>13</sup>CO abundance ratio is  $\sim 10^{-3}$  in the warm envelope,  $\sim 10^{-2}$  in the gaseous disc better traced with ALMA and PdBI observations, and  $\sim 1$  in the close vicinity of the star where NIR molecular absorptions arise. This behaviour of the HCN abundance, peaking at  $R \sim 2-3$  au from the star is consistent with chemical models predictions (see e.g. Walsh et al. 2010; Agúndez et al. 2018). These models also predict that a high abundance of H<sub>2</sub>O is expected in this inner region, coeval with HCN.

Spitzer observations revealed that GV Tau N is specially rich in organic species (Gibb et al. 2007, 2008; Doppmann et al. 2008; Bast et al. 2013). The interaction of the jet with the circumbinary gaseous disc is expected to produce intense shocks that could contribute to enhance the abundance of organic species in this disc. High spatial resolution interferometric observations are required to disentangle the interplay between chemistry and dynamics in this prototypical system.

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## DATA AVAILABILITY

The data underlying this article will be shared on reasonable request to the corresponding author.

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## APPENDIX A: SUPPORTING MATERIAL

**Table A1.** Gaussian fits to FIR observations.

Line	Freq (GHz)	Beam (arcsec)	Area (K km s <sup>-1</sup> )	V <sub>lsr</sub> (km s <sup>-1</sup> )	FWHP (km s <sup>-1</sup> )	T <sub>mb</sub> (K)
o-H <sub>2</sub> O 1 <sub>10</sub> → 1 <sub>01</sub>	556.94	40	0.93 (0.02)	7.52 (0.09)	7.92 (0.23)	0.11
p-H <sub>2</sub> O 1 <sub>11</sub> → 0 <sub>00</sub>	1113.34	18.5	2.14 (0.05)	7.58 (0.08)	7.42 (0.20)	0.27
p-H <sub>2</sub> O 2 <sub>02</sub> → 1 <sub>11</sub>	987.93	20	1.79 (0.06)	7.32 (0.09)	6.64 (0.29)	0.25
<sup>13</sup> CO 10→9	1101.35	18.5	1.18 (0.03)	6.72 (0.03)	2.97 (0.08)	0.37
CN 5→4	566.95	40	0.04 (0.01)	7.25 (0.13)	1.65 (0.25)	0.02
HCN 7→6	620.30	35	0.14 (0.03)	7.02 (0.37)	3.71 (1.14)	0.04

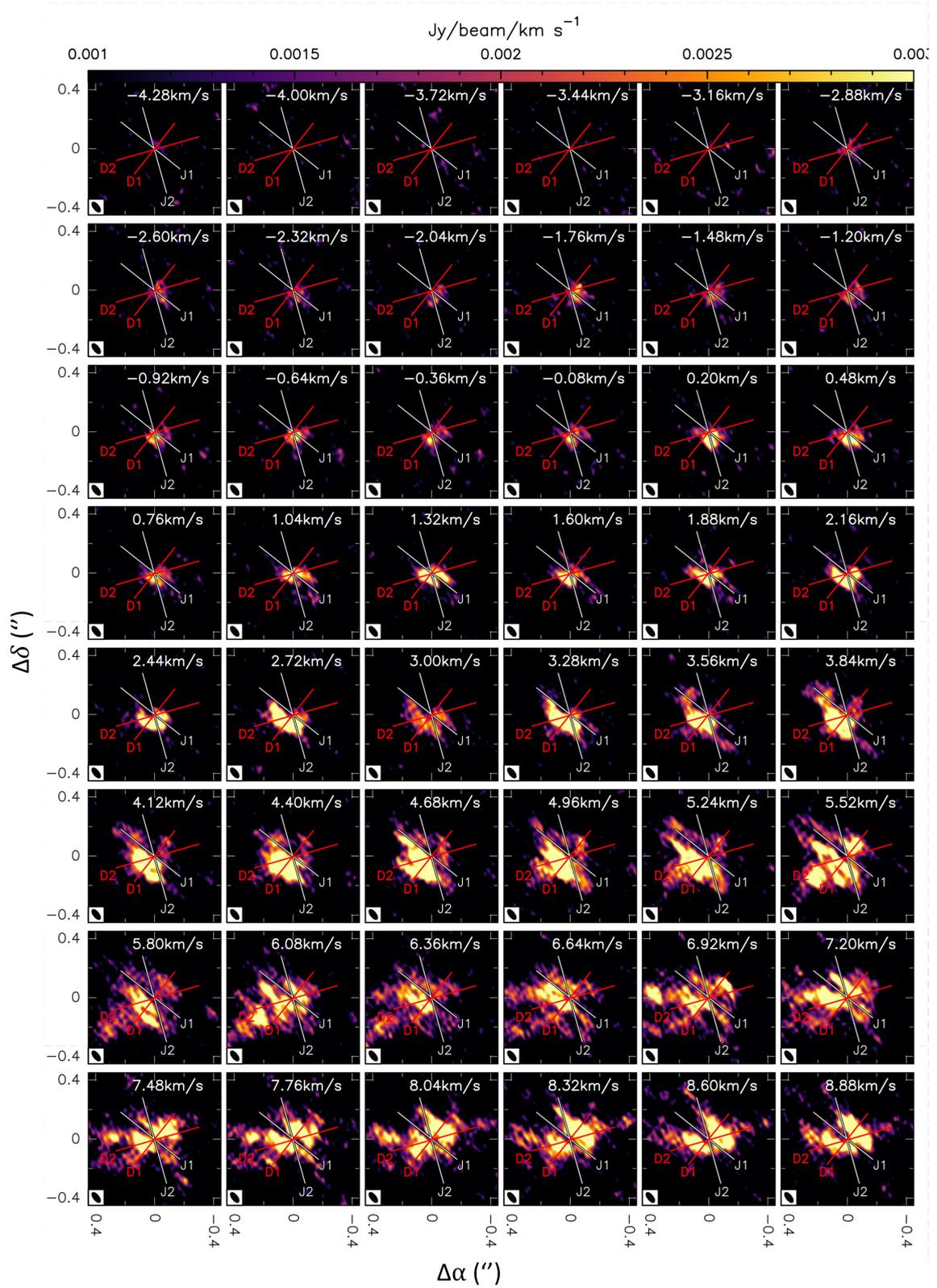


Figure A1. Channel maps of the HCN 3→2 line as observed with ALMA. The solid lines indicate the directions of *Jet 1*, *Jet 2*, *D1*, and *D2*.

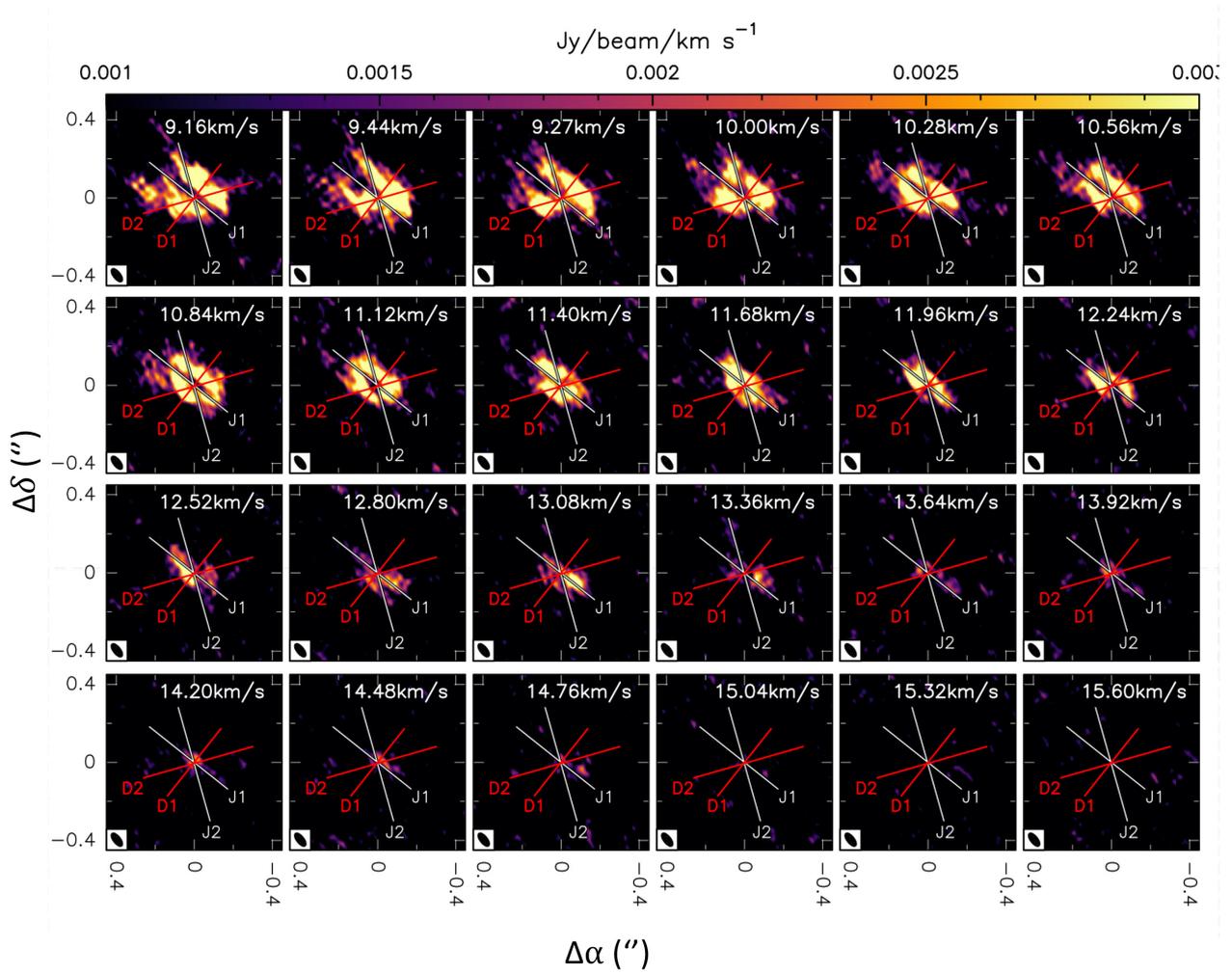


Figure A2. The same as Fig. A1.

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