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# PHANGS–*HST*: star cluster spectral energy distribution fitting with CIGALE

Jordan A. Turner <sup>1</sup>★, Daniel A. Dale,<sup>1</sup> Janice C. Lee,<sup>2</sup> Médéric Boquien,<sup>3</sup> Rupali Chandar,<sup>4</sup> Sinan Deger,<sup>2</sup> Kirsten L. Larson,<sup>2</sup> Angus Mok,<sup>4</sup> David A. Thilker <sup>5</sup>, Leonardo Ubeda,<sup>6</sup> Bradley C. Whitmore,<sup>6</sup> Francesco Belfiore <sup>7</sup>, Frank Bigiel,<sup>8</sup> Guillermo A. Blanc,<sup>9,10</sup> Eric Emsellem,<sup>11,12</sup> Kathryn Grasha <sup>13</sup>, Brent Groves,<sup>13</sup> Ralf S. Klessen <sup>14,15</sup>, Kathryn Kreckel <sup>16</sup>, J. M. Diederik Kruijssen <sup>16</sup>, Adam K. Leroy,<sup>17</sup> Erik Rosolowsky <sup>18</sup>, Patricia Sanchez-Blazquez,<sup>19,20</sup> Eva Schinnerer,<sup>21</sup> Andreas Schruba,<sup>22</sup> Schuyler D. Van Dyk <sup>2</sup> and Thomas G. Williams <sup>21</sup>

<sup>1</sup>Department of Physics & Astronomy, University of Wyoming, Laramie, WY 82071, USA

<sup>2</sup>Caltech/IPAC, Pasadena, CA 91125, USA

<sup>3</sup>Universidad de Antofagasta, Centro de Astronomía, Avenida Angamos 601, Antofagasta 1270300, Chile

<sup>4</sup>Department of Physics & Astronomy, University of Toledo, Toledo, OH 43606, USA

<sup>5</sup>Department of Physics & Astronomy, Bloomberg Center for Physics and Astronomy, Johns Hopkins University, Baltimore, MD 21218, USA

<sup>6</sup>Space Telescope Science Institute, Baltimore, MD 21218, USA

<sup>7</sup>INAF – Osservatorio Astrofisico di Arcetri, Largo E. Fermi 5, I-50157 Firenze, Italy

<sup>8</sup>Argelander-Institut für Astronomie, Universität Bonn, Auf dem Hügel 71, D-53121 Bonn, Germany

<sup>9</sup>The Observatories of the Carnegie Institution for Science, 813 Santa Barbara St., Pasadena, CA 91101, USA

<sup>10</sup>Departamento de Astronomía, Universidad de Chile, Casilla 36-D, Santiago, Chile

<sup>11</sup>European Southern Observatory, Karl-Schwarzschild Straße 2, D-85748 Garching bei München, Germany

<sup>12</sup>Univ. Lyon, Univ. Lyon1, ENS de Lyon, CNRS, Centre de Recherche Astrophysique de Lyon UMR5574, F-69230 Saint-Genis-Laval, France

<sup>13</sup>Research School of Astronomy and Astrophysics, Australian National University, Canberra ACT 2611, Australia

<sup>14</sup>Zentrum für Astronomie, Institut für Theoretische Astrophysik, Universität Heidelberg, Heidelberg 69120, Germany

<sup>15</sup>Interdisziplinäres Zentrum für Wissenschaftliches Rechnen, Universität Heidelberg, Heidelberg 69120, Germany

<sup>16</sup>Zentrum für Astronomie der Universität Heidelberg, Astronomisches Rechen-Institut, Heidelberg 69120, Germany

<sup>17</sup>Department of Astronomy, The Ohio State University, Columbus, OH 43210, USA

<sup>18</sup>4-183 CCIS, University of Alberta, Edmonton, Alberta, Canada

<sup>19</sup>Departamento de Física de la Tierra y Astrofísica, Universidad Complutense de Madrid, E-28040 Madrid, Spain

<sup>20</sup>Instituto de Física de Partículas y del Cosmos IPARCOS, Universidad Complutense de Madrid, E-28040 Madrid, Spain

<sup>21</sup>Max-Planck-Institut für Astronomie, Königstuhl 17, D-69117 Heidelberg, Germany

<sup>22</sup>Max-Planck-Institut für Extraterrestrische Physik, Giessenbachstraße 1, D-85748 Garching, Germany

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## ABSTRACT

The sensitivity and angular resolution of photometric surveys executed by the *Hubble Space Telescope* (*HST*) enable studies of individual star clusters in galaxies out to a few tens of megaparsecs. The fitting of spectral energy distributions (SEDs) of star clusters is essential for measuring their physical properties and studying their evolution. We report on the use of the publicly available Code Investigating GALaxy Emission (CIGALE) SED fitting package to derive ages, stellar masses, and reddenings for star clusters identified in the Physics at High Angular resolution in Nearby Galaxies–*HST* (PHANGS–*HST*) survey. Using samples of star clusters in the galaxy NGC 3351, we present results of benchmark analyses performed to validate the code and a comparison to SED fitting results from the Legacy Extragalactic Ultraviolet Survey. We consider procedures for the PHANGS–*HST* SED fitting pipeline, e.g. the choice of single stellar population models, the treatment of nebular emission and dust, and the use of fluxes versus magnitudes for the SED fitting. We report on the properties of clusters in NGC 3351 and find, on average, the clusters residing in the inner star-forming ring of NGC 3351 are young ( $< 10$  Myr) and massive ( $10^5 M_{\odot}$ ) while clusters in the stellar bulge are significantly older. Cluster mass function fits yield  $\beta$  values around  $-2$ , consistent with prior results with a tendency to be shallower at the youngest ages. Finally, we explore a Bayesian analysis with additional physically motivated priors for the distribution of ages and masses and analyse the resulting cluster distributions.

**Key words:** methods: data analysis – galaxies: individual: NGC 3351 – galaxies: star clusters: general.

\* E-mail: jturne19@uwo.edu

## 1 INTRODUCTION

The measurement of the physical properties of stellar populations using broad-band photometry is an arduous task. Even at the level of an individual single-aged star cluster, various phenomena are at play. As starlight interacts with dust, a fraction is absorbed and then re-emitted at longer wavelengths. Additionally, the dust obscures the view of the star clusters leading to extinction and reddening of the stellar emission. Young clusters contain massive stars that ionize the surrounding gas, leading both to the rise of a nebular continuum and to the appearance of a series of emission lines. This nebular emission can be quite bright, and may represent a non-negligible fraction of the flux captured by broad-band filters (e.g. Anders & Fritze-v. Alvensleben 2003; Groves et al. 2008; Boquien et al. 2010; Reines et al. 2010). In addition to the effects from dust and ionized gas, there are dependencies on the initial mass function (IMF) of the stellar population where the IMF can be fully sampled or stochastically sampled (e.g. Bararo & Bertelli 1977; Girardi & Bica 1993). To further our understanding of star cluster formation and evolution, it is necessary to carefully account for these effects to accurately measure fundamental cluster properties such as the age, mass, and reddening.

These measurements of stellar cluster properties provide an essential tool for understanding the mechanisms which drive, regulate, and extinguish star formation at small scales. In turn, catalogues of star clusters can be combined with observations of gas in nearby galaxies to chart the cycling of gas into stars, allowing us to study the dependence of star formation on environmental parameters on galactic scales. Previously developed catalogues of stellar cluster ages and masses have enabled the study of cluster mass and age functions (e.g. Foesneau et al. 2014; Ashworth et al. 2017; Chandar et al. 2017; Linden et al. 2017; Mok, Chandar & Fall 2019), star formation in different environments (e.g. Whitmore et al. 2014; Chandar et al. 2017; Leroy et al. 2018), and star formation efficiencies and time-scales (e.g. Grasha et al. 2018, 2019). Most recently, the Legacy ExtraGalactic Ultraviolet Survey (LEGUS) project (Calzetti et al. 2015) has derived the ages, masses, and reddenings of star clusters using the SED modelling techniques described in Adamo et al. (2017).

Over the past decades, various codes have been developed to model the panchromatic emission from galaxies. Combined with Bayesian techniques, they can provide deeper insight into the properties of star cluster populations in conjunction with simple  $\chi^2$  minimization for constraining physical parameters (e.g. da Cunha, Charlot & Elbaz 2008; Franzetti et al. 2008; Han & Han 2012; Moustakas et al. 2013; Chevillard & Charlot 2016). Although this machinery has enabled significant advances for the interpretation of galaxy spectral energy distributions (SEDs), application to stellar clusters has been more limited.

The availability of public codes under permissive licenses offers an opportunity to bring the benefits enjoyed by galaxy studies to studies of stellar clusters. Here, we report on an augmentation of the publicly available SED fitting package CIGALE: Code Investigating GALaxy Emission<sup>1</sup> (Burgarella, Buat & Iglesias-Páramo 2005; Noll et al. 2009; Boquien et al. 2019) to derive ages, stellar masses, and reddenings for star clusters identified in the Physics at High Angular resolution in Nearby Galaxies–*HST* (PHANGS–*HST*) survey (Lee et al. 2021).<sup>2</sup>PHANGS<sup>3</sup> is a panchromatic collaboration comprised of: PHANGS–ALMA, a large CO(2–1) mapping program aimed

at covering a representative sample of  $\sim 74$  nearby galaxies (Leroy et al., in preparation); PHANGS–MUSE, a Very Large Telescope (VLT) imaging program of  $\sim 20$  of the PHANGS galaxies with the MUSE optical integral field unit (IFU) instrument (Emsellem et al., in preparation; see first results in Kreckel et al. 2016, 2019); and PHANGS–*HST* which aims to chart the connections between molecular clouds and young star clusters/associations throughout a range of galactic environments by imaging the 38 galaxies from the PHANGS sample best suited<sup>4</sup> for study of resolved stellar populations. PHANGS–*HST* observations, which begun in 2019 April and are scheduled to conclude in mid-2021, are expected to yield *NUV–U–B–V–I* photometry for tens of thousands of stellar clusters and associations.

CIGALE is designed for speed, ease of use, and adaptability. It is based on a Bayesian approach for estimating physical properties and the corresponding uncertainties. It has the necessary flexibility for handling different stellar evolution tracks, star formation histories (SFHs), dust attenuation curves, and options for including nebular emission. A  $\chi^2$  minimization option is available with CIGALE to simplify comparisons with prior work. It is well suited for supporting stellar cluster studies, as well as for self-consistent modelling of both single-aged and composite stellar populations. Such an approach is needed for the characterization of structures across the full star formation hierarchy, beyond the densest peaks (i.e. stellar clusters), as will be investigated by PHANGS.

For this specific work, CIGALE has been expanded to handle the modelling of pure single-aged stellar populations, and the YGGDRASIL stellar populations models (Zackrisson et al. 2011) have been added to allow for easy comparison with previous studies, complementing the Bruzual & Charlot (2003, hereafter BC03) populations that are provided by default.<sup>5</sup> CIGALE fits photometry in linear flux units by default, but a modification to enable the fitting of photometry in units of magnitudes has also been implemented to allow for easy comparison with previous star cluster SED modelling studies. These changes are available in dedicated branches, *SSP* and *SSPmag*, respectively, of the public git repository<sup>6</sup> of CIGALE. When employing these modified versions, one still has the option to run CIGALE in its default operating mode of modelling the SFH rather than to fit the SSP track directly.

This paper is part of a series which documents the major components of the overall PHANGS–*HST* data products pipeline: survey design and implementation (Lee et al. 2021); source detection and selection of compact star cluster candidates (Thilker et al., in preparation); aperture correction and quantitative morphologies of star clusters (Değer et al., in preparation); star cluster candidate classification (Whitmore et al., in preparation); neural network classification proof-of-concept demonstration (Wei et al. 2020); stellar association identification and analysis (Larson et al., in preparation); and constraints on galaxy distances through analysis of the Tip of the Red Giant Branch as observed in the PHANGS–*HST* parallel pointings (Anand et al. 2020).

<sup>4</sup>Galaxies that are relatively face-on, avoid the Galactic plane, and have robust molecular cloud populations to facilitate joint analysis of resolved stellar populations and molecular clouds (Anand et al. 2020; Lee et al. 2021).

<sup>5</sup>There are a number of SFH and SSP modules available by default within CIGALE. The SFH modules included are a double exponential, a delayed SFH with optional exponential burst or constant burst/quench, a periodic SFH, and a user specified SFH read in from an input file. The SSP models included by default are Bruzual & Charlot (2003) and Maraston (2005).

<sup>6</sup><https://gitlab.lam.fr/cigale/cigale.git>

<sup>1</sup><https://cigale.lam.fr>

<sup>2</sup><https://phangs.stsci.edu>

<sup>3</sup><http://www.phangs.org>

Here, we focus on the methodology for fitting the  $NUV-U-B-V-I$  photometry for star clusters with CIGALE. In Section 2, we review the data utilized in this work. In Section 3, we present the results of a benchmark analysis to validate the code.<sup>7</sup> We use CIGALE to fit mock cluster photometry and examine the accuracy of the recovered properties. We also use CIGALE to fit photometry for stellar clusters in NGC 3351 published by the LEGUS project, and compare with results from their proprietary stellar cluster SED fitting code (Adamo et al. 2017). In Section 4, we use new imaging data obtained by PHANGS-*HST* for a larger region of NGC 3351 to help establish procedures for the PHANGS-*HST* SED fitting pipeline using CIGALE [e.g. single stellar population (SSP) models to be adopted, treatment of nebular emission and dust], and quantify model dependencies in the results. In Section 5, we take an initial look at the stellar cluster age, mass, and reddening results for NGC 3351 and explore spatial dependencies as well as the mass functions. Additionally, we explore the application of physically motivated Bayesian priors. Finally, in Section 6, we discuss possible future additions and tweaks to the SED modelling pipeline. We end by summarizing our findings in Section 7.

## 2 DATA

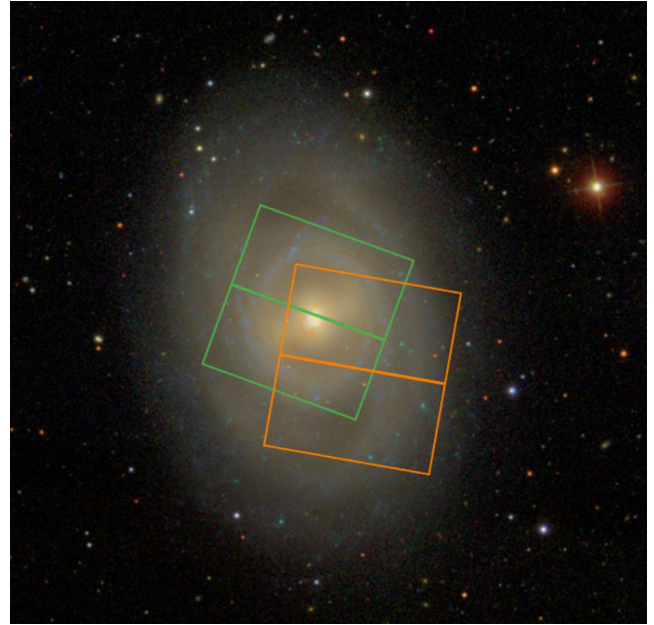
### 2.1 NGC 3351 *HST* imaging

To exercise our SED fitting procedures, we use photometry of stellar clusters measured from *Hubble Space Telescope* (*HST*)  $NUV-U-B-V-I$  ( $F275W-F336W-F438W-F555W-F814W$ ) imaging of the nearby galaxy NGC 3351. NGC 3351 is an Sb spiral at 10.0 Mpc (Freedman et al. 2001) with an approximately solar metallicity (Moustakas et al. 2010). Measurements using PHANGS-MUSE H II regions confirm this, and also show a flat radial metallicity gradient across our field of view (Kreckel, private communication), which simplifies the interpretation of results based upon the standard assumption of a single metallicity model for the entire cluster population of a galaxy.

Though 80 per cent of the PHANGS-*HST* galaxy sample had no existing *HST* wide-field imaging prior to the start of the program, NGC 3351 is one of the exceptions. It was observed by LEGUS in 2014 and was one of the first targets to be observed in the PHANGS program in 2019 May. While LEGUS observations were taken to maximize radial coverage of the galaxy including the nucleus, PHANGS-*HST* seeks to maximize the coverage of available PHANGS-ALMA CO(2-1) mapping, which leads to overlapping but complementary observations as shown in Fig. 1. Exposure times are given in Table 1. Star cluster photometry along with masses, ages, and reddening from SED fitting for sources detected in the area observed by LEGUS (Adamo et al. 2017) are publicly available through *MAST* from LEGUS.<sup>8</sup> Not only does this make NGC 3351 an excellent choice for a benchmark analysis, but the galaxy has been well studied due to its circumnuclear star formation (e.g. Buta 1988; Elmegreen et al. 1997), and it is also in the SINGS sample (Kennicutt et al. 2003). Hence, there is a wealth of ancillary data available, which provide independent constraints on parameters of interest here, such as dust reddening and metallicity, albeit at lower resolutions.

<sup>7</sup>See also Hunt et al. (2019) for a recent benchmark study involving CIGALE for the modelling of galaxy SEDs.

<sup>8</sup><https://legus.stsci.edu>



**Figure 1.** WFC3 observation footprints from PHANGS-*HST* (green) and LEGUS (orange) overlaid on a Sloan Digital Sky survey  $g-r-i$  image of NGC 3351 (David W. Hogg, Michael R. Blanton, and the Sloan Digital Sky Survey Collaboration). The WFC3 field of view is 162 arcsec  $\times$  162 arcsec.

**Table 1.** *HST* integrations for NGC 3351.

Name	$F275W$ (s)	$F336W$ (s)	$F438W$ (s)	$F555W$ (s)	$F814W$ (s)
PHANGS	2190	1110	1050	670	830
LEGUS	2361	1062	908	1062	908

*HST* WFC3 UVIS imaging exposure times corresponding to the footprints illustrated in Fig. 1. New data for NGC 3351 were obtained by PHANGS-*HST* in all five filters, in order to better cover the region mapped in CO with ALMA.

NGC 3351 is also an interesting subject for a first look at the variation of cluster properties in different environments using the PHANGS-*HST* data set, again due to its morphological and dynamical structure. This barred spiral galaxy exhibits a large range in optical surface brightness and levels of obscuration by dust between its inner and outer star-forming rings (the latter of which is now completely sampled with the addition of new PHANGS-*HST* imaging as shown in Fig. 1).

### 2.2 NGC 3351 star cluster photometry

For our SED fitting analysis, we use  $NUV-U-B-V-I$  photometry of star clusters which have been identified by the PHANGS-*HST* pipeline, visually inspected and classified. The PHANGS-*HST* methodology for source detection, candidate selection, and cluster classification builds upon the process developed by LEGUS. The LEGUS procedure is described in detail in Grasha et al. (2015) and Adamo et al. (2017), and we give a brief overview here. Source detection was performed with SEXTRACTOR on the  $V$ -band image. Clusters selected for analysis and whose properties were derived via SED fitting were those objects which (1) had a concentration index (CI) greater than 1.3 (i.e. a difference in magnitudes between circular apertures with radii of 3 and 1 pixels, which indicates that the object is more extended than a point source); (2) had  $M_V < -6$  Vega

**Table 2.** SED fit choices. (A) Parameter ranges adopted in the benchmark comparison between CIGALE and LEGUS SED fits. The extinction curve adopted for this comparison is from Cardelli et al. (1989) appropriate for the Milky Way. (B) Parameters adopted for PHANGS–*HST* SED modelling.

(A) LEGUS benchmark testing	
Star formation history	Instantaneous burst
Reddening and extinction	$E(B - V) = [0 : 1.5]$ mag; $\Delta = 0.01$ mag; $R_V = A_V/E(B - V) = 3.1$
SSP model	Yggdrasil Padova-AGB
Metallicity	$Z = 0.02$ (i.e. solar metallicity)
IMF	Kroupa; $[0.1 : 100]$ $M_\odot$ ; fully sampled
Gas covering fraction	0.5
(B) PHANGS– <i>HST</i>	
Star formation history	Instantaneous burst
Reddening and extinction	$E(B - V) = [0 : 1.5]$ mag; $\Delta = 0.01$ mag; $R_V = A_V/E(B - V) = 3.1$
Ages	$[1 : 10]$ Myr with $\Delta T = 1.0$ Myr; $[11 : 13\,750]$ Myr with $\Delta \log(T/\text{Myr}) = 0.3$
SSP model	Bruzual & Charlot (2003)
Metallicity	$Z = 0.02$
IMF	Chabrier (2003); $[0.1 : 100]$ $M_\odot$ ; fully sampled
Gas covering fraction	0.0

mag (after aperture correction) and were detected in at least four of five filters; (3) had been visually inspected and classified as either a class 1: symmetric compact cluster, class 2: asymmetric compact cluster, class 3: multi-peaked compact association.

In the PHANGS–*HST* and LEGUS footprints, we find a total of 468 clusters (with 133, 166, 166 visually inspected to be class 1, class 2, and class 3 objects using the same classification criteria as LEGUS). Of these, 136 are within the new area of NGC 3351 observed by PHANGS–*HST* (with 18, 56, and 61 visually inspected to be class 1, class 2, and class 3 objects). Details of the PHANGS–*HST* cluster identification methodology is presented in Thilker et al. (in preparation) and Whitmore et al. (in preparation).

A detailed comparative analysis of the PHANGS–*HST* and LEGUS clusters catalogues is presented in Thilker et al. (in preparation), and comparison with the stellar association catalogue is presented in Larson et al. (in preparation). Overall, the PHANGS–*HST* and LEGUS catalogues for the imaging observations originally obtained by LEGUS for NGC 3351 contain a comparable number of clusters to  $M_V < -6$  Vega magnitude (after aperture correction) with  $\sim 75$  per cent overlap.

Photometry is performed on each cluster with an aperture radius of 4 pixels which, at the distance of NGC 3351, corresponds to a physical scale of 7.7 pc. To account for extended emission beyond the 4 pixel radius, an aperture correction of 0.68 mag in the *V* band is applied (*NUV*: 0.87 mag, *U*: 0.80 mag, *B*: 0.71 mag, *I*: 0.80 mag) independent of cluster profile. Foreground extinction to NGC 3351 due to the Milky Way is computed following Schlafly & Finkbeiner (2011), which adopt Schlafly, Finkbeiner & Davis (1998) reddening maps and the Fitzpatrick (1999) reddening law with  $R_V = 3.1$ . The details of the aperture correction derivation is presented in Değ̃er et al. (in preparation). The median signal-to-noise ratio for a cluster in the *V* band is  $\sim 45$  (min:  $\sim 8$ , max:  $\sim 310$ ).

### 3 BENCHMARK TESTING CIGALE

#### 3.1 CIGALE: basic considerations

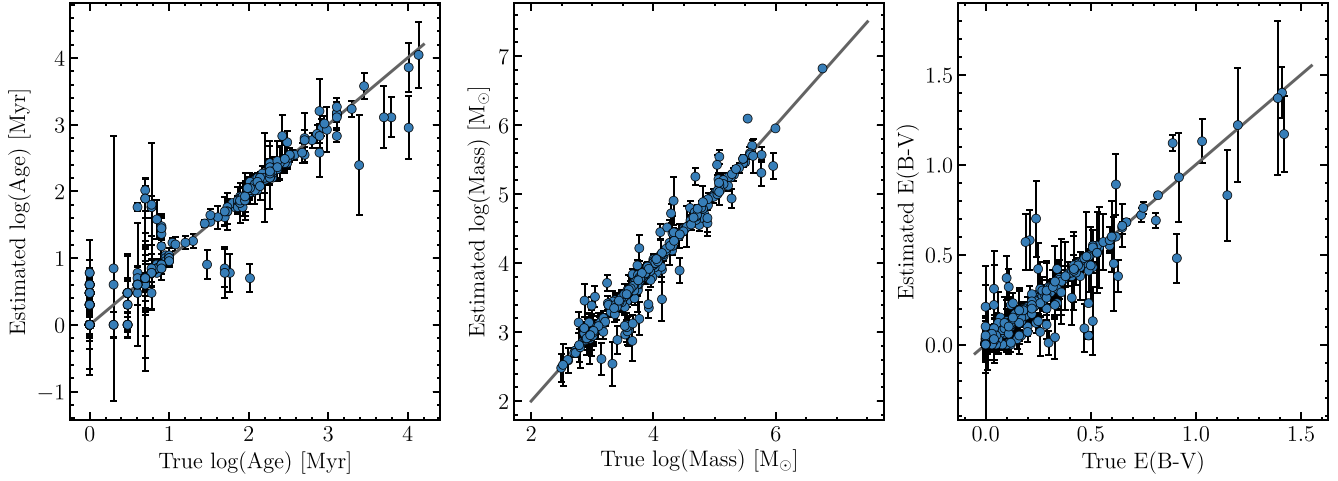
CIGALE operates by generating a grid of models based on the user’s input parameters. In our case, the model grid samples two free parameters: age and reddening based on our chosen single-aged stellar population (SSP) model, e.g. the BC03 models. The sampling

of the age and reddening grids is chosen by the user. For the age grid, we have ten linearly spaced models for 1–10 Myr ( $\Delta T = 1$  Myr, the highest precision available to CIGALE) and 100 evenly log-spaced models for 11–13 750 Myr [ $\Delta \log(T/\text{Myr}) \approx 0.3$ ]. Although cluster mass is an output of the SED modelling, it is not treated as a third dimension of the model grid. The masses corresponding to a particular model on the age–reddening grid are determined directly from the chosen IMF and SFH. The masses are normalized to  $1 M_\odot$  at birth and once CIGALE fits a cluster’s SED based on the age and reddening, the mass is appropriately scaled based on the cluster’s luminosity. A fully sampled IMF is assumed. The effect of a stochastically sampled IMF is discussed in Section 5.5. Model assumptions are listed in Table 2B.

CIGALE compares the cluster’s photometry with each model of the grid and calculates the  $\chi^2$  value to determine the goodness of fit. The  $\chi^2$  value is converted into a likelihood via  $\exp(-\chi^2/2)$ . Once each model has been tested, CIGALE estimates the best-fitting parameters in two ways: simple  $\chi^2$  minimization and a likelihood-weighted mean (see Boquien et al. 2019, Section 4.3). The model with the lowest  $\chi^2$  is the ‘best-fitting’ result and allows for easy comparison with other SED fitting procedures commonly used in the past (e.g. Chandar et al. 2010; Adamo et al. 2017).  $1\sigma$  uncertainties can be calculated by the difference between the best-fitting model and the models with  $\chi_{\text{reduced}}^2$  values of  $1 + \chi_{\text{reduced, min}}^2$ . The likelihood-weighted mean of all the models on the grid is computed which is used as a Bayesian estimate for the physical properties.  $1\sigma$  uncertainties are determined by the likelihood-weighted standard deviation of all the models.<sup>9</sup> Note that while the figures of this paper present the ages and masses in logarithmic units, CIGALE performs all analyses in linear units.

For the first part of our analysis we examine results based on the best-fitting ( $\chi^2$  minimized) values, which facilitates comparison to the body of previous star cluster work. Later in Section 5.6, we explore differences when additional Bayesian priors are imposed on the age and mass distributions and how to estimate the cluster properties from the posterior probability distribution functions (PDFs).

<sup>9</sup>In other words, we calculate the likelihood-weighted mean of the marginalized PDF with only a flat, bounded prior. The  $1\sigma$  uncertainties describe the width of the peak of the PDF.



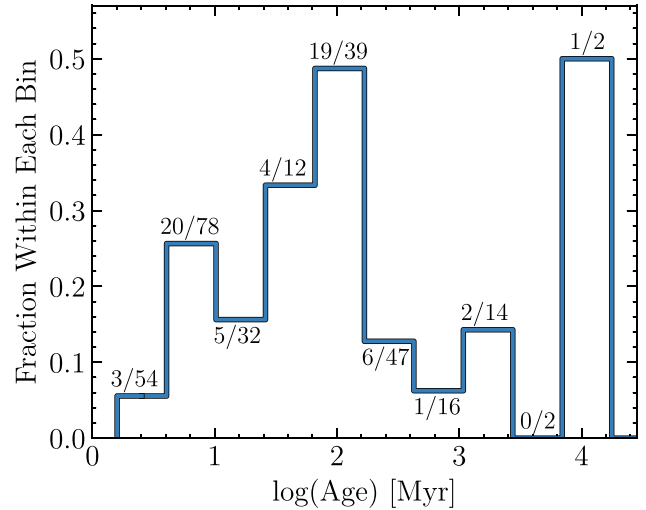
**Figure 2.** CIGALE best-fitting values compared to the ‘true’ input values.  $1\sigma$  error bars are given for each cluster’s estimated property and account for most of the scatter about unity. Cluster ages are well recovered across the sample except for clusters with ‘true’ ages at 1 Myr (0.25 dex scatter), at around 10 Myr (0.34 dex scatter), and at the oldest ages (0.39 dex scatter). Both the masses and reddenings are well recovered across the sample.

### 3.2 Recovery of mock clusters

First, to determine how well cluster properties can be constrained, we generate mock cluster photometry for known cluster ages, masses, and reddenings to estimate how well CIGALE recovers these cluster properties. As our starting point, we adopt the 296 model SEDs (i.e. model *NUV*, *U*, *B*, *V*, *I* fluxes, ages, masses, and reddenings) that best fit ( $\chi^2$  minimized) the PHANGS–*HST* visually classified class 1 and 2 star clusters in NGC 3351 which have photometric detections in all five bands. The mock cluster fluxes are then produced by randomly selecting a flux from a Gaussian distribution centred on the model flux with a standard deviation based on the median photometric uncertainty in each band of our PHANGS–*HST* cluster catalogue. The median uncertainties (for photometry within a 4 pixel radius aperture; see Thilker et al., in preparation) are 5.02 per cent for *F275W*, 5.12 per cent for *F336W*, 3.51 per cent for *F438W*, 2.24 per cent for *F555W*, and 3.21 per cent for *F814W*. We then run CIGALE on this mock catalogue and compare the resulting best-fitting values with the ‘true’ input values as shown in Fig. 2.

As Fig. 2 shows, cluster ages are well recovered across the sample within the uncertainties except at 1 Myr, at just below 10 Myr, and for the oldest clusters. From 10 to  $\sim 30$  Myr, the evolutionary tracks loop back on to themselves (see Fig. 5) which introduces degeneracies of the available SEDs which map to the SSP models at those ages. Across the entire sample, the standard deviation of the difference between the ‘true’ and recovered ages, masses, and reddenings are 0.31 dex, 0.18 dex, and 0.09 mag, respectively. There is no significant systematic offset between the true and recovered ages.

The clusters with large residuals are found to be mostly those with PDFs which exhibit bimodality where the young models with high reddening and older models with less reddening are both likely. We find  $\sim 20$  per cent of the estimated cluster sample to have bimodal PDFs. Fig. 3 shows the distribution of recovered ages for the bimodal cases as a fraction of the total number of mock clusters across ten age bins. The grouping of clusters with true ages around 100 Myr and underestimated age results have PDFs which are all bimodal. In this case, the  $\chi^2$  minimization chose the young mode when the old mode was the true age. A majority of the clusters with true ages at 10 Myr and overestimated age results are bimodal as well. In this case, the opposite is true; the  $\chi^2$  minimization chose the old mode when the young mode was the true age. The bimodality causes an increase



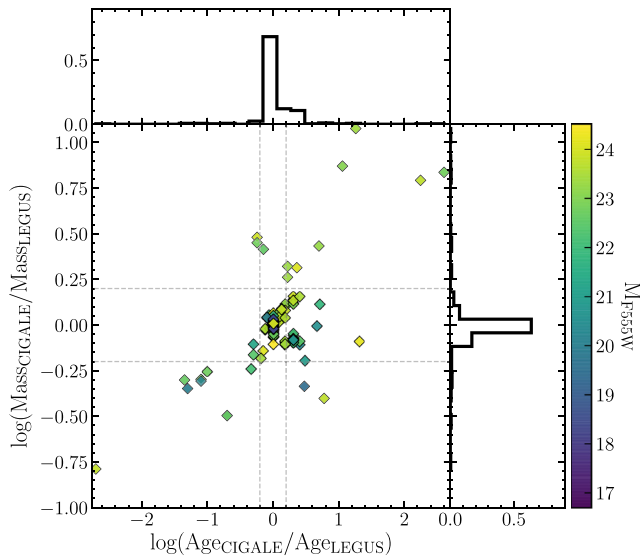
**Figure 3.** Histogram of the recovered ages of the bimodal cases as a fraction of the total number of mock clusters within each bin. The fraction of bimodal cases to the total number clusters is given at the top of each bin. The majority of the bimodal cases are concentrated at around 10 and 100 Myr.

in the scatter for the estimated  $E(B - V)$  versus the true  $E(B - V)$  values. Aside from these cases, for the most part, the reddenings are recovered well with a median difference from the ‘true’ reddening value of 0.0 mag with a dispersion of 0.09 mag. The cluster masses are also recovered well across the mock catalogue sample within the uncertainties. The median logarithmic mass ratio is 0.01 dex with a dispersion of 0.18 dex. Overall, we can expect CIGALE to behave consistently within our photometric uncertainties with the exception of the bimodal cases which are discussed in depth in later sections.

### 3.3 Comparison with LEGUS SED modelling

As a second method of benchmark testing, we use the published photometry<sup>10</sup> for the 292 clusters identified by LEGUS in their *HST*

<sup>10</sup>[https://archive.stsci.edu/prepds/legus/cluster\\_catalogs/ngc3351.html](https://archive.stsci.edu/prepds/legus/cluster_catalogs/ngc3351.html)



**Figure 4.** The residuals in CIGALE versus LEGUS stellar cluster ages and masses. Both approaches utilize the YGGDRASIL stellar tracks with a 50 per cent covering fraction (Table 2A) and both the CIGALE and LEGUS values are based on  $\chi^2$  minimization. Dotted lines mark  $\pm 0.2$  dex around perfect agreement. Data are coloured by  $V$ -band ( $F555W$ ) Vega magnitude. The histogram above the scatter plot shows the distribution of the age ratios and the histogram to the right shows the distribution of the mass ratios. The median age ratio is  $0.001 \pm 0.027$  dex. The median mass ratio is  $0.003 \pm 0.011$  dex. Clusters with poor agreement are generally found to be faint with large photometric uncertainties.

imaging of NGC 3351, fit the photometry with CIGALE, and compare to results from the SED modelling performed by LEGUS as described in Adamo et al. (2017). For the benchmark testing, we focus only on the 289 LEGUS clusters with photometric measurements in all five bands. We adopt the same assumptions used to produce the LEGUS ‘reference’ catalogues (Adamo et al. 2017); Table 2A summarizes the parameters adopted for the benchmark comparison. Briefly, Adamo et al. (2017) utilized the Padova–AGB stellar evolution isochrones and the YGGDRASIL population synthesis code (Zackrisson et al. 2011) to generate single-aged stellar population models. The LEGUS work assumes a Kroupa (2001) IMF from 0.1 to  $100 M_{\odot}$ ; the Cardelli, Clayton & Mathis (1989) Milky Way extinction law; flux from the nebular continuum and emission lines with a fixed covering fraction of 0.5; and solar metallicity isochrones. The minimum reddening is  $E(B - V) = 0$  mag and the maximum is set to be  $E(B - V) = 1.5$  mag with steps of 0.01 mag. The models are reddened before being fitted to the observed photometry.

Fig. 4 provides the residuals when comparing our CIGALE-derived stellar cluster ages and masses to those given by LEGUS. The agreement between the LEGUS and CIGALE SED fitting results is good. The median value of the age ratios, defined as  $\log(\text{Age}_{\text{CIGALE}}/\text{Age}_{\text{LEGUS}})$ , is  $0.001 \pm 0.017$  dex. The median value of the mass ratios, defined as  $\log(\text{Mass}_{\text{CIGALE}}/\text{Mass}_{\text{LEGUS}})$ , is  $0.003 \pm 0.011$  dex. Looking at the best-fitting SED models of the outliers, we do not find poor fitting models (i.e. large  $\chi^2$  values) or inaccurate/unreliable photometry, but we find that the outliers tend to be faint as seen in Fig. 4. In short, CIGALE is able to derive similar physical properties as LEGUS for the vast majority of clusters in the test case galaxy NGC 3351, establishing CIGALE as a reliable star cluster SED modelling code and consistent with previous star cluster SED fitting codes.

#### 4 DERIVATION OF PHYSICAL PROPERTIES FOR PHANGS–HST STAR CLUSTERS

In this section, we review the process for deriving the masses, ages, and reddening for PHANGS–HST star clusters with CIGALE. We use the aperture and foreground Milky Way extinction corrected photometry and propagate an additional 5 per cent uncertainty within CIGALE into the photometric errors to account for systematic errors in the flux calibration.

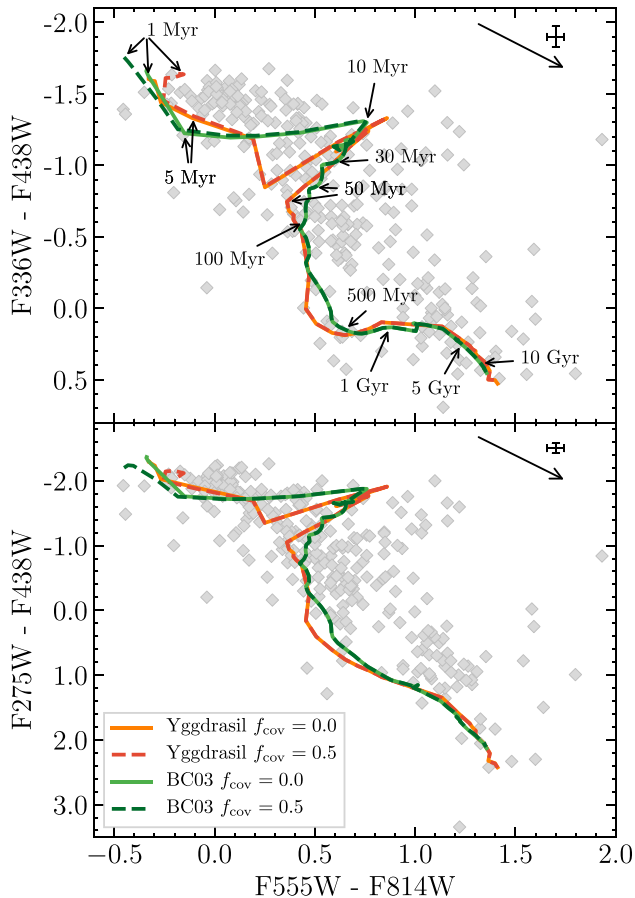
Throughout the rest of the paper, our focus is on comparisons with the class 1 and 2 clusters (i.e. potentially bound), rather than the class 3 compact associations. Cluster classes are defined in the same way as in LEGUS – class 1 clusters are symmetric and compact, class 2 clusters are asymmetric and compact, and class 3 clusters are multi-peaked compact associations. We note that, while we still include class 3 compact associations detected by our cluster pipeline in our catalogues, and provide visual classifications (Thilker et al., in preparation; Whitmore et al., in preparation) and SED fitting for those objects, we will use a more systematic hierarchical approach (watershed algorithm) for identification of younger star-forming associations, which is distinct from the cluster pipeline. This is detailed in Larsen et al. (in preparation) and will also include comparisons between age estimates using CIGALE for stellar associations identified with a watershed algorithm and the LEGUS and PHANGS–HST class 3 objects. The focus on class 1 and 2 clusters for this paper makes this cluster sample ‘exclusive’ (see Krumholz, McKee & Bland-Hawthorn 2019; Adamo et al. 2020, for discussions on exclusive versus inclusive cluster samples).

The set of SEDs available in the model grid is determined by the choice of input parameters. Table 2B summarizes our choices and the following sections discuss these choices. We study single-aged stellar populations, as the clusters do not have complex SFHs and, for the purposes of PHANGS science, can be effectively modelled by an ‘instantaneous burst’ of star formation. With the functionality available in CIGALE, we implement this SFH by utilizing the double-exponential SFH module `sfh2exp` and inputting a very short e-folding time for the stellar population (a microburst) and a zero mass fraction of the second/late starburst population. This effectively acts as an instantaneous burst of star formation. We tested this method against directly fitting to the SSP models themselves and verified that the results are consistent. We also adopt solar metallicities as the PHANGS–HST galaxies are all approximately solar (Section 2). By adopting a fully sampled IMF, Chabrier (2003) in our case, the SED modelling is ‘deterministic’.

As given in Section 3.1, the model grid of ages consists of 10 linearly spaced models from 1 to 10 Myr with  $\Delta T = 1$  Myr and 100 log-spaced models from 11 to 13 750 Myr with  $\Delta \log(T/\text{Myr}) \approx 0.3$ . We check if any differences arise when assuming a fully linearly sampled age grid and find no significant changes to the resulting cluster properties even for the bimodal PDF cases.

##### 4.1 Single-aged stellar population models and nebular emission

We plan to consistently use the same single-aged stellar population models as the baseline for PHANGS–HST stellar cluster analyses. In this section, we compare two common models: BC03 and YGGDRASIL (Zackrisson et al. 2011), the latter of which were adopted for the SED fitting of the LEGUS star clusters (Adamo et al. 2017). Tracks from both models are shown in Fig. 5 overplotted on the PHANGS–HST class 1 and 2 cluster photometry. Included in the figure are the model tracks with and without nebular emission, specified by a parameter describing the gas covering fraction. The gas covering



**Figure 5.** Top:  $F336W - F438W$  versus  $F555W - F814W$  ( $U - B$  versus  $V - I$ ) colour-colour diagram for a collection of theoretical single-aged stellar population models, along with the PHANGS-*HST* stellar clusters for NGC 3351. Magnitudes are in the Vega system. The stellar clusters have been aperture corrected as well as corrected for foreground Milky Way extinction. The BC03 (green) and YGGDRASIL (orange) tracks are shown both with (dashed) and without (solid) nebular emission. The y-axis has been flipped so the top left corner of the diagram is the ‘bluest’ and the bottom right corner is the ‘reddest’. The top panel annotates various ages along the SSP tracks; ages where the tracks deviate significantly from each other are pointed out with two arrows. The class 1 and 2 star clusters are plotted in grey. Both panels include a reddening vector for  $A_V = 1$  mag as well as a typical error bar for the cluster photometry computed from the median uncertainty in the colours. Bottom: Same as the top panel except with  $F275W$  (*NUV*) instead of  $F336W$  (*U*).

fraction, opposite of the escape fraction, determines the fraction of the Lyman continuum photons that ionize the surrounding gas. Nebular continuum and line emission can significantly contribute to the observed fluxes within broad-band filters, especially for spatially integrated observations of star-forming galaxies (e.g. Anders & Fritzev. Alvensleben 2003; Groves et al. 2008; Boquien et al. 2010; Reines et al. 2010; Pellegrini et al. 2020). PHANGS-*HST* observations, however, benefit from the combination of *HST*’s angular resolution and relative proximity of PHANGS galaxies, such that star clusters and surrounding H II regions can often be spatially disentangled, except for (1) crowded regions where it is still often unclear which star clusters are responsible for ionizing a given patch of H  $\alpha$  emission and (2) clusters with compact H  $\alpha$  morphologies (Hannon et al. 2019). The YGGDRASIL tracks with zero covering fraction (i.e. no nebular emission) are relatively consistent with the BC03 tracks

with no nebular emission; the main difference is in between 5 and 10 Myr where the YGGDRASIL tracks dip to much redder colours than the BC03 tracks. The addition of nebular emission is insignificant for most ages in both YGGDRASIL and BC03. As expected, only for the youngest ages ( $\sim 0-5$  Myr) is there a difference: a hook-like feature towards redder colours in the YGGDRASIL track with a covering fraction of a half while the BC03 track with nebular emission branches towards slightly bluer colours.

While both the YGGDRASIL and BC03 tracks show good agreement with the photometry in colour-colour space as is apparent in Fig. 5, we decide to adopt the BC03 tracks because they were originally designed to match galactic star clusters as well as reproduce the colours of star clusters in the Magellanic Clouds. The YGGDRASIL stellar population models were developed for constraining high-redshift galaxies (Zackrisson et al. 2011).<sup>11</sup>

We check how the inclusion of a nebular emission component into the SED fitting affects the results. In Fig. 5, we see the BC03 model tracks only differ slightly at very young ages when including the nebular emission. Fig. 6 shows the best-fitting SED models for three example PHANGS-*HST* clusters. The grey dashed line is the best-fitting SED model while including a nebular emission component. In this case, we assume an ionization parameter  $\log U = -2.0$ , line width of  $300 \text{ km s}^{-1}$ , and a covering fraction  $f_{\text{cov}} = 0.5$  (i.e. 50 per cent of the Lyman continuum photons ionize the surrounding gas). The first two clusters in Fig. 6 are young, blue clusters where the nebular emission could have the largest impact on the SED fitting. The first cluster’s SED fit with nebular emission gives a larger reduced  $\chi^2$  value (8.2) than compared to the non-nebular fit (2.44). For the second cluster, we find similar reduced  $\chi^2$  values between the nebular and non-nebular SED fits. As a contrast to these young clusters, the best-fitting SED for an old, red cluster is also given in the bottom panel of Fig. 6. The difference between SEDs with and without nebular emission is negligible. Therefore, for the PHANGS-*HST* sample, we use SED fitting without a nebular emission component. We note that Krumholz et al. (2015) perform a similar analysis and find the resulting cluster properties are robust to choice of evolutionary track and the inclusion of a nebular emission.

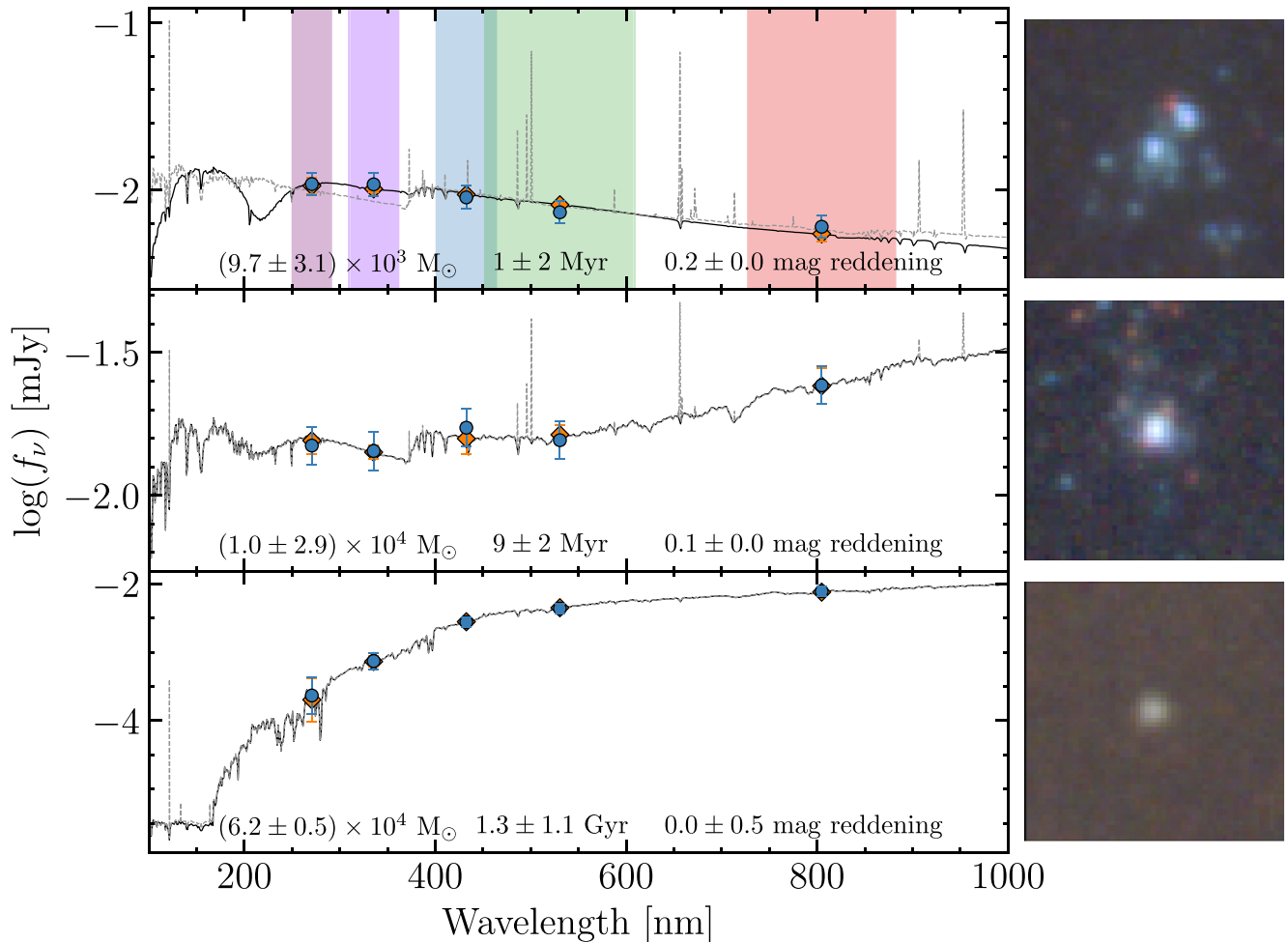
## 4.2 Dust

For the PHANGS-*HST* model grids, we make use of a Milky Way-like extinction curve from Cardelli et al. (1989) with  $R_V = 3.1$ . Krumholz et al. (2015) find that Milky Way-like extinction curves provide better results over other models for solar metallicity, face-on spiral galaxies when working with deterministic models. We allow the internal reddening to range from  $E(B - V) = 0$  to 1.5 mag in 0.01 mag steps – the same range adopted by LEGUS.<sup>12</sup> For reference, the drift scan spectroscopy for NGC 3351 from Moustakas et al. (2010) indicates  $E(B - V)$  for nuclear (aperture of  $\sim 6 \text{ arcsec}^2$ ), circumnuclear ( $400 \text{ arcsec}^2$ ), and large-scale radial strips ( $\sim \text{few arcmin}^2$ ) of  $0.03 \pm 0.05$ ,  $0.64 \pm 0.02$ , and  $0.55 \pm 0.07$  mag, respectively. Though the regions sampled by Moustakas et al. (2010) exhibit smaller values of reddening than  $E(B - V)_{\text{max}} = 1.5$  mag, we have chosen to keep the maximum value since sightlines to some individual clusters will inevitably exceed the large-scale spatial averages. As a sanity check, we perform a test SED fitting run

<sup>11</sup>We note that both the BC03 and YGGDRASIL models do not account for possible binary stars within the clusters.

<sup>12</sup>The cluster photometry is corrected for foreground Milky Way extinction before the SED fitting.





**Figure 6.** Example CIGALE fits (black line) for three PHANGS clusters (two young, one old) in NGC 3351. Blue circles mark the observed fluxes and orange diamonds indicate the fluxes extracted from the best-fitting model. Error bars denote  $3\sigma$  uncertainties on the observation and model fluxes. In some cases, the error bars are smaller than the data markers. Coloured strips in the top panel illustrate the filter bandpass rectangular widths. The black solid line is the best-fitting SED model using the parameters outlined in Table 2B; the grey dashed line is the same SED model while including the nebular emission component ( $f_{\text{cov}} = 0.5$ ). Best-fitting star cluster masses, ages, and reddenings are given for each cluster. A 1 arcsec  $\times$  1 arcsec false-colour ( $B - V - I$ ) image of the cluster is given to the right of each SED. The same image stretch is used for all three images to allow for easy visual comparison.

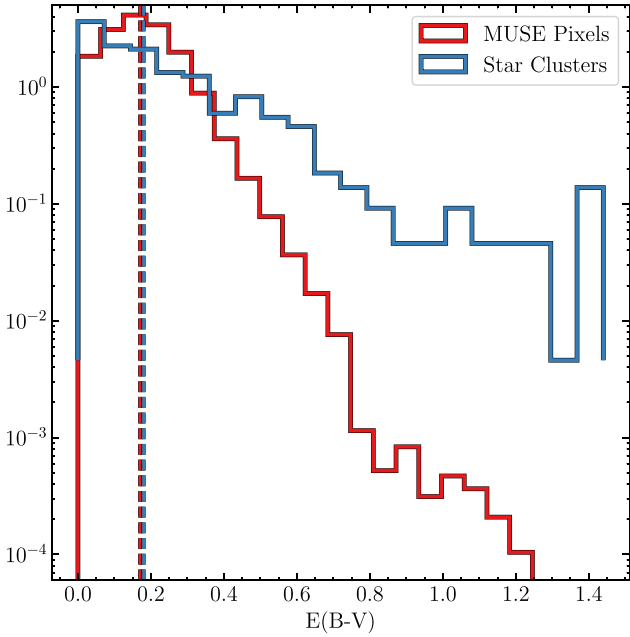
with a maximum reddening of 3.0 mag allowed and find only two clusters with a reddening greater than 1.5 mag (1.72 and 1.74 mag). Therefore, we decide to use a maximum value of  $E(B - V) = 1.5$  mag over smaller values to allow for individual sightlines with large reddening.

It could be possible to derive accurate maximum reddening values for the youngest ( $\lesssim 10$  Myr) stellar populations using the PHANGS–MUSE IFU data and Balmer decrement measurements. The IFU maps cover an area of the galaxies similar to our *HST* footprint but, the ground-based MUSE  $\sim 1$  arcsec angular resolution is too coarse to be directly applicable to our *HST*-resolution maps ( $\sim 0.08$  arcsec). Still, the  $A_V$  maps generated from Balmer decrement measurements from PHANGS–MUSE (Emsellem et al., in preparation) can be used as an independent check of our allowable reddening values.  $E(B - V)$  values for each pixel in the PHANGS–MUSE map are computed by assuming  $R_V = A_V/E(B - V) = 4.05$  (Calzetti et al. 2000). The distribution of the reddening values by pixel is given in Fig. 7 which shows a roughly power-law distribution with a median value of 0.171 mag and a maximum value of 1.24 mag. This distribution aligns fairly well with the chosen reddening parameters for the SED fitting.

A well-known problem in all star cluster SED fitting work is how to deal with dust extinction to break the age–reddening degeneracy. A star cluster can appear to be red due to age or due to reddening by dust. In our sample, we find  $\sim 70$  (20 per cent) of the clusters appear to be bimodal in the age and reddening PDFs. There are a number of possible ways to move towards breaking the degeneracy. In Section 5.6, we make use of Bayesian priors as an attempt to help limit the number of degenerate cases. See Section 6 for a discussion on additional methods we might employ in the future to break the age–reddening degeneracy in the PHANGS–*HST* sample.

### 4.3 Use of fluxes versus magnitudes in SED fitting

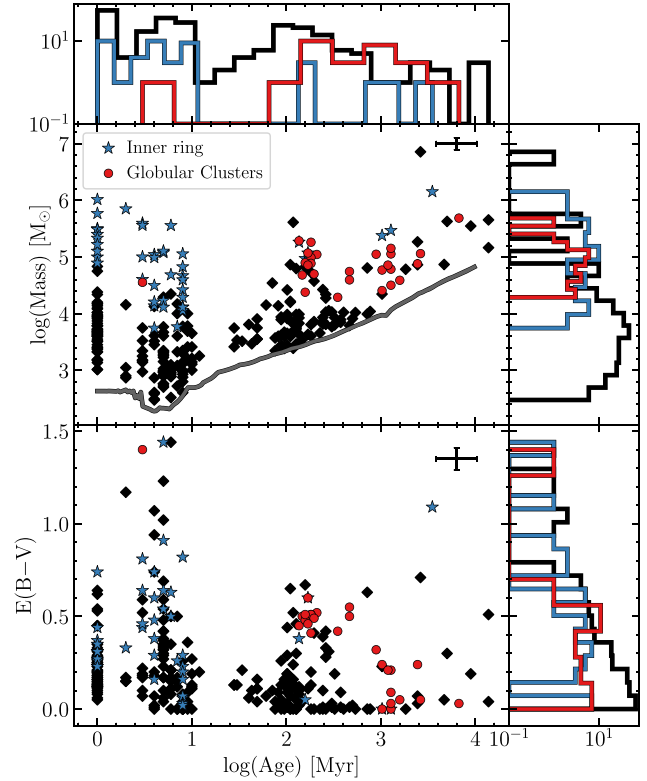
SED fitting can be carried out inputting either linear flux units (e.g. da Cunha et al. 2008) or logarithmic magnitude units (e.g. Chandar et al. 2010). We explore here if significant differences arise in key extracted output parameters (cluster age, mass, and reddening) when utilizing fluxes versus magnitudes as input to the SED fits. To perform the SED fitting in logarithmic magnitude units, we use the CIGALE branch *SSPmag* which expects the input photometry



**Figure 7.** Normalized logarithmic distributions of  $E(B - V)$  values for the pixels (0.2 arcsec in size) across the PHANGS–MUSE map of NGC 3351 as determined by Balmer decrement measurements (red) and the individual star cluster best-fitting reddenings as found by our SED fitting (blue). Only the reddening values of star clusters found where the Balmer decrement has been measured are used. Coloured dashed lines mark the median values for each distribution: 0.171 mag for the MUSE pixels and 0.18 mag for the star clusters.

to be of the form  $2.5 \log(f_\nu)$ . The uncertainties must be properly converted from linear fluxes. Comparing the best-fitting results from fluxes to those from magnitudes reveals they are quantitatively very similar. The median age residual, defined as  $\log(\text{Age}_{\text{flux}}/\text{Age}_{\text{mag}})$ , is  $0.000 \pm 0.026$  dex. The median mass residual is  $-0.001 \pm 0.011$  dex, and the median reddening residual is  $0.000 \pm 0.009$  mag. The main difference between the two results is that the magnitude-based fits have, on average, much larger uncertainties for the ages, masses, and reddenings.

Understanding how the uncertainties in our flux measurements are distributed – preferentially Gaussian in linear or logarithmic units – is essential for determining if fitting in flux or magnitudes is more accurate. However, we find the distribution of the uncertainties to be similarly non-Gaussian for both fluxes and magnitudes for all five filters; all distributions display similar positive skewness with long tails out to high uncertainties. We cannot come to a conclusion on using fluxes or magnitudes based on their uncertainty distributions. A more thorough study of the photometric uncertainties is presented in Thilker et al. (in preparation). An advantage of working in linear fluxes is in the case of non-detections (i.e. insignificant positive flux measurements or even slightly negative measurements). In linear fluxes, meaningful information can still be fed into the SED fitting using flux upper limits which is not as easily accomplished with magnitudes. Additionally, at low signal to noise, errors in magnitude units are not expected to be symmetric and the  $\chi^2$  minimization assumes errors to be symmetric and Gaussian (Hogg, Bovy & Lang 2010). Hence, these assumptions will not hold when using logarithmic units in the low signal-to-noise regime. Given these results, we choose to perform all the SED fitting in linear fluxes.



**Figure 8.** Best-fitting cluster masses and reddenings versus ages for the visually classified class 1 and 2 clusters of NGC 3351. Clusters of the inner star-forming ring are shown as blue stars and visually identified globular cluster candidates are shown as red circles. Logarithmic histograms beside each axis give the distributions of the cluster properties. Typical error bars, computed from the median of property uncertainties, are given in the corner of the two panels. The inner ring star clusters are found to be mostly young and massive with a large range in reddening. All but one of the globular cluster candidates are found to have ages  $>100$  Myr and to be massive following along the  $M_V = -6$  Vega mag observational limit (grey line).

#### 4.4 Summary of SED modelling assumptions

We summarize our SED modelling assumptions here and in Table 2B. The age model grid consists of ten linearly spaced models from 1 to 10 Myr with  $\Delta T = 1$  Myr and 100 log-spaced models from 11 to 13 750 Myr with  $\Delta \log(T/\text{Myr}) \approx 0.3$ . We assume a Milky Way-like extinction curve (Cardelli et al. 1989) with  $R_V = 3.1$ . The reddening  $E(B - V)$  model grid spans from 0 mag reddening to 1.5 mag with  $\Delta E(B - V) = 0.01$  mag. This gives a final age-reddening model grid with 16 610 models. Each model on the grid has a corresponding mass based on the BC03 SSP track and the fully sampled Chabrier (2003) IMF. We assume solar metallicity and a gas covering fraction of zero.

## 5 RESULTS AND DISCUSSION

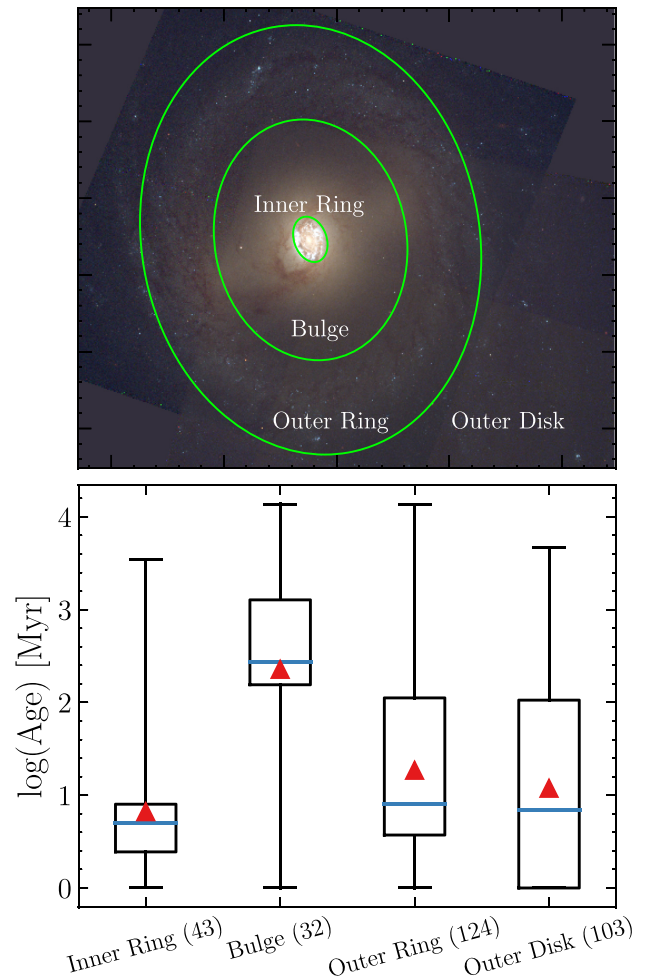
We provide a first presentation of the star cluster ages, masses, and reddenings based on the  $\chi^2$  minimized best-fitting results for NGC 3351 using PHANGS–*HST* photometry and the CIGALE SED fitting as described in Section 4. Fig. 8 shows the derived cluster masses and reddening values as a function of the cluster ages as well as the distributions of the cluster properties. We find a large population of young clusters at 1 Myr with reddening values ranging from 0.05 to 0.74 mag. The most highly reddened clusters are within

ages from  $\sim 3$  to 10 Myr. A noticeable dearth of clusters at ages from 10 and 30 Myr is found due to the inherent degeneracy of the SSP model at these ages where clusters of similar colours can be found to be either above or below 10 Myr with corresponding reddening. The  $\chi^2$  minimization seems to prefer the younger models with slightly higher extinctions. We also find a trend of the older clusters to be more massive, an expected selection effect, due to the absolute magnitude limit of the observations of around  $M_V = -6$  mag. As star clusters grow older, they will become less luminous due to the evolution of the stellar population.

Given the morphology of NGC 3351, we can break down the galaxy into distinct regions, allowing us to probe different star-forming environments within a single galaxy. Environmental masks have been developed by the PHANGS collaboration (Querejeta et al., in preparation) identifying discs and bulges using the Spitzer Survey of the Stellar Structure in Galaxies (S<sup>4</sup>G) pipeline as well as rings, bars, and lenses from Herrera-Endoqui et al. (2015). The box and whisker plot in Fig. 9 shows the breakdown of the star cluster ages in four distinct regions, moving radially outward from the centre of the galaxy: the inner star-forming ring with 43 clusters, the inter-ring or stellar bulge region with 32 clusters, the outer star-forming ring with 124 clusters, and the outer disc region beyond the outer ring with 103 clusters. Median best-fitting cluster  $\log(\text{Age}/\text{Myr})$  for the four regions are 0.69, 2.43, 0.90, and 0.84 and the means are 0.82, 2.36, 1.27, and 1.08. We find all four regions to have a similar dynamic range of ages but the bulge is home to significantly older star clusters on average. Over time, star clusters tend to migrate away from their natal gas clouds and disperse more uniformly with age over the galaxy (Gieles & Bastian 2008; Bastian et al. 2009; Davidge, Puzia & McConnell 2011; Kruijssen et al. 2011; Grasha et al. 2019). The bulge has few signatures of current star formation, hence any star clusters found in that area must be old enough to have had the time to migrate there or have been left in place by past star formation.

We show the spatial distribution of the clusters and their properties across the galaxy in Fig. 10. The majority of the clusters found in the outer ring and disc have masses less than  $10^{4.5} M_\odot$ . Higher mass clusters are generally found to be in the stellar bulge and inner star-forming ring. The majority of clusters in the inner ring are young as well. Kruijssen (2014) and Reina-Campos & Kruijssen (2017) predict higher mass clusters at small galactocentric radii due to higher ISM pressure which our result supports. These young, massive clusters can also be formed from the funnelling of gas into the centre by the galactic bar (Swartz et al. 2006). Approximately 70 per cent of the clusters have reddenings less than 0.3 mag ( $A_V < 0.93$  mag). There is most likely a population of faint, low-mass clusters in the inner ring which we are missing due to the bright background level. For the more heavily reddened and dust-obscured clusters, a high-resolution study with *James Webb Space Telescope* would provide a more robust examination of the dust. We note that the most massive cluster is found in the outer ring.

There are valid concerns regarding the ages of intermediate and older ( $>100$  Myr) clusters due to the age-reddening degeneracy. In Section 5.3, we will discuss the how the SED fitting handles clusters which are likely old globular clusters. Here, we check on the intermediate-aged (100 Myr to 1 Gyr) clusters based on their location in colour–colour space (Fig. 5) and location within the galaxy. A cut in colour–colour space is made to isolate the clusters redder than (i.e. below) the BC03 track at  $\sim 50$  Myr. This gives 135 clusters with 25 clusters ( $\sim 20$  per cent) found to have ages of 10 Myr and younger and corresponding high reddening values. Nine of these clusters are found within the dusty inner ring of the galaxy meaning they are

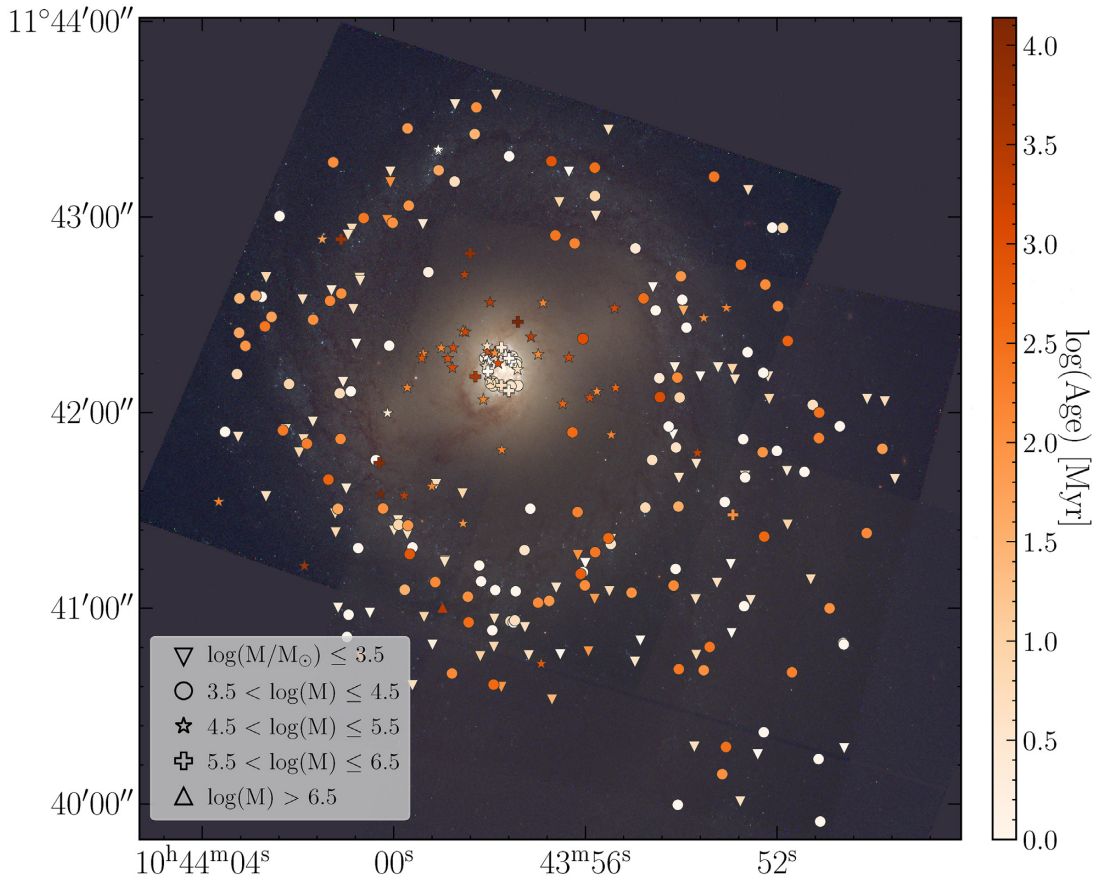


**Figure 9.** Top: Colour image of NGC 3351 with the four regions outlined and labelled. Bottom: Box and whisker plot showing the distribution of the star cluster ages within the four regions moving radially outward from the centre: the inner star-forming ring (43 clusters), the inter-ring or bulge region (32 clusters), the outer star-forming ring (124 clusters), and the outer disc beyond the outer ring (103 clusters). The blue lines in each box mark the median best-fitting cluster age and the red triangles mark the mean. Boxes extend from the lower to upper quartiles and the whiskers extend to the minimum and maximum ages found in each region.

likely highly reddened as the SED fitting results suggest. One cluster is found in the bulge and the remaining 15 are found in the outer ring and disc. In these regions, it is unlikely to have the high reddening as returned by the SED fitting. Overall, this represents only 5 per cent of the clusters (16 out of the total 302) which are likely intermediately aged (based on the location in the colour–colour diagram) but given incorrect young ( $<10$  Myr) ages. See Section 6 for a discussion on how we may resolve this issue.

## 5.1 Comparison with LEGUS

There are 199 visually classified class 1 and 2 PHANGS–*HST* clusters that are also found in the LEGUS cluster catalogue. On average, our best-fitting ages agree quite well with a median logarithmic age ratio of  $0.000 \pm 0.067$  dex. We find our cluster masses to be slightly larger compared to LEGUS with a median logarithmic mass ratio of  $0.035 \pm 0.035$  dex and our cluster reddenings following the same



**Figure 10.** Colour image of NGC 3351 with the location of the clusters overlotted. Clusters are colour-coded by best-fitting  $\log(\text{Age}/\text{Myr})$ . Data point symbols differentiate between five  $\log(\text{Mass}/M_{\odot})$  bins. The majority of clusters found in the outer ring and disc fall in the two lowest mass bins (downward triangles and circles). Higher mass clusters (stars and crosses) are generally found to be in the stellar bulge and the inner star-forming ring. The most massive cluster (upward triangle) is located in the outer ring.

trend with a median  $E(B - V)$  difference of  $0.030 \pm 0.020$  mag. There is no correlation of the residuals with ages, masses, or reddenings.

We can also examine agreement in the four different regions as discussed in the previous section. We find better agreement of the best-fitting cluster ages in the outer ring (median residual =  $0.050 \pm 0.125$  dex;  $N = 70$  clusters) and outer disc ( $0.0 \pm 0.077$  dex;  $N = 67$ ) than compared to the inner ring ( $-0.125 \pm 0.131$  dex;  $N = 33$ ) and bulge ( $0.313 \pm 0.217$  dex;  $N = 29$ ). For the bulge clusters, the best-fitting reddening is slightly lower compared to LEGUS with an median residual of  $-0.040 \pm 0.070$  mag, while the inner ring ( $0.070 \pm 0.039$  mag), the outer ring ( $0.005 \pm 0.037$  mag), and outer disc ( $0.060 \pm 0.021$  mag) show good agreement if not slightly higher. The masses are well matched with LEGUS across all regions of the galaxy.

We note that in a recent study of LEGUS star clusters in NGC 4449, Whitmore et al. (2020) show issues with LEGUS cluster ages and reddenings. In particular, LEGUS finds young ages for spectroscopically confirmed old globular clusters. This issue should be considered while comparing our cluster results with those from LEGUS.

## 5.2 Clusters with H II regions

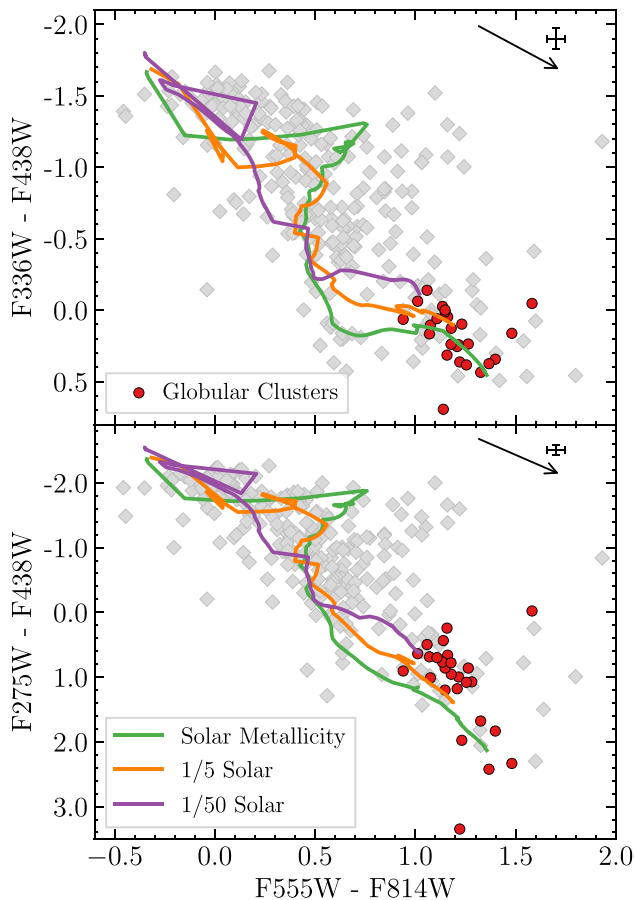
*HST*  $H\alpha$  imaging was obtained by  $H\alpha$ -LEGUS for the LEGUS pointing of NGC 3351 (PI: R. Chandar; Hannon et al., in preparation). We adopt the following classification scheme for the  $H\alpha$  emission

for each star cluster: (i)  $H\alpha$  emission directly on top of cluster, (ii)  $H\alpha$  emission in a ring around the cluster, (iii)  $H\alpha$  emission possibly associated with the cluster, or (iv) no  $H\alpha$  detected. We check our ability to accurately recover the ages of these young clusters by looking at the clusters which have strong  $H\alpha$  emission directly on top of the cluster or in a ring around the cluster. There are 60 class 1 and 2 clusters with such  $H\alpha$  emission which implies that they must be young ( $< 10$  Myr). The PHANGS-*HST* SED modelling recovers best-fitting ages of 10 Myr and younger for 93 per cent of these clusters which bodes well for the validity of our results for young clusters.

## 5.3 Globular clusters

We have visually identified a sub-sample of candidate globular clusters within our star cluster sample for NGC 3351 in order to check if the results provided by our SED fitting are as expected for globular clusters, i.e. old ( $\sim 10$  Gyr) with little reddening. This is done by selecting clearly resolved objects with similar red colours in the bulge of the galaxy (see Fig. 9). We note that nearly all the objects in this region are red in colour with almost no blue (i.e. young) objects in the vicinity.

Fig. 11 shows the location of the globular cluster candidates in the  $UBVI$  colour-colour diagram. The globular cluster candidates all group together in a tight region near the old end of the SSP track which demonstrates that they are indeed likely to be globular clusters.



**Figure 11.** Colour–Colour diagram similar to Fig. 5 highlighting the visually identified globular clusters candidates (red circles). The green track is the BC03 SSP model with solar metallicity used in the SED fitting. The orange track is the BC03 SSP model with 1/5 solar metallicity, and the purple track is 1/50 solar metallicity. Reddening vectors are given in each panel along with typical error bars on the cluster colours computed by the median uncertainties of the photometry. As expected, globular cluster candidates are found at the oldest end of the SSP track. However, the sub-solar metallicity tracks provide a better match to the globular cluster candidates’ colours. This is especially apparent in the bottom panel where the grouping of globular cluster candidates occupy the region at the very end of the 1/50 solar metallicity track.

We note that the SED fitting derived ages for most of the candidate globular clusters are younger than the expected  $\sim 10$  Gyr (Fig. 8). This is a well-known problem with most SED fitting approaches to age-dating cluster populations when the focus is on the young population. It is caused by two effects: (1) using high-metallicity model isochrones appropriate for young, but not old, populations and (2) the age-reddening degeneracy which allows the SED fitting to find a better fit with large reddening values even in cases where it is not physically reasonable (e.g. there is very little dust in the bulge region used to select the candidate globular clusters). These two effects can compound each other since the use of the wrong isochrones for low-metallicity objects results in a gap between the main grouping of points in the colour–colour diagram (most clearly seen in the bottom panel of Fig. 11) and where the solar-metallicity isochrone would predict a 10 Gyr cluster should be. This increases the number of cases where a better fit is found using a large reddening value (i.e. backtracking down the reddening vector) to an age of around 100 Myr on the isochrone (see Fig. 5). These issues are discussed in

more detail in Whitmore et al. (2020) where the same phenomena of understanding ages for old globular clusters is found for sample of spectroscopically age-dated globular clusters in NGC 4449.

Fig. 11 also shows sub-solar metallicity BC03 SSP tracks. We find roughly half of the globular cluster candidates benefit (i.e. return ages  $\sim 10$  Gyr) from the use of a lower metallicity model. The other half of the candidates must be suffering from the age-reddening degeneracy. We test this by enforcing the reddening  $E(B - V)$  to zero (following Whitmore et al. 2020) and the globular cluster candidates with ages originally around 100 Myr are shifted 1 Gyr and older. Given these results, we must be cautious using the SED derived ages of globular clusters due to both the metallicity and age-reddening degeneracy effects.

#### 5.4 Mass functions

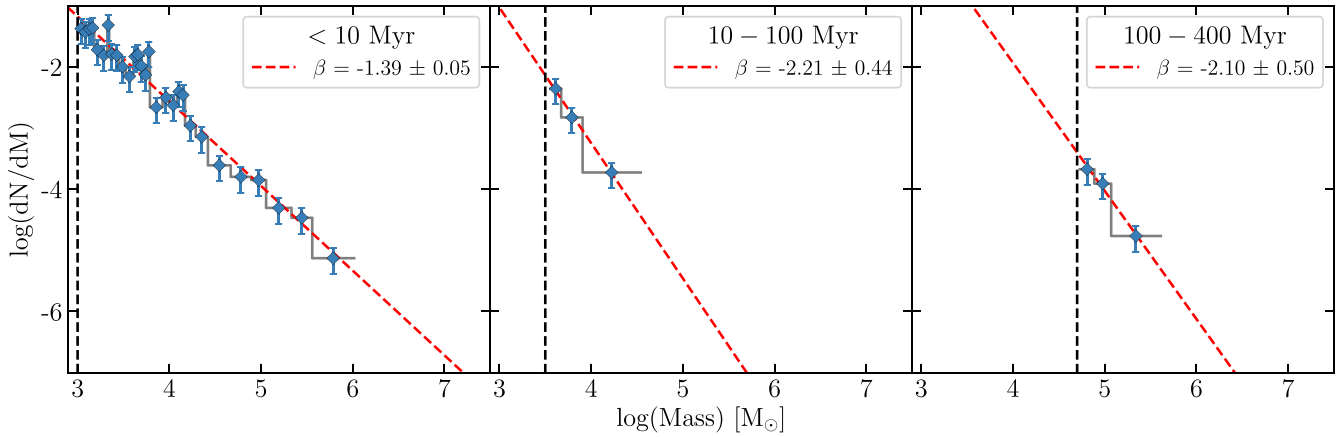
The shape of the star cluster mass function provides important information on the cluster formation and evolution. These distributions can be described to first order by a power law,  $dN/dM \propto M^\beta$ , where  $\beta \approx -2$  has been found for the young cluster populations in a number of galaxies (e.g. Lada & Lada 2003; Portegies Zwart, McMillan & Gieles 2010; Krumholz et al. 2019). There have also been claims for an exponential-like downturn or cut-off at the high end of the mass function (e.g. Larsen 2009; Adamo et al. 2015; Johnson et al. 2017; Messa et al. 2018; Mok et al. 2019; Adamo et al. 2020). This upper mass cut-off is often modelled by a Schechter function,  $dN/dM \propto M^\beta \exp(-M/M_*)$ , where  $M_*$  is the cut-off mass. Here, we are interested in addressing two questions: What is the power-law index  $\beta$  for young clusters in NGC 3351 and is there evidence for an upper mass cut-off?

We perform a least-squares fit of the form:  $\log dN/dM = \beta \log M + \text{const}$  to the distributions of the visually classified class 1 and 2 clusters of NGC 3351 which is shown in Fig. 12 as the red, dashed line and recorded in each panel. The results for  $\beta$  range from about  $-1.4$  to  $-2.2$ , consistent with the range found for cluster populations in other nearby galaxies (e.g. Fall & Chandar 2012), with the possible tendency to be somewhat shallower at the youngest ages. We also find a best-fitting value of  $\beta = -1.75 \pm 0.23$  for clusters in the age interval of 1–200 Myr, which was used by LEGUS (Adamo et al. 2017).

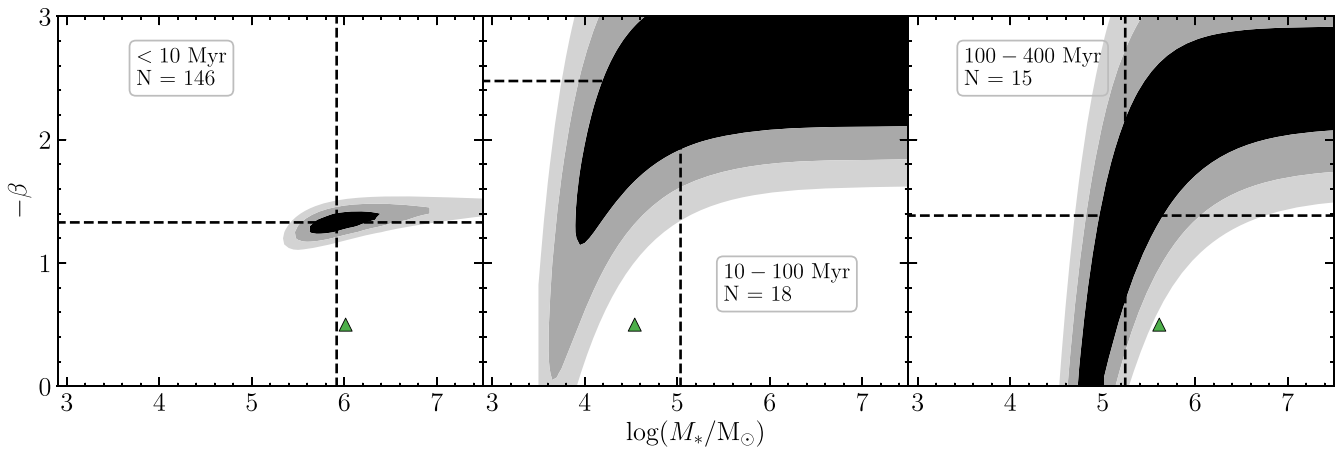
We also perform maximum likelihood fits of the Schechter function to determine if the cluster mass functions are better fit by a Schechter function than a power law, i.e. show statistically significant evidence for an upper mass cut-off. These fits follow the methodology described in Mok et al. (2019), and have the advantage of not using binned or cumulative distributions. In Fig. 13, we plot the resulting  $1\sigma$ ,  $2\sigma$ , and  $3\sigma$  contours for the  $M_*$  and  $\beta$  from our maximum likelihood fits for the same three age intervals as before. For the two older age bins, the contours, including that at  $1\sigma$ , remain open all the way up to the maximum tested mass at the right edge of the plots. This indicates that no Schechter-like cut-off mass is detected, and the best-fitting value of  $M_*$  is simply a lower limit. In the youngest age bin (ages between 1 and 10 Myr), the  $1\sigma$  and  $2\sigma$  contours are closed but the  $3\sigma$  contour remains open to the mass limit. In this case, there is weak evidence, at the  $\sim 2\sigma$  level, for a possible upper mass cut-off. However, since such a cut-off is not seen in the older cluster populations, we consider it unlikely that this is a physical feature present in the mass function of the NGC 3351 clusters.

#### 5.5 Stochasticity in sampling the IMF

We use a deterministic approach to SED fitting which is justified for stellar populations that completely sample the stellar IMF. However,



**Figure 12.** Cumulative stellar cluster mass functions for three age bins: less than 10 Myr (left column), 10–100 Myr (middle column), and 100–400 Myr (right column). The red dashed lines are the best-fitting power laws for each mass function with the power-law exponent  $\beta$  given in the top right of each panel. The black vertical dashed lines mark the completeness limits.



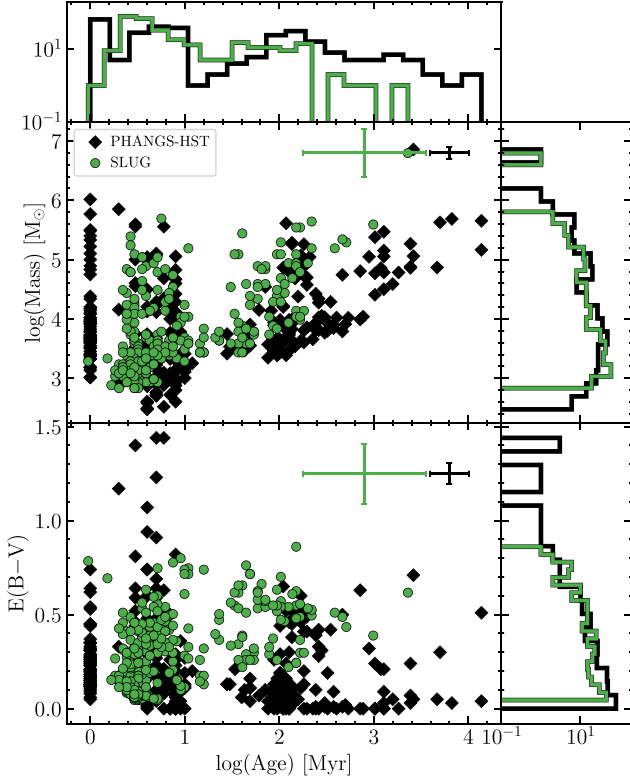
**Figure 13.** The  $1\sigma$ ,  $2\sigma$ , and  $3\sigma$  confidence contours from the maximum likelihood fits for the mass functions in Fig. 12. The green triangles denote the position of the most massive cluster in each respective sample.

there is a low-mass cluster regime ( $<10^4 M_\odot$ ) where the IMF may not be fully sampled. It becomes more likely that a low-mass stellar cluster will not contain high-mass stars and the IMF will appear to be truncated. The stellar population will also be fainter than a higher mass population which fully populates the IMF. This effect is called sampling stochasticity and it results in a broad range of possible luminosities and colours for these low-mass clusters (Barbaro & Bertelli 1977; Girardi & Bica 1993; Lançon & Mouhcine 2000; Bruzual 2002; Cerviño & Luridiana 2006; Deveikis et al. 2008; Foesneau et al. 2012; Hannon et al. 2019). In colour–colour space, clusters, which may be stochastically sampling the IMF, can be found in the region bluer in  $V-I$  than the SSP tracks at 1–5 Myr. These clusters have a deficiency of post-main-sequence stars versus the fully populated IMF, i.e. the clusters appear bluer than expected from a lack of red post-main-sequence stars (Foesneau et al. 2012). On the other hand, when an excess of post-main-sequence stars occurs, clusters are then found in the region redder than the SSP tracks. However, this is degenerate with reddening effects due to dust extinction, and can lead to higher SED fit reddening values (Hannon et al. 2019).

Recently, work has been done to explore the modelling of stellar populations which do not fully sample the IMF. One such model

is SLUG: Stochastically Lighting up Galaxies (da Silva, Fumagalli & Krumholz 2012; Krumholz et al. 2015). As a sanity check, we run CLUSTER\_SLUG<sup>13</sup> on our PHANGS–*HST* cluster catalogue while including the fiducial prior function  $p_{\text{prior}} = M^{-2.0}T^{-0.5}$  used by Krumholz et al. (2015). We use the SLUG star cluster model library modp020\_chabrier\_MW with a Milky Way extinction curve, Padova–AGB stellar evolution isochrones, solar metallicity, STARBURST99 v7 stellar atmospheres, and the Chabrier (2005) IMF. Fig. 14 shows the comparison between the PHANGS–*HST* best-fitting results and the SLUG results. We find SLUG tends to characterize clusters as slightly less massive, younger, and with larger reddenings than our PHANGS–*HST* results. SLUG does not return any old clusters with little reddening which is in direct contrast with our results where we find a majority of the clusters beyond 100 Myr have low reddening values at or close to 0 mag. For our sample, the largest changes in cluster ages (from older to younger when using SLUG) occur in clusters with masses of less than about  $10^4 M_\odot$  which makes up 60 per cent of our cluster sample. The stochastic sampling of the IMF will have the largest impact when the youngest, most massive stars

<sup>13</sup><http://www.slugsp.com/>



**Figure 14.** Comparison of the PHANGS–*HST* best-fitting results (black diamonds) to the SLUG results (green circles) for the class 1 and 2 clusters of NGC 3351. Logarithmic histograms on each axis show the distributions for both the PHANGS–*HST* and SLUG results. Typical error bars, computed from the median uncertainties on the properties, are shown in the corner of each panel. SLUG tends to find fewer old clusters and higher reddening values for the clusters older than 10 Myr. With SLUG, the majority of the clusters are found to be between 2 and 10 Myr and only a single cluster is found at 1 Myr. The mass distributions are similar, however, the PHANGS–*HST* results peak at slightly higher masses than SLUG.

are missing. Since SLUG accounts for this, it naturally finds more lower mass and younger star clusters.

## 5.6 Application of Bayesian priors

In the previous sections, we have focused on the  $\chi^2$  minimized SED fitting results for the clusters of NGC 3351. In the following analysis, we explore the application of astrophysically motivated Bayesian priors to infer the star cluster properties. CIGALE generates a  $\chi^2$  value [and subsequent likelihood value equal to  $\exp(-\chi^2/2)$ ] for each model on the model grid. The priors scale the likelihoods of each model thus reshaping the PDF. Krumholz et al. (2015) tested priors of the form

$$p_{\text{prior}} \propto M^\beta T^\gamma, \quad (1)$$

where  $M$  is the model mass and  $T$  is the model age. The priors are a reflection of the expected observed clusters’ mass and age distributions. Young star cluster observations have consistently found a mass distribution of  $dN/dM \propto M^{-2}$  giving a  $\beta$  value of  $-2$  (e.g. Williams & McKee 1997; Zhang & Fall 1999; Bik et al. 2003; de Grijs et al. 2003; Bastian et al. 2011; Fall & Chandar 2012; Foesneau et al. 2012; Adamo et al. 2017). We find a consistent  $\beta$  value for our sample of star clusters (Section 5.4). Krumholz et al. (2015) tested three values for  $\gamma$ :  $\gamma = 0$  corresponding to a distribution that is flat in

linear age;  $\gamma = -1.0$  corresponding to a distribution flat in  $\log(\text{age})$  i.e. 90 per cent decline each decade of linear time (e.g. Whitmore, Chandar & Fall 2007); and  $\gamma = -0.5$  fiducial prior acting as a compromise and is most similar to observations for spiral galaxies. In this section, we test these values of  $\beta$  and  $\gamma$  on our sample.

We first show how the priors are implemented in the analysis and how the cluster properties are inferred. The priors scale the likelihoods which gives the posterior probability distribution

$$p_{\text{post}}(M, T, A_V) \propto M^\beta T^\gamma \exp(-\chi^2/2), \quad (2)$$

where  $A_V = 3.1/E(B - V)$ . To derive a Bayesian estimate for a cluster’s age, mass, and extinction, we calculate the expectation value of the posterior probability distribution. The expectation value of, for example, a cluster’s age is defined as

$$\langle T \rangle = \frac{1}{C} \int dM \int dT \int dA_V T p_{\text{post}}(M, T, A_V), \quad (3)$$

where

$$C = \int dM \int dT \int dA_V p_{\text{post}}(M, T, A_V). \quad (4)$$

However, given the discrete nature of the model grid in  $M$ ,  $T$ ,  $A_V$  space, a quadrature approximation is calculated to give the expectation value

$$\begin{aligned} \langle T \rangle &= \frac{1}{C} \sum_i (M_{i+1} - M_i) \sum_j (T_{j+1} - T_j) \sum_k (A_{V_{k+1}} - A_{V_k}) \\ &\quad \times T_j p_{\text{post}}(M_i, T_j, A_{V_k}), \end{aligned} \quad (5)$$

where

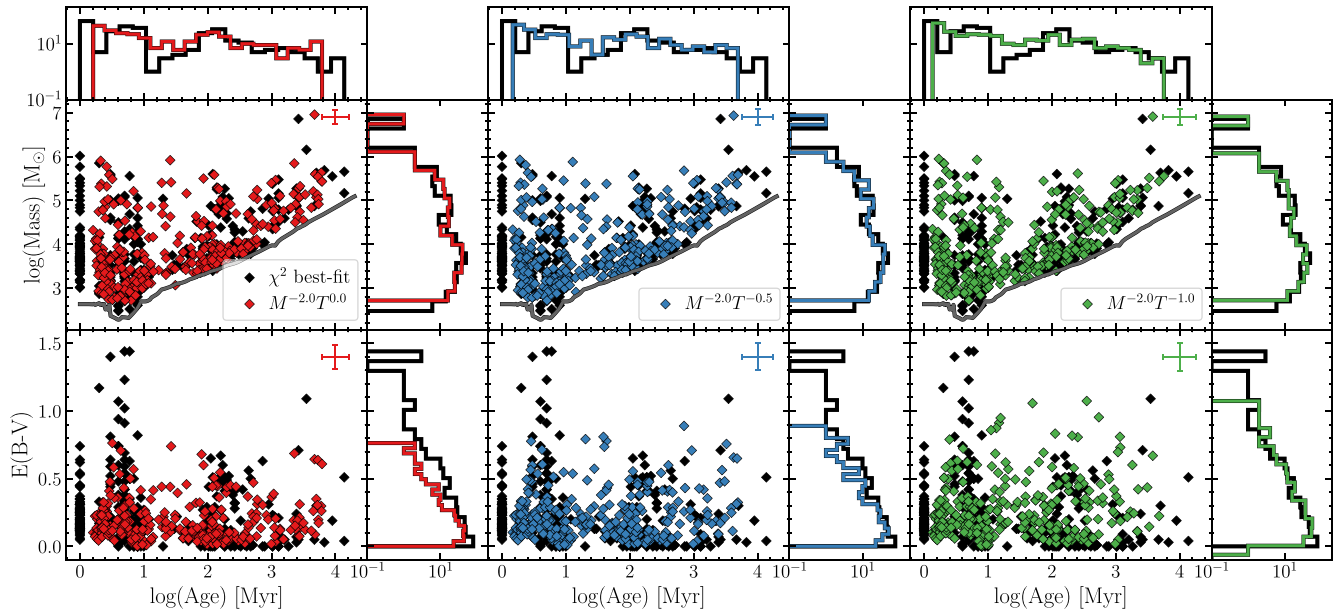
$$\begin{aligned} C &= \sum_i (M_{i+1} - M_i) \sum_j (T_{j+1} - T_j) \sum_k (A_{V_{k+1}} - A_{V_k}) \\ &\quad \times p_{\text{post}}(M_i, T_j, A_{V_k}). \end{aligned} \quad (6)$$

We then calculate the dispersion as the square root of the second moment of the posterior probability distribution:

$$\sigma_T^2 = \sum_j (T_j - \langle T \rangle)^2 p_{\text{post}}. \quad (7)$$

However, this dispersion can be a poor estimation of the error on the Bayesian parameter estimate in cases with long-tailed or bimodal posterior PDFs. For such cases, the 16th to 84th percentile range is a more robust estimator of the error. For this paper, we will continue to use the dispersion as the uncertainty estimator but the 16th and 84th percentiles, along with the dispersion, will be included in the PHANGS–*HST* cluster catalogues. All calculations are performed in linear  $M$  and  $T$  space. The expectation value (equation 5) with corresponding uncertainty (equation 7) acts as the Bayesian estimates for the cluster ages, masses, reddening.

A comparison of the Bayesian parameter estimates is given in Fig. 15. For all three priors tested,  $\beta$  is set to  $-2$ . The steeper distribution with  $\gamma = -1.0$  (green diamonds) shows fewer older clusters than the fiducial results (blue diamonds), which is to be expected. With  $\gamma = 0.0$ , clusters are found to extend to older ages which is, again, to be expected. Without sufficient physically motivated criteria to evaluate which output age and reddening distribution depending on the chosen  $\gamma$  value is more likely than another, it is difficult to definitively decide on the ‘best’ prior. Instead, we quantify the differences in the estimated physical properties compared to the fiducial prior. With  $\gamma = 0.0$ , we find the median difference in cluster ages and masses to be  $-0.048$  and  $-0.005$  dex, respectively. The median difference in reddening is 0.014 mag. A



**Figure 15.** Bayesian estimated mass and reddening versus age for three different prior functions (blue, red, and green diamonds) compared to the  $\chi^2$  best-fitting results (black diamonds) for PHANGS–*HST* photometry of star clusters in NGC 3351. The grey line in the upper panels indicates the magnitude limit  $M_V = -6$  Vega mag. Logarithmic histograms beside each plot show the distribution of the ages, masses, and extinctions in each case. We adopt  $p_{\text{prior}} = M^{-2.0}T^{-0.5}$  (middle panels) as the fiducial prior. Typical error bars are computed as the median of the uncertainties and are given in corner of each panel. For the fiducial prior, the median uncertainties on the Bayesian age, mass, and reddening estimates are 0.24 dex, 0.18 dex, and 0.10 mag, respectively.

negative difference in ages and a positive difference in reddening confirms that the  $\gamma = 0.0$  prior returns older cluster ages with less reddening. With  $\gamma = -1.0$ , the median difference in cluster ages and masses is 0.057 and 0.003 dex, respectively. The median difference in reddening is  $-0.017$  mag. A positive difference in ages and a negative difference in reddening confirms that the  $\gamma = -1.0$  prior returns younger clusters with more reddening. In both cases, the cluster properties are fairly robust to the choice of prior. Given these results, we adopt the fiducial prior  $p_{\text{prior}} = M^{-2.0}T^{-0.5}$  using the method outlined above in order to compute the Bayesian inferred cluster property estimates. The quantified differences stated here can be treated as systematic uncertainties within the Bayesian estimates.

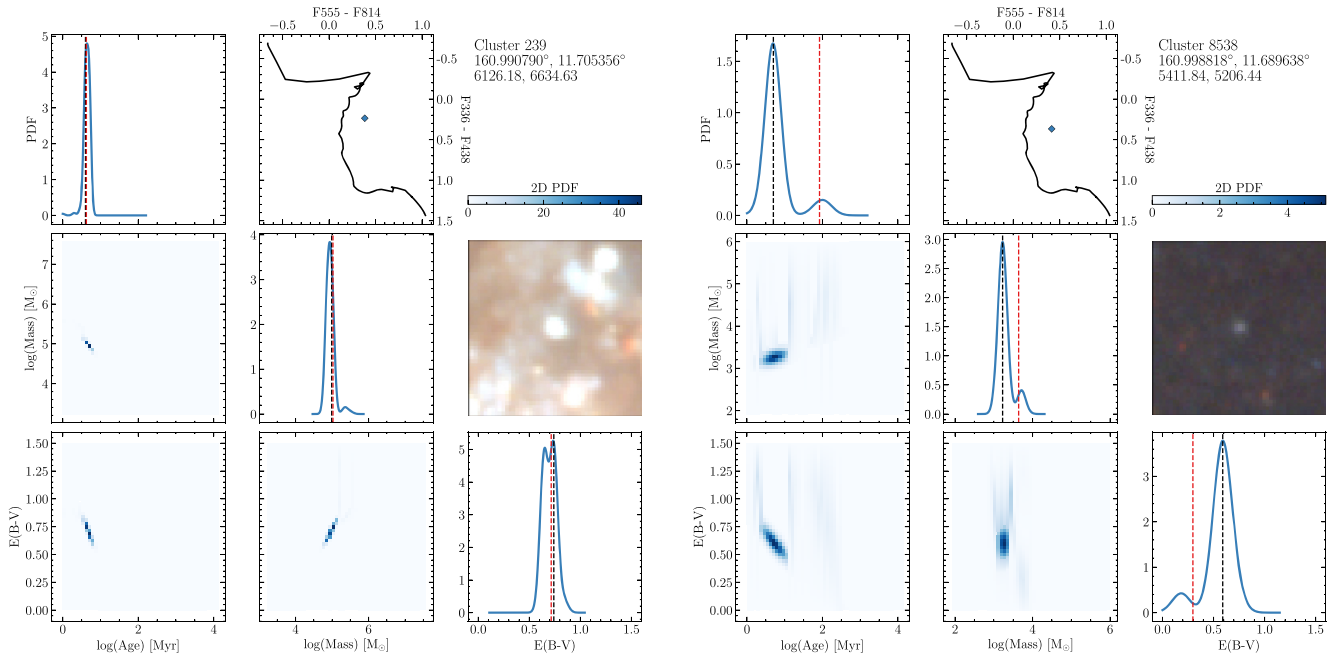
Fig. 16 shows example corner plots produced during the Bayesian analysis with the fiducial prior applied. For the corner plots, the posterior PDFs are computed by Gaussian kernel density estimation (which is only used for a visual representation of the PDFs and has no bearing on the computation of the expectation values). In the first case, cluster 239, the PDFs show singly peaked distributions indicating a well-constrained Bayesian estimate. This case also demonstrates good agreement between the minimized  $\chi^2$  results (black dashed lines) and the Bayesian estimates (red dashed lines). The second case, cluster 8538, demonstrates the age-reddening degeneracy where two cases are likely: a young, reddened cluster or an old cluster with little reddening. This is a case where Bayesian estimate disagrees with the  $\chi^2$  result. We inspect each cluster’s age and reddening PDFs and find  $\sim 70$  clusters with bimodality before the application of the Bayesian prior. After the application of the Bayesian prior, we find  $\sim 30$  bimodal cases ( $\sim 10$  per cent of the entire cluster sample) when inspecting the newly modified PDFs. The Bayesian analysis proves to be advantageous to help break degeneracies but not for all of the cases (as with cluster 8538). More information is needed to further reduce the number of bimodal cases. One option is to visually inspect each cluster in a 3-colour image to

look for obvious signs (to humans) of dust extinction. However, this is not feasible for the automated PHANGS–*HST* pipeline of thousands of star clusters without a well-taught and tested machine learning algorithm. In Section 6, we discuss other possible solutions for breaking the age-reddening degeneracy.

In a similar vein, in this paper, we are able to identify bimodal cases simply by inspecting each cluster’s corner plot. However, for the full, automated PHANGS–*HST* pipeline, this will not be possible. The bimodal cases generally find disagreement between the  $\chi^2$  minimized result and the Bayesian estimate so clusters with large disagreements could be flagged as tentatively bimodal within the pipeline. For these cases, instead of calculating the expectation value, the PDFs could be split between the two modes and new Bayesian estimates could be computed for each of the modes. Before this can be implemented, we will need to decide how to consistently split apart the two peaks in the PDFs and how the Bayesian estimates from the two modes are handled with regards to the final cluster properties used in the cluster catalogues. Other methods could be used to extract a Bayesian estimate from the PDF such as taking the median with 16th and 84th percentiles or finding the model with the maximum likelihood (i.e. the peak of the posterior PDF).

Fig. 17 compares the cluster properties as derived from the expectation value and maximum likelihood methods. The expectation value results reveal a few interesting, and possibly worrisome, trends. The method fails to find very young clusters (youngest cluster is 1.5 Myr) due to 1 Myr being the boundary. All the available models are older than 1 Myr which skews the expectation value (a weighted average) to be older than 1 Myr. We find a slight trend for older clusters to have larger reddenings which is counter-intuitive. This is in contrast to the negative correlation between age and extinction found by Bastian et al. (2005) for the star clusters of M51. However, Grasha et al. (2018) find a poor correlation between star cluster age and extinction in NGC 7793 based on SED fitting from the





**Figure 16.** Corner plots of two example star clusters showing the age, mass, and reddening 1D and 2D kernel density estimated posterior PDFs after the application of the Bayesian fiducial prior. The black dashed lines are the minimized  $\chi^2$  best-fitting result and the red dashed lines are the Bayesian estimates. The top middle panel of each corner plot is a colour–colour diagram marking the location of the cluster and the BC03 model track. A postage stamp 3-colour image of the cluster is given in the middle-right panel. Cluster 239 on the left is an example of a single-peaked age PDF with good agreement between the  $\chi^2$  result and the Bayesian estimate. Cluster 8538 on the right shows a bimodal case where the Bayesian estimate and the  $\chi^2$  result disagree.

LEGUS program. Further, Grasha et al. (2018) find that star clusters not associated with any GMCs have a slightly higher extinction than those still semi-embedded in their natal gas clouds, which may be unexpected. The same result is found for the star clusters of M51 (Grasha et al. 2019). This serves as a testament to the complexities of star cluster SED modelling and the degeneracy of age and reddening. We see a maximum reddening of around 0.9 mag with a low dispersion in reddening for the youngest clusters and a larger dispersion in reddening for clusters from 10 Myr to 100 Myr. The majority of the clusters in this age range appear to be the ones identified as bimodal which makes sense given the high probability of age–reddening degeneracy at this location of the SSP model track (Fig. 5).

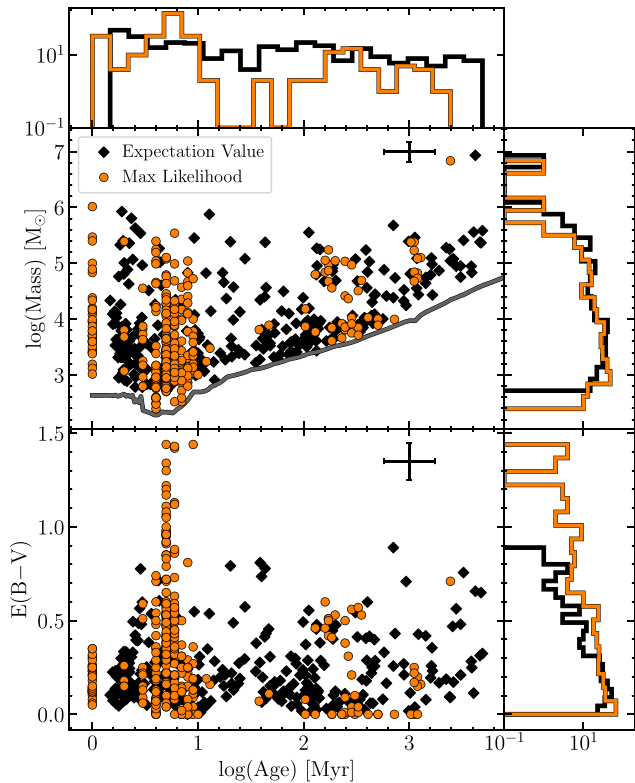
The maximum likelihood results show a few contrasting trends. We see a quantization of cluster ages, especially at 10 Myr and younger. This is to be expected as the models available for the method to choose from are those following the model grid. We do find clusters at 1 Myr using this method. We find a significant portion of clusters at ages from 5 to 9 Myr with a virtual lack of clusters between 10 and 100 Myr. The clusters found to be around 20 Myr with the expectation value method appear to be shifting all to ages younger than 10 Myr with the maximum likelihood method. We also find roughly half of the clusters older than 100 Myr to have very low and even zero reddenings. We check if these are clusters identified as globular cluster candidates and find only one of them to be a globular cluster candidate. The remaining globular cluster candidates are found in the grouping at around 1 Gyr with 0.2 mag reddening, a few at younger ages around 200 Myr with 0.5 mag reddening, and, most troubling, 30 per cent of the globular clusters are found at 5 Myr old. In contrast, the expectation value method finds the youngest globular cluster candidate at 130 Myr and the next youngest at 180 Myr. The remaining ones are at 200 Myr up to 4 Gyr with an average reddening of 0.33 mag.

While the expectation value method returns possibly unreliable results for the bimodal cases (about 10 per cent of our NGC 3351 sample), the maximum likelihood method is not a perfect solution. By choosing the single model with the highest likelihood, information within the PDF is ignored. The maximum likelihood model may be surrounded by unlikely models while the second mode of the PDF, at a lower relative likelihood, could be encompassed by equally likely models which gives better confidence that the second mode may actually provide a better estimate of the cluster’s physical properties. This also makes it difficult to characterize the uncertainties on the maximum likelihood result. With the expectation value, the uncertainties on the estimate can be calculated as the square root of the second moment of the posterior PDF which describes the width of the peak in the PDF and therefore the uncertainty of the estimate. By ignoring the information available in the PDF, there is no way to measure the uncertainty on the maximum likelihood result.

More work is needed to adequately explore these trends and effects before the Bayesian analysis can be fully implemented into the PHANGS–*HST* SED fitting pipeline. We have examined the expectation value as the Bayesian estimate but have highlighted potential issues with these estimates, particular for those objects with multimodal PDFs, in this section. In the following section, we discuss future efforts with the goal to resolve the outstanding issues with the Bayesian analysis, the age–reddening degeneracy, and other areas of the SED fitting which need further exploration.

## 6 FUTURE WORK

There are still a number of tweaks that can be made to the SED modelling procedures discussed here with the goal of providing the most accurate and robust star cluster properties across the



**Figure 17.** Bayesian estimates from the expectation values (black diamonds) compared to maximum likelihood Bayesian estimates (orange circles) assuming the fiducial prior. Logarithmic histograms show the distributions of the cluster properties for the two cases along each axis. The expectation value method finds a large dispersion in reddening for the clusters with ages from 10 to 100 Myr while the maximum likelihood method only finds two clusters in this age range; the clusters are shifted to young ages which leads to the large spike in reddening at the young ages. The maximum likelihood method is able to find clusters at 1 Myr which is not found with the expectation value method. The mass distribution is robust between the two Bayesian estimate methods.

PHANGS-*HST* sample as possible. Here, we outline future work that can be done to help reach this goal.

One of the biggest obstacles for all star cluster SED fitting endeavours is the age-reddening degeneracy. The answer usually involves the incorporation of more data and how to best apply those data. For the PHANGS-*HST* sample, one option is to supplement the cluster photometry with high-resolution  $H\alpha$  observations. Whitmore et al. (2020) detail methodology for using *HST*  $H\alpha$  observations to improve upon the ages provided by star cluster SED fitting. This could be applied to our PHANGS-*HST* pipeline if such observations are available for our sample. Currently, there are less than a dozen galaxies within the PHANGS-*HST* sample that have *HST*  $H\alpha$  imaging, and archival observations do not necessarily cover the same footprint as PHANGS-*HST* as is the case for NGC 3351. Ground-based  $H\alpha$  could in principle be used, but the impacts of their significantly coarser resolution ( $\sim 1$  arcsec) would first need to be understood. The PHANGS collaboration has PHANGS-MUSE  $\sim 1$  arcsec  $H\alpha$  maps available for 19 of the PHANGS-*HST* galaxy sample. Ground-based  $H\alpha$  maps from narrow-band photometry have been compiled by the PHANGS collaboration (Razza et al., in preparation) which cover all PHANGS galaxies.

One additional option to potentially break the degeneracy is to apply a different prior for the reddening  $E(B - V)$ . Currently, we

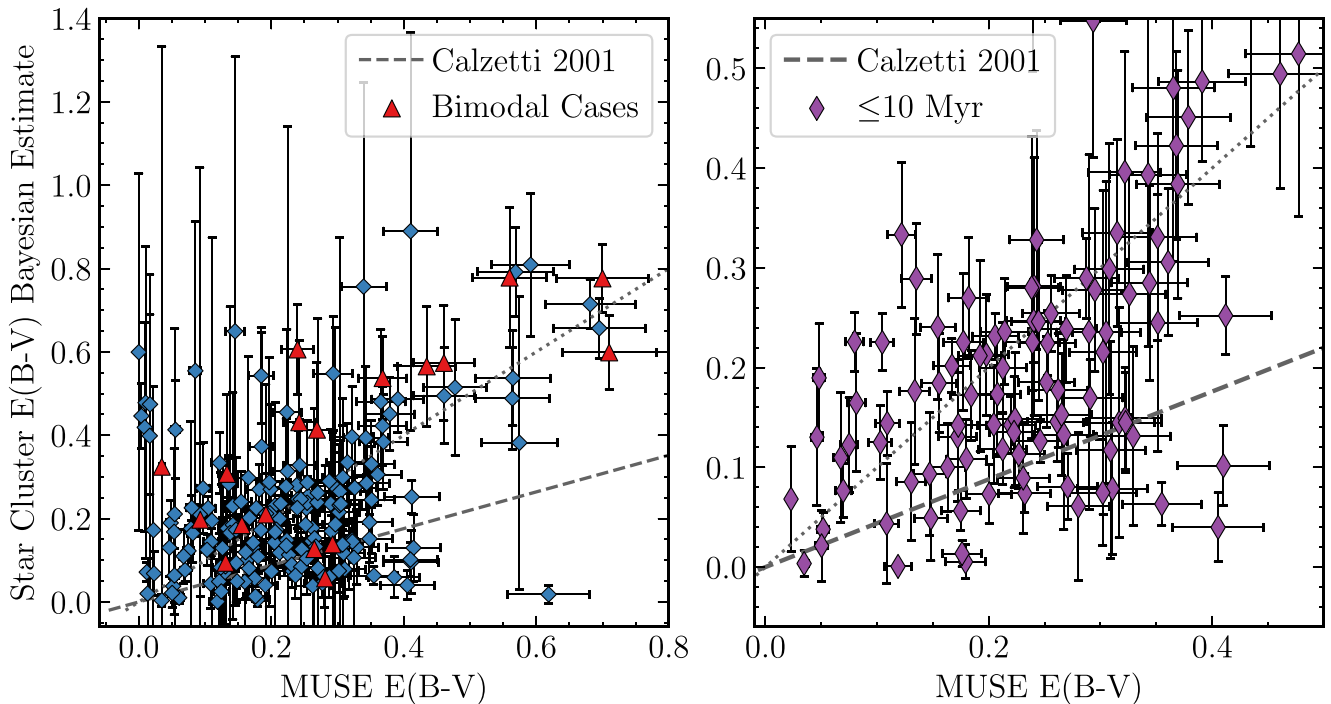
assume a flat prior where all reddening models are equally likely. However, a non-flat prior is an option, particularly a lognormal distribution on the basis that the gas column density PDF, which is completely decoupled from star cluster work, follows a lognormal distribution. With LEGUS star clusters, Ashworth et al. (2017) produce an observed extinction distribution that strongly disfavors high  $A_V$ . This result indicates that a non-flat prior for the reddening could be necessary.

Another possibility for breaking the age-reddening degeneracy is to take advantage of the PHANGS high-resolution ALMA CO maps. Using a gas-to-dust ratio to convert the CO maps into dust maps, we could match the star clusters with the dust map to identify how much dust is present at the cluster location. This information would then inform which mode of the bimodal age-reddening distribution to choose as the best estimate for the age and reddening. However, there are line-of-sight uncertainties with this simple matching. It is unknown whether the cluster resides in front of, within, or behind the dust. Regardless, a lack of CO emission would provide a strong prior that there is little extinction. Additionally, the gas-to-dust conversion introduces uncertainties and similar to the ground-based  $H\alpha$  data, there is also a mismatch in resolution, and the impacts of which would need to be examined since the CO maps have resolutions of  $\sim 1$  arcsec. Despite these drawbacks, the ALMA CO maps are still a promising direction in which to pursue, as they are, by design, available for the full PHANGS-*HST* sample. In conjunction with this method, we can identify dust lanes in the PHANGS-*HST* optical imaging observations and apply this dust information to the SED fitting as well.

As discussed in Sections 4.2, PHANGS-MUSE observations produce Balmer decrement maps as well as stellar  $E(B - V)$  maps from stellar continuum fitting at resolutions similar to ALMA (1 arcsec). Pellegrini et al. (2020) find Balmer decrement measurements work well for de-reddening  $H\alpha$  fluxes in model galaxies using the population synthesis model WARPFIELD. This is a promising result which supports the use of PHANGS-MUSE Balmer decrement measurements to inform our SED modelling.

Here, as a first step, we can check if the Bayesian analysis from Section 5.6 is inferring reddening values that are in agreement with  $E(B - V)$  values from the PHANGS-MUSE data. The PHANGS-MUSE  $E(B - V)$  measurements are derived from  $A_V$  maps from Balmer decrement measurements. We match the location of each star cluster to the closest pixel in the PHANGS-MUSE map and measure the mean  $E(B - V)$  value of a three by three pixel grid ( $0.36$  arcsec $^2$ ) centred on that pixel. We compare our star cluster Bayesian estimates to the MUSE measurements in Fig. 18 and find a weak correlation between the two  $E(B - V)$  measurements. In the left-hand panel of Fig. 18, 18 of the bimodal cases which have MUSE Balmer decrements at their locations are highlighted. The majority of bimodal cases agree with the MUSE measurements within uncertainties. A few of the bimodal cases could possibly be resolved by including the information provided by the MUSE data. More of the star clusters not identified as bimodal could possibly be improved by using the MUSE measurements. However, more attention beyond the scope of this paper will need to be devoted to such analysis before it can be successfully integrated into the PHANGS-*HST* star cluster catalogue pipeline.

The right-hand panel of Fig. 18 focuses on a potentially encouraging result. All the clusters with Bayesian estimates of 10 Myr and younger are plotted along with the relation  $E(B - V)_{\text{star}} = 0.44E(B - V)_{\text{gas}}$  (Calzetti 2001, equation 9) which gives roughly a factor of 2 greater reddening in gas compared to stars (or star clusters in our case). Clusters at these young ages should still have



**Figure 18.** Star cluster Bayesian estimates for  $E(B - V)$  versus the MUSE  $E(B - V)$  measurements at the location the star clusters. MUSE  $E(B - V)$  measurements are derived from Balmer decrements to calculate  $A_V$  assuming  $R_V = 4.05$ . 266 clusters lie within the PHANGS–MUSE footprint but 75 clusters have no Balmer decrement measurement at their location. Clusters with bimodal PDFs are highlighted as red triangles. The dotted line shows unity and the dashed line marks the factor of 1/0.44 times greater reddening of gas compared to stars (Calzetti 2001, equation 9).  $1\sigma$  error bars are given for the Bayesian estimates and 10 per cent error bars are given for the MUSE reddening values. The right-hand panel shows a subset of the left-hand panel highlighting only the clusters which are found to be 10 Myr and younger.

$H\alpha$  associated with them and the MUSE reddening measurements (derived from Balmer decrements) depend on  $H\alpha$  detection. We see a potential correlation along the Calzetti (2001) line for some of the young star clusters which may explain the deviation from unity for the clusters. Thus, it is encouraging that our star cluster reddening estimates are consistent (within the large scatter) with the values from MUSE for these young clusters along the expected relation.

In Section 4.1, we discuss the choice to not include a nebular emission component in our SED fitting. However, there is a case to be made to include nebular continuum and line emission for clusters that lie bluer than the SSP tracks at 1 Myr to 5 Myr in colour–colour space (see Fig. 5). In our NGC 3351 star cluster sample, we find three clusters in this region. For such clusters, including a nebular emission component in the fitting could be justified. Additionally, other clusters could intrinsically have colours in this same region but have been reddened by dust and therefore lie to the right of the SSP track in colour–colour space. This approach will need to be tested to see if it makes a significant difference in the SED fitting results, since the solar metallicity tracks at those ages are very similar, to warrant implementing it into our pipeline in the future.

In Section 5.6, we discuss the application of Bayesian priors to modify the likelihood values of each model on the grid. A final Bayesian parameter estimate is calculated as the expectation value or finding the model with the maximum likelihood. We show these methods only work well for certain populations and fail for others (e.g. bimodal cases). Going beyond these simple methods, we could incorporate each cluster’s entire PDF into computing the mass and age functions of the cluster population. This will circumvent the need

to fix every single bimodal case while utilizing the information-rich PDFs.

## 7 CONCLUSIONS

The PHANGS–*HST* project aims to study individual star clusters in 38 nearby galaxies making use of the publicly available SED fitting code CIGALE to characterize their physical properties.

(i) We test CIGALE’s ability to recover known cluster properties from a mock catalogue within the photometric uncertainties and find good recovery of ages (standard deviation of the difference between the ‘true’ and recovered ages of 0.31 dex), masses (standard deviation of the difference of 0.18 dex), and reddenings (standard deviation of the difference of 0.09 mag). The largest age residuals are found at 1 Myr, at around 10 Myr, and at the very oldest ages. Degeneracies in the SEDs of the SSP models at 5–50 Myr are the major cause of the large age residuals. We find no change in this result when using finer or coarser age grid sampling. We benchmark CIGALE’s ability to recover the same SED fitting results as the LEGUS stellar cluster catalogues and can successfully recover the LEGUS results. The median value of the ratios between the CIGALE and LEGUS ages is  $0.001 \pm 0.017$  dex and the median of the mass ratios is  $0.003 \pm 0.011$  dex.

(ii) Using the visually classified class 1 and 2 subset of the PHANGS–*HST* star cluster catalogue for NGC 3351, we test the SED modelling dependencies. We consider the differences between two SSP models (Bruzual & Charlot (2003) and YGGDRASIL) and the inclusion of nebular emission within those models. We explore how to treat extinction and reddening from dust in the SED modelling

and evaluate our chosen modelling parameters against IFU data from PHANGS–MUSE.

(iii) We test how fitting in linear fluxes or logarithmic magnitudes affects the SED results and find no significant changes in the resulting distributions of the ages, masses, and reddenings.

(iv) Based on the results of our tests in this paper, we choose to adopt for our SED fitting: linear fluxes, the BC03 SSP model, solar metallicity, an instantaneous burst SFH, a fully sampled Chabrier (2003) IMF, intrinsic reddening varying from 0 to 1.5 mag in 0.01 mag steps, and no nebular emission component. We sample age models linearly in 1 Myr intervals for 1–10 Myr and logarithmically for 11 Myr to 13.75 Gyr [ $\Delta \log(\text{Age}/\text{Myr}) \approx 0.3$ ]. This gives an age–reddening model grid with 16,610 models.

(v) We apply this SED modelling approach to derive the best-fitting star cluster ages, masses, and reddenings for NGC 3351. We find that the star clusters within the inner ring at the centre of NGC 3351 to be young and more massive than the rest of the clusters, which may be due to the funnelling of gas into the centre by the galactic bar. This funnelling triggers and maintains ongoing star formation.

(vi) The clusters present in the stellar bulge region are found to be much older on average, consistent with expectations from previous work.

(vii) The SED fitting results are checked against visually identified globular cluster candidates and, although they are found to be older relative than the rest of the cluster sample, the ages fall short of the expected age of 10 Gyr. This is most likely due to the solar-metallicity SSP model being used for an old, low-metallicity cluster population as well as the preference for the SED fitting to choose younger, higher reddened models (see Whitmore et al. 2020). Thus, caution should be exercised when studying globular cluster ages derived from broadband SED fitting. Clusters with visually identified H II regions are also checked, and the SED ages are below 10 Myr as is expected.

(viii) We study the cluster mass functions and find power-law slopes of  $\beta \sim -2$ , consistent with the literature. Maximum likelihood fits find no evidence of an upper mass cut-off for the clusters of NGC 3351 down to our completeness limit at  $10^{3.5} M_{\odot}$  for the 10–100 Myr age bin and  $10^{4.8} M_{\odot}$  for the 100–400 Myr age bin.

(ix) We explore the application a Bayesian prior of the form  $p_{\text{prior}} \propto M^{\beta} T^{\gamma}$  to modify the marginalized PDFs and derive Bayesian estimates for the cluster properties. Without sufficient physically motivated criteria to evaluate which values of  $\beta$  and  $\gamma$  provide the most-likely output distributions, we choose to adopt  $\beta = -2$  and  $\gamma = -0.5$  as our fiducial prior, and quantify the differences from this prior. We find that our Bayesian estimates perform poorly for clusters at 1 Myr and clusters with bimodal PDFs. Further testing of the Bayesian analysis is needed to diagnose and mitigate these issues.

(x) We find agreement with large scatter between our star cluster reddening Bayesian estimates and the MUSE reddening measurements (based on Balmer decrements). For clusters 10 Myr and younger, the disagreement between our measurement and MUSE may be explained by the Calzetti (2001) relation  $E(B - V)_{\text{star}} = 0.44E(B - V)_{\text{gas}}$ .

The SED fitting methodology detailed in this paper will be the basis for estimating star cluster physical properties in the forthcoming publicly available PHANGS–HST star cluster catalogues. The methodology will also be applied to the stellar associations identified in the PHANGS–HST sample which will be detailed in Larson et al. (in preparation). In short, a watershed algorithm is used to select the stellar associations of star clusters which provides a better way of

identifying the youngest and least-massive clusters, and yields a more complete picture than detection of the difficult-to-model asymmetric class 3 clusters. Applying the same SED fitting methodology with possible modifications for more complex SFHs will provide self-consistency across both the PHANGS–HST star cluster catalogue and the PHANGS–HST stellar association catalogue.

The physical properties derived by SED fitting for the PHANGS–HST cluster and stellar association catalogues will be crucial for studying the connections between molecular clouds and young star clusters across diverse galactic environments. The catalogues will provide both the  $\chi^2$  minimized best-fitting results, the Bayesian parameter estimates for star cluster ages, masses, and reddenings, and the associated uncertainties. These catalogues, including the catalogue for NGC 3351 used in this paper, will be made publicly available through the *Mikulski Archive for Space Telescopes (MAST)*.

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## DATA AVAILABILITY

The PHANGS–*HST* star cluster catalogues for NGC 3351 and the remaining PHANGS–*HST* galaxies will be made publicly available through the *Mikulski Archive for Space Telescopes (MAST)* in the near future. In the meantime, the NGC 3351 cluster catalogue used in this article will be shared on reasonable request to the corresponding author.

## REFERENCES

- Adamo A., Kruijssen J. M. D., Bastian N., Silva-Villa E., Ryon J., 2015, *MNRAS*, 452, 246
- Adamo A. et al., 2017, *ApJ*, 841, 131
- Adamo A. et al., 2020, *Space Sci. Rev.*, 216, 69
- Anand G. S. et al., 2020, *MNRAS*, 501, 3621
- Anders P., Fritze-v. Alvensleben U., 2003, *A&A*, 401, 1063
- Ashworth G. et al., 2017, *MNRAS*, 469, 2464
- Barbaro C., Bertelli C., 1977, *A&A*, 54, 243
- Bastian N., Gieles M., Efremov Y. N., Lamers H. J. G. L. M., 2005, *A&A*, 443, 79
- Bastian N., Trancho G., Konstantopoulos I. S., Miller B. W., 2009, *ApJ*, 701, 607
- Bastian N. et al., 2011, *MNRAS*, 417, L6
- Bik A., Lamers H. J. G. L. M., Bastian N., Panagia N., Romaniello M., 2003, *A&A*, 397, 473
- Boquien M., Burgarella D., Roehlly Y., Buat V., Ciesla L., Corre D., Inoue A. K., Salas H., 2019, *A&A*, 622, A103
- Boquien M., Duc P. A., Galliano F., Braine J., Lisenfeld U., Charmandaris V., Appleton P. N., 2010, *AJ*, 140, 2124
- Bruzual A. G., 2002, *IAUS*, 207, 616
- Bruzual G., Charlot S., 2003, *MNRAS*, 344, 1000 (BC03)
- Burgarella D., Buat V., Iglesias-Páramo J., 2005, *MNRAS*, 360, 1413
- Buta R., 1988, *ApJS*, 66, 233
- Calzetti D., 2001, *PASP*, 113, 1449
- Calzetti D., Armus L., Bohlin R. C., Kinney A. L., Koornneef J., Storchi-Bergmann T., 2000, *ApJ*, 533, 682
- Calzetti D. et al., 2015, *AJ*, 149, 51
- Cardelli J. A., Clayton G. C., Mathis J. S., 1989, *ApJ*, 345, 245
- Cerviño M., Luridiana V., 2006, *A&A*, 451, 475
- Chabrier G., 2003, *PASP*, 115, 763
- Chabrier G., 2005, in Corbelli E., Palla F., Zinnecker H., eds, *The Initial Mass Function 50 Years Later*. Springer Netherlands, Dordrecht, p. 41
- Chandar R., Fall S. M., Whitmore B. C., Mulia A. J., 2017, *ApJ*, 849, 128
- Chandar R. et al., 2010, *ApJ*, 719, 966
- Chevallard J., Charlot S., 2016, *MNRAS*, 462, 1415
- da Cunha E., Charlot S., Elbaz D., 2008, *MNRAS*, 388, 1595
- da Silva R. L., Fumagalli M., Krumholz M., 2012, *ApJ*, 745, 145
- Davidge T. J., Puzia T. H., McConnachie A. W., 2011, *ApJ*, 728, L23
- de Grijs R., Anders P., Bastian N., Lynds R., Lamers H. J. G. L. M., O’Neil E. J., 2003, *MNRAS*, 343, 1285
- Deveikis V., Narbutis D., Stonkutė R., Bridžius A., Vansevicius V., 2008, *Baltic Astron.*, 17, 351
- Elmegreen D. M., Chromey F. R., Santos M., Marshall D., 1997, *AJ*, 114, 1850
- Fall S. M., Chandar R., 2012, *ApJ*, 752, 96
- Fitzpatrick E. L., 1999, *PASP*, 111, 63
- Fouesneau M., Lançon A., Chandar R., Whitmore B. C., 2012, *ApJ*, 750, 60
- Fouesneau M. et al., 2014, *ApJ*, 786, 117
- Franzetti P., Scodreggio M., Garilli B., Fumana M., Paioro L., 2008, in Argyle R. W., Bunclark P. S., Lewis J. R., eds, *ASP Conf. Ser. Vol. 394, Astronomical Data Analysis Software and Systems XVII*. Astron. Soc. Pac., San Francisco, p. 642
- Freedman W. L. et al., 2001, *ApJ*, 553, 47
- Gieles M., Bastian N., 2008, *A&A*, 482, 165
- Girardi L., Bica E., 1993, *A&A*, 274, 279
- Grasha K. et al., 2015, *ApJ*, 815, 93
- Grasha K. et al., 2018, *MNRAS*, 481, 1016
- Grasha K. et al., 2019, *MNRAS*, 483, 4707
- Groves B., Dopita M. A., Sutherland R. S., Kewley L. J., Fischera J., Leitherer C., Brandl B., van Breugel W., 2008, *ApJS*, 176, 438
- Hannon S. et al., 2019, *MNRAS*, 490, 4648
- Han Y., Han Z., 2012, *ApJ*, 749, 123
- Herrera-Endoqui M., Díaz-García S., Laurikainen E., Salo H., 2015, *A&A*, 582, A86
- Hogg D. W., Bovy J., Lang D., 2010, preprint ([arXiv:1008.4686](https://arxiv.org/abs/1008.4686))
- Hunt L. K. et al., 2019, *A&A*, 621, A51
- Johnson L. C. et al., 2017, *ApJ*, 839, 78
- Kennicutt, Jr. R. C. et al., 2003, *PASP*, 115, 928
- Kreckel K., Blanc G. A., Schinnerer E., Groves B., Adamo A., Hughes A., Meidt S., 2016, *ApJ*, 827, 103
- Kreckel K. et al., 2019, *ApJ*, 887, 80
- Kroupa P., 2001, *MNRAS*, 322, 231
- Kruijssen J. M. D., 2014, *Class. Quantum Gravity*, 31, 244006
- Kruijssen J. M. D., Pelupessy F. I., Lamers H. J. G. L. M., Portegies Zwart S. F., Icke V., 2011, *MNRAS*, 414, 1339
- Krumholz M. R., McKee C. F., Bland-Hawthorn J., 2019, *ARA&A*, 57, 227
- Krumholz M. R. et al., 2015, *ApJ*, 812, 147
- Lada C. J., Lada E. A., 2003, *ARA&A*, 41, 57
- Lançon A., Mouhcine M., 2000, in Lançon A., Boily C. M., eds, *ASP Conf. Ser. Vol. 211, Massive Stellar Clusters*. Astron. Soc. Pac., San Francisco, p. 34
- Larsen S. S., 2009, *A&A*, 494, 539
- Lee J. C. et al., 2021, preprint ([arXiv:2101.02855](https://arxiv.org/abs/2101.02855))
- Leroy A. K. et al., 2018, *ApJ*, 869, 126
- Linden S. T. et al., 2017, *ApJ*, 843, 91
- Maraston C., 2005, *MNRAS*, 362, 799
- Messa M. et al., 2018, *MNRAS*, 477, 1683
- Mok A., Chandar R., Fall S. M., 2019, *ApJ*, 872, 93
- Moustakas J., Kennicutt, Jr. R. C., Tremonti C. A., Dale D. A., Smith J.-D. T., Calzetti D., 2010, *ApJS*, 190, 233
- Moustakas J. et al., 2013, *ApJ*, 767, 50
- Noll S., Burgarella D., Giovannoli E., Buat V., Marcellac D., Muñoz-Mateos J. C., 2009, *A&A*, 507, 1793
- Pellegrini E. W., Reissl S., Rahner D., Klessen R. S., Glover S. C. O., Pakmor R., Herrera-Camus R., Grand R. J. J., 2020, *MNRAS*, 498, 3193
- Portegies Zwart S. F., McMillan S. L. W., Gieles M., 2010, *ARA&A*, 48, 431
- Reina-Campos M., Kruijssen J. M. D., 2017, *MNRAS*, 469, 1282
- Reines A. E., Nidever D. L., Whelan D. G., Johnson K. E., 2010, *ApJ*, 708, 26
- Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
- Schlegel D. J., Finkbeiner D. P., Davis M., 1998, *ApJ*, 500, 525
- Swartz D. A., Yukita M., Tennant A. F., Soria R., Ghosh K. K., 2006, *ApJ*, 647, 1030
- Wei W. et al., 2020, *MNRAS*, 493, 3178
- Whitmore B. C., Chandar R., Fall S. M., 2007, *AJ*, 133, 1067
- Whitmore B. C. et al., 2014, *ApJ*, 795, 156
- Whitmore B. C. et al., 2020, *ApJ*, 889, 154
- Williams J. P., McKee C. F., 1997, *ApJ*, 476, 166
- Zackrisson E., Rydberg C.-E., Schaerer D., Östlin G., Tuli M., 2011, *ApJ*, 740, 13
- Zhang Q., Fall S. M., 1999, *ApJ*, 527, L81

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