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ESO 137-002: a large spiral undergoing edge-on ram-pressure stripping with little star formation in the tail

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ABSTRACT

Ram pressure stripping (RPS) is an important mechanism for galaxy evolution. In this work, we present results from *HST* and *APEX* observations of one RPS galaxy, ESO 137-002 in the closest rich cluster Abell 3627. The galaxy is known to host prominent X-ray and H α tails. The *HST* data reveal significant features indicative of RPS in the galaxy, including asymmetric distribution of dust in the galaxy, dust filaments, and dust clouds in ablation generally aligned with the direction of ram pressure, and young star clusters immediately upstream of the residual dust clouds that suggest star formation (SF) triggered by RPS. The distribution of the molecular gas is asymmetric in the galaxy, with no CO upstream and abundant CO downstream and in the inner tail region. A total amount of $\sim 5.5 \times 10^9 M_{\odot}$ of molecular gas is detected in the galaxy and its tail. On the other hand, we do not detect any active SF in the X-ray and H α tails of ESO 137-002 with the *HST* data and place a limit on the SF efficiency in the tail. Hence, if selected by SF behind the galaxy in the optical or UV (e.g. surveys like GASP or using the *Galex* data), ESO 137-002 will not be considered a ‘jellyfish’ galaxy. Thus, galaxies like ESO 137-002 are important for our comprehensive understanding of RPS galaxies and the evolution of the stripped material. ESO 137-002 also presents a great example of an edge-on galaxy experiencing a nearly edge-on RPS wind.

Key words: galaxies: clusters: individual: Abell 3627 – galaxies: evolution – galaxies: individual: ESO 137-002 – galaxies: star clusters: individual: star formation – galaxies: interactions.

1 INTRODUCTION

Galaxy clusters are the largest gravitationally collapsed structures in the Universe, filled with hot diffuse plasma with a typical temperature of 10^7 – 10^8 K and electron number density of 10^{-4} – 10^{-1} cm $^{-3}$ (e.g. Kravtsov & Borgani 2012). They contain a large fraction of early-type galaxies mostly concentrated around the cluster core and have a relatively small fraction of spirals/irregulars (e.g. Dressler 1980). Cluster galaxies are preferentially red with little cold gas and suppressed star formation (SF) activity. Conversely, field galaxies are mostly spirals that are preferentially blue, exhibit more cold gas, and have active SF. These characteristics suggest that galaxy evolution is strongly influenced by the environment (Boselli & Gavazzi 2006). Galaxy clusters are cosmic labs for studying the effect of the environment on galaxy morphology and evolution.

Ram pressure stripping (RPS) (Gunn & Gott 1972) is one of the most important mechanisms for evolution of gas-rich galaxies in clusters by removing hot and cold interstellar medium (ISM) from the galaxy. As the galaxy moves through the intracluster medium (ICM), SF can be suppressed as the ISM is lost. There can also be a momentarily initial starburst, as thermal instabilities and turbulent motions trigger the collapse of the molecular clouds before SF is eventually quenched (e.g. Bekki & Couch 2003). The fate of the stripped ISM is also an interesting question. It eventually should get mixed with the ICM to enrich its metallicity (e.g. Schindler & Diaferio 2008). On the other hand, it is now known that some fraction of the stripped ISM can turn into stars in the galactic halo and the intracluster space, even though the quenching happens within the galaxy (e.g. Sun, Donahue & Voit 2007; Yagi et al. 2007; Smith et al. 2010; Yagi et al. 2010; Jáchym et al. 2014; Poggianti et al. 2017).

Observational evidence of RPS has been revealed in many different bands. Stripped gaseous tails and star-forming clumps have been

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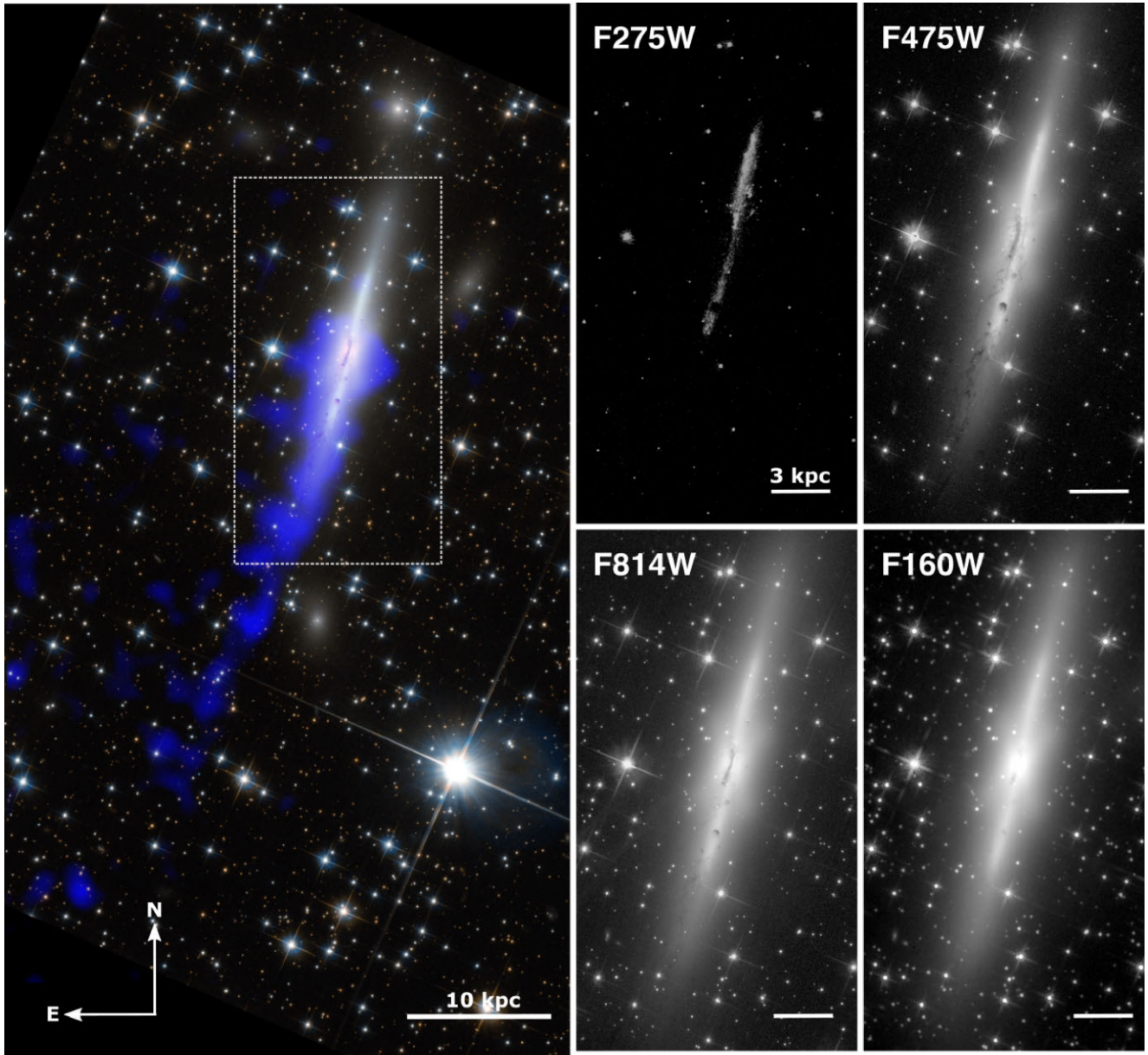


Figure 1. Left-hand panel: Composite image of ESO 137-002 with the *HST* F475W/F814W data (Credit: ESA/Hubble & NASA). We also overlay the 0.6–2 keV X-ray emission from *Chandra* in blue from Zhang et al. (2013) to show the X-ray tail. The galaxy region inside the white box in the dotted line is highlighted in the right-hand panels. Right-hand panel: A close-up view of ESO 137-002 in the four *HST* bands (see Table 2 for the detail of these four bands). While the upstream of the galaxy is dust free, many dusty clouds and filaments are observed in the downstream. One can also observe the X-shaped bulge, especially on the red bands.

detected in UV and $H\alpha$ (e.g. Gavazzi et al. 2001; Cortese et al. 2007; Sun et al. 2007; Yagi et al. 2007; Yoshida et al. 2008; Hester et al. 2010; Smith et al. 2010; Yagi et al. 2010; Kenney et al. 2014; Fossati et al. 2016). H I tails in some late-type galaxies in the Virgo cluster and other clusters have also been discovered (e.g. Oosterloo & van Gorkom 2005; Chung et al. 2007). Stripped tails have also been revealed in X-ray (e.g. Wang, Owen & Ledlow 2004; Sun & Vikhlinin 2005; Sun et al. 2006, 2010) and radio (e.g. Giovanelli & Haynes 1985; Gavazzi & Jaffe 1987; Chen et al. 2020). Jáchym et al. (2014, 2017, 2019) detected a large amount of cold molecular gas, traced by the CO emission, in the stripped tails behind ESO 137-001 and D100, embedded in the hot ICM. Also, cold molecular gas has

been detected in the tails of other galaxies undergoing RPS (Verdugo et al. 2015; Moretti et al. 2018, 2020). Warm molecular H_2 gas clouds have also been revealed in the tails of ESO 137-001 and other RPS galaxies (Sivanandam, Rieke & Rieke 2010, 2014). With the detection of active SF in many RPS galaxies, these galaxies are often called ‘jellyfish galaxies’.

RPS has also been studied extensively using numerical simulations (e.g. Abadi, Moore & Bower 1999; Quilis, Moore & Bower 2000; Schulz & Struck 2001; Vollmer et al. 2001; Roediger & Brüggén 2006; Jáchym et al. 2007; Kapferer et al. 2009; Tonnesen & Bryan 2012; Ruszkowski et al. 2014). These simulations demonstrate important effects of RPS on galaxy evolution like disc truncation, SF

Table 1. Properties of ESO 137-002.

| Parameter | ESO 137-002 |
|--|-----------------------|
| Heliocentric velocity (km s ⁻¹) ^a | 5691 (+820) |
| Distance (Mpc) ^b | 69.6 |
| Offset (kpc) ^c | 110 |
| Position Angle | ~13.1° |
| Inclination | ~83° |
| W1 (Vega mag) ^d | 10.21 |
| L _{W1} (10 ⁹ L _⊙) ^d | 3.39 |
| W1 – W4 (Vega mag) ^d | 4.25 |
| m _{F160W} (AB mag) ^e | 11.39 |
| Half light semimajor axis (kpc) ^e | 3.30 |
| M _* (10 ⁹ M _⊙) ^f | 32–39 |
| M _{mol} (10 ⁹ M _⊙) ^g | ~5.5 |
| L _{FIR} (10 ¹⁰ L _⊙) ^h | 1.22 |
| SFR (M _⊙ yr ⁻¹) ⁱ | 0.94 |
| Tail length (kpc) ^j | 40 (X-ray), >20 (H α) |

Notes. ^aThe heliocentric velocity from Wegner et al. (2003). The velocity value in parentheses is the radial velocity relative to that of Abell 3627.

^bFor consistency, we use the cluster redshift ($z = 0.0163$) and the luminosity distance used in Sun et al. (2010). 1 arcsec = 0.327 kpc.

^cThe projected offset of the galaxy from the X-ray centre of A3627

^dThe *WISE* 3.4 μm magnitude, luminosity, and the *WISE* 3.4–22 μm colour. The Galactic extinction was corrected with the relation from Indebetouw et al. (2005).

^eThe total magnitude and the half-light semimajor axis at the F160W band.

^fThe total stellar mass estimated from Sun et al. (2010).

^gThe total amount of molecular gas detected from ESO 137-002 from this work

^hThe total FIR luminosity from the *Herschel* data (see Section 7.1)

ⁱThe average value from the first estimate (1.08) based on the *Galax* NUV flux density and the total FIR luminosity from *Herschel* with the relation from Hao et al. (2011), and the second estimate (0.80) based on the *WISE* 22 μm flux density with the relation from Lee, Hwang & Ko (2013). The Kroupa IMF is assumed.

^jTail length from Zhang et al. (2013)

quenching, central bulge build-up, formation of flocculent arms, the transformation of dwarf galaxies, and long filamentary structures. The overall results of these simulations have brought to a more realistic comparison with the observation data.

In this paper, we focus on ESO 137-002 (Fig. 1), which is a large and bright late-type galaxy in the closest rich cluster Abell 3627 and hosts a Seyfert2-like active galactic nucleus (AGN) (Sun et al. 2010; Zhang et al. 2013). It is at a projected distance of only ~110 kpc (or ~0.06 r_{200}) from the cluster's X-ray centre and has a radial velocity of ~+820 km s⁻¹ with respect to that of the cluster. Table 1 lists some properties of ESO 137-002. It is ~5 times more massive than its neighbour RPS galaxy ESO 137-001 (Sun et al. 2010) but with a similar SFR and a similar optical size. While stripping in ESO 137-001 is close to face-on, ESO 137-002 is experiencing a near edge-on stripping, which makes it an important comparison with ESO 137-001 for detailed studies on RPS and galaxy evolution. With an axial ratio of 0.3 from the *H*-band image and assuming a morphological type from Sa to Sb (Sab), using the classical Hubble formula, the inclination angle, measured between the line-of-sight and the disc axis, is ~88°–78° (81°) so we are viewing the galaxy nearly edge-on (an average inclination angle of 83° is then adopted in this work). Downstream of ESO 137-002, there is a ~40 kpc narrow X-ray tail (Sun et al. 2010; Zhang et al. 2013) with a constant width of ~3 kpc (Zhang et al. 2013). ESO 137-002 also features double H α tails extending ~20 kpc from the nucleus without any H II regions identified (Sun et al. 2010). The secondary tail (~12 kpc)

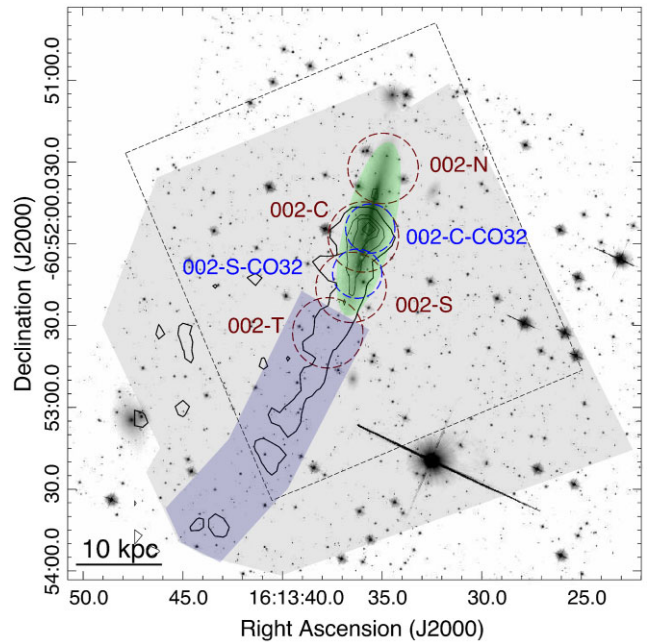


Figure 2. Regions of interest – tail, galaxy, and control regions shown by a lightblue polygon, a green ellipse, and a grey polygon (excluding the first two regions), respectively, on the F475W image. The centre and axial ratio of the ellipse is given in Section 4.1. The tail region encloses the X-ray and H α tails shown in Zhang et al. (2013) and the X-ray contours are also shown. The choice of the control region is mainly decided by the FOV of the F275W data. The dashed line shows the FOV of the F160W data. The source populations in these three regions are studied in Section 5. The red dashed circles show *APEX* pointings with ¹²CO(2-1) beams (FWHM = 26'' 7 = 8.7 kpc) and the blue dashed circles show *APEX* pointings with ¹²CO(3-2) beams (FWHM = 17'' 8 = 5.8 kpc). The CO spectra in these regions are shown in Fig. 12, Fig. 13, and Fig. 14 and discussed in section 6.

resides at a distance of ~7.5 kpc from the nucleus at a ~23° to the main tail. The H α main tail is spatially coincident with the X-ray tail. The sharp H α leading edge is also at the same position of the X-ray leading edge. As emphasized in Sun et al. (2010) and Zhang et al. (2013), the existing H α data of ESO 137-002 are quite shallow and its X-ray data are also not deep for its proximity to the bright cluster core. The true extent of ESO 137-002's X-ray and H α tails can be much longer than what is revealed from the current data.

This paper presents new results on ESO 137-002 from the *Hubble Space Telescope* (*HST*) and the *Atacama Pathfinder EXperiment* telescope (*APEX*). We focus on the main body and tail of ESO 137-002 (regions of interests defined in Fig. 2). Sections 2 and 3 give details about the observations and data reduction for *HST* and *APEX*, respectively. Section 4 presents the *HST* results on ESO 137-002 itself. Section 5 studies the *HST* source population in ESO 137-002's tail. Section 6 presents the CO emission in the galaxy and its tail. Section 7 is the discussion and we present the final conclusions in Section 8.

2 *HST* OBSERVATIONS AND DATA ANALYSIS

ESO 137-002 was observed in the *HST* program 12372 (PI:Sun), with two *HST* imaging instruments – Advance Camera for Surveys (ACS) and Wide Field Camera 3 (WFC3). Table 2 lists details of the observations. The F475W and F814W high-throughput filters in the ACS/Wide Field Channel (WFC) allow the detection of faint

Table 2. *HST* Observations of ESO 137-002 (PI: Sun).

| Filter (mean λ /FWHM) | Instrument | Mode | Dither ^a | Date | Exp (s) |
|-------------------------------|------------|------------|---------------------|------------|-------------------|
| F275W (2719/418 Å) | WFC3/UVIS | ACCUM | 3 (2.4 arcsec) | 07/17/2011 | 3×1010.0 |
| F475W (4802/1437 Å) | ACS/WFC | ACCUM | 2 (3.01 arcsec) | 07/17/2011 | 2×864.0 |
| F814W (8129/1856 Å) | ACS/WFC | ACCUM | 2 (3.01 arcsec) | 07/17/2011 | 2×339.0 |
| F160W (15436/2874 Å) | WFC3/IR | MULTIACCUM | 2 (0.636 arcsec) | 07/17/2011 | 2×399.2 |

Note. ^anumber of dither positions (and offset between each dither)

features (e.g. faint star clusters in the tail) and a deep colour map of the galaxy. The F275W filter in the WFC3/UVIS channel adds important information on the SF in the last several 10^7 yr. On the other hand, the F160W filter in the WFC3/IR channel is the band least affected by the dust extinction. When combined, they can be used to constrain the stellar age of star clusters found in the tail and the galaxy. The foreground Milky Way (MW) extinction is high in the direction of ESO 137-002 as it lies near the Galactic plane, 1.096 mag for F275W, 0.653 mag for F475W, 0.305 mag for F814W, and 0.102 mag for F160W, with the extinction law and the dust map from Fitzpatrick (1999) and Schlafly & Finkbeiner (2011), respectively. The $3\text{-}\sigma$ detection thresholds for point sources are 29.4, 30.8, 30.2, and 31.1 mag for F275W, F475W, F814W, and F160W, respectively.

The first step of the data reduction involves alignment of different frames in each filter using Tweakreg (Gonzaga & et al. 2012). For absolute astrometry, we used the Guide Star Catalog II (GSC2; Lasker et al. 2008), which aligns different frames within an error of ~ 0.05 arcsec. After alignment, we ran AstroDrizzle to combine different frames for the same filter. The data were also corrected for charge transfer efficiency (CTE) to ensure maximum accuracy in the photometry measurement. The pixel scale was fixed to 0.03 arcsec (or 9.81 pc).

The combined images in the F275W, F475W, and F814W bands still have many CRs left, especially at the ACS and WFC3 gaps, and edges. The Laplacian Edge Detection algorithm by van Dokkum (2001) has been used for further CR removal. The algorithm assumes that a real source such as a star or a galaxy has a smooth gradient light profile whereas a spurious source like CRs have discontinuous gradient. This algorithm usually works very well for ground optical data but requires extra caution and tests for the *HST* data for its very sharp point spread function (PSF). We also ran SExtractor (Bertin & Arnouts 1996) with a low detection threshold for the sole purpose of identifying any spurious objects like CRs. Sources detected in only one *HST* band are considered candidate CRs. We visually inspected all these candidates to verify their nature as CRs and subsequently removed confirmed CRs from the image.

3 APEX CO-OBSERVATIONS

The observations of ESO 137-002 were carried out with the *APEX* 12 m antenna in 2011 September with the program ID of 088.B-0934(A) and in 2014 September with the program ID of 094.B-0766(A). The observations were done at the frequency of the $^{12}\text{CO}(2-1)$ ($\nu_{\text{rest}} = 230.538$ GHz) using the APEX-1 receiver of the Swedish Heterodyne Facility Instrument (SHFI), and partially also at the $^{12}\text{CO}(3-2)$ frequency ($\nu_{\text{rest}} = 345.796$ GHz) using the APEX-2. The eXtended Fast Fourier Transform Spectrometer (XFFTS) backend was used with a total bandwidth of 2.5 GHz divided into 32 768 channels. The corresponding velocity resolution is about 0.1 km s^{-1} . The XFFTS consisted of two units with a 1 GHz overlap region. It thus covered the entire IF bandwidth of the SHFI. At the observed CO(2-1) and CO(3-2) frequencies, the full width at half-maximum

Table 3. APEX CO observations of ESO 137-002 (PI: Jáchym).

| Pointing (transition) | R.A. (J2000) | Dec. (J2000) | T_{ON} (min) |
|-----------------------|-----------------|-----------------|--------------------------|
| 002-C (2-1) | 16:13:35.83 | -60:51:58.7 | 26 |
| 002-S (2-1) | 16:13:36.47 | -60:52:16.8 | 43 |
| 002-N (2-1) | 16:13:34.89 | -60:51:32.3 | 89 |
| 002-T (2-1) | 16:13:37.70 | -60:52:33.6 | 77 |
| 002-C-CO32 (3-2) | 16:13:35.62 | -60:51:54.6 | 42 |
| 002-S-CO32 (3-2) | 16:13:36.26 | -60:52:10.9 | 138 |

(FWHM) of the primary beam of the telescope is ~ 26.7 and ~ 17.8 arcsec, respectively (following $\theta_{\text{MB}} \approx 1.17\lambda/D$), which for the adopted distance of the Norma Cluster corresponds to ~ 8.7 and 5.8 kpc, respectively.

The observations were done in a symmetric Wobbler switching mode with the maximum throw of 300 arcsec in order to avoid with OFF positions the tail if oriented in azimuth. Three integration points were selected over the main body of the galaxy, one (002-C') aiming nearly at the centre of ESO 137-002 (~ 4.4 arcsec = 1.4 kpc southwards off the centre), another ('002-S') covering the southern part of the disc, where H α and X-ray tail connects to the galaxy, and the last covering the northern disc part (see Fig. 2). The pointings 002-C and 002-S were selected to cover most of the brightest X-ray emission. The 002-N pointing is located symmetrically to the 002-S one (with respect to the optical galaxy centre) in order to reveal possible asymmetries in the CO-emitting gas distribution. Another integration point ('002-T') was aimed over the inner tail region, about ~ 15 kpc from its optical centre, already outside of the disc. Two regions were also observed at the $^{12}\text{CO}(3-2)$ frequency, over the optical centre of the galaxy and an adjacent location to the south. Coordinates of the observed positions are given in Table 3, together with the actual on-source observing times. The receiver was tuned to the $^{12}\text{CO}(2-1)$ frequency shifted in each position to its respective optical radial velocity. Observing conditions were good with PWV $\lesssim 1$ mm. The system temperatures were typically about 140 K.

The data were reduced according to the standard procedure using CLASS from the GILDAS software package developed at IRAM. Bad scans were flagged and emission line-free channels in the total width of about 1000 km s^{-1} were used to subtract first-order baselines. The corrected antenna temperatures, T_{A}^* , provided by the *APEX* calibration pipeline (Dumke & Mac-Auliffe 2010), were converted to main-beam brightness temperature by $T_{\text{mb}} = T_{\text{A}}^*/\eta_{\text{mb}}$, using a main beam efficiency of about $\eta_{\text{mb}} = 0.75$ at 230 GHz and 0.73 at 345 GHz. The rms noise levels typically of $1\text{--}2 \text{ mK}$ per 12.7 km s^{-1} channels were obtained in the main body pointings. Gaussian fits were used to measure peak T_{mb} , width, and position of the detected CO lines. Flux densities can be obtained using the conversion factor $S_{\nu}/T_{\text{mb}} = 39 \text{ Jy beam}^{-1}/\text{K}$ for the APEX-1 receiver at 230 GHz, and $41 \text{ Jy beam}^{-1}/\text{K}$ for the APEX-2 receiver at 345 GHz.

Table 4. GALFIT fits on the F160W image of ESO 137-002.

| Parameter | Single | Double | |
|-----------------------|----------------------------|--------|-------|
| | ($\chi^2_{\nu} = 5.693$) | Bulge | Disc |
| Total magnitude (mag) | 11.53 | 12.49 | 12.12 |
| r_e (kpc) | 4.09 | 2.05 | 5.17 |
| Sérsic index | 1.41 | 2.43 | 0.74 |
| Axis ratio | 0.23 | 0.66 | 0.13 |
| PA (deg) | -13.1 | -13.9 | -13.3 |

Note. The axial ratio is the ratio between the minor axis and the major axis. The position angle is measured relative to the north and counterclockwise.

4 HST VIEW OF ESO 137-002

4.1 Morphology and light profiles

Fig. 1 shows the *HST* composite and individual images of ESO 137-002 in four bands, also with the X-ray image from *Chandra* shown. ESO 137-002 is a large, edge-on spiral galaxy with a boxy bulge. While the upstream side of the galaxy is nearly dust-free, there are many dust features downstream, which suggests that the RPS has nearly cleared the north half of the galaxy and the stripping is ongoing within 1–2 kpc of the nucleus and the southern part. We attempted to fit the F160W image of ESO 137-002 (least affected by the dust extinction) with the 2D image fitting algorithm – GALFIT (Peng et al. 2002). Single and double Sérsic models are applied. The results are listed in Table 4. A double Sérsic model (one for the bulge and the other for the disc) can fit the galaxy image reasonably well. The resulting Sérsic index of each component is also reasonable with the trend found in large surveys like *GAMA* (e.g. Lange et al. 2015).

The galaxy has a boxy bulge (Fig. 3). After subtracting the double Sérsic model as shown in Table 4, the X/peanut-shaped bulge is clear, which implies the existence of a bar with a length of ~ 5 kpc in the bulge. The X-shaped structure can be the consequence of dynamical instabilities in the stellar bars creating thick vertical layers, often seen in massive galaxies (e.g. Erwin & Debattista 2017; Laurikainen & Salo 2017).

To quantitatively examine the galaxy structure, we derive the surface brightness profiles in all four bands, along the major axis and in elliptical annuli centred at the nucleus (Fig. 4). The galaxy centre is set at the nuclear position presented in Section 4.4. The major axis has a position angle (PA) of -13.1° (measured from the north and running counterclockwise). The total F160W light of the galaxy is measured to be 11.39 mag (without correction on the intrinsic extinction) and the half-light semimajor axis is 3.30 ± 0.05 kpc at F160W. Both the total magnitude and the half-light size are similar to the best-fitting results from the single-Sérsic fit from GALFIT.

The light profiles along the major axis are measured in 25 boxes, each with a dimension of 2.0 arcsec (width) \times 2.5 arcsec (length). The elliptical annuli have an axial ratio of 0.23. Fig. 4 shows that the NIR light profiles in F814W and F160W are more symmetric and smoother than the blue light profiles in F275W and F475W, where the effect of dust extinction and SF is more pronounced. The strong dust lane around the nucleus affects all light profiles. The F275W light profile is the least smooth one. We also searched for tidal features around the galaxy, especially in the F814W and F160W bands. No such features are detected from the current data.

We further applied PROFILER (Ciambur 2016) to decompose the 1D light profile to verify the non-axisymmetric components like an X/peanut-shaped bulge in ESO 137-002. The input isophote table was generated by the IRAF task *ISOFIT* (Ciambur & Graham

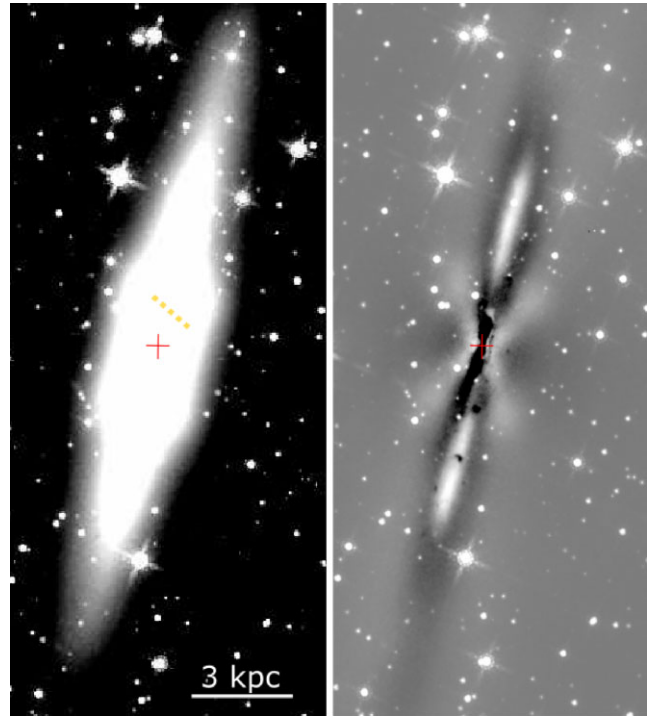


Figure 3. Left-hand panel: the F160W image to show the boxy bulge. The yellow dashed line shows the approximate position of the stripping front, determined from the X-ray and $H\alpha$ data (Zhang et al. 2013). Note this projected stripping front is really the outer edge of the curved, 3D stripping front. Right-hand panel: the residual image after removing the best-fitting double Sérsic model with GALFIT (see Table 4 for the detail of the fit). The central bulge shows an X-shaped residual, typical for galaxies with a boxy bulge and a bar. The red cross in each panel shows the position of the nucleus.

2016) which is appropriate for the edge-on disc galaxies. We fit the F160W image with a single Sérsic component while keeping the ellipticity and PA free. The best-fitted values for the Sérsic index and the half-light radius are 1.24 and 4.36 kpc, respectively, in good agreement with the GALFIT results. Fig. 5 shows the full light profile along with the residual, PA, and two harmonic namely the fourth cosine harmonic amplitude (B_4) and the sixth cosine harmonic amplitude (B_6). The B_4 profile shows that the galaxy appears boxy within $R_{\text{maj}} \sim 3$ arcsec and becomes discy at large radii. The B_6 profile shows the significance of the X/peanut-shaped component at $R_{\text{maj}} \sim 5$ –6 arcsec.

4.2 Dust features

To better show the dust features in the galaxy, we applied unsharp masking on the F475W image, by subtracting the original image smoothed with 1 pixel from the same image smoothed with 20 pixels (Fig. 6). It shows many dust features around the nucleus and downstream, extending to at least 13.5 kpc to the SE. A prominent dust lane is observed around the nucleus and there are many dust filaments and clouds downstream. One can also see the upstream end of the nuclear dust lane (~ 1.5 kpc from the nucleus) is pushed to the east side, presumably by the ram pressure. The downstream part of the nuclear dust is also pushed to the east side (but to a less extent compared with the upstream part). Assuming a galactic rotation speed of 100 km s^{-1} at 1.5 kpc radius from the nucleus, the downstream and upstream dust lanes and the associated gas clouds would switch positions in ~ 46 Myr. Is the difference on the curvature of the dust

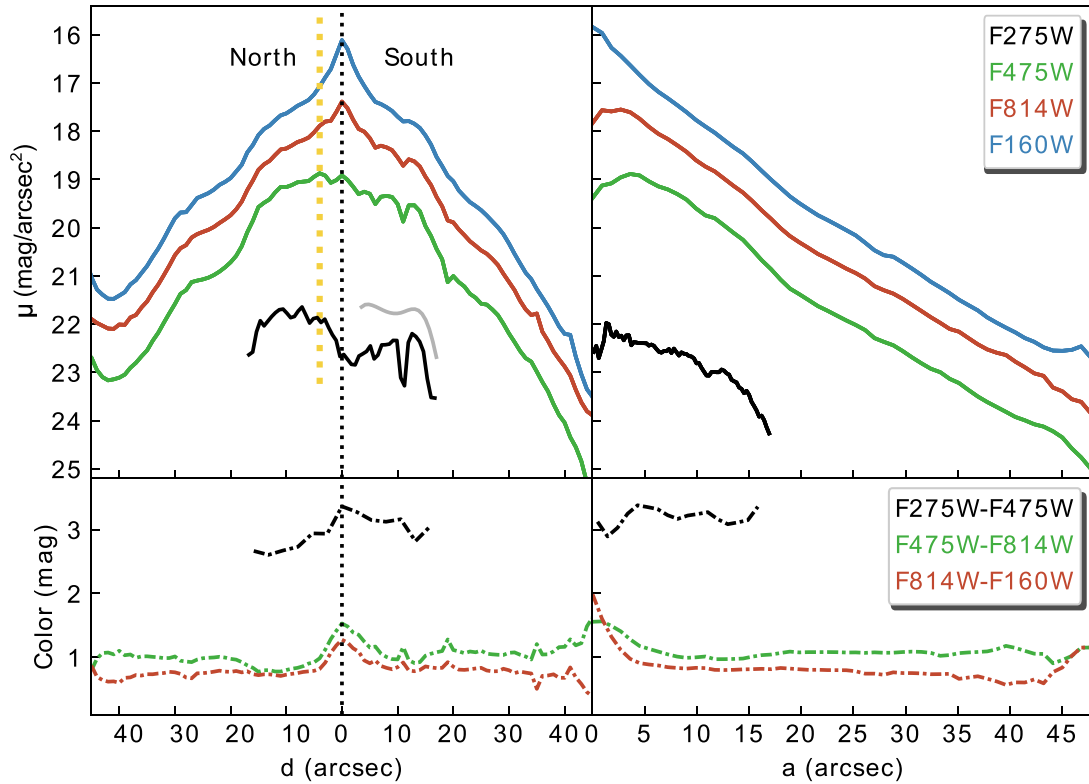


Figure 4. The surface brightness profiles measured along the major axis (upper left-hand panel, with $d = 0$ at the nucleus) and the elliptical annuli centred at the nucleus (upper right-hand panel, a as the semimajor axis), in all four bands. The colour profiles are shown in the lower panel. The elliptical annuli have a PA of -13.1° and an axial ratio 0.23. One can observe the significant effect of dust extinction at short wavelengths, while the F160W (or H -band) profiles are smoother and more symmetric than the profiles of the other three bands. As shown in Fig. 1, the F275W light truncates at ~ 16 arcsec from the nucleus. The yellow dashed line marks the position of the stripping front. The grey profile in the upper left-hand panel is the downstream F275W profile after correction of the intrinsic extinction with the F475W and F814W data (see Section 4.2 for detail), which suggests that the galaxy would have been more symmetric in F275W without the intrinsic extinction (see Fig. 1).

lanes upstream and downstream caused by the cloud fallback once moved to the downstream that is less affected by ram pressure? Future spectroscopic data with e.g. *MUSE* or *ALMA* can examine the ISM kinematic under the ram pressure impact.

We also used GALFIT again to produce a smooth model of the galaxy, this time on the F475W image. The final GALFIT model has four Sérsic components, two in the disc and two in the bulge, along with the Fourier modes (mode = 4, 6, and 8) (Peng et al. 2010). The residual image is shown in the middle panel of Fig. 6. It should be noted that the dark features upstream of the galaxy are caused by the oversubtraction of the model and not caused by dust. Its smooth and extended distribution is clearly different from those of local dust clouds. We also highlight dust features in zoom-in insets. Inset (a) shows the strong dust lane in the bulge, with the nucleus buried in the position with the strongest dust extinction. Inset (b) goes further downstream and encloses an interesting ‘X’-shaped feature with a length of ~ 0.8 kpc. Inset (c) shows numerous dust threads, arcs, and filaments further downstream, to at least 13.5 kpc from the nucleus. These dust threads have a typical length of ~ 0.8 kpc and width of ~ 0.1 kpc, which can be compared with dust filaments in the leading edge of NGC 4921 (Kenney, Abramson & Bravo-Alfaro 2015) with 0.5–1 kpc in length and 0.1–0.2 kpc in width. The dust filaments in the downstream are roughly aligned with the RPS direction, determined from the X-ray/ $H\alpha$ tails. This reflects the effect of RPS on dust ablation from dense ISM clouds. Most dust filaments around the nucleus and downstream have angles of $\sim 20^\circ$ from the major axis of

the galaxy, which is taken as the 2D wind angle. The orientation of these dust filaments may also be affected by magnetic field and disc rotation in ESO 137-002. The significance of these dust features also suggests that they mostly lie in the front side of the galaxy closer to us. The existence of long, coherent dust filaments suggests that magnetic field plays an important role here. If a similar dust-to-gas ratio as in the Milky Way is assumed (Güver & Özel 2009), the $H\text{I} + H_2$ column density of the strongest dust filament is $\sim 13.7 M_\odot \text{pc}^{-2}$ at the downstream region.

Interestingly, several dust clouds far from the major axis (e.g. one marked by a red arrow and an X-shaped feature in Fig. 6) are not near the disc mid-plane but are extraplanar on the ‘wrong side’ (upstream side), as the ICM wind has a west component. With an inclination angle of $\sim 83^\circ$ and a PA of $\sim -13.1^\circ$, the cloud marked by the red arrow in Fig. 6 is at ~ 0.8 kpc from the disc mid-plane. These clouds likely present evidence of gas fallback as predicted in RPS simulations (e.g. Tonnesen & Bryan 2012). Clouds are initially accelerated upwards from the disc. However, without achieving escape velocity, at least some will ultimately fall back, ending up momentarily on the upstream side of the disc. One expects dense clouds to be more prone to fallback since they are accelerated less with the same ram pressure.

To quantify the dust extinction in the galaxy, we derive the $E(B-V)$ map of the galaxy, with the F475W and F814W data. A key assumption in this analysis is that the F475W – F814W colour of the galaxy is symmetric around the nucleus. Since the upstream appears

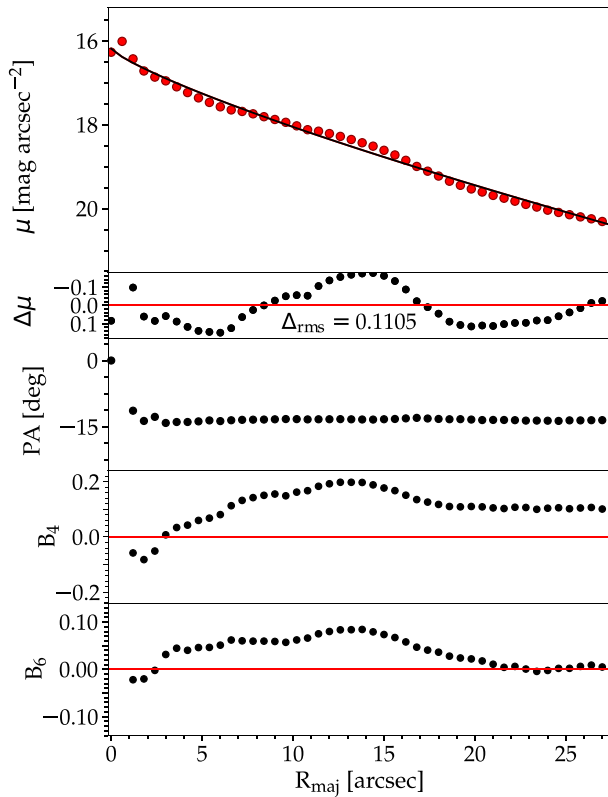


Figure 5. The results from PROFILER on the F160W image of ESO 137-002. The top panel shows the brightness profile (red circle) with the semimajor axis (R_{maj}). The profile is modelled with a single Sérsic component (black solid line) which yields a Sérsic index of 1.24 and a half-light radius of 4.36 kpc, in good agreement with the GALFIT result (see Table 4). Directly underneath is the residual profile followed by the PA profile, the fourth cosine harmonic amplitude B_4 and the sixth cosine harmonic amplitude B_6 profiles. The B_4 profile indicates the presence of a boxy bulge at $R_{\text{maj}} < 3$ arcsec while the B_6 profile confirms the existence of an X/peanut-shaped bulge.

dust-free (e.g. beyond 5 arcsec from the nucleus in Fig. 4), we can use the upstream colour as the colour in the mirror position of the downstream without intrinsic extinction. We employ the Voronoi binning method (Cappellari & Copin 2003) to adaptively bin the F475W image with a minimum S/N of 50 and also bin the F814W image with the same choice of bins. The new binned images were then converted to an $E(B-V)$ map using the equation

$$E(B-V) = -\frac{(F475W - F814W) - (F475W - F814W)_o}{k_{F475W} - k_{F814W}} \quad (1)$$

$F475W - F814W$ is the measured colour downstream, while $(F475W - F814W)_o$ is the measured colour in the corresponding upstream position, taken as the colour without the intrinsic extinction. $A_\lambda = k_\lambda E(B-V)$ is the extinction law. Since ESO 137-002 is a L_* spiral galaxy like our Milky Way, we use the extinction law of Fitzpatrick (1999) with $R_V = 3.1$. The results are shown in Fig. 6, with indeed enhanced extinction at the positions of dust clouds.

With this correction, it appears that the F275W light distribution becomes more symmetric between downstream and upstream regions. Thus, the SF is not strongly enhanced in the downstream regions, although some evidence of localized SF triggered by RPS is observed downstream as shown in the next subsection.

4.3 Triggered SF in the galaxy

While the galaxy is much fainter in F275W than in F475W (Figs 1 and 4), there are some compact knots and clumps revealed in F275W with blue colours as shown in Fig. 7. Several blue sources bright in F275W also have enhancement in the *SOAR* $H\alpha$ image, likely $H\text{II}$ regions. These are most likely young star clusters formed recently. Interestingly, all of these blue sources have dust clouds nearby and some of them are immediately upstream of dust clouds (Fig. 7). This association suggests that the star formation of at least some of these young star clusters (if not all) is triggered by the compression of ram pressure. Not all dust clouds have associated blue star clusters, which may reflect their different evolution stages.

On the other hand, these young star clusters may just shine in F275W because ram pressure has cleared dust around them, without actually triggering their formation. To test this scenario, we can move the F275W sources atop the cloud into the cloud to see whether they can still be detected. In principle, when a uniform dust cloud is put in front of a young star cluster and its surroundings, the brightness of the young star cluster and the surrounding r.m.s. will be reduced by the same factor so the significance of the young star cluster remains the same in this ideal situation. We further examined two sources, one in Figs 7(c) and another as the brightest one in (d). These two sources are detected at 3.8σ and 6.0σ in F275W, respectively. With the $E(B-V)$ map in the galaxy derived in Section 4.2, we can examine the expected significance of these two sources if moving them to the centre of their associated clouds. After adjusting the difference on extinction, these two sources would have been detected at 3.0σ and 4.3σ , respectively, in F275W at the centre of their associated clouds. As no UV-bright star clusters are detected around the centre of dust clouds (even those with low $E(B-V)$ values) and few of them are detected in regions completely free of dust, the observed correlation between some young star clusters and the dust clouds is better explained by the triggered SF from ram pressure. Similar RPS-triggered SF was also observed in e.g. NGC 4921 (Kenney et al. 2014), NGC 4402 and NGC 4522 (Abramson & Kenney 2014; Abramson et al. 2016), UGC 6697 (Consolandi et al. 2017), and IC 3476 (Boselli et al. 2021).

Most of these young clusters are downstream of the stripping front, as the upstream or north side of the galaxy has been cleared by ram pressure. The stripping front is defined by the $H\alpha$ and X-ray emission front as the region upstream of the stripping front is free of $H\alpha$ and X-ray emission, also with few dust features.¹ However, there are still some young star clusters upstream of the front, including several around the only remaining dust cloud and a blue stream just beyond the stripping front (the first two zoom-in boxes in Fig. 7). The large number of young star clusters immediately upstream of the stripping front also supports the scenario of SF triggered by ram pressure.

Source detection and aperture photometry were performed in F275W, F475W, and F814W with SExtractor, with an aperture radius of 0.5 arcsec. The ‘dual image mode’ was used, with the F275W data as the detection. The AB magnitude system is used in this work. Diagrams of the F275W – F475W colour versus the F475W – F814W colour, as well as the F275W – F475W versus the F475W magnitude for these sources, are shown in Fig. 8. As a comparison, we also measured the colours of the diffuse galactic disc detected in F275W, by defining 11 boxes along the major axis of the disc, each with

¹This is really a projected stripping front as the dimension of the front along the line of sight is substantial. Moreover, the front can be porous to allow ram pressure, at a direction somewhat tilted from the disc plane, to reach downstream regions.

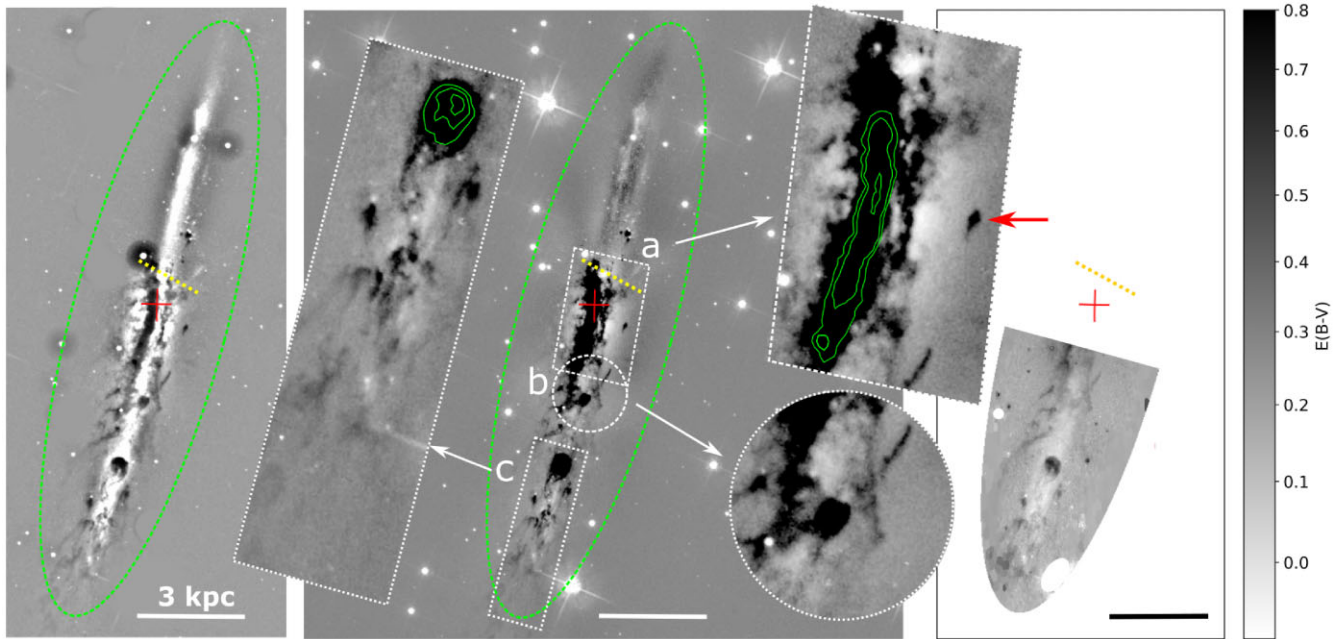


Figure 6. Left-hand panel: The unsharp masked F475W image of ESO 137-002, after subtracting the original image smoothed with 1 pixel (or 0.03 arcsec) from the same image smoothed with 20 pixels. Prominent dust lanes and filaments are visible which extend downstream from the nucleus. The red cross shows the position of the nucleus (the same for the other two panels) and the green ellipse shows the region for the $E(B-V)$ analysis (the same for the middle panel). The yellow dashed line is the same stripping front as shown in Fig. 3 (also the same for the other panels). Middle: the residual F475W image, after subtracting the GALFIT model image obtained with four Sérsic components with the Fourier mode (see Section 4.2). Similar dust features as shown on the left are also visible in this residual image and highlighted in zoomed images from a to c (the large and smooth dark features upstream are from oversubtraction of the model, instead of real dust features). Inset (a) shows the main dust lane around the nucleus, with the green contours show the levels of deficits. Inset (b) shows some dust filaments trailing the SE of the nuclear region. Inset (c) shows a large dust cloud (with a radius of ~ 0.22 kpc) in the middle of downstream (green contours also show the levels of deficits), as well as many dust clouds and filaments further downstream that can be traced up to ~ 13.5 kpc from the nucleus. The isolated dust cloud pointed by the red arrow in the inset a and the X-shaped feature in the inset (b) are likely clouds in fallback (see the text for discussion). Right-hand panel: The $E(B-V)$ map derived from the F475W and F814W images, assuming that the upstream region is dust free and an extinction law from Fitzpatrick (1999).

a size of $3.0 \text{ kpc} \times 2.6 \text{ kpc}$ excluding all unresolved sources. The nomenclature of these boxes has #1 as the northernmost one (most upstream), #6 at the nucleus, and #11 as the southernmost one. Their colours are also shown in Fig. 8. It is clear that the compact sources that are bright in F275W are bluer than the diffuse galactic disc. On the other hand, the upstream portion of the diffuse disc is bluer than the downstream portion, with the nuclear region the reddest. This is caused by the intrinsic extinction around the nucleus and downstream, as the upstream is nearly clear of gas and dust.

We can compare colours of these sources with simple stellar population (SSP) models. Here we use the Starburst99 (SB99 hereafter) model (Leitherer et al. 1999) with the Kroupa initial mass function (IMF) (Kroupa 2001) and Geneva 2012 zero rotation tracks (Levesque et al. 2012; Leitherer et al. 2014) to derive the track of star clusters, with a metallicity of 0.014 (or 0.7 solar, appropriate for ESO 137-002). Instantaneous SF is assumed. Pysynphot6 (STScI Development Team 2013; Blagorodnova et al. 2014) was used. Fig. 8 shows two tracks with and without intrinsic extinction overlaid on the sources.

As some young star clusters in the disc may be quite young, their nebular emission can be significant. To account for the contribution of the nebular lines into the broad *HST* bands, we employed the development version of the photoionization code Cloudy, last reviewed by Ferland et al. (2017), to add nebular emission to the stellar component of the radiation field reported by SB99. Cloudy does a full ab initio simulation of the emitting plasma, and solves self-consistently for the thermal and ionization balance of a cloud, while

transferring the radiation through the cloud to predict its emergent spectrum. We assumed a nebula of density 100 cm^{-3} , and metallicity of 0.7 solar, surrounding the stellar source and extending out to 1 kpc from it. For the inner radius of the cloud, we experimented with two values (1 pc and 10 pc), but found that the predicted colours do not depend on that choice. We also imposed a lower limit of 1 per cent on the electron fraction to let the calculation extend beyond the H II region, into the photo-dissociation region. The Cloudy modification is only important for star clusters younger than 10 Myr (Fig. 8) but does help to explain the blue colour of F475W - F814W for some sources.

As shown in Fig. 8, the blue sources highlighted in Fig. 7 have ages of several Myr to about 200 Myr, depending on the amount of intrinsic extinction. There is also no significant difference between sources upstream and downstream of the stripping front, which may suggest different timescales of SF after the initial ram pressure compression for different clouds, on top of different intrinsic extinction. With an age range of 3–100 Myr, the estimated total mass of these young star clusters are $\sim 4.3 \times 10^5 M_{\odot}$ and the corresponding SFR is $\sim 1.4 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$ which is significantly lower than the total SFR of the galaxy. When the intrinsic extinction with $E(B-V) = 0.25$ mag (still assuming the extinction law from Fitzpatrick 1999) is included, both the total mass and the SFR would increase by ~ 3 times.

In contrast, the galactic disc holds an older population of stars (500–1000 Myr) that evolved before the galaxy plunged into the dense ICM. It is also noted that no compact F275W sources

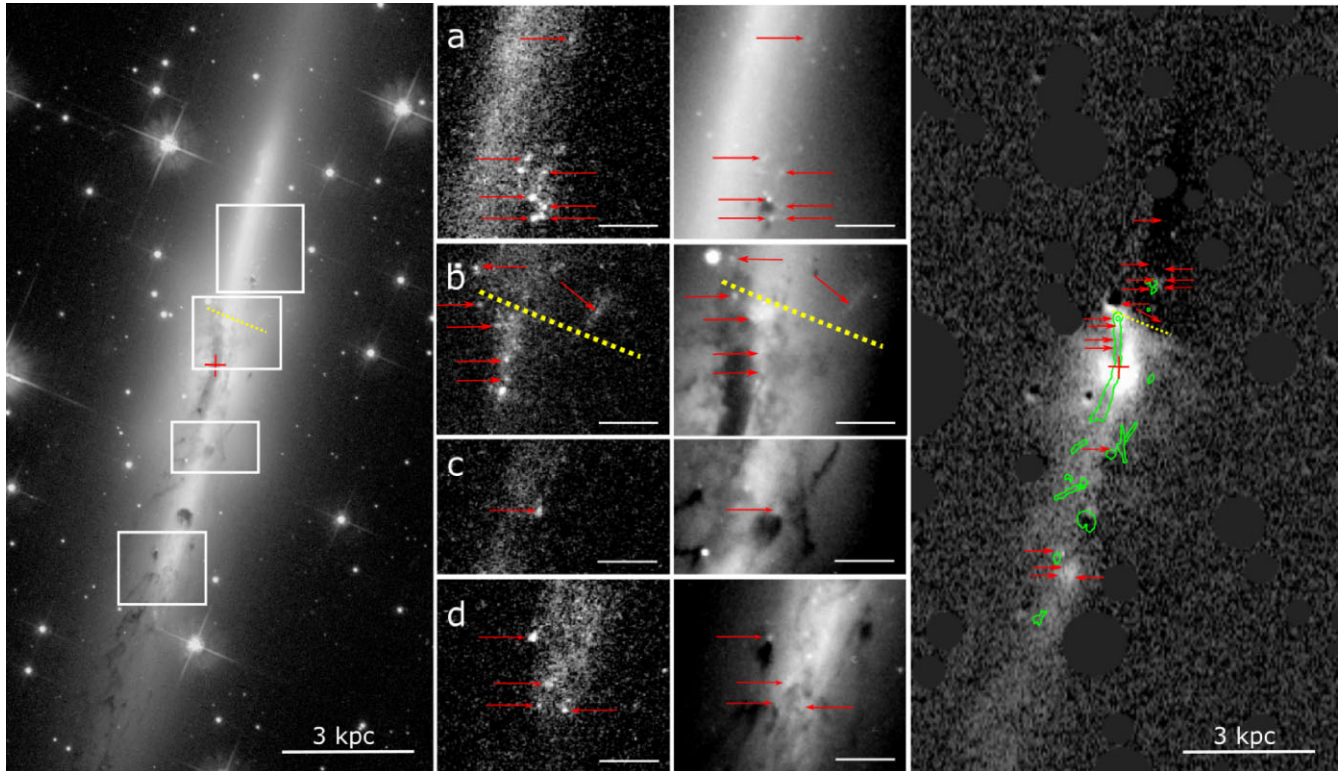


Figure 7. UV-bright young star clusters in the galaxy. Left-hand panel: The F475W image of ESO 137-002 with four white boxes to zoom in around young star clusters revealed in the F275W band. The stripping front is shown by the yellow dashed line (the same as in Fig. 3). Most F275W-bright compact sources are found around the stripping front. Middle: The four zoom-in regions (a, b, c, and d) to show the young star clusters in the disc (highlighted by red arrows) in the F275W (left-hand) and F475W (right-hand) images. These young star clusters are bright in F275W but faint in F475W and F814W. Some are located immediately upstream of dusty clouds, which implies that the SF there was triggered by the compression from ram pressure. The white scalebar in each small panel is 0.5 kpc. Right-hand panel: The net $H\alpha$ image of ESO 137-002 from *SOAR* (Sun et al. 2010), with regions around bright stars masked and also overlaid with contours of significant dust filaments in green. The same UV-bright sources marked in the middle panels are also marked here, as well as the stripping front and the nucleus. The upstream of the galaxy is also free of the $H\alpha$ emission. There is a faint $H\alpha$ source at the position of the young star cluster association just upstream of the stripping front as shown in the small panel a.

are detected beyond ~ 1.2 kpc upstream of the stripping front, which constrains the speed of the stripping front. If the age of the star cluster association in region a of Fig. 7 is 10–30 Myr, the stripping front would have moved with a speed of 120–40 km s^{-1} , if SF happens immediately after ram pressure compression. Since SF is almost certainly delayed after the initial compression, the actual speed of the front will be smaller.

4.4 Nucleus

ESO 137-002’s nucleus hosts an obscured, Seyfert2-like AGN from the X-ray data (Sun et al. 2010; Zhang et al. 2013). The nuclear position is determined from the *Chandra* data, at (16:13:35.763, $-60:51:54.74$) with an uncertainty of ~ 0.5 arcsec. The *HST* data reveal that the nuclear region is indeed heavily obscured by dust, without a nuclear point-like source in the optical and NIR (Fig. 1, Fig. 6, and Fig. 7). The average $E(B-V)$ values around the nucleus is ~ 0.5 .

5 *HST* SOURCE POPULATION IN THE TAIL

We also want to examine the source population in the X-ray/ $H\alpha$ tail and search for young star clusters as found in many ‘jellyfish’ galaxies. As shown in Fig. 2, we defined three regions of interest – tail, control, and galaxy regions. The tail region encloses the X-ray and $H\alpha$ tails as shown in Zhang et al. (2013). The choice of the

control region is mainly determined by the common FOV between the F275W and the F475W/F814W data. The galaxy region encloses sources in the galactic disc. The region area, after excluding bright stars, is 0.747, 4.97, and 0.281 arcmin^2 for the tail, control, and galaxy region, respectively.

Source detection and aperture photometry were again performed in all bands with SExtractor, with an aperture radius of 0.5 arcsec. The F475W data were used as the detection.² Before the final colour–colour and colour–magnitude diagrams, we filtered out spurious sources, weak sources, very red sources, and foreground stars with the following steps. Bright foreground stars are excluded with magnitude cuts of < 20.30 mag in F475W and < 20.06 mag in F814W, using GSC2. We also exclude sources with a colour error greater than one magnitude. Colour limits of $F275W - F475W > 5$ and $F475W - F814W > 3.5$ were applied to exclude very red sources. Saturated stars and artefacts are also excluded by visual inspection. Many sources detected in F475W are not detected in F275W so only the F275W - F475W limits are shown for those sources. In the end, there are 369 sources with 265 as upper limits in the control region and 38 sources with 18 as upper limits in the tail region. The colour–colour diagram ($F275W - F475W$ versus $F475W - F814W$) for all

²If the F275W data were taken as the detection, far fewer sources are detected but our conclusion stays the same.

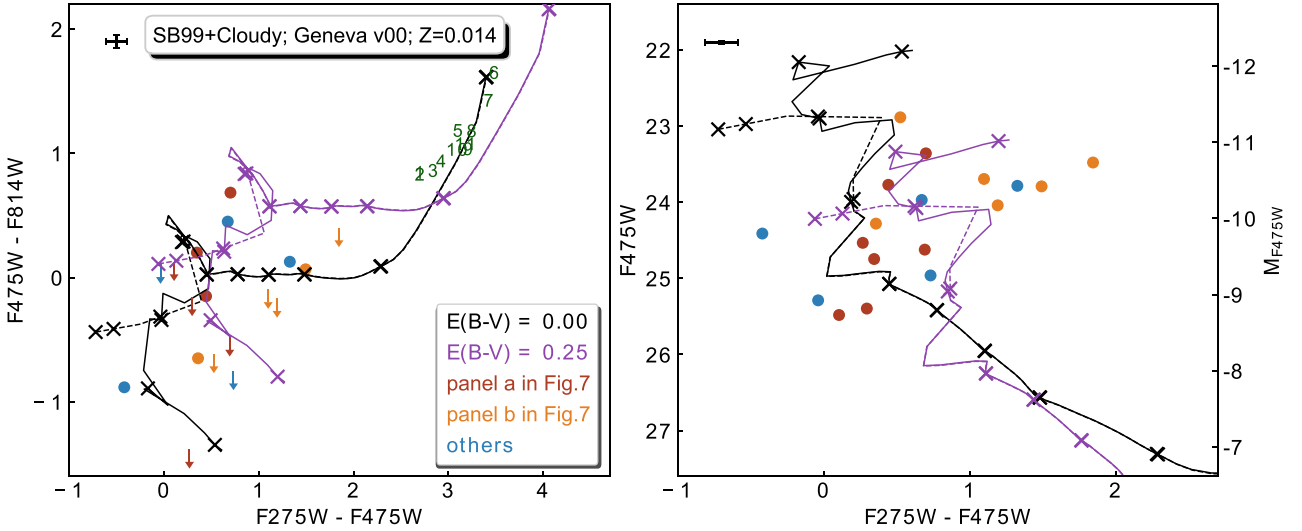


Figure 8. Left-hand panel: The colour–colour diagram for blue sources in the disc identified in Fig. 7, with the typical errors of colours are shown on the upper left-hand corner. Sources in the region b are generally redder than those in other regions, which may be at least partially due to the enhanced intrinsic extinction around the nuclear region. Sources in the region a and other regions (c and d) have similar colours. A SB99 model with the Geneva v00 track is shown by the dashed line for two values of the intrinsic extinction, $E(B-V) = 0$ and 0.25. We also ran Cloudy to include nebular emission that is important at ages of < 10 Myr, and added the nebular emission on the SB99 tracks as shown by the solid line. The markers on the track indicate ages of 1, 3, 5, 10, 30, 50, 100, 200, 500 and 1000 Myr, starting from the blue end (dashed and solid lines are the same at ages > 10 Myr). The galactic disc shown in the F275W image (Fig. 1) is divided into 11 boxy regions (a length of 3.0 arcsec \times a width of 2.6 arcsec), starting from #1 in the north (upstream) to #6 at the nucleus and #11 in the south (downstream). The colour in each box is shown by green number. It is clear that these blue sources selected in F275W are bluer than the diffuse galactic disc also defined in F275W. These tracks suggest that the blue sources have ages of several Myr to ~ 200 Myr, while the diffuse galactic disc has an age of 500–1000 Myr if fit with a single-burst model. Right-hand panel: The colour–magnitude diagram for the same sources with the absolute magnitude also shown on the right-hand side, with the typical errors of colours are shown on the upper left-hand corner. The same SB99 tracks are also shown (for a total mass of $10^4 M_{\odot}$) for two $E(B-V)$ values and with/without the nebular emission from Cloudy. The markers on tracks are also the same but only run to 500 Myr. From these tracks, the mass of these young star clusters is $10^3 - 10^4 M_{\odot}$ if younger than ~ 100 Myr.

detected sources with and without upper limits are shown in Figs 9 and 10, respectively. We also use 11 boxy regions (a length of 3.0 arcsec \times a width of 2.6 arcsec) along the galactic disc detected in the F275W image (Fig. 1) to study the colour of the galactic disc, which can be compared with the light and colour profiles in Fig. 4.

To explore whether there is an excess of blue sources in the tail, we compare the colour–colour and colour–magnitude diagrams for the sources in the tail and control regions. Particularly, histograms (the discrete form or the KDE form) for both colours and the F475W magnitude are compared. As shown in Figs 9 and 10, there is no evidence for any excess of blue sources (bright or faint) in the tail. We also show the colour–magnitude diagram for sources in the tail and control regions in Fig. 11. Again, there is no evidence for a significant excess of blue sources in the tail at any magnitude range. Thus, there is no evidence of active SF in the stripped tail of ESO 137-002, which is consistent with the lack of compact H α sources in the tail (Sun et al. 2010; Zhang et al. 2013).

We also plot the SB99+Cloudy tracks, one without intrinsic extinction and another with an intrinsic extinction of $E(B-V) = 0.161$ on the colour–colour and colour–magnitude diagrams in Fig. 10 and Fig. 11 to show the expected positions of young star clusters. This particular intrinsic extinction value comes from the median intrinsic extinction of $A_V = 0.5$ mag derived for GASP clumps from Poggianti et al. (2019) and the assumed R_V value of 3.1. Since our analysis suggests no excess of blue sources in the tail, we assume a F475W upper limit of 25.7 mag for any faint star clusters in the tail (Fig. 11). Based on the track in Fig. 11, we can put an upper limit on the mass of young star clusters in the tail at $\sim 2.1 \times 10^3 M_{\odot}$ at an age of 10 Myr when no intrinsic extinction is applied. For even younger star

clusters, the upper mass limit will be smaller. For star clusters with an age of 100 Myr, the upper mass limit becomes $\sim 1.3 \times 10^4 M_{\odot}$. When intrinsic extinction is applied, the upper mass limit becomes $\sim 4.2 \times 10^3 M_{\odot}$ at an age of 10 Myr and $\sim 2.6 \times 10^4 M_{\odot}$ at an age of 100 Myr, respectively.

6 CO EMISSION FROM THE GALAXY AND THE TAIL

The presence of the bright X-ray and H α emission over the centre and the southern side of the disc, and its absence in the northern disc side, suggest that (diffuse) ISM has been displaced by ram pressure. It is thus of interest to inspect how the dense ISM component has been affected. As we will show in this section, about $5.5 \times 10^9 M_{\odot}$ of molecular gas was detected in total, distributed asymmetrically in the galaxy and the inner tail.

6.1 Centre of ESO 137-002

We first searched for CO emission in the 002-C pointing covering the central part of the disc, where the gravitational restoring force is the strongest. CO(2-1) emission is strongly detected there (see Fig. 12 and Table 5). We calculate the CO luminosity from the standard relation of Solomon & Vanden Bout (2005)

$$L'_{\text{CO}} = 3.5 \times 10^7 S_{\text{CO}} \Delta v v_{\text{obs}}^{-2} D_L^2 (1+z)^{-3}, \quad (2)$$

where L'_{CO} is the CO(2-1) line luminosity in $\text{K km s}^{-1} \text{pc}^2$, $S_{\text{CO}} \Delta v$ is the CO velocity integrated line flux in Jy km s^{-1} , v_{obs} is the CO(2-1) line observed frequency, and D_L is the distance in Mpc. The

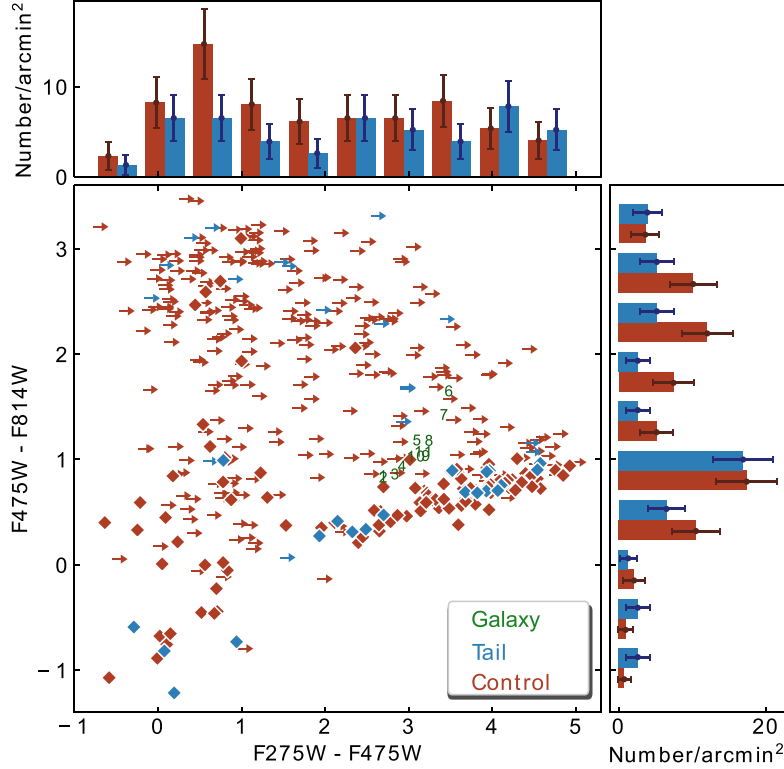


Figure 9. The colour–colour diagram for sources detected in F275W, F475W, and F814W (diamonds) and sources only detected in F475W and F814W (arrows), in tail (blue) and control (red) regions. The initial detection was made in F475W. Most sources are undetected in F275W so F275W – F475W colours are mostly limits. The tail/control regions are defined in Fig. 2. The source selection is discussed in Section 5. Colours of the galactic disc are again shown in green numbers as in Fig. 8. The median F475W – F814W errors are 0.02 mag and 0.06 mag for the tail and control regions, respectively. The histograms are normalized to the region area (0.747 arcmin² for the tail region and 4.97 arcmin² for the control region). Error bars of the source counts are also plotted. As the distributions of both colours in the tail region are consistent with those in the control region, there is no evidence of young star clusters in the tail region.

CO luminosity in the 002-C pointing is $\sim 8.2 \times 10^8 \text{ K km s}^{-1} \text{ pc}^2$. Following

$$M_{\text{H}_2} [M_{\odot}] = 5.4 L'_{\text{CO}(2-1)} [\text{K km s}^{-1} \text{ pc}^2], \quad (3)$$

where we assume a standard Galactic CO/H₂ conversion factor $\alpha_{\text{CO}} = 4.3 M_{\odot}/(\text{K km s}^{-1} \text{ pc}^2)$ (e.g. Leroy et al. 2009; Kennicutt & Evans 2012), and a typical value for the CO(2-1)/CO(1-0) ratio of 0.8, the corresponding molecular gas mass in the 002-C region is about $4.4 \times 10^9 M_{\odot}$. The formula includes a factor of 1.36 to account for the contribution of Helium. We assume a typical CO(3-2)/CO(1-0) ratio of 0.6.

A Gaussian fit to the 002-C CO(2-1) line (shown in Fig. 12) has a width of $\sim 430 \text{ km s}^{-1}$ (Table 5) which is consistent with ESO 137-002’s circular velocity of $\sim 200 \text{ km s}^{-1}$ that is estimated from the galaxy’s *K*-band luminosity and using the velocity–luminosity relation of Courteau et al. (2007). While the CO line seems to have a classical two-horned profile caused by the disc rotation, it is asymmetric with (a) more pronounced low-velocity side and (b) strong peak in the middle of the line. The former is likely due to pointing the APEX 002-C aperture about 1.4 kpc southwards off the (optical) galaxy centre where it encloses more gas from the southern disc side that rotates towards the observer, i.e. at smaller radial velocities. The latter may suggest the existence of a gas with a low radial velocity component relative to the galaxy, as the peak occurs at the galaxy’s systemic radial velocity of 5691 km s^{-1} . Its origin is likely in a compact central source, such as a circumnuclear disc (e.g. Israel, van der Werf & Tilanus 1999). Many observations

have shown that molecular gas is more centrally concentrated in galaxies with bars, suggesting that a bar can play an important role in gas-fueling toward the central region, thus feeding starbursts or AGN activities (Kuno et al. 2007). The average central gas surface density of barred spirals is known to be about three times higher than that of unbarred galaxies (Sheth et al. 2005). ESO 137-002 indeed has a bar as discussed in Section 4.1 and it hosts a Seyfert-2-like nucleus (Sun et al. 2010; Zhang et al. 2013).

6.2 Asymmetric CO distribution in the disc

To cover the continuation of the bright X-ray and H α emission in the southern part of the disc, the 002-S APEX beam was pointed $\sim 7.8 \text{ kpc}$ south from the optical disc centre, adjacent to the 002-C region (see Fig. 2). A rather strong CO(2-1) line is detected with an S/N ~ 12 (see Fig. 12), defined as $S/N = I_{\text{CO}}/(\sigma_{\text{rms}} \sqrt{\Delta v_{\text{CO}} dv})$, i.e. a ratio of the integrated intensity and the noise integrated over the channels covered by the line, where Δv_{CO} is the channel width and dv the integration range approximated by the FWHM of the line. The corresponding molecular gas mass is $\sim 1.0 \times 10^9 M_{\odot}$. The linewidth measured from a Gaussian fit is $\sim 302 \text{ km s}^{-1}$, by about 130 km s^{-1} less than in the central 002-C pointing. The line is asymmetric, with more emission coming from the low-velocity side, which is consistent with detecting the gas that is rotating towards the observer. The high-velocity wing of the line exceeds the galaxy’s systemic radial velocity, suggesting that the kinematics of the gas in the southern disc part may be altered by the effects of

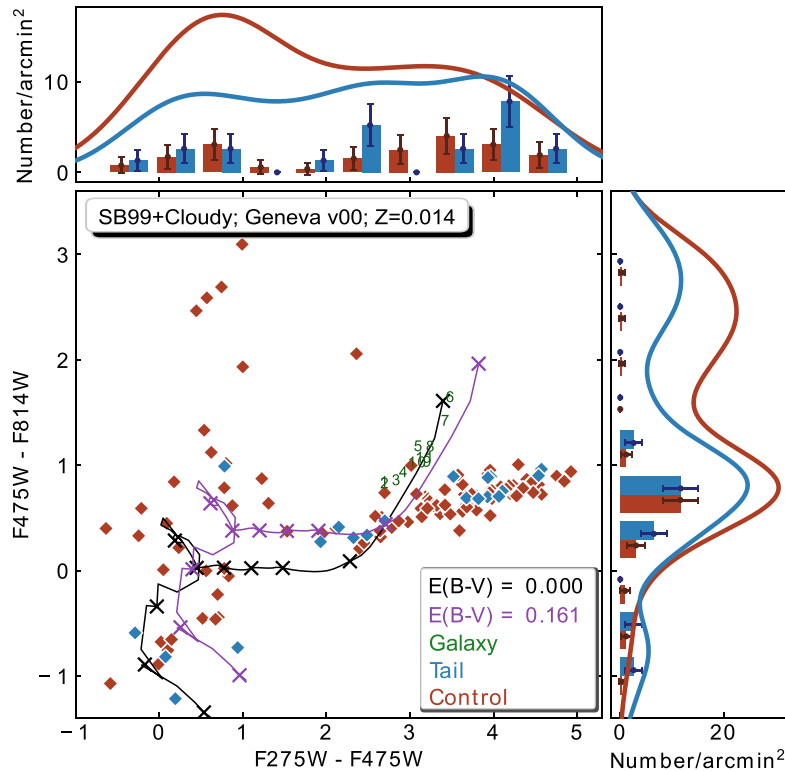


Figure 10. Similar to Fig. 9 but only for detections in these three bands. Colours of the galactic disc are again shown in green numbers as in Fig. 8. The median $F275W - F475W$ error is 0.27 mag and 0.31 mag for the tail and control region, respectively. The median $F475W - F814W$ errors are 0.01 mag in both regions. An evolutionary track of a star cluster, after an initial starburst, is derived from SB99, with the nebular emission added from Cloudy simulations and plotted. The specific track is the Geneva V00 track. The markers on the track indicate ages of 1, 3, 5, 10, 30, 50, 100, 200, 500, and 1000 Myr, starting from the blue end. A track with an intrinsic extinction of $E(B-V) = 0.161$ is also shown. The source histogram in each region, normalized to the region area, is again shown. For comparison, the Gaussian kernel density estimation (KDE) of the histogram in Fig. 9 is also shown. Again there is no evidence of active SF in the tail region.

ram pressure. However, contamination by the secondary beam lobes that may register emission from the receding (northern) disc side is also possible.

We further observed the region 002-N located in the northern, upstream disc side, symmetrically to the southern 002-S pointing with respect to the optical centre of the disc (see Fig. 2). No obvious line emission appears in the 002-N spectrum. From the measured rms noise σ_{rms} we calculate 3σ upper limit on the CO flux density

$$S_{\text{CO}} < 3\sigma_{\text{rms}} \sqrt{\Delta v_{\text{CO}} d\nu}, \quad (4)$$

where $d\nu$ is the channel size and Δv_{CO} is the mean FWHM linewidth. Assuming a linewidth of 300 km s^{-1} , the rms corresponds to a 3σ mass sensitivity of $\sim 2.1 \times 10^8 M_{\odot}$. In the northern, windward part of the disc of ESO 137-002, there is thus at least 5-times less molecular gas than in its southern, downstream part. This indicates a strong asymmetry in the molecular component distribution in the disc, likely due to the effects of ram pressure stripping.

6.3 Inner tail of ESO 137-002

We also observed the inner tail region (002-T) (see Fig. 2). The integration was deep to reach an rms sensitivity of $\sim 0.7 \text{ mK}$ in 51 km s^{-1} channels. Fig. 13 shows the spectrum where the CO line is detected with an ‘integrated’ $S/N > 6$. The FWHM of the line is considerably smaller than in the 002-S region - only $\sim 149 \text{ km s}^{-1}$ (by about a factor of 2.2). The line luminosity corresponds to an H_2 mass of $\sim 2.2 \times 10^8 M_{\odot}$, and the integrated intensity corresponds

to a column density of $\sim 2M_{\odot} \text{ pc}^{-2}$ averaged over the beam. At a higher resolution, the emission would likely mostly split into a set of compact regions with higher column densities. The observed linewidth of $\sim 150 \text{ km s}^{-1}$ could correspond to a superposition of many molecular clouds with typical values of the velocity dispersion of $\sim 5-15 \text{ km s}^{-1}$. Dust features, which are presumably good tracers of molecular gas, show mostly compact morphologies in the *HST* image. At the same time, previous observations of other RPS galaxies indicated from comparison of interferometric and single-dish fluxes that there might be an important extended component of the CO emission on top of a rich distribution of compact CO sources (e.g. Jáchym et al. (2019) in ESO 137-001 or Moretti et al. (2020) in JW100). It is thus possible that in ESO 137-002, a certain fraction of the molecular gas is also in the form of a more diffuse component. However, these observations suggested that the fraction of the extended emission increases with the distance along the tail and is low in the main galaxy and the inner parts of the RPS tails.

In another Norma cluster galaxy, ESO 137-001, large amounts of molecular gas were detected in the tail. Since ESO 137-002 is stripped (and observed) edge-on, the tail overlies the disc, before it extends from the galaxy. In a face-on stripping configuration, the tail is extra-planar. The 002-T region thus should be compared to the 001-tailB region in ESO 137-001, where about 3-times more H_2 was detected (Jáchym et al. 2014). The CO line is centred at a velocity close to the galaxy’s systemic velocity ($\Delta v \sim 40 \text{ km s}^{-1}$). This may indicate that a substantial fraction of the orbital velocity of

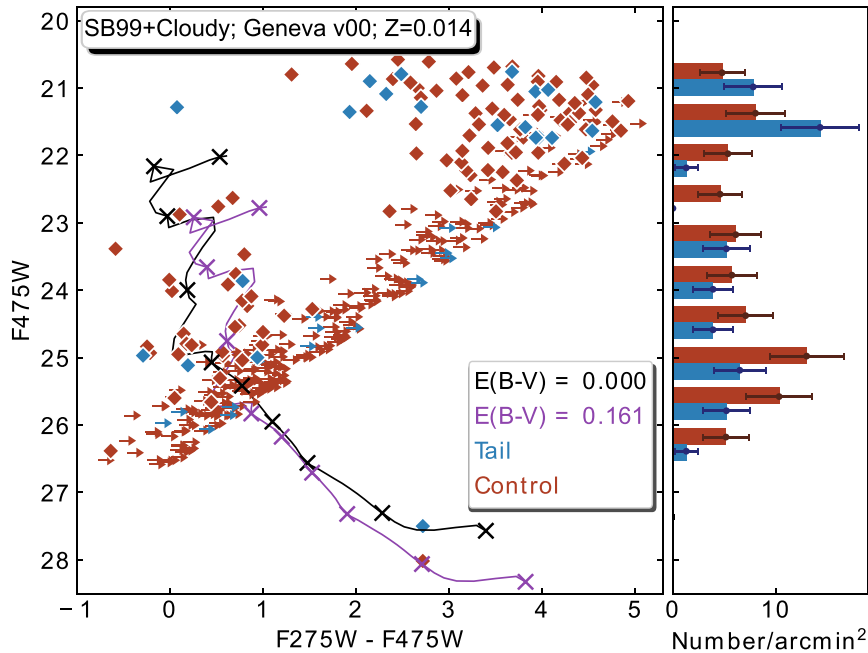


Figure 11. The F275W – F475W colour – F475W magnitude diagram for sources detected in F275W, F475W, and F814W (diamonds) and sources only detected in F475W and F814W (arrows), in tail (blue) and control (red) regions. The SB99 tracks (for a total mass of $10^4 M_{\odot}$), with the nebular emission added from Cloudy simulations, are also shown with $E(B-V) = 0$ and 0.161. The markers on the track are same as in Fig. 10. A histogram of the source surface number densities in these two regions is shown on the right, which suggests no significant difference in the F475W magnitude distribution between sources in the tail and control regions. The somewhat higher number of bright sources in the tail region are almost all red ones (e.g. F275W – F475W > 1 mag) and the only bright blue tail source is likely a Galactic star for its lack of H α emission and bright NIR emission from F160W. Combined with Fig. 9 and Fig. 10, the source population in the tail region has no significant difference from that in the control region and there is no evidence for active SF in the tail region.

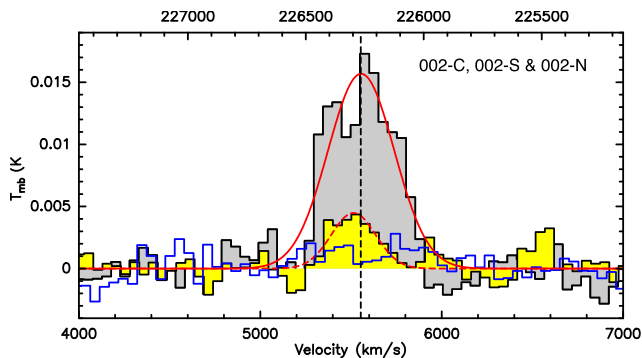


Figure 12. The CO(2-1) spectra of ESO 137-002 in the central pointing 002-C (grey), the southern off-centre position 002-S (yellow), and the symmetrical northern position 002-N (blue) (see Fig. 2 for the positions of CO regions). Spectra are smoothed to 50.8 km s^{-1} channels. Parameters of the Gaussian fits to the 002-C and 002-S lines are given in Table 5. The central velocity 5554 km s^{-1} of the fit in central point is shown with the dashed line, while the systemic radial velocity of ESO 137-002 is 5585 km s^{-1} (both assuming radio definition of the velocity and the latter velocity value corresponds to 5691 km s^{-1} in the optical definition of the velocity).

the galaxy is happening in the plane of the sky, and the stripped gas has an important tangential velocity component.

6.4 CO(3-2) emission and CO(3-2)-to-CO(2-1) line ratios

The central and southern regions of the ESO 137-002 disc were further observed in CO(3-2). The locations of the two pointings were slightly offset with respect to the CO(2-1) regions 002-C and 002-S

in order to better cover the asymmetric X-ray and H α emission with the smaller CO(3-2) beams (see Table 3 and Fig. 2). The corresponding CO luminosities and H $_2$ masses are given in Table 5. In the central region the CO(3-2) line is double-horned with some emission also at the central line velocity, similarly to the CO(2-1) 002-C line profile (see Fig. 14). The CO(3-2) lines in the central and southern regions have slightly larger FWHM than the corresponding CO(2-1) lines. The difference, however, is minimum and well within the error of the Gaussian line fits. Other factors may play a role, such as noise, differences in beam sizes, beam offsets, and possible source extension.

Following equation (2), the line luminosity ratios, $r_{32} = L'_{\text{CO}(3-2)}/L'_{\text{CO}(2-1)}$ in the two regions are 0.51 and 0.43, not corrected for the different beam sizes (~ 27 arcsec versus 17 arcsec). The galaxy is edge-on and it can be expected that most of the molecular gas is concentrated in the disc plane. Also, the ram pressure operating on the galaxy is nearly edge-on. The difference in the beam sizes in the direction perpendicular to the disc-plane thus does not play much role and it is mostly the difference along the disc plane that matters. We thus simply apply a 1D geometrical beam correction factor of $26.7 \text{ arcsec}/17.8 \text{ arcsec} = 1.5$ that increases the r_{32} line ratios in the 002-C-CO32 and 002-S-CO32 regions to 0.77 and 0.65, respectively. A more precise way for calculating a beam correction factor would be to follow the assumption that the IR dust emission is a good tracer of the cold molecular gas and measure the flux within different apertures in the IR images. But this approach is beyond the scope of this paper.

In low-redshift normal star-forming galaxies, typical mean values of the CO line ratios are $r_{31} = 0.55 - 0.61$ (Mao et al. 2010; Saintonge et al. 2017), and $r_{21} = 0.79 - 0.81$ (Leroy et al. 2009; Saintonge et al. 2017). This translates to the typical r_{32} ratio of ~ 0.7 . Our

Table 5. APEX results: The 1σ rms, the parameters of single Gaussian fits, the measured integrated intensity, and molecular gas mass are given for CO(2-1) emission detected in the observed positions. Temperatures are given in T_{mb} scale. First order (third order for 002-S) baselines were subtracted in the velocity range 4000–7200 km s^{-1} .

| Source | Line | rms ^a (mK) | Velocity ^b (km s^{-1}) | FWHM (km s^{-1}) | T_{peak} (mK) | $I_{\text{CO,fit}}$ (K km s^{-1}) | L_{CO} ($10^8 \text{ K km s}^{-1} \text{ pc}^2$) | M_{mol} ($10^8 M_{\odot}$) |
|------------|---------|--------------------------|---|--------------------------------|---------------------------|---|--|--|
| 002-C | CO(2-1) | 1.7 | 5659.3 ± 8.8 | 429.1 ± 17.8 | 15.7 | 7.2 ± 0.3 | 8.2 | 44.3 |
| 002-C-CO32 | CO(3-2) | 3.1 | 5687.9 ± 14.4 | 443.0 ± 29.8 | 16.7 | 7.9 ± 0.5 | 4.21 | 30.2 |
| 002-S | CO(2-1) | 1.4 | 5617.4 ± 23.1 | 301.7 ± 49.3 | 5.0 | 1.6 ± 0.2 | 1.8 | 9.8 |
| 002-S-CO32 | CO(3-2) | 1.7 | 5681.9 ± 25.2 | 321.0 ± 56.0 | 4.3 | 1.46 ± 0.23 | 0.78 | 5.6 |
| 002-N | CO(2-1) | 1.2 | – | ~300 | – | <0.32 ^c | <0.39 ^c | <2.1 ^c |
| 002-T | CO(2-1) | 0.9 | 5651.7 ± 21.3 | 148.9 ± 39.7 | 2.2 | 0.35 ± 0.09 | 0.4 | 2.2 |

Note. ^a in 25 km s^{-1} channels (or 27 km s^{-1} channels for CO(3-2) lines); ^b heliocentric central velocity (optical definition); ^c upper limit for $3\sigma_{\text{rms}}$

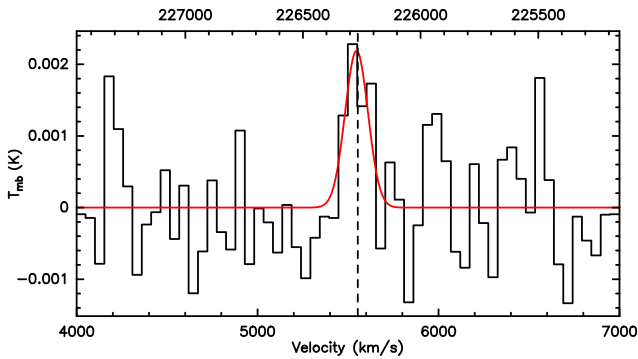


Figure 13. The CO(2-1) spectrum measured in the (inner) tail of ESO 137-002 (002-T in Fig. 2). Spectrum is smoothed to $\sim 51 \text{ km s}^{-1}$ channels. Parameters of the Gaussian fits are given in Table 5. The central velocity of the line in the 002-C region is indicated with the dashed vertical line.

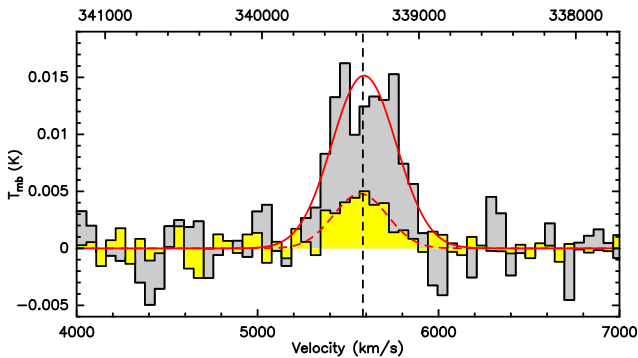


Figure 14. The CO(3-2) spectra measured in the central (grey) and southern (yellow) positions over ESO 137-002 (see positions in Fig. 2). Spectra are smoothed to $\sim 50 \text{ km s}^{-1}$ channels. Parameters of the Gaussian fits are given in Table 5. The central velocity of the line in the 002-C region is indicated with the dashed vertical line. The spectrum of the central position is double-peaked with a velocity separation of $\sim 300 \text{ km s}^{-1}$.

measurements in the two regions in ESO 137-002 are thus consistent within the uncertainties with the typical values in local star-forming galaxies. The line luminosity ratio can be interpreted as an indicator of the gas density. The critical density of the CO(3-2) transition calculated under the optically thin assumption is $3.6 \times 10^4 \text{ cm}^{-3}$. The r_{32} line ratio may thus be considered a proxy for the ratio of relatively dense to more diffuse molecular gas. Many studies have searched in galaxies for correlations between the CO line ratios and star formation rates or star formation efficiencies (SFE), finding that the r_{31} ratio tends to increase with SFE, suggesting that galaxies with

a higher fraction of dense molecular gas tend to have higher SFE (e.g. Lamperti et al. 2020). Despite a large scatter in the correlation (see fig. 3 in Lamperti et al. 2020), it is clear that the values of the line ratios measured in the two regions in ESO 137-002 would in a normal star-forming galaxy correspond to a larger SFE. In other words, despite its large amounts of molecular gas and the relatively large fraction of the denser molecular gas traced by the CO(3-2) emission, ESO 137-002 has a rather low SFE: $\text{SFE} = \text{SFR}/M_{\text{H}_2} \approx 0.89 \text{ M}_{\odot} \text{ yr}^{-1} / 5.5 \times 10^9 \text{ M}_{\odot} \sim 1.6 \times 10^{-10} \text{ yr}^{-1}$. The effects of ram pressure of the surrounding ICM, as well as the active nucleus (i.e. a source of X-rays) of the galaxy, may be responsible for its measured low SFE.

7 DISCUSSION

7.1 Star formation in the galaxy

The total FIR luminosity of the galaxy was derived from the *Herschel* data. We used the *Herschel* source catalogue, particularly the PACS Point Source Catalog and the SPIRE Point Source Catalogue. ESO 137-002 is significantly detected in all six *Herschel* bands. We used the python code MBB_EMCEE to fit modified blackbodies to photometry data using an affine invariant Markov chain Monte Carlo (MCMC) method, with the *Herschel* passband response folded (Dowell et al. 2014). Assuming that all dust grains share a single temperature T_{d} , that the dust distribution is optically thin, and neglecting any power-law component towards shorter wavelengths, the fit results in a temperature of $T_{\text{d}} = (27.8 \pm 0.2) \text{ K}$, a luminosity $L_{8-1000 \mu\text{m}} = (1.22 \pm 0.02) \times 10^{10} L_{\odot}$, and a dust mass of $M_{\text{d}} = (6.7 \pm 0.3) \times 10^6 M_{\odot}$ for $\beta = 1.5$. For $\beta = 2$, $T_{\text{d}} = (24.7 \pm 0.2) \text{ K}$, $L_{8-1000 \mu\text{m}} = (1.21 \pm 0.02) \times 10^{10} L_{\odot}$ and $M_{\text{d}} = (1.2 \pm 0.1) \times 10^7 M_{\odot}$. We also performed a fit using a free β parameter, which results in $T_{\text{d}} = 26.6^{+0.9}_{-0.7} \text{ K}$, $L_{8-1000 \mu\text{m}} = (1.22 \pm 0.02) \times 10^{10} L_{\odot}$, $M_{\text{d}} = 8.7^{+0.9}_{-1.6} \times 10^6 M_{\odot}$, and $\beta = 1.69^{+0.09}_{-0.14}$. With the measured FIR and CO luminosities, we can put ESO 137-002 on the well-known CO–FIR correlation (e.g. Solomon & Vanden Bout 2005) and its position is consistent with normal spirals on the relation.

The total SFR of ESO 137-002 is $1.08 \text{ M}_{\odot} \text{ yr}^{-1}$, from the *Galex* NUV flux density and the total *Herschel* FIR luminosity from the relation of Hao et al. (2011). If using the *WISE* 22 μm flux density and the relation from Lee et al. (2013), the estimated total SFR is $0.80 \text{ M}_{\odot} \text{ yr}^{-1}$. The Kroupa IMF is assumed in both cases. The Lee et al. (2013) work assumed the Salpeter IMF so we multiply its SFR relation by 0.62 to convert to the Kroupa IMF. For the total amount of molecular gas in the galaxy, $\sim 5.5 \times 10^9 \text{ M}_{\odot}$, the molecular gas depletion time is $\sim 5.9 \text{ Gyr}$.

ESO 137-002 is ~ 5 times more massive than ESO 137-001 but with a comparable SF activity. The specific SFR of ESO 137-002 is ~ 7 times less than that of ESO 137-001. This is also consistent with their different $W1 - W4$ colours, 4.25 for ESO 137-002 versus 6.99 for ESO 137-001, as $W1 - W4$ is a proxy for the specific SFR. Because of its prominent bulge and higher surface brightness in optical and NIR, ESO 137-002 is in fact a more compact galaxy than ESO 137-001.

The upstream part of the galaxy is nearly dust-free and gas free (Sun et al. 2010; Zhang et al. 2013) so the current SF is mainly around the nucleus and the downstream. Fig. 9 and Fig. 10 show that there is little SF in the tail. However, the disc hosts some blue star clusters with recent SF most likely triggered by ram pressure (Fig. 7). The colour–colour diagram (Fig. 8) shows that these star clusters are indeed young, although the total SFR of these young star clusters is very small compared with the total SFR of the galaxy.

7.2 Star formation efficiency in the tail

Sun et al. (2010) presented the $H\alpha$ data of ESO 137-002, from the narrow-band imaging observations with *SOAR*. While the *SOAR* $H\alpha$ observations are not deep (three exposures of 800 s each), the data were taken under the excellent seeing condition (0.45–0.65 arcsec). The continuum data were taken from another narrow-band filter centred at $\sim 87\text{\AA}$ bluewards the $H\alpha$ filter. Similar to what we did for ESO 137-001 with the *SOAR* data in Sun et al. (2007), the source colours from the two narrow band filters are examined to select emission-line sources in the X-ray/ $H\alpha$ tail of ESO 137-002. None is detected and the $3-\sigma$ upper limit of the $H\alpha$ luminosity on the compact sources is $7 \times 10^{37} \text{ erg s}^{-1}$, within an aperture radius of 0.6 arcsec and assuming $[N II] \lambda 6584 / H\alpha = 0.2$ and $[N II] \lambda 6548 / [N II] \lambda 6584 = 1/3$. This limit is four times smaller than the median $H\alpha$ luminosity of $4 \times 10^{38} \text{ erg s}^{-1}$ for over 500 $H\alpha$ clumps in the tails of GASP galaxies (Poggianti et al. 2019), even after the correction for the intrinsic extinction (with the median A_V of 0.5 mag from Poggianti et al. 2019). This limit is also smaller than ~ 90 per cent of $H II$ regions detected by *SOAR* in ESO 137-001’s tail (Sun et al. 2007). With the SFR- $H\alpha$ calibration from Hao et al. (2011) for the Kroupa IMF, the corresponding SFR limit is $3.9 \times 10^{-4} M_\odot \text{ yr}^{-1}$. Thus, the SF in ESO 137-002’s tail is weak at most.

The *HST* data also do not reveal any significant young star clusters in the tail, even with its superior angular resolution. Young star clusters with ages of less than 10 Myr have strong $H\alpha$ emission and will be detected as $H II$ regions. Observationally, compact UV emission that traces young star clusters correlates well with compact $H\alpha$ emission that traces $H II$ regions (e.g. George et al. 2018). With the *HST* data, we set an upper limit of mass of $M_* \approx 4.2 \times 10^3 M_\odot$ (with $A_V = 0.5$ mag intrinsic extinction) at an age of 10 Myr. The corresponding upper limit on the SFR in the tail is,

$$\text{SFR}_{\text{tail}} < \frac{4.2 \times 10^3 M_\odot}{1 \times 10^7 \text{ yr}} < 4.2 \times 10^{-4} M_\odot \text{ yr}^{-1} \quad (5)$$

Note that the resulting SFR limit would be smaller if using an age of 100 Myr and assuming no intrinsic extinction. The inner tail (002-T) of ESO 137-002 from the *APEX* data corresponds $\sim 2.2 \times 10^8 M_\odot$ of H_2 . With the larger SFR limit derived earlier, we estimate the SFE as,

$$\text{SFE}_{\text{tail}} < \frac{4.2 \times 10^{-4} M_\odot \text{ yr}^{-1}}{2.2 \times 10^8 M_\odot} < 1.9 \times 10^{-12} \text{ yr}^{-1} \quad (6)$$

Taking the reciprocal of the SFE, the gas depletion time in 002-T is 526 Gyr. The SFE upper limit is more than 19 times lower compared

to the ESO 137-001 tail (SFE $\approx 3.6 \times 10^{-11} \text{ yr}^{-1}$) (Jáchym et al. 2014). It is also smaller than the SFE in D100’s tail ($\sim 6.0 \times 10^{-12} \text{ yr}^{-1}$). The SFE typically falls with the distance from the main body to the outer tail as in the case of other RPS galaxies (e.g. Boissier et al. 2012; Jáchym et al. 2014).

7.3 Stripping history of ESO 137-002

We can compare the amount of gas at different phases (cold molecular, warm ionized, and X-ray emitting) detected in the galaxy and its tail.

(i) Cold molecular gas

We detected $\sim 5.5 \times 10^9 M_\odot$ of the cold molecular gas from the galaxy and at least $\sim 2 \times 10^8 M_\odot$ from the tail. Deeper observations plus wider coverage of the tail may reveal more cold molecular gas there. For the total stellar mass and the sSFR of ESO 137-002, the cold gas fraction of the galaxy is ~ 20 per cent from the scaling relations by Boselli et al. (2014) and Catinella et al. (2018), although the big scatter can change the result by a factor of several. This fraction translates to $\sim 7 \times 10^9 M_\odot$ of the cold gas initially in the galaxy. Thus, ESO 137-002 may still retain most of its initial cold ISM, especially since the cold atomic gas is not included yet for the lack of sensitive $H I$ data. This would suggest that the stripping in ESO 137-002 is likely still in the early stage, in term of the ISM content. On the other hand, the distribution of the ISM in the galaxy may have been significantly modified, as shown also by the dust distribution.

(ii) Warm ionized gas

Zhang et al. (2013) gave a total mass of $\sim 1.9 \times 10^7 M_\odot$ for the $H\alpha$ emitting gas in the tail, if a filling factor of 0.2 per cent is adopted. Even if the filling factor is 10 times higher, the mass only increases by 3.2 times. The mass of the warm ionized gas in the galaxy is hard to estimate, for the contribution from star forming regions and intrinsic extinction, but it is unlikely to be comparable to the mass of the cold molecular gas.

(iii) Hot X-ray emitting gas

Zhang et al. (2013) estimated a total mass of $\sim 2 \times 10^8 M_\odot$ for the X-ray emitting gas in the tail, if the filling factor is close to one. The amount of the X-ray emitting gas in the galaxy is less due to its smaller volume.

Beside the total gas content, this work also reveals a strong asymmetry in the distribution of the CO gas in the disc, similar to that of the $H\alpha$ and X-ray emitting gas. While the galaxy may still retain a significant portion of its initial cold ISM, the ISM distribution has been significantly modified by ram pressure and narrow stripping tails have been formed behind the galaxy.

As discussed in Sun et al. (2010) and Zhang et al. (2013), ESO 137-002 is experiencing a near edge-on stripping. This is also supported by the *HST* images of dust filaments in the disc, mostly aligned with the disc plane. The positions of young star clusters and their associated dust clouds also suggest near edge-on stripping. On the other hand, the orientations of most dust filaments, the $H\alpha$ and X-ray tails suggest a small projected wind angle (W_{2D}) of $\sim 20^\circ$ between the ICM wind direction and the disc plane (Fig. 15). Since much of the wind is projected in the plane of sky, it creates vertical distribution of the ISM. ESO 137-002 also has a significant line-of-sight motion with the cluster, so there must be a wind component perpendicular to the plane of the sky. For such case, the 3D disc-wind angle (W_{3D}) between the galactic disc and the ICM wind direction is less than 20° (using the technique of Abramson et al. 2011). This low inclined angle likely has a contribution to pushing the disc ISM from the

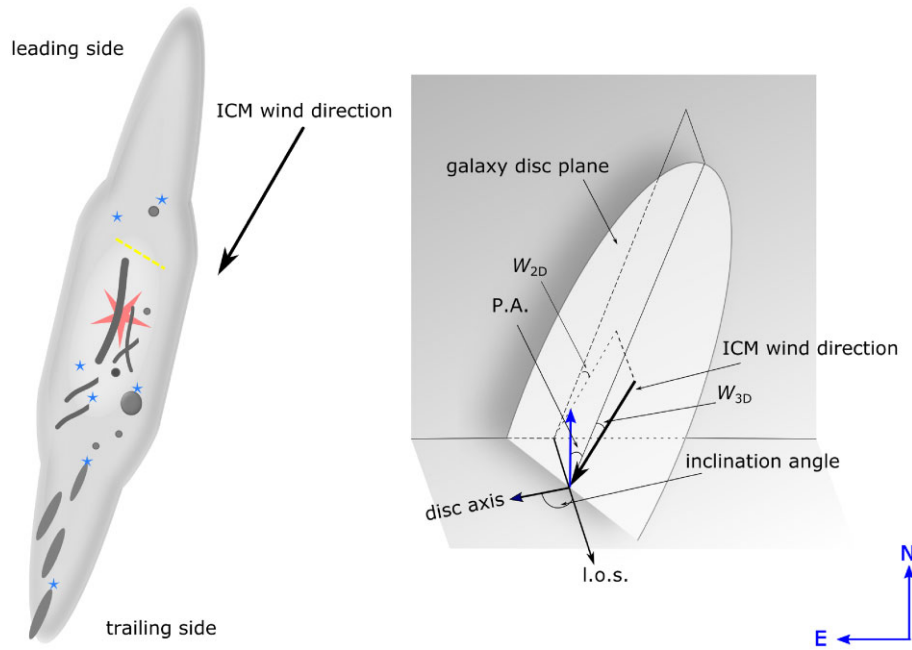


Figure 15. Left-hand panel: A cartoon representation of ESO 137-002, currently undergoing ram-pressure stripping. The ICM wind is projected to the galaxy from NW, dragging dust clouds, and forming filaments at $\sim 20^\circ$ to the east from the south. It is likely that the ablated dust clouds are located between the bright galactic disc and us to make the dust features more significant in the optical. This is consistent with the galaxy falling to the cluster with a redshift (or moving away from us). The model depicts dark clouds, along with star clusters in blue (some at the tip of ablated dust clouds), and a strong AGN (in red) at the nuclear position, plus a yellow stripping front separating the upstream and downstream regions. Note that the stripping front is defined by the $H\alpha$ and X-ray emission front, while stripping proceeds throughout the disc because the ICM wind is tilted relative to the disc plane ($\sim 20^\circ$) and the ISM is porous. Right-hand panel: A 2D representation of wind angle projected on the plane of sky. The direction of the ICM wind is nearly edge-on along the galactic disc. ESO 137-002 likely has some line-of-sight motion with cluster such that the 3D disc-wind angle (W_{3D}) is less than the 2D projected disc-wind angle (W_{2D}). On the other hand, the asymmetric dust distribution in the galaxy suggests that the angle between the line-of-sight and the ICM wind cannot be too small (or too different from 90 deg).

leading side. Simulations done on near edge-on stripping cases show the low efficiency of stripping (e.g. Vollmer et al. 2001; Roediger & Brüggén 2006). It is also noted that the secondary $H\alpha$ tail also has the same orientation with respect to the disc plane (Zhang et al. 2013).

The rich dust features revealed by *HST*, including filaments and large clouds, likely disclose the distribution of the cold ISM clouds experiencing ram pressure. The dust lane exhibits numerous dust threads with a typical length of ~ 0.8 kpc and width of ~ 0.1 kpc. NGC 4921 (Kenney et al. 2015), another galaxy undergoing near edge-on stripping in the Coma cluster, but with face-on view ($i \lesssim 25^\circ$), also reveal similar dust features indicating ongoing RPS. The ~ 20 kpc dust front on the leading (NW) side of NGC 4921 shows different types of filamentary structures with 0.5–1 kpc in length and 0.1–0.2 kpc in width, approximately in the wind direction. Similar dust cloud filaments were also observed in NGC 4402 and NGC 4522 (Abramson & Kenney 2014; Abramson et al. 2016), mostly aligned with the projected wind direction. Kenney et al. (2015) proposed that dust clouds with different shapes reflect spatial variations in ram pressure acceleration, due principally to variations in gas density, leading to the partial decoupling of the densest clouds. The complex morphologies suggest the interplay of magnetic field, galactic rotation, and turbulence.

As shown by e.g. Bekki (2014), indeed edge-on stripping mostly compresses the ISM instead of removing them from the galaxy. This compression can result in enhanced SF in the galaxy and the SF can continue for ~ 1 Gyr after the pericentre passage. Since the colours of these young star clusters spread over a range (Fig. 8), it is possible that the RPS triggered SF proceeds with different time-scales in different

clouds so there is not a single burst-like event in the disc and some clouds do not have associated star clusters at the moment, besides the likely different intrinsic extinction for different clouds. As the ages of these young star clusters are \sim several Myr – 200 Myr, ESO 137-002 would have moved for ~ 200 kpc in the sky if its velocity in the plane of the sky is 1000 km s^{-1} . Thus, the RPS triggered SF likely started when the galaxy began to plunge into the central, dense part of the cluster ICM.

ESO 137-002 presents a great case of an edge-on galaxy experiencing a nearly edge-on ram pressure wind with a large component in the plane of the sky. It is especially useful for showing the vertical ISM behaviour during the strong, nearly edge-on stripping, and for getting a nearly edge-on view of the stripping features. There are other galaxies undergoing edge-on stripping, but many are viewed face-on, e.g. NGC 4921, NGC 4654 (Lizée et al. 2021), NGC 4501 (Vollmer et al. 2008), and others are probably not only affected by RPS, e.g. UGC 6697 (Consolandi et al. 2017). ESO 137-002 thus renders a perfect prototype for studying the edge-on stripping because of the favourable combination of 3 angles – nearly edge-on 3D disc-wind angle, disc viewed nearly edge-on, and much of the ICM wind projected in the plane of sky.

7.4 RPS galaxies with little SF in the stripped tail

ESO 137-002 is a great example of RPS galaxies with little SF in the stripped ISM. It has been known that active SF is not always present in the stripped tails, e.g. some galaxies in the Coma and A1367 samples of RPS galaxies selected by $H\alpha$ diffuse tail by Yagi et al. (2010, 2017) and the Virgo sample studied by Boissier et al.

(2012). D100 (Yagi et al. 2007; Jáchym et al. 2017), another RPS galaxy with a remarkably long and narrow $H\alpha$ tail (~ 60 kpc) also has weak SF in the tail. The total SFR in the tail is found to be $\sim 6 \times 10^{-3} M_{\odot} \text{ yr}^{-1}$ (Cramer et al. 2019). At least 1/3 of RPS galaxies in Yagi et al. (2010, 2017) have little to weak SF in their stripped tails, e.g. GMP 3071, RB 219, IC 3949, NGC 4853, CGCG 097-073, and CGCG 097-079. Boselli et al. (2016) studied NGC 4569, the most massive spiral galaxy in the Virgo cluster, and revealed the presence of a diffuse gas tail but without any associated young stellar components.

Because of the lack of active SF in the stripped medium, the diffuse gas tails of the above RPS galaxies can be faint and missed in the optical and UV imaging. As described in Poggianti et al. (2017), the key selection criterion of the GASP sample is the stripping signatures shown in the B -band morphologies of these galaxies. The majority of GASP galaxies shows bright star-forming clumps in the disc and/or along the tails. The fraction of diffuse gas (not star forming) in the tails³ may increase for the galaxies with significant gaseous tail features. Poggianti et al. (2019) studied 16 GASP galaxies with clear stripping tails beyond the galaxy boundary and found that SF contributes to 64 per cent – 94 per cent of the $H\alpha$ emission in the tails, which is considered as the dominating ionization mechanism there. By separating the diffuse gas and $H\alpha$ clumps (the places where SF occurs) with the $H\alpha$ surface brightness map, they found that, only in four galaxies, the diffuse ionized gas can contribute over 60 per cent of the $H\alpha$ emission in the tails. Note that Poggianti et al. (2019) used the stellar continuum surface brightness to define the galaxy boundary. Adopting D_{25} (the isophotal level of 25 mag arcsec⁻² in the B -band) as a boundary of the galaxy main body, we also performed an independent inspection on the released $H\alpha$ surface brightness maps of 114 GASP galaxies and estimated the $H\alpha$ flux fraction of the diffuse gas in their tail regions beyond D_{25} . For 12 galaxies with clear tail features beyond D_{25} , we apply SExtractor to select H II region candidates from the $H\alpha$ surface brightness map by requesting point-like sources (CLASS_STAR > 0.9) with a low ellipticity ($e < 0.2$) (e.g. Fossati et al. 2016). Four galaxies in the full sample of 114 GASP galaxies have a fraction of diffuse gas larger than 50 per cent in the tail. Only one of these 114 galaxies has a fraction of diffuse gas larger than 60 per cent, JO204 (~ 75 per cent) and its $H\alpha$ tail is only detected to ~ 32 kpc from the nucleus, while the galaxy is ~ 30 per cent more luminous than ESO 137-002 at the *WISE* W1 band. Thus, indeed the GASP sample mainly covers the RPS galaxies with active SF downstream in the stripped tails. On the contrary, the $H\alpha$ flux fraction of the diffuse gas in ESO 137-002's tail is over 95 per cent for the lack of H II regions. The fraction of diffuse gas that we discussed above is estimated from the $H\alpha$ surface brightness map, which may be underestimated, since the exact ionization mechanism of the $H\alpha$ clumps and H II region candidates need to be considered with more information (e.g. emission-line ratios, see Tomičić et al. (2021) as an example). Multiwavelength imaging and IFU observations can provide further information to effectively distinguish the star-forming from diffuse gas and identify the stripping tails undergoing different evolutionary states.

³The diffuse gas in the tails is different from the concept 'diffuse ionized gas (DIG)'. DIG is a warm ($\sim 10^4$ K) and low density ($\sim 0.1 \text{ cm}^{-3}$) gas phase in the ISM of galaxies (see Haffner et al. 2009 for a review). In the stripping tails, while the diffuse gas surrounding the H II regions may partly include the DIG, the diffuse gas far away from the H II regions is most likely the result of mixing between the stripped ISM and the ICM.

It is still unclear what determines the SF efficiency in the stripped tails. Little to weak SF in the stripped tails may happen at the early, late or even intermediate stage of stripping. Considering the whole evolution process of the stripped gas tails, the RPS galaxies with/without star-forming trails are both important subgroups of the full population of the RPS galaxies. If selected by SF tails behind the galaxy in the optical or UV (e.g. surveys like GASP or using the *Galex* data), ESO 137-002 and NGC 4569 will not be considered as a 'jellyfish' galaxy. Future deep and wide-field surveys in $H\alpha$, HI, radio continuum and X-rays, insensitive to SF, will help to build unbiased sample of RPS galaxies and further explore how the stripped ISM mixing with the hot ICM and its evolution path.

8 CONCLUSIONS

We present detailed analysis of ESO 137-002, a large edge-on spiral with a boxy bulge currently undergoing nearly edge-on RPS, with the *HST* and the *APEX* data. The main results of this paper are:

(i) The galaxy has an undisturbed stellar body but a strongly disturbed ISM distribution, as shown by the $H\alpha$ emission and the dust distribution (Figs 1 and 6). The upstream of ESO 137-002 is nearly dust-free while the downstream is full of rich dust features, suggesting RPS has nearly cleared the north half of the galaxy and still continue to strip around the nucleus and the southern part. This asymmetric distribution also constrains that the ICM wind direction should not be close to the line of sight. At the downstream side, there are dust filaments mostly aligned with the ICM wind direction and large dust clouds being ablated by the ICM wind. These results suggest a significant motion of the galaxy in the plane of sky. Even though the radial velocity of the galaxy is $+820 \text{ km s}^{-1}$ to the velocity of the cluster, for Abell 3627 likely in a major merger, the actual radial component of ESO 137-002's motion relative to the surrounding ICM can be very different.

(ii) Some UV bright, compact sources are discovered in the galaxy, mostly upstream of nearby dust clouds (Figs 7 and 8). We suggest that they are young star clusters triggered by the compression of ram pressure, which offers evidence of RPS-triggered SF in RPS galaxies. On the other hand, the total SFR from these young star clusters is much smaller than the total SFR of the galaxy.

(iii) A large amount of the molecular gas ($5.5 \times 10^9 M_{\odot}$) has been revealed from the galaxy and the inner tail region (Figs 12 and 13), corresponding to a molecular gas fraction of ~ 16 per cent. On the other hand, the distribution of the molecular gas is asymmetric, with no significant CO emission detected from upstream and abundant CO emission downstream. CO emission is also detected in an inner tail region beyond the optical disc. This again shows the impact of RPS on the ISM distribution. ESO 137-002 likely still retains most of its ISM so the stripping is probably still at the early stage, also aided by the slow development of stripping under the edge-on configuration.

(iv) Despite its long X-ray/ $H\alpha$ tail and the detection of the molecular gas in the inner tail region, no SF is revealed in the tail region from the *HST* data (Figs 9 and 10), which is consistent with the lack of compact $H\alpha$ sources in the tail (Sun et al. 2010; Zhang et al. 2013). An upper limit of the SFE ($\sim 1.9 \times 10^{-12} \text{ yr}^{-1}$) is put on the inner tail region.

(v) ESO 137-002 is a galaxy undergoing RPS with long gaseous X-ray/ $H\alpha$ tails. However, for the lack of SF downstream and the lack of asymmetry in the optical blue light, it would have been excluded from the 'jellyfish' samples based on the ground optical data or UV survey data from e.g. *Galex*. Thus, ESO 137-002 presents a great example to demonstrate the diversity of RPS galaxies and their

observation signature. Samples based on different selection criteria should be combined for a more comprehensive understanding of RPS, their impact on galaxy evolution and the fate of the stripped ISM.

As demonstrated by this work and previous works (e.g. Cortese et al. 2007; Kenney et al. 2015; Abramson et al. 2016; Cramer et al. 2019), the *HST* data are important for such kind of detailed analysis. More analysis with the *HST* data and future wide-field survey data from *Euclid* and *Nancy Grace Roman Space Telescope* will allow us to understand the young stellar population and SFE in the RPS tails better, also with the CO data from e.g. *ALMA* and optical spectroscopic data from e.g. *MUSE*.

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DATA AVAILABILITY

The *HST* raw data used in this paper are available to download at the The Barbara A. Mikulski Archive for Space Telescopes.⁴ The APEX raw data are available to download at the ESO Science Archive Facility.⁵ The reduced data underlying this paper will be shared on reasonable requests to the corresponding author.

REFERENCES

Abadi M. G., Moore B., Bower R. G., 1999, *MNRAS*, 308, 947
 Abramson A., Kenney J. D. P., 2014, *AJ*, 147, 63
 Abramson A. et al., 2011, *AJ*, 141, 164
 Abramson A. et al., 2016, *AJ*, 152, 32
 Bekki K., 2014, *MNRAS*, 438, 444
 Bekki K., Couch W. J., 2003, *ApJ*, 596, L13
 Bertin E., Arnouts S., 1996, *A&AS*, 117, 393
 Blagorodnova N. et al., 2014, *MNRAS*, 442, 327
 Boissier S. et al., 2012, *A&A*, 545, A142
 Boselli A., Gavazzi G., 2006, *PASP*, 118, 517
 Boselli A. et al., 2014, *A&A*, 564, A66
 Boselli A. et al., 2016, *A&A*, 587, A68
 Boselli A. et al., 2021, *A&A*, 646, A139
 Cappellari M., Copin Y., 2003, *MNRAS*, 342, 345

Catinella B. et al., 2018, *MNRAS*, 476, 875
 Chen H. et al., 2020, *MNRAS*, 496, 4654
 Chung A. et al., 2007, *ApJ*, 659, L115
 Ciambur B. C., 2016, *Publ. Astron. Soc. Austr.*, 33, e062
 Ciambur B. C., Graham A. W., 2016, *MNRAS*, 459, 1276
 Consolandi G. et al., 2017, *A&A*, 606, A83
 Cortese L. et al., 2007, *MNRAS*, 376, 157
 Courteau S. et al., 2007, *ApJ*, 671, 203
 Cramer W. J. et al., 2019, *ApJ*, 870, 63
 Dowell C. D. et al., 2014, *ApJ*, 780, 75
 Dressler A., 1980, *ApJ*, 236, 351
 Dumke M., Mac-Auliffe F., 2010, in Silva D. R., Peck A. B., Soifer B. T., eds, Proc. SPIE Conf. Ser. Vol. 7737, Observatory Operations: Strategies, Processes, and Systems III. SPIE, Bellingham, p. 77371J
 Erwin P., Debattista V. P., 2017, *MNRAS*, 468, 2058
 Ferland G. J. et al., 2017, *Rev. Mex. Astron. Astrofis.*, 53, 385
 Fitzpatrick E. L., 1999, *PASP*, 111, 63
 Fossati M. et al., 2016, *MNRAS*, 455, 2028
 Gavazzi G., Jaffe W., 1987, *A&A*, 186, L1
 Gavazzi G. et al., 2001, *ApJ*, 563, L23
 George K. et al., 2018, *MNRAS*, 479, 4126
 Giovanelli R., Haynes M. P., 1985, *ApJ*, 292, 404
 Gonzaga S. et al., 2012, *The DrizzlePac Handbook*, Baltimore, STScI
 Gunn J. E., Gott J. R., 1972, *ApJ*, 176, 1
 Güver T., Özel F., 2009, *MNRAS*, 400, 2050
 Haffner L. M. et al., 2009, *Rev. Mod. Phys.*, 81, 969
 Hao C.-N. et al., 2011, *ApJ*, 741, 124
 Hester J. A. et al., 2010, *ApJ*, 716, L14
 Indebetouw R. et al., 2005, *ApJ*, 619, 931
 Israel F. P., van der Werf P. P., Tilanus R. P. J., 1999, *A&A*, 344, L83
 Jáchym P. et al., 2007, *A&A*, 472, 5
 Jáchym P. et al., 2014, *ApJ*, 792, 11
 Jáchym P. et al., 2017, *ApJ*, 839, 114
 Jáchym P. et al., 2019, *ApJ*, 883, 145
 Kapferer W. et al., 2009, *A&A*, 499, 87
 Kenney J. D. P. et al., 2014, *ApJ*, 780, 119
 Kenney J. D. P., Abramson A., Bravo-Alfaro H., 2015, *AJ*, 150, 59
 Kennicutt R. C., Evans N. J., 2012, *ARA&A*, 50, 531
 Kravtsov A. V., Borgani S., 2012, *ARA&A*, 50, 353
 Kroupa P., 2001, *MNRAS*, 322, 231
 Kuno N. et al., 2007, *PASJ*, 59, 117
 Lamperti I. et al., 2020, *ApJ*, 889, 103
 Lange R. et al., 2015, *MNRAS*, 447, 2603
 Lasker B. M. et al., 2008, *AJ*, 136, 735
 Laurikainen E., Salo H., 2017, *A&A*, 598, A10
 Lee J. C., Hwang H. S., Ko J., 2013, *ApJ*, 774, 62
 Leitherer C. et al., 1999, *ApJS*, 123, 3
 Leitherer C. et al., 2014, *ApJS*, 212, 14
 Leroy A. K. et al., 2009, *AJ*, 137, 4670
 Levesque E. M. et al., 2012, *ApJ*, 751, 67
 Lizée T. et al., 2021, *A&A*, 645, A111
 Mao R.-Q. et al., 2010, *ApJ*, 724, 1336
 Moretti A. et al., 2018, *MNRAS*, 480, 2508
 Moretti A. et al., 2020, *ApJ*, 889, 9
 Oosterloo T., van Gorkom J., 2005, *A&A*, 437, L19
 Peng C. Y. et al., 2002, *AJ*, 124, 266
 Peng C. Y. et al., 2010, *AJ*, 139, 2097
 Poggianti B. M. et al., 2017, *ApJ*, 844, 48
 Poggianti B. M. et al., 2019, *MNRAS*, 482, 4466
 Quilis V., Moore B., Bower R., 2000, *Science*, 288, 1617
 Roediger E., Brüggem M., 2006, *MNRAS*, 369, 567
 Ruszkowski M. et al., 2014, *ApJ*, 784, 75
 Saintonge A. et al., 2017, *ApJS*, 233, 22
 Schindler S., Diaferio A., 2008, *Space Sci. Rev.*, 134, 363
 Schlafly E. F., Finkbeiner D. P., 2011, *ApJ*, 737, 103
 Schulz S., Struck C., 2001, *MNRAS*, 328, 185
 Sheth K. et al., 2005, *ApJ*, 632, 217
 Sivanandam S., Rieke M. J., Rieke G. H., 2010, *ApJ*, 717, 147

⁴<https://archive.stsci.edu/hst/>

⁵<http://archive.eso.org/cms.html>

- Sivanandam S., Rieke M. J., Rieke G. H., 2014, *ApJ*, 796, 89
 Smith R. J. et al., 2010, *MNRAS*, 408, 1417
 Solomon P. M., Vanden Bout P. A., 2005, *ARA&A*, 43, 677
 STScI Development Team, 2013, ppsynphot: Synthetic photometry software package, record ascl:1303.023
 Sun M., Vikhlinin A., 2005, *ApJ*, 621, 718
 Sun M. et al., 2006, *ApJ*, 637, L81
 Sun M., Donahue M., Voit G. M., 2007, *ApJ*, 671, 190
 Sun M. et al., 2010, *ApJ*, 708, 946
 Tomičić N. et al., 2021, *ApJ*, 907, 22
 Tonnesen S., Bryan G. L., 2012, *MNRAS*, 422, 1609
 van Dokkum P. G., 2001, *PASP*, 113, 1420
 Verdugo C. et al., 2015, *A&A*, 582, A6
 Vollmer B. et al., 2001, *ApJ*, 561, 708
 Vollmer B. et al., 2008, *A&A*, 483, 89
 Wang Q. D., Owen F., Ledlow M., 2004, *ApJ*, 611, 821
 Wegner G. et al., 2003, *AJ*, 126, 2268
 Yagi M. et al., 2007, *ApJ*, 660, 1209
 Yagi M. et al., 2010, *AJ*, 140, 1814
 Yagi M. et al., 2017, *ApJ*, 839, 65
 Yoshida M. et al., 2008, *ApJ*, 688, 918
 Zhang B. 2013, *ApJ*, 777, 122

APPENDIX A: BLUE SOURCES IN THE TAIL REGION

There are six blue sources in the tail region as shown in Fig. A1. They are #1 at (16:13:39.70, -60:53:21.17), #2 at (16:13:40.56, -60:53:02.66), #3 at (16:13:42.43, -60:53:36.55), #4 at (16:13:37.55, -60:52:33.05), #5 at (16:13:36.59, -60:52:31.57), and #6 at (16:13:38.44, -60:52:51.09). The brightest two sources, #1 and #2, are likely Galactic stars with significant F160W emission and no H α emission. The morphology of the source #3 suggests it

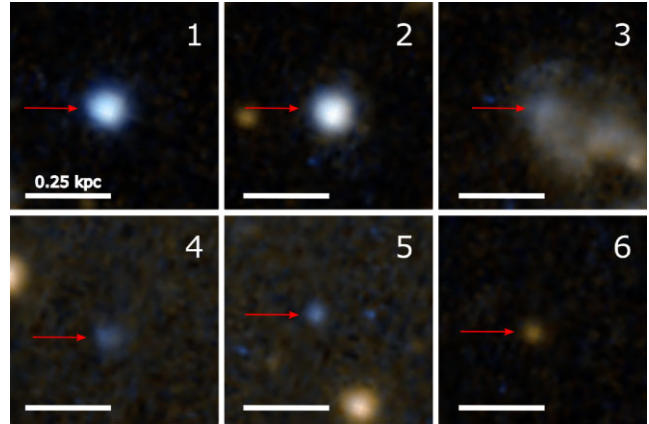


Figure A1. The composite image of the six sources in the tail region (zoom-in of Fig. 1) with $F275W - F475W \leq 2$ and $F475W - F814W \leq 1$. Sources are numbered in order of the $F475W$ flux. The brightest two objects are likely Galactic stars with significant $F160W$ detections. Source #3 is likely a background object.

is a background galaxy. If we exclude the source #1 (also see its position in Fig. 11), the total $F275W$ flux of these five sources is $1.49 e^-/s$, while the total $F275W$ flux of sources in the control region selected with the same colour cuts is $11.57 e^-/s$. After correcting the difference on the sky area, the net $F275W$ flux in the tail region is $-0.25 e^-/s$, again consistent with no recent SF present in the tail.

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