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












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The multichord stellar occultation on 2019 October 22 by the trans-Neptunian object (84922) 2003 VS₂★

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ABSTRACT

Context. Stellar occultations have become one of the best techniques to gather information about the physical properties of trans-Neptunian objects (TNOs), which are critical objects for understanding the origin and evolution of our Solar System.

Aims. The purpose of this work is to determine, with better accuracy, the physical characteristics of the TNO (84922) 2003 VS₂ through the analysis of the multichord stellar occultation on 2019 October 22 and photometric data collected afterward.

Methods. We predicted, observed, and analyzed the multichord stellar occultation of the Second *Gaia* Data Release (*Gaia* DR2) source 3449076721168026624 ($m_v = 14.1$ mag) by the plutino object 2003 VS₂ on 2019 October 22. We performed aperture photometry on the images collected and derived the times when the star disappeared and reappeared from the observing sites that reported a positive detection. We fit the extremities of such positive chords to an ellipse using a Monte Carlo method. We also carried out photometric observations to derive the rotational light curve amplitude and rotational phase of 2003 VS₂ during the stellar occultation. Combining the results and assuming a triaxial shape, we derived the 3D shape of 2003 VS₂.

Results. Out of the 39 observatories involved in the observational campaign, 12 sites, located in Bulgaria (one), Romania (ten), and Serbia (one), reported a positive detection; this makes it one of the best observed stellar occultations by a TNO so far. Considering the rotational phase of 2003 VS₂ during the stellar occultation and the rotational light curve amplitude derived ($\Delta m = 0.264 \pm 0.017$ mag), we obtained a mean area-equivalent diameter of $D_{A_{\text{eq}}} = 545 \pm 13$ km and a geometric albedo of 0.134 ± 0.010 . By combining the rotational light curve information with the stellar occultation results, we derived the best triaxial shape for 2003 VS₂, which has semiaxes $a = 339 \pm 5$ km, $b = 235 \pm 6$ km, and $c = 226 \pm 8$ km. The derived aspect angle of 2003 VS₂ is $\theta = 59^\circ \pm 2^\circ$ or its supplementary $\theta = 121^\circ \pm 2^\circ$, depending on the north-pole position of the TNO. The spherical-volume equivalent diameter is $D_{V_{\text{eq}}} = 524 \pm 7$ km. If we consider large albedo patches on its surface, the semi-major axis of the ellipsoid could be ~ 10 km smaller. These results are compatible with the previous ones determined from the single-chord 2013 and four-chord 2014 stellar occultations and with the effective diameter and albedo derived from *Herschel* and *Spitzer* data. They provide evidence that 2003 VS₂'s 3D shape is not compatible with a homogeneous triaxial body in hydrostatic equilibrium, but it might be a differentiated body and/or might be sustaining some stress. No secondary features related to rings or material orbiting around 2003 VS₂ were detected.

Key words. Kuiper belt objects: individual: 2003 VS₂ – methods: observational – techniques: photometric

1. Introduction

Trans-Neptunian objects (TNOs) offer a unique opportunity to better understand the origins and the chemical, dynamical, and collisional evolution of the outer Solar System. Possibly dispatched to further distances from the Sun than that of Neptune due to gravitational perturbations after their formation (Gomes et al. 2005; Levison et al. 2008), their global composition has

been virtually unaffected by solar irradiation, keeping it very similar to that of the primitive nebula (Morbidelli et al. 2008).

In the last decade, stellar occultations by TNOs have proved to be one of the best techniques to determine the size and shape of these objects, to show features such as satellites (Sickafoose et al. 2019) or rings (Braga-Ribas et al. 2014; Ortiz et al. 2015, 2017), and to reveal possible atmospheres (Hubbard et al. 1988; Stern & Trafton 2008; Meza et al. 2019). If we combine this technique with light-reflection measurements, we can also derive their geometric albedo. Furthermore, we can calculate their density if we assume hydrostatic equilibrium or if the body is part of a binary system or has a satellite, in which case we can obtain its mass.

* The photometric data used to obtain the rotational light curve of (84922) 2003 VS₂ are only available at the CDS via anonymous ftp to cdsarc.u-strasbg.fr (130.79.128.5) or via <http://cdsarc.u-strasbg.fr/viz-bin/cat/J/A+A/663/A121>

† Deceased.

Table 1. Orbital elements and physical characteristics of 2003 VS₂.

a (au)	q (au)	e	i (°)	H_V	P (h)	Δm (mag)	$D_{\text{area,eq}}^{(a)}$ (km)	$D_{\text{vol,eq}}$ (km)	$p_V^{(b)}$
39.249	36.386	0.073	14.83	4.14 ± 0.07	7.41753 ± 0.00001	0.141 ± 0.009	553^{+36}_{-33}	$548.3^{+29.5}_{-44.6}$	$0.131^{+0.024}_{-0.013}$

Notes. Orbital elements of 2003 VS₂ from the Jet Propulsion Laboratory (JPL) Small-Body Database Browser (<https://ssd.jpl.nasa.gov/sbdb.cgi>): a , semi-major axis; q , perihelion distance; e , eccentricity; and i , inclination. Absolute magnitude H_V is from Alvarez-Candal et al. (2016) and private communication. Rotation period P is the preferred value from Santos-Sanz et al. (2017). The rotational light curve amplitude (Δm), spherical volume-equivalent diameter ($D_{\text{vol,eq}}$), and albedo (p_V) are from Benedetti-Rossi et al. (2019). The mean area-equivalent diameter ($D_{\text{area,eq}}$) was calculated from the values given in Benedetti-Rossi et al. (2019) and taking into account the rotational phase using Eq. (1). ^(a)Mommert et al. (2012) obtained an area-equivalent diameter of $D = 523^{+35.1}_{-34.4}$ km and ^(b)an albedo of $p_V = 0.147^{+0.063}_{-0.043}$.

Although this technique has been increasingly used in the past few years, it is still challenging to predict and obtain positive results from multichord stellar occultations (Ortiz et al. 2020), mainly due to the uncertainties in the orbits of TNOs. The Second *Gaia* Data Release (*Gaia* DR2; *Gaia* Collaboration 2016b,a, 2018) eases the process by providing peerless accurate positions and proper motions of more than one billion sources. However, the large orbital periods of TNOs and the short time span of the observations result in non-negligible uncertainties in their orbital elements, making astrometric observations close to the stellar occultation date indispensable for a successful prediction.

The object (84922) 2003 VS₂ orbits in the 3:2 mean motion resonance (MMR) with Neptune, which makes it a plutino. This TNO presents a double-peaked rotational light curve with an accurately determined rotation period of 7.41753 ± 0.00001 h (Santos-Sanz et al. 2017). No satellites nor secondary features have been discovered so far around 2003 VS₂. Its near-infrared spectrum reveals the presence of exposed water ice (Barkume et al. 2008) which, according to Mommert et al. (2012), might explain the increase in 2003 VS₂'s albedo from the typical value of 0.07 for the plutinos. The orbital elements and the most relevant physical characteristics of 2003 VS₂ can be found in Table 1.

A recent work was published with results from a four-chord and two single-chord stellar occultations by 2003 VS₂ in 2013 and 2014 (Benedetti-Rossi et al. 2019). In that work, the authors reconstructed the 3D shape of 2003 VS₂ by combining the data from the multichord stellar occultation and the rotational light curve obtained. The principal semiaxes that provided the best fit had values of $a = 313.8 \pm 7.1$ km, $b = 265.5^{+8.8}_{-9.8}$ km, and $c = 247.3^{+26.6}_{-43.6}$ km, with these measurements being inconsistent with a Jacobi ellipsoid (Chandrasekhar 1987). This solution was highly affected by the rotational light curve amplitude derived in that paper, which was 0.141 ± 0.009 mag and significantly smaller than the previously moderately large ones reported by Ortiz et al. (2006), Sheppard (2007), and Thirouin et al. (2010). In this work we show that such a rotational light curve amplitude cannot be correct based on new data and other reasoning. Benedetti-Rossi et al. (2019) also reported the presence of a putative secondary feature detected from one of the observing stations but, without further data, they could not discard the possibility that it could be due to a companion star or instrumental effects.

In this work we present the results of the multichord stellar occultation by 2003 VS₂ on 2019 October 22. A total of 39 observatories were involved in the campaign, from which 12 positive chords were obtained, which is a considerable improvement with respect to the previous occultations by this TNO. Combining the collected data from the stellar occultation and ensuing photometric measurements, we derived the 3D shape of 2003 VS₂.

2. Observations

This section presents the observations carried out to improve the prediction of the stellar occultation, the observations of the actual event, and the observations performed shortly afterward to obtain the rotational light curve of 2003 VS₂.

2.1. Occultation predictions

The stellar occultation on 2019 October 22 was singled out from the systematic searches for TNO occultation candidate stars carried out by the European Research Council (ERC) Lucky Star¹ project collaboration (Desmars et al. 2015, 2018). The Lucky Star's NIMA² ephemeris thus gave the initial prediction³ for the 2019 October 22 stellar occultation, and the candidate star was identified in the *Gaia* DR2 catalog (source ID: 3449076721168026496; UCAC4⁴ identifier 616-023624). Relevant information about the star, such as its coordinates, proper motions, parallax, and their uncertainties, as well as the star's G , B , V , and K magnitudes, can be found in Table 2.

To reduce the uncertainty on 2003 VS₂'s orbit and narrow down the predicted location of the shadow path, we performed two observing runs with two different telescopes a few days before the event. The first observing run was carried out on 2019 October 5, with the charge-coupled device (CCD) Andor ikon-L camera of the 1.5-m telescope at the Sierra Nevada Observatory (OSN) in Granada, Spain. This camera provides a field of view (FOV) of $7'92 \times 7'92$, with an image scale of $0''.232 \text{ pixel}^{-1}$. The 15 images of this set were acquired in 2×2 binning mode, with no filter and an integration time of 400 s. The average seeing was $1''.84$. Bias and flat-field frames were taken for standard calibration, which was done afterward, following the steps in Fernández-Valenzuela et al. (2016).

The second run was taken on 2019 October 8, with the IO:O camera of the Liverpool 2-m Telescope at the Roque de los Muchachos Observatory in La Palma, Spain. This instrument has a FOV of $10' \times 10'$ and an image scale of $0''.15 \text{ pixel}^{-1}$. The ten images of this set were acquired in 2×2 binning mode, with the Sloan R -filter and 300 s of integration time. The average seeing was $1''.2$. Bias and flat-field frames were also obtained for standard calibration.

With the OSN data, the obtained offsets with respect to the Jet Propulsion Laboratory (JPL) #30 orbit were (-360 ± 36) mas in right ascension (RA) and $(+4 \pm 25)$ mas in declination (Dec). The Liverpool data yielded offsets of (-368 ± 12) mas in RA and $(+24 \pm 11)$ mas in Dec.

¹ <https://lesia.obspm.fr/lucky-star/index.php>

² Numerical Integration of the Motion of an Asteroid.

³ <https://lesia.obspm.fr/lucky-star/occ.php?p=32895>

⁴ The fourth U.S. Naval Observatory CCD Astrograph Catalog.

Table 2. Main information of the occulted star (*Gaia* DR2 3449076721168026496 ^(a); UCAC4 identifier 616-023624 ^(a)).

RA (ICRF ^(b))	errRA (mas)	Dec (ICRF ^(b))	errDec (mas)	pmRA (mas yr ⁻¹)	pmDec (mas yr ⁻¹)	Plx (mas)	<i>G</i> (mag)	<i>B</i> (mag)	<i>V</i> (mag)	<i>K</i> (mag)
05h 30m 38.0442s	0.0401	+33°07′01″.748	0.0330	-0.34 ± 0.06	-1.76 ± 0.05	0.2084 ± 0.0533	14.1625	15.030	14.150	10.968

Notes. Star coordinates in RA and Dec were propagated to the occultation epoch (2458779.35870) and respective errors (errRA, errDec), proper motion in RA and Dec and respective errors, absolute stellar parallax with error, and magnitude *G* from *Gaia* DR2 (Gaia Collaboration 2016b,a, 2018). Magnitudes *B*, *V*, and *K* are from the Naval Observatory Merged Astrometric Dataset (NOMAD) catalog (Zacharias et al. 2004). ^(a)From Vizier: <https://bit.ly/3vUjwkZ>, ^(b)International Celestial Reference System.

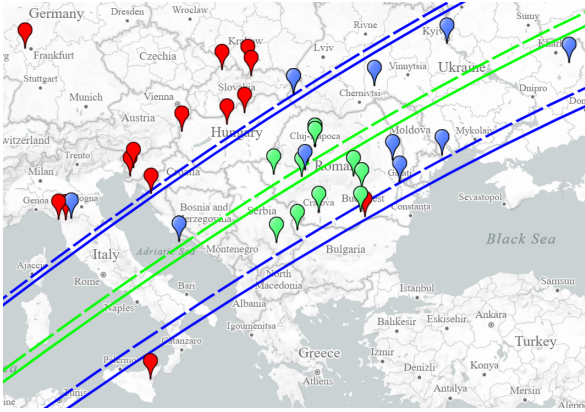


Fig. 1. Predicted (solid lines) and observed (dashed lines) shadow paths for the 2003 VS₂ stellar occultation on 2019 October 22 through *Gaia* DR2 source 3449076721168026496. The prediction was made updating the JPL #30 ephemeris with the offsets obtained with data from the Liverpool 2-m Telescope in the Roque de los Muchachos Observatory (La Palma, Spain). The green line represents the middle of the shadow path and the blue lines indicate the limits of the shadow (width of the shadow path is 479 km, from JPL). The observing sites involved in the event are also marked: in green, the ones that reported a positive detection; in red, those that reported a negative detection; and, in blue, those that could not observe due to bad weather or technical problems, see Tables 3, 4, and A.1.

We made two different predictions using the OSN and the Liverpool data, although the result was roughly identical (the Liverpool prediction being ~ 20 mas north of the OSN). We used the predicted shadow path obtained by updating 2003 VS₂'s ephemeris from JPL with the Liverpool data (shown in Fig. 1), as it was taken closer to the event date and with better seeing.

2.2. Stellar occultation observations

On 2019 October 22, 39 observational stations spread across 11 countries (including both professional telescopes and amateur observers, Fig. 1) were all set to observe *Gaia* DR2 source 3449076721168026496 around the predicted occultation time. Out of the total participating stations, 12 reported a positive detection, 14 reported a negative detection (among which two were very close to the shadow path and provide constraints to the shape fitting), and 13 could not observe due to bad weather or technical difficulties. Detailed information about all the participant observatories, divided between positive detections, negative detections, and observations with technical problems or overcast, can be found in Tables 3, 4, and A.1, respectively. The detailed instrumental settings at the sites with positive detections are also given in Table 3. All the teams that could observe collected series of flexible image transport system (FITS) images, except for one site in Romania that took data in tagged image file format (TIFF)

and required a different analysis, as explained in Sect. 3.1. The time span of the collected images includes several minutes before and after the stellar occultation.

In this kind of event, it is crucial to save, within the images' header, the individual acquisition time and all participating stations must be synchronized. Clock synchronizations were used to accomplish this: all the observing teams used internet network time protocol (NTP) servers, except for two negative sites that used global positioning system video time inserters (GPS-VTI). Three positive sites had synchronization issues that are addressed in Sect. 3.1.

We performed synthetic aperture photometry of the occulted star and nearby comparison stars to correct for seeing effects and atmospheric transparency fluctuations. The synthetic aperture photometry was done using our own DAOPHOT-based routines coded in interactive data language (IDL), following the procedures described in Fernández-Valenzuela et al. (2016). In this case, though, we fixed the centroid relative to the nonocculted stars in the FOV so that it did not change when the TNO occulted the star. We used several apertures and annuli and chose those that resulted in the least scatter in the photometry. In Fig. 2 we show the resulting normalized flux (the blended flux of the occulted star plus TNO divided by its mean value outside the occultation moment) versus the time from the positive sites, which show deep drops in flux around the predicted occultation time. We note that 2003 VS₂ is too faint to be seen with the available equipment, so the star's flux drops to zero during the occultation. We derived the error bars in Fig. 2 from Poisson noise calculations, but the results were scaled so that their standard deviation matches that of the data outside the main drop in occultation (column seven in Table 5). We do not see secondary flux drops other than the one corresponding to the main body, which could indicate that there is not a sufficiently wide and dense ring orbiting around 2003 VS₂ that could have produced detectable flux drops; however, due to the small signal-to-noise ratio (S/N), we cannot discard it. We took special care when performing the photometry of the data from the places that missed the occultation but were very close to the shadow path, as they put constraints to 2003 VS₂'s final shape.

2.3. Rotational light curve observations

Given that 2003 VS₂ has a double-peaked rotational light curve and is large enough to allow a hydrostatic equilibrium shape (see a detailed explanation in Benedetti-Rossi et al. 2019), we can assume that the body is a triaxial ellipsoid (Chandrasekhar 1987). If this were the case, one would need at least three body projections at different rotational phases to correctly derive the actual 3D shape of 2003 VS₂.

However, we can overcome this by combining the occultation information with time-series photometric data taken closer to the

Table 3. Details of the observing stations of the 2003 VS₂ multichord stellar occultation on 2019 October 22 with positive detection.

Chord number	Site name location MPC Code	Latitude (N) longitude (E) altitude (m)	Telescope aperture (cm) detector/instrument	Exposure time Cycle time ^(a) (s)	Observers
1	ROASTERR-1 Observatory (1) Romania L04	46°49'15".6 23°35'47" 391 ^(b)	30 KAF8300	3 5.18	L. Hudin
2	ROASTERR-1 Observatory (2) Romania L04	46°49'15".6 23°35'47" 391 ^(b)	28 ATIK-414ex	2 3.01	M. Boaca V. Inceu
3	Romania –	46°42'37".557 23°35'35".647 783.405	30.5 SBIG STT-1603ME	2 3	V. Turcu D. Moldovan L. Mircea
4	Romania –	46°42'37".452 23°35'36".94 787.56	40.6 SBIG STT-1603ME	2 3	V. Turcu D. Moldovan L. Mircea
5	Romania –	45°42'11".94 21°26'12".78 92 ^(b)	27.94 ZWO ASI 224 CMOS	5 5.14	L. Stoian A. Juravle
6	Berthelot Observatory Romania L54	45°36'59".4 22°53'19".7 400	36.83 SBIG STL11000M CCD	2.5 8.64	D. A. Nedelcu A. Sonka
7	Stardust Observatory Romania L13	45°38'30" 25°37'19" 597	20.32 CCD Atik 383L	8.5 10.28	L. Curelaru
8	Romania –	44°19'24".5 23°47'19".0 109 ^(b)	35.6 ASI 1600	3.01 3.34	M. Predatu
9	Stardreams Observatory Romania L16	45°12'13".3 26°02'44".2 382 ^(b)	20 ATIK 460ex	4 5.52	R. Gherase
10	Astronomical Station Vidojevica Serbia C89	43°08'24".6 21°33'20".4 1150	140 Andor ikonL	0.8 1.55	D. Ilic O. Vince
11	Belogradchik Observatory Bulgaria –	43°37'22" 22°40'30" 650	60 FLI PL9000	6 6.86	R. Bachev E. Semkov
12	Romania –	44°19'19".14 25°59'8".69 70 ^(b)	35.5 ASI 1600	2 2.50	M. Teodorescu

Notes. The sites are sorted by their distance to the center of the predicted shadow path. ^(a)The cycle time is the sum of the exposure time plus dead time. ^(b)Altitudes are from Google Earth.

occultation date. To do so, we observed 2003 VS₂ for two consecutive nights, two days after the event (2019 October 24 and 25), with the 1.23-m telescope at the Calar Alto Observatory in Almería (Spain). This telescope is equipped with a 4k × 4k CCD DLR-MKIII camera, which provides a FOV of 21'.5 × 21'.5 with an image scale of 0".32 pixel⁻¹. This wide FOV allowed us to keep the same stellar field both nights, making it possible to choose the same reference stars during the run to minimize systematic photometric errors.

A total of 143 images were acquired in 2 × 2 binning mode and with an integration time of 400 s, using no filter to

maximize the S/N. The moonshine was 14% for the first night and 7% for the second night. The average seeing at Calar Alto from the Differential Image Motion monitor was 1".6 for the second night; unfortunately, the Calar Alto seeing tracker did not work for the first night. The average measured full width at half maximum (FWHM) was ~2.5 pixels (~1".6) for the first night and ~2.3 pixels (~1".5) for the second. The S/N was between 30 and 70 during the first night and between 20 and 80 during the second night. We also took bias and flat-field frames each night for standard calibration, which was performed using our own specific routines written in IDL.

Table 4. Details of the observing stations of the 2003 VS₂ multichord stellar occultation on 2019 October 22 with negative detection.

Site name location MPC Code	Latitude (N) longitude (E) altitude (m)	Telescope aperture (cm)	Observers
Bremthal Germany –	50°08′17″.4 08°21′50″.4 256	25.4	O. Klös
PDlink Observatory Cadca, Slovakia –	49°24′15″.20 18°42′09″.47 680	40	P. Delincak
Mt. Suhora Observatory Poland –	49°34′09″ 20°04′03″ 1009	60	M. Drozd W. Ogloza
Skalnate Pleso Slovakia 056	49°11′21″.8 20°14′02″.1 1786	130	T. Pribulla M. Husárik P. Sivanič
Massa, Italy –	44°01′16″.96 10°07′56″.65 30	30	P. Baruffetti
Gothard Observatory Szombathely, Hungary –	47°15′29″.83 16°36′15″.67 220 ^(a)	80	G. M. Szabó A. Derekas L. Szigeti
Črni Vrh Observatory Slovenia –	45°56′45″.1 14°04′16″.4 726	60	J. Skvarč
Monte Agliale Observatory Borgo a Mozzano, Italy 159	43°59′43″.0 10°30′53″.7 756	50	F. Ciabattari
Stazione Osservativa di Basovizza Osservatorio Astronomico di Trieste, Italy 038	45°38′33″ 13°52′23″ 400	35	P. Di Marcantonio G. Iafrate I. Coretti, V. Baldini
Piszkéstető Station Konkoly Obs., Hungary 561	47°55′6″.0 19°53′41″.7 924 ^(a)	100	A. Pal R. Szakats C. Kiss
Konkoly Observatory Normafa, Hungary 053	47°29′59″.28 18°57′51″.12 468 ^(a)	60	A. Pal R. Szakats C. Kiss
Vratnik, Croatia –	49°24′15″.20 18°42′09″.47 775	20	H. Mikuž A. Mohar
Romania –	44°07′56″.8 26°13′04″.3 88 ^(a)	20	V. Dumitrescu
Mount Etna Observatory Osservatorio Astrofisico di Catania, Italy 156	37°41′05″ 14°58′04″ 1735	91	J. Alonso-Santiago A. Frasca

Notes. The sites are sorted by their distance to the center of the predicted shadow path. ^(a)Altitudes are from Google Earth.

3. Data reduction

3.1. Time synchronization and extraction

Robust clock synchronization is essential to have the absolute acquisition time written on each frame header to obtain the actual star disappearance and reappearance times from each

site and, hence, the projected chords' relative position. In this regard, some of the images collected required a different treatment to obtain their absolute acquisition time, as we detail below. We saw no evidence for wrong time synchronization for the remaining positive chords, given their relative positions (see Fig. 2).

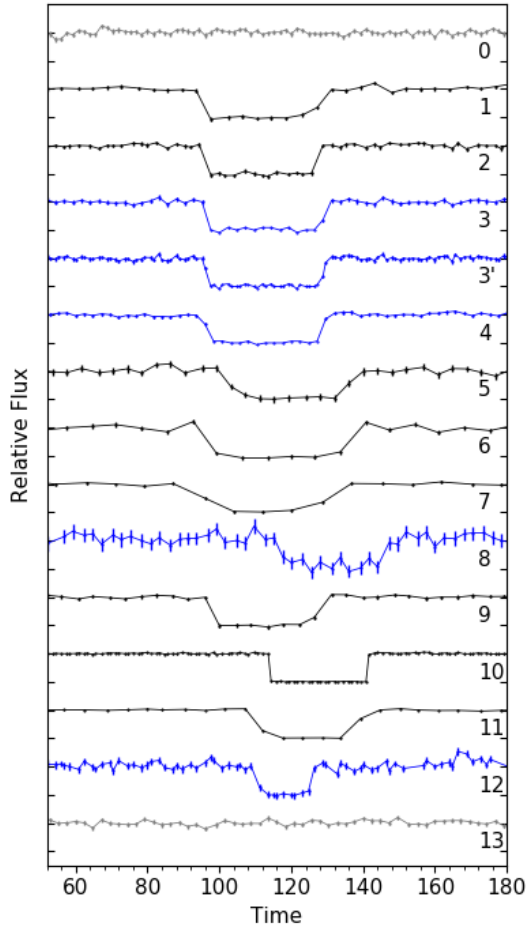


Fig. 2. Normalized light curves from the positive detections of the stellar occultation by 2003 VS₂ on 2019 October 22 and the two closest negatives. The relative flux of the occulted star with respect to the comparison chosen stars is plotted against time, given in seconds after 2019 October 22 20:40:00 UT. The uncertainty bars of the flux were plotted for all the chords, although some have the size of the points and are not visible. The light curves have been displaced in flux for better visualization and they follow the same order as in Table 3. Chords 0 and 13 (top and bottom chords, in gray) correspond to the negative detections from observers H. Mikuž and V. Dumitrescu, respectively (see Table 4.) Light curves plotted in blue required special consideration regarding time synchronization, see Sect. 3.1. Chord 3' is the result of merging chords 3 and 4, see the text for details.

Chords 3 and 4

There was an intentionally applied 1 s difference between the images acquired from the two telescopes at this site, aiming to cover the whole event avoiding the dead time of the CCDs. Since the two telescopes used the same instrument and took FITS images with the same exposure time, we merged the data from these two telescopes to obtain one single chord. By doing this, we doubled the sampling of the light curve and, therefore, we obtained the immersion and emersion times from this site with a smaller uncertainty. Hereafter we refer to this as a single chord (chord 3').

Chord 8

Although this chord nearly overlapped with chord 9, its center was shifted by more than 12 s with respect to the linear fit of the rest of the centers (see Fig. 2 and Table 5). Later tests performed

a month after the occultation event showed a difference of more than one minute between the acquisition system and a synchronized clock, but we were not able to determine the exact time offset during the occultation event. All of this suggested that the time synchronization at this site was not correctly applied, so we could only use the chord length and not the absolute time of the chord.

Chord 12

The images from this site were recorded in a data cube in TIFF format. We converted them to FITS format using the Planetary Imaging PreProcessor (PIPP), but then all of the images had the same time written on the header. The observer provided us with the acquisition times of the images, but their accuracy only reached the order of seconds (as this is the accuracy reached by the software used to save the images). We used the mid-integration times for the light curve and then increased the uncertainties of the immersion and emersion times by the exposure time (two seconds) since it was not possible to know if the acquisition times we have corresponded to the beginning, middle, or end of the images' integration.

3.2. Occultation times and chord lengths

From every site that reported a positive detection, the stellar occultation starting and ending times (disappearance and reappearance of the star, respectively) were determined by fitting each light curve to a sharp-edge square-well model convolved by the following: (1) Fresnel diffraction by a point-like source at the distance of 2003 VS₂ from the observer, (2) the CCD bandwidth, (3) the finite stellar diameter projected at the object's distance, in kilometers, and (4) the finite integration time (see, e.g., Widemann et al. 2009; Braga-Ribas et al. 2013).

During the stellar occultation, 2003 VS₂ was at a geocentric distance of $\Delta = 36.1$ au = 5.40×10^9 km. This implies that for a typical wavelength of $\lambda = 0.65$ μ m, the Fresnel scale $F = \sqrt{\lambda D/2}$ has a value of 1.33 km.

We estimated the projected stellar diameter using the formulae from van Belle (1999) and its B , V , and K apparent magnitudes, obtained from the NOMAD catalog (Zacharias et al. 2004, Table 2). The diameters obtained are 0.86 km if considering a super giant or 0.91 km if considering a main sequence star.

Finally, we converted the shortest integration time among the positive observations (0.8 s, chord 10) into the distance in the sky-plane traveled by 2003 VS₂ between two adjacent data points; since the velocity of 2003 VS₂'s shadow path was 13.01 km s⁻¹, that distance was 10.41 km. As a result, our light curves are mainly dominated by the integration times and not by Fresnel diffraction or the stellar diameter.

The fitting procedure then searches for the times of disappearance and reappearance of the star that minimize a classical χ^2 function, as explained in Sicardy et al. (2011, supplementary information). The uncertainty bars were estimated by varying the occultation times to increase χ^2 to $\chi^2 + 1$. Since these uncertainty bars depend on the fit of the photometry and as the typical errors in the photometry are very small, the uncertainty in the derived time of ingress and egress is also small, that is to say a small fraction of the integration time. Figure 3 shows an example of the best fit to one of the stellar occultation light curves (chord 2). All the derived disappearance and reappearance times and chord lengths are listed in Table 5.

Table 5. Star disappearance and reappearance UT times on 2019 October 22, chord lengths, time shifts, and dispersion (σ) of the light curves outside the occultation.

Chord number	Disappearance (s)	Reappearance (s)	Chord size (s)	Chord size (km)	Shift (s) ^(a)	σ
1	20:41:34.800 \pm 1.155	20:42:14.995 \pm 0.225	40.2 \pm 1.2	523 \pm 16	+0.7	0.06
2	20:41:35.54 \pm 0.46	20:42:14.47 \pm 0.55	38.9 \pm 0.7	506 \pm 9	+0.6	0.06
3' ^(b)	20:41:35.609 \pm 0.109	20:42:16.585 \pm 0.077	40.98 \pm 0.13	533.1 \pm 1.7	+0.03	0.06
5	20:41:44.375 \pm 0.525	20:42:24.95 \pm 0.40	40.6 \pm 0.7	528 \pm 9	-0.06	0.11
6	20:41:37.325 \pm 1.100	20:42:24.062 \pm 1.287	46.7 \pm 1.7	608 \pm 22	+1.6	0.08
7	20:41:35.420 \pm 0.425	20:42:17.411 \pm 0.382	42.0 \pm 0.6	546 \pm 8	+0.07	0.05
8	20:42:02.05 \pm 0.73	20:42:37.76 \pm 0.76	35.7 \pm 1.1	464 \pm 14	-14	0.23
9	20:41:37.491 \pm 0.590	20:42:14.897 \pm 0.160	37.4 \pm 0.6	487 \pm 8	+1.1	0.06
10	20:41:57.914 \pm 0.110	20:42:31.944 \pm 0.136	34.03 \pm 0.17	443 \pm 2	+0.3	0.02
11	20:41:54.60 \pm 0.15	20:42:28.76 \pm 0.12	34.16 \pm 0.19	444 \pm 2	-0.7	0.02
12	20:41:53.551 \pm 2.186 ^(c)	20:42:12.949 \pm 2.387 ^(c)	19 \pm 3	247 \pm 39	-2	0.12

Notes. Chords' shifts are explained in Sect. 4.1. ^(a)A positive shift means a shift toward the east, and a negative one indicates a shift toward the west. ^(b)Chord 3' is the result of merging chords 3 and 4 in Table 3, see Sect. 3.1 for details. ^(c)We have added two seconds to the uncertainty bar given by the square-well fit for chord 12, as explained in Sect. 3.1.

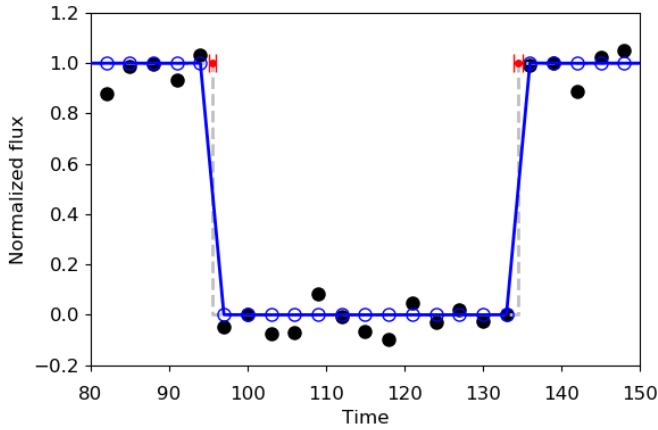


Fig. 3. Best fit of the data from ROASTERR-1 Obs.(2) to the convolved square-well model. The flux from the occulted star plus 2003 VS₂ is plotted with black dots, normalized to the flux of the star while unocculted. Time is in seconds after 2019 October 22–20:40:00 UT. The gray-dashed line represents the initial square-well model; the solid blue line represents the final fit, after convolving the square-well model by Fresnel diffraction, the exposure time, and the stellar diameter, see Sect. 3.2 for details. The open blue circles show the expected flux from this model. The stellar occultation starting and ending times derived are plotted in red, with their uncertainties. For all the positive chords, the star disappearance and reappearance times and chord lengths derived are listed in Table 5.

3.3. Rotational light curve

We performed aperture photometry on 2003 VS₂ and 16 reference stars with good photometric behavior using the routines and procedures mentioned in Fernández-Valenzuela et al. (2016). We tested different synthetic aperture radii and sky annuli to maximize the S/N of the object while minimizing the dispersion of the residuals to the rotational light curve fit. We chose different aperture parameters for each observing day, selecting the same reference stars, and then combined the best photometry results.

The obtained rotational light curve is shown in Fig. 4 and the data used are available at the CDS. We folded the relative flux of 2003 VS₂ versus time using its well-known rotation period of 7.41753 ± 0.00001 h (Santos-Sanz et al. 2017), and we used a fourth-order Fourier function to fit the folded data. The

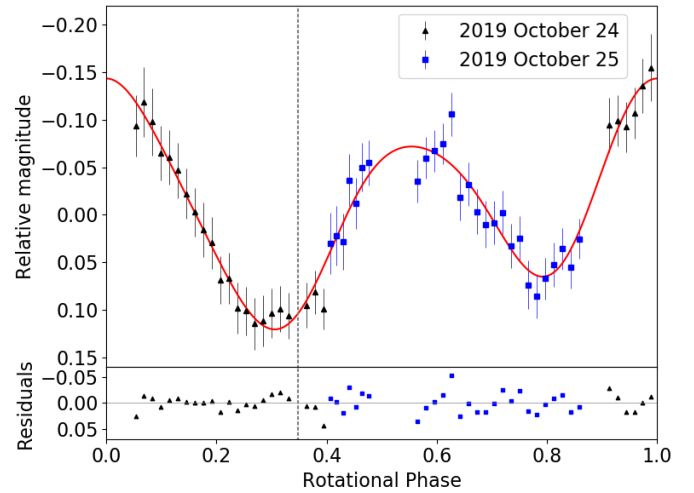


Fig. 4. *Top:* rotational light curve of 2003 VS₂ from data collected on 2019 October 24 (black triangles) and 25 (blue squares). The data were folded using the rotation period of 7.41753 h (Santos-Sanz et al. 2017). The fourth-order Fourier fit is shown as a solid red line. The vertical black-dashed line indicates the rotational phase of 2003 VS₂ at the time of the stellar occultation. The plot has been arbitrarily shifted to make the minimum of the fit (maximum brightness) correspond to rotational phase 0. *Bottom:* differences between the observational data and the fit. Julian dates have not been corrected for light travel time.

obtained peak-to-peak amplitude of the rotational light curve was $\Delta m = 0.264 \pm 0.017$ mag; the nominal value is the one that best fits the data in terms of minimization of the sum of squared residuals and the uncertainty is given as the standard deviation of a Monte Carlo distribution. This amplitude is larger but consistent within the error bars with other values found in the literature (0.23 ± 0.07 mag, 0.21 ± 0.02 mag, 0.21 ± 0.03 mag; Ortiz et al. 2006; Sheppard 2007; Thirouin et al. 2010, respectively). However, the greatest disagreement appears when comparing this work's amplitude with the last 0.141 ± 0.009 mag value found in Benedetti-Rossi et al. (2019), which was obtained with data from 2014. Our confidence is stronger for this work's rotational light curve since the dispersion of the residuals to the fit is lower than the one in Benedetti-Rossi et al. (2019) and the obtained amplitude is more compatible with the ones previously published

in the literature. The slight increase in the rotational light curve amplitude throughout the years may be explained by a change in the object's aspect angle, as found in the case of the TNO Varuna (Fernández-Valenzuela et al. 2019).

We calculated the rotational phase of 2003 VS₂ at the time of the stellar occultation⁵, which was 0.32 with respect to the absolute brightness maximum (see the vertical black-dashed line in Fig. 4). We did not correct the Julian dates for light travel time since this source of error is negligible due to the closeness in time of all data. The rotational phase obtained implies that the apparent surface area of 2003 VS₂ was near its minimum during the occultation event. We note that the rotational phase at the time of the occultation is not influenced by the amplitude of the rotational light curve, but by the rotation period. Given the precision of the rotation period and the time span between the data taken for the occultation event and the rotational light curve, the error calculating the rotational phase is negligible.

4. Results analysis

4.1. Limb fitting

The shape generally considered for a TNO's limb fitting is an ellipse (see, e.g., Ortiz et al. 2017; Benedetti-Rossi et al. 2019), even though the object might not project a perfect regular elliptical shape. In this work we do not have enough positive chords to account for topographic features or deviations from a true elliptical shape, so we can only use the simplest and most general model for the limb fitting.

Five adjustable parameters characterize the considered ellipse: the body's center coordinates relative to the star in the sky plane (f_c , g_c); the apparent semi-major axis a' ; the apparent semi-minor axis b' ; and the tilt angle of the ellipse PA, which is the position angle of the semi-minor axis from celestial north and positive to the west. If we assume the *Gaia* DR2 star position to be correct, the coordinates (f_c , g_c) give the offsets in RA and Dec, respectively, to be applied to the object's adopted ephemeris.

We considered two configurations of the positive chords for the limb fitting. For the first configuration, we kept the relative positions of the original chords but had to correct the absolute time of chord 8 (see Sect. 3.1). In this regard, we performed a weighted⁶ least-squares polynomial fit of degree one to the centers of the chords in Table 5, not using chord 8 for that fit. We then shifted chord 8 until its center laid on said linear regression and thus performed the elliptical fitting to the chords' extremities (Fig. 5a). For the second configuration, we aligned the centers of all the chords using the aforementioned linear regression and then performed the elliptical limb fitting to the chords' extremities (see Fig. 5b).

The alignment of the centers is a condition for the parallel chords of an ellipse, and the needed shifts are not within the centers' uncertainties. Although synchronization via NTP servers provides theoretical accuracies of 0.01 s, uncertainties of up to tenths of a second have been reported to arise from using different operating systems (Barry et al. 2015), camera software,

⁵ To do this, we considered 2458779.36248264 to be the Julian date of the occultation event, as it was the closest recorded date to the average value of the mid-occultation times from all sites. However, choosing a different Julian date close to the stellar occultation would not change the results given the object's rotation period.

⁶ The weights of the chords' centers were calculated as the inverse of $\frac{1}{2} \sqrt{\sigma_{\text{ingress}}^2 + \sigma_{\text{egress}}^2}$.

or even due to delays in the shutter opening. However, adding a nominal average error to all the extremities would be unwise since that error should be the same for the ingress and egress of each chord, and we would be overestimating the error. On the other hand, the necessity of correcting unexplained time shifts has also been reported (Elliot et al. 2010; Braga-Ribas et al. 2013). So by aligning the chords, we account for possible systematic timing errors and small topography that could have decentered the chords.

We obtained the best elliptical fit to the extremities of the chords via minimization of the sum of squared residuals $\sum_i (d_{i,\text{obs}} - d_{i,\text{fit}})^2$, with d_i being the shortest distance between the i -extremity and the evaluated ellipse following the direction of the chord. The uncertainties of the elliptical parameters were determined using a Monte Carlo method. To do this, we generated 10 000 random sets of extremities of the chords for each configuration, sampled from within the uncertainty bars, and then searched for the ellipse that minimizes the aforementioned equation. The Monte Carlo distributions obtained for each of the ellipse parameters are included in Appendix B. In Table 6 we show the nominal value and the standard deviation of each distribution.

The instantaneous area-equivalent diameter of 2003 VS₂ is also given in Table 6. Given that the rotational phase of 2003 VS₂ was near its brightness minimum during the stellar occultation (see Fig. 4), the projected area was also near its minimum, and thus this diameter is a lower limit for the real equivalent diameter of 2003 VS₂. If we take the rotational phase and the rotational light curve amplitude into account, we can derive the mean area-equivalent diameter as follows:

$$D_{\text{mean}} = D_{\text{occ}} \times 10^{(m_{\text{occ}} - m_{\text{mean}})/5}, \quad (1)$$

where m_{occ} is the relative magnitude of 2003 VS₂ during the stellar occultation ($m_{\text{occ}} = 0.104 \pm 0.010$ and $m_{\text{mean}} = 0$, see Fig. 4), and D_{occ} is its instantaneous area-equivalent diameter.

The derived mean area-equivalent diameter of 545 ± 13 km is slightly greater than the value obtained from Herschel radiometric data ($D = 523_{-34.4}^{+35.1}$ km, Mommert et al. 2012), but it is in agreement within the error bars. It is, however, slightly smaller than the area-equivalent diameter of 553_{-33}^{+36} km calculated with the data from Benedetti-Rossi et al. (2019) and applying Eq. (1) to take the rotational phase at the occultation time into account, yet it is still compatible within the uncertainty.

4.2. Geometric albedo

The geometric albedo at the V band of 2003 VS₂ at the time of the stellar occultation can be derived from the following:

$$p_V = 10^{0.4(m_{V,\text{Sun}} - H_{V,\text{occ}})} / (A/\pi), \quad (2)$$

with $m_{V,\text{Sun}}$ being the V magnitude of the Sun ($m_{V,\text{Sun}} = -26.74$ mag), $H_{V,\text{occ}}$ being the instantaneous absolute magnitude of 2003 VS₂ in the V band at the time of the stellar occultation, and A being the projected area of 2003 VS₂ during the event in astronomical units squared.

To obtain $H_{V,\text{occ}}$ we corrected 2003 VS₂'s rotationally averaged H_V (4.14 ± 0.07 mag; Alvarez-Candal et al. (2016) and private communication) by adding the theoretical relative magnitude of 2003 VS₂ during the stellar occultation, which is given by the Fourier fit to the rotational light curve and has a value of 0.104 ± 0.010 mag at the 1σ level.

The obtained geometric albedo of 0.134 ± 0.010 (Table 6) is slightly smaller but compatible with the one derived from

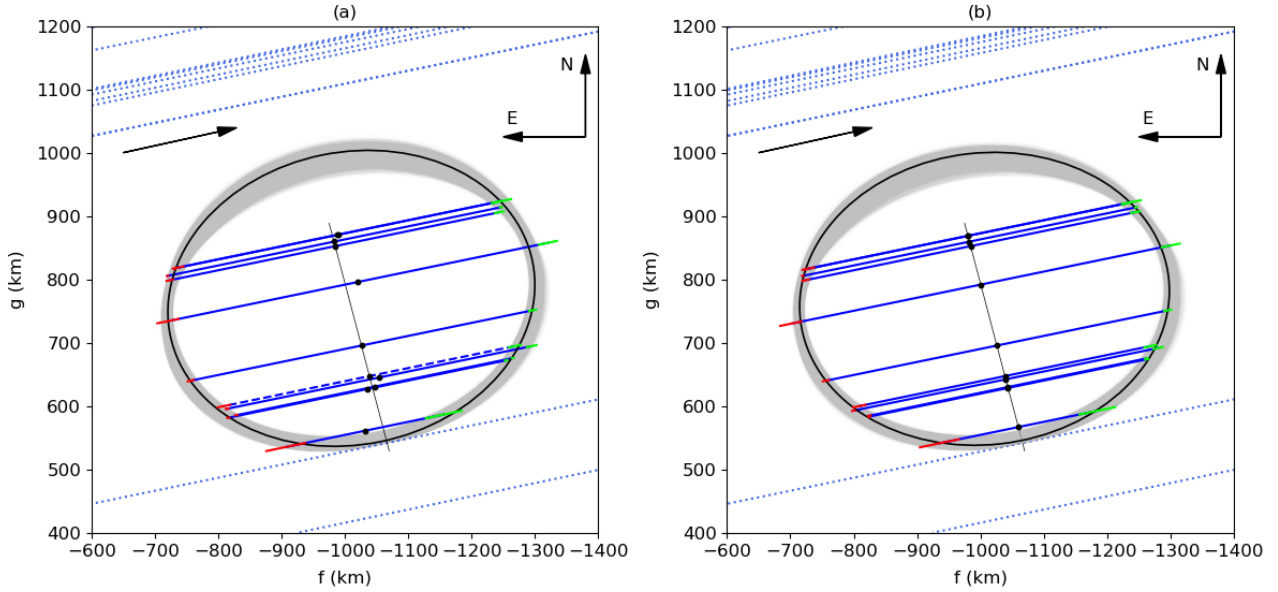


Fig. 5. Elliptical fit to the chords of the stellar occultation for the two considered configurations: (a) original distribution of the chords, with chord 8 (dashed line) already shifted; and (b) final distribution of the chords after aligning their centers using a least squares linear fit. In both plots, the positive chords are shown in solid blue and the negative chords are in dotted blue; uncertainties of the star’s disappearance time are shown in green and those of the reappearance time are in red; and the black dots show the center of the chords. From top to bottom, the chords follow the same order as in Table 3. The limiting negative chord in the north corresponds to observer H. Mikuž and the chord limiting in the south corresponds to observer V. Dumitrescu, see Table 4. The black arrows in the top left of each plot show the direction of the shadow motion. The best elliptical fit to the extremities of the chords is shown in black and all the ellipses from the Monte Carlo distribution are plotted in gray, see Sect. 4.1 for details.

Table 6. Results of the elliptical limb fitting.

Parameter	Original chords ^(a)	Aligned chords
Semi-major axis a' (km)	292 ± 3	293 ± 3
Semi-minor axis b' (km)	231 ± 6	230 ± 6
Oblateness $\epsilon' = (a' - b')/a'$	0.21 ± 0.03	0.21 ± 0.03
Tilt angle ($^\circ$)	-11 ± 2	-6 ± 2
(f_c, g_c) (km) ^(b)	$(-1010 \pm 2, 770 \pm 5)$	$(-1007 \pm 2, 769 \pm 5)$
(f_c, g_c) (mas) ^(b)	$(-38.57 \pm 0.09, 29.40 \pm 0.19)$	$(-38.43 \pm 0.09, 29.37 \pm 0.17)$
Area-equivalent diameter (km)	519 ± 10	519 ± 10
Mean area-equivalent diameter (km)	545 ± 13	545 ± 13
Geometric albedo p_V	0.134 ± 0.010	0.134 ± 0.010

Notes. The nominal values were obtained by minimization of a least squares function and the uncertainties are the standard deviation of the Monte Carlo distributions, see Sect. 4.1. The corresponding distributions can be seen in Appendix B. The uncertainty bars of the equivalent diameters and albedo were obtained analytically from their respective mathematical expressions. ^(a)Chord 8 is the only chord that has been shifted for this case, see Sect. 4.1. ^(b)Offsets with respect to the JPL #33 orbit.

the combination of Herschel thermal measurements and Spitzer data ($0.147^{+0.063}_{-0.043}$, Lellouch et al. 2013). It is also in agreement with the last derived albedo from Benedetti-Rossi et al. (2019) ($0.131^{+0.024}_{-0.013}$), see Table 1. If we use the same absolute magnitude as Lellouch et al. (2013) ($H_V = 4.11 \pm 0.38$ mag) to calculate the albedo, it has a value of 0.14 ± 0.05 , which is still smaller but compatible within the error bars.

4.3. Size and shape

The double-peaked rotational light curve of 2003 VS₂ suggests that it is either a rotating triaxial ellipsoid or an oblate spheroid with a significant irregularity or a large albedo variation along its surface; this second case is unlikely, especially considering 2003 VS₂’s large light curve amplitude. In addition, the two maxima and minima per rotation cycle are different,

which usually indicates a triaxial shape with some albedo spots or small topographic features, since we do not expect a perfectly symmetrical and homogeneous body orbiting in space. This is the case for many observed TNOs with large rotational light curve amplitudes, such as Haumea (Lacerda 2010), Varuna (Fernández-Valenzuela et al. 2019), and 2008 OG₁₉ (Fernández-Valenzuela et al. 2016). On the other hand, if not triaxial, the rotation period of 2003 VS₂ would be 3.6 h, which is probably too fast for a TNO (no TNO is known to rotate that fast). Hence, assuming a triaxial ellipsoidal shape for 2003 VS₂, we searched for the axes $a > b > c$ of a triaxial body, rotating around its c axis, which would give the observed elliptical projection during the stellar occultation (Fig. 5) while also showing an amplitude of $\Delta m = 0.264 \pm 0.017$ mag on its rotational light curve.

To do this, we used the procedures described in Gendzwil & Stauffer (1981) to generate ellipsoids characterized by three

Table 7. Results of the 3D fit of 2003 VS₂.

Parameter	Observed Δm		$\Delta m_{\min} = 0.18$ mag	
	Original chords ^(a)	Shifted chord	Original chords ^(a)	Shifted chords
Semimajor axis a (km)	339 ± 5	338 ± 5	327 ± 5	330 ± 5
Semimajor axis b (km)	235 ± 6	233 ± 7	233 ± 6	235 ± 9
Semimajor axis c (km)	226 ± 8	227 ± 7	228 ± 8	222 ± 8
Aspect angle θ (°)	59 ± 2	57 ± 3	51 ± 2	51 ± 6
Rotational light curve amplitude (mag)	0.2497 ± 0.0008	0.2489 ± 0.0009	0.187 ± 0.003	0.184 ± 0.003
a/b	1.44 ± 0.06	1.45 ± 0.06	1.41 ± 0.06	1.40 ± 0.07
b/c	1.040 ± 0.017	1.022 ± 0.018	1.02 ± 0.02	1.06 ± 0.03
Spherical volume equivalent diameter (km)	524 ± 7	523 ± 7	517 ± 7	516 ± 8

Notes. We present the ellipsoids obtained using the observed rotational light curve amplitude ($\Delta m = 0.264 \pm 0.017$ mag) and the ones allowing for bigger albedo spots, see Sect. 4.3. The nominal values of the parameters are the ones obtained from the optimal ellipses in Table 6. The uncertainties are given as the standard deviation of the Monte Carlo distributions shown in Appendix B for the two chords' configurations considered for the limb fitting. ^(a)Chord 8 is the only chord that has been shifted for this case, see Sect. 4.1.

semi-major axes and three orientation angles with respect to the Cartesian system, and to then project these ellipsoids into a plane perpendicular to the line of sight in order to compare that projection with the instantaneous limb of 2003 VS₂ during the stellar occultation (Fig. 5). An additional constraint to the model is the observed rotational light curve amplitude of 2003 VS₂ (Fig. 4). To include it, we implemented the well-known relation between the peak-to-peak variation of the rotational light curve of a small body and its three principal semiaxes (Binzel et al. 1989):

$$\Delta m = 2.5 \log\left(\frac{a}{b}\right) - 1.25 \log\left(\frac{a^2 \cos^2(\theta) + c^2 \sin^2(\theta)}{b^2 \cos^2(\theta) + c^2 \sin^2(\theta)}\right), \quad (3)$$

where θ is the polar aspect viewing angle, that is, the angle between the rotation axis (c , in this case) and the line of sight.

For each of the ellipses obtained via Monte Carlo for the limb fitting (Sect. 4.1, Fig. 5), we searched for the ellipsoid that would give the most similar projection in terms of least squares minimization. Although the limb fit obtained in Sect. 4.1 has virtually the same values for both considered chord configurations, we decided to search for the corresponding ellipsoid of both solutions to see if the slight difference in the tilt angle would give different 3D shapes. The obtained distributions for the ellipsoid's semiaxes, aspect angle, and the corresponding rotational light curve amplitude derived from Eq. (3) are shown in Appendix B. In Table 7 we show the best 3D fit, which is the one corresponding to the best limb fit in Table 6; also the uncertainty bars are given as the standard deviation of the Monte Carlo distributions. According to the results, the obtained triaxial ellipsoid would produce a variation of 0.24 mag during its rotation so, as stated before, the remaining observed rotational light curve amplitude might be due to albedo spots or topographic features that cannot be studied with the available data; we would either need more chords from the stellar occultation to study topographic features, or rotational light curves obtained with different filters to study albedo variations. The differences in the rotational light curve amplitude due to albedo spots would be on the order of a few cents of a magnitude; this is the range of variability due to albedo variations observed in large TNOs such as Varuna, Haumea, 2008 OG₁₉. Considering this and given that a model of a triaxial body without albedo spots to explain the light curve amplitude is a simplification, we also computed the ellipsoid accepting solutions that give a rotational light curve amplitude of 0.06 mag less than the observed value, so as to allow for large albedo

spots. These results are also included in Table 7 and Appendix B. Finally, the spherical volume equivalent diameter ($D_{V_{eq}}$) derived from these parameters is also presented in Table 7.

Both obtained solutions are almost compatible with the one presented in Benedetti-Rossi et al. (2019) within the errors. The differences might arise from the fact that, in the previous work, the calculations were simplified by considering that 2003 VS₂ was on its brightness maximum during the 2014 occultation, when the actual rotational phase at the time of the occultation was +0.07 with respect to the brightness maximum. Moreover, the rotational light curve amplitude they used was smaller than the one used in this work. However, we checked that the two amplitudes cannot be explained simultaneously by a change in the aspect angle.

We derived the ecliptic coordinates of the pole (λ_p, β_p) by simultaneously minimizing the difference between the observed rotational light curve amplitude and the one derived using Eq. (3), and the difference between the aspect angle δ derived in Sect. 4.3 and the one derived via the following:

$$\delta = \frac{\pi}{2} - \arcsin\left[\sin\beta_e \sin\beta_p + \cos\beta_e \cos\beta_p \cos(\lambda_e - \lambda_p)\right], \quad (4)$$

with (λ_e, β_e) being the ecliptic coordinates of the sub-Earth point in the object-centered reference frame at the time of this work's stellar occultation. The pole's ecliptic coordinates (λ_p, β_p) derived are (228°, 39°) (for a more detailed explanation, see Fernández-Valenzuela (2022) and references therein). Combining the pole's coordinates, 2003 VS₂'s ephemeris, and the 3D axes obtained in Sect. 4.3, we obtained the theoretical variation of the rotational light curve amplitude through the years and compared it to the values in the literature. The results are plotted in Fig. 6. It can be seen that the previously published results were obtained during the minimum of the rotational light curve amplitude, but it has been increasing since 2005. The published value found in Benedetti-Rossi et al. (2019) is much lower than theoretically expected, but the remaining measurements do agree with the model. We suspect that the rotational light curve in Benedetti-Rossi et al. (2019) may have been contaminated by a background star or had some unidentified technical problem as of yet or a reduction artifact because there are no physical scenarios that can explain a sudden decrease in the amplitude, except perhaps a sudden brightening due to dust release through sublimation activity or through a collision, but this is very unlikely.

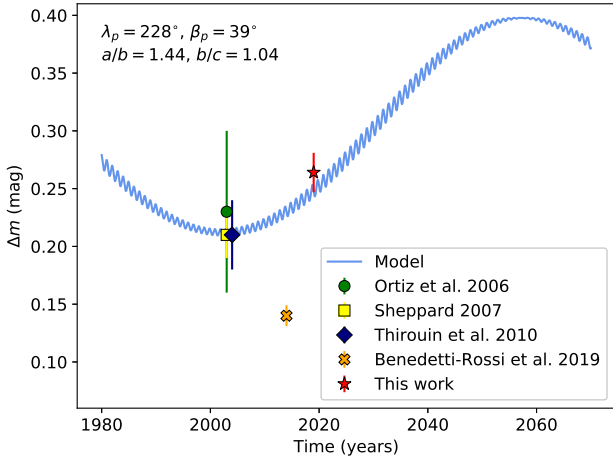


Fig. 6. Theoretical long-term variability of 2003 VS₂'s rotational light curve amplitude and observational results found in the literature.

The projection of the ellipsoids obtained from the original-chords' configuration at the rotational phase in which the stellar occultation in [Benedetti-Rossi et al. \(2019\)](#) happened (+0.07 with respect to the brightness maximum) gives ellipses with semiaxes of $a \times b = 305 \pm 4 \times 230 \pm 6$ km (for the observed rotational light curve amplitude) and $a \times b = 302 \pm 4 \times 230 \pm 6$ km (for a minimum rotational light curve amplitude of 0.18 mag). We note that they are in agreement with the limb fit reported in [Benedetti-Rossi et al. \(2019\)](#).

To conclude, we can compare these results with the theoretical axes ratios of a triaxial ellipsoid with the rotation period of 2003 VS₂ and a homogeneous density to see if the derived shape is compatible with a hydrostatic-equilibrium figure. Using the formalism from [Chandrasekhar \(1987\)](#), we searched for the theoretical values of the axes ratios a/b and b/c (solid and dashed lines in [Fig. 7](#), respectively) of a triaxial ellipsoid with a rotation period of 7.41753 ± 0.00001 h and homogeneous densities between 600 and 1000 kg m^{-3} . We also plotted in [Fig. 7](#) the axes ratios' bands derived from the stellar occultation (their values are presented in [Table 7](#)). In the figure, one can see that there is an overlapping region between the theoretical and observed axes ratio a/b , but that is not the case for the axes ratio b/c . Therefore, we can conclude that the derived shape of 2003 VS₂ is not consistent with the hydrostatic-equilibrium figure of a homogeneous body with the rotation period of 2003 VS₂ for any density value. This result suggests that, similar to the case of Haumea ([Ortiz et al. 2017](#)), we might need to consider granular physics to explain the body's shape, because in this case, a differentiated body is less plausible due to 2003 VS₂'s smaller size. In fact, the differentiation obtained by [Loveless et al. \(2022\)](#) for icy bodies with radii larger than 200 km, though not negligible, is so small that it would probably not produce a significant deviation from the equilibrium shape for a homogeneous body of the size of 2003 VS₂. However, that model excludes or simplifies some of the physical and chemical processes and considers pure spherical bodies, so we cannot discard some differentiation. The possibility that 2003 VS₂ might be sustaining some stress seems more plausible but, in reality, we can only speculate that the deviation of its shape from that of pure hydrostatic equilibrium could be due to one or a combination of both scenarios.

5. Conclusions

On 2019 October 22, we observed the stellar occultation of the *Gaia* source 3449076721168026496 ($m_V = 14.1$ mag) caused by

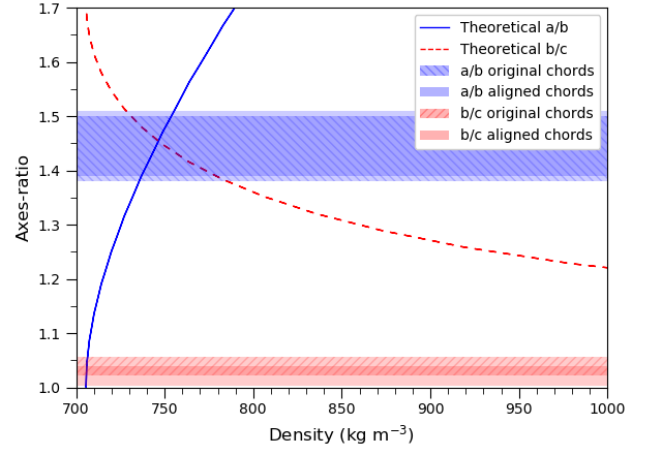


Fig. 7. Theoretical axes ratios a/b (solid blue line) and b/c (dashed red line) of a 3D body in hydrostatic equilibrium, rotating with a period of 7.41753 ± 0.00001 h, for different densities ([Chandrasekhar 1987](#)). The axes ratios from the observed data are plotted as colored bands for both of the considered cases. The theoretical b/c ratios from hydrostatic equilibrium do not agree with the observations for any possible density of a homogeneous body because there is no intersection of the dashed red line with the pink band.

the plutino (84922) 2003 VS₂. Out of the 39 participant observing sites, 12 reported a positive detection, located in Bulgaria, Romania, and Serbia. Two positive chords were combined forming a single one, giving 11 effective positive chords. This is one of the best observed stellar occultations by a TNO so far.

The projected shape of 2003 VS₂ was fitted to an ellipse considering two configurations of the positive chords. For the first one, we fit the extremities of the original chords obtained from the stellar occultation after shifting only one chord due to a problem with its absolute time. For the second case, we shifted all the positive chords to align their centers. The best solution for the instantaneous limb of 2003 VS₂ was obtained by minimizing the sum of squared residuals, and the uncertainties were derived via a Monte Carlo method and correspond to the standard deviation of the obtained distributions. The result has virtually the same value for both configurations, as can be seen in [Table 6](#); for the original configuration, it has semiaxes of $a'_1 = 292 \pm 3$ km and $b'_1 = 231 \pm 6$ km, and a tilt angle of $-11^\circ \pm 2^\circ$.

We carried out photometric observations of 2003 VS₂ two days after the stellar occultation and derived an amplitude of $\Delta m = 0.264 \pm 0.017$ mag from the obtained rotational light curve. This value is slightly greater than the ones previously published for this object, but it can be explained via a change in the aspect angle of the body. Taking this and the rotational phase during the stellar occultation into account, the mean area-equivalent diameter of 2003 VS₂ is $D_{\text{Aeq}} = 545 \pm 13$ km. The derived geometric albedo during the stellar occultation is $p_V = 0.134 \pm 0.010$. These values are in agreement with the ones published in [Benedetti-Rossi et al. \(2019\)](#), and with the diameter derived from thermal models and the radiometric albedo obtained with Herschel and Spitzer data within the error bars ($D = 523^{+35.1}_{-34.4}$ km and $p_V = 0.147^{+0.063}_{-0.043}$, [Mommert et al. 2012](#)).

We derived the 3D shape of 2003 VS₂ by combining its rotational light curve and the instantaneous projection and assuming a triaxial ellipsoidal shape. The obtained ellipsoid has principal semiaxes of $a = 339 \pm 5$ km, $b = 235 \pm 6$ km, and $c = 226 \pm 8$ km, with an aspect angle of either $\theta_1 = 59^\circ \pm 2^\circ$ or its supplementary $\theta_2 = 121^\circ \pm 2^\circ$, depending on the considered sense of rotation. If we allow for a 0.1 mag variability due to albedo spots

in the 3D model, the resulting ellipsoid has a semi-major axis ~ 10 km smaller and fully compatible with the projected shape in [Benedetti-Rossi et al. \(2019\)](#). These values give a spherical-volume equivalent diameter of $D_{V_{\text{eq}}} = 524 \pm 7$ km. This solution is not compatible with a homogeneous body in hydrostatic equilibrium rotating with the known period of 2003 VS₂, requiring differentiation or, most likely, an internal structure that can sustain stress to some degree. Finally, we found no evidence of a dense ring or debris material orbiting around 2003 VS₂ of the type seen in Chariklo ([Braga-Ribas et al. 2014](#)).

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Appendix A: Observing sites involved in the occultation campaign that could not observe

In this appendix we include the list of contacted observing sites that could not observe due to bad weather or technical problems.

Table A.1: Details of the observing sites involved in the 2003 VS₂ 2019 October 22 occultation that could not observe.

Site name Location MPC Code	Latitude (N) Longitude (E)	Telescope aperture (cm)	Observers	Problem
Galați Observatory (×2) Romania C73	45° 25' 7".9 28° 1' 57"	40 20	O. Tercu A. M. Stoian G. Neagu D. Zlat	Bad Weather
Astroclubul București (×2) Romania (mobile telescope) –	–	40 20	D. Berteșteanu M. H. Naiman Z. Deak	Technical problems
Romania (mobile telescope) –	–	20	B. Dumitru	Technical problems
Bârlad Observatory Romania L22	46° 13' 54".1 27° 40' 10".2	20	C. Vantdevara	Bad weather
G. V. Schiaparelli Observatory Varese, Italy 204	45° 52' 04" –08° 46' 15"	84	L. Buzzi	Bad weather
San Marcello Pistoiese Italy 104	44° 3' 46".9 10° 48' 15".1	60	P. Bacci	Technical problems
ISON Uzhgorod Observatory Derenivka, Ukraine K99	48° 33' 48".2 22° 27' 12".6	40	V. Kudak	Technical problems
Odessa-Mayaki Ukraine 583	46° 23' 50".2 30° 16' 18".1	80	Y. Krugly I. Belskaya	Bad weather
Zalistci Ukraine L18	48° 50' 53".9 26° 43' 5".9	30	Y. Krugly I. Belskaya	Bad weather
Uzhgorod-Derenivka, Ukraine K99	48° 33' 48".2 22° 27' 12".599	40	Y. Krugly I. Belskaya	Bad weather
Kyiv-Comet Station Ukraine 585	50° 17' 52".8 30° 31' 28".6	70	Y. Krugly I. Belskaya	Bad weather
Kharkiv-Chuguev Station Ukraine 121	49° 38' 34".6 36° 56' 12".8	70	Y. Krugly I. Belskaya	Bad weather
Hvar Observatory Hvar, Croatia –	43° 10' 42" 16° 26' 54"	106	D. Ruzdjak S. Cikota	Bad weather

Appendix B: Monte Carlo distributions

In this appendix we present the obtained Monte Carlo distributions for the limb fitting and the 3D fitting of 2003 VS₂ discussed in Sects. 4.1 and 4.3, respectively, for the two considered distributions of the positive chords.

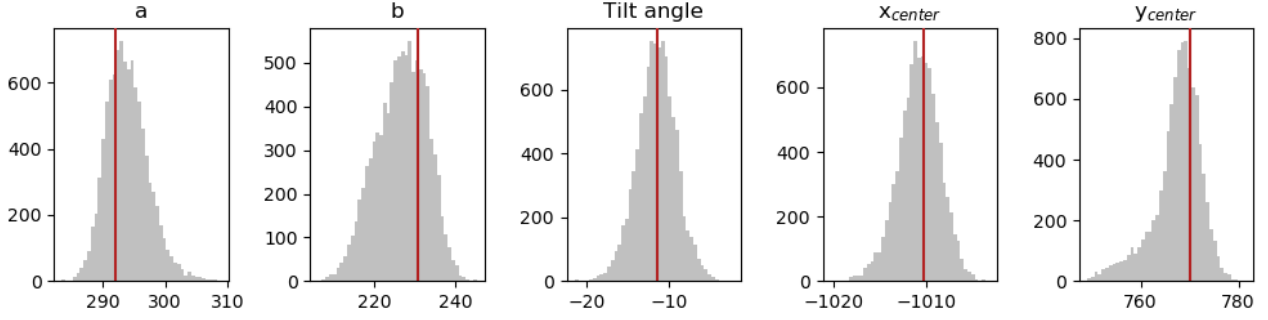


Fig. B.1: Monte Carlo distributions for the semiaxes a and b (in km), the tilt angle (in degrees), and the coordinates (x,y) of the center of the ellipse, from the elliptical fit to the unshifted chords (with only chord 8 shifted) in Fig. 5a, see Sect. 4.1. The vertical red lines show the value of the parameters for the best elliptical fit via minimization of the sum of squared residuals, see Table 6.

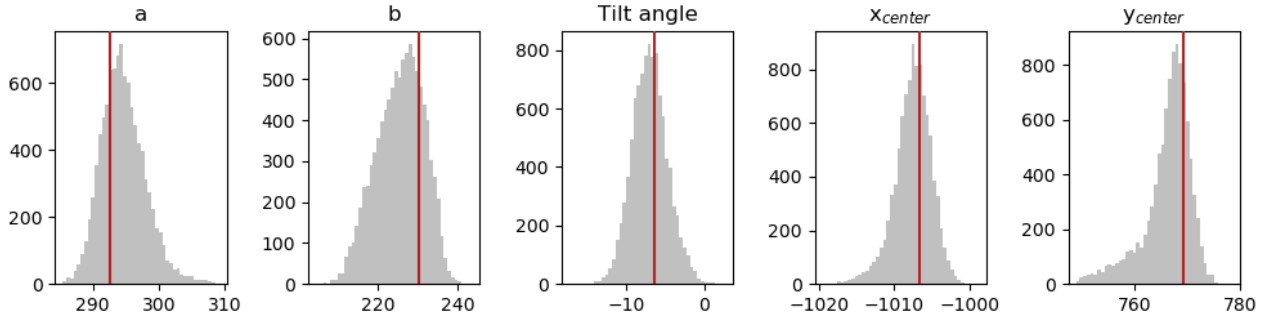


Fig. B.2: Monte Carlo distributions for the semiaxes a and b (in km), the tilt angle (in degrees), and the coordinates (x,y) of the center of the ellipse, from the elliptical fit to the aligned chords in Fig. 5b, see Sect. 4.1. The vertical red lines show the value of the parameters for the best elliptical fit via minimization of the sum of squared residuals, see Table 6.

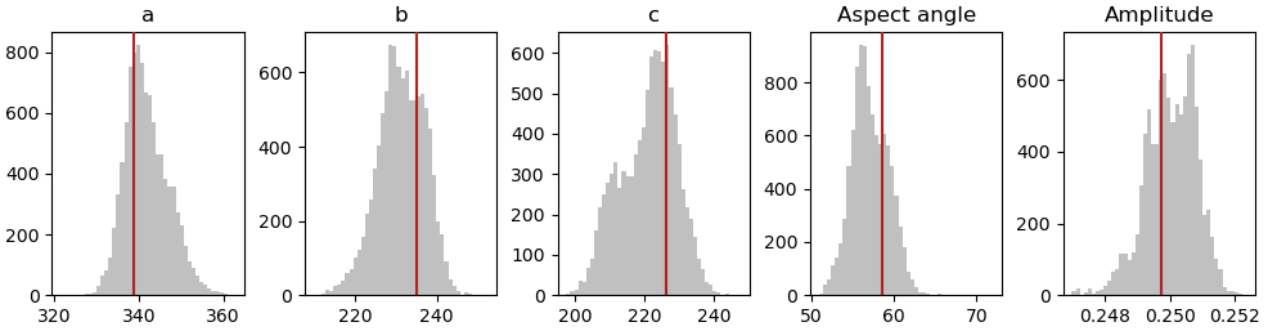


Fig. B.3: Monte Carlo distributions for the semiaxes a , b , and c (in km) of a triaxial ellipsoid compatible with the occultation observation and observed rotational light curve amplitude, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of the original chords with only chord 8 shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

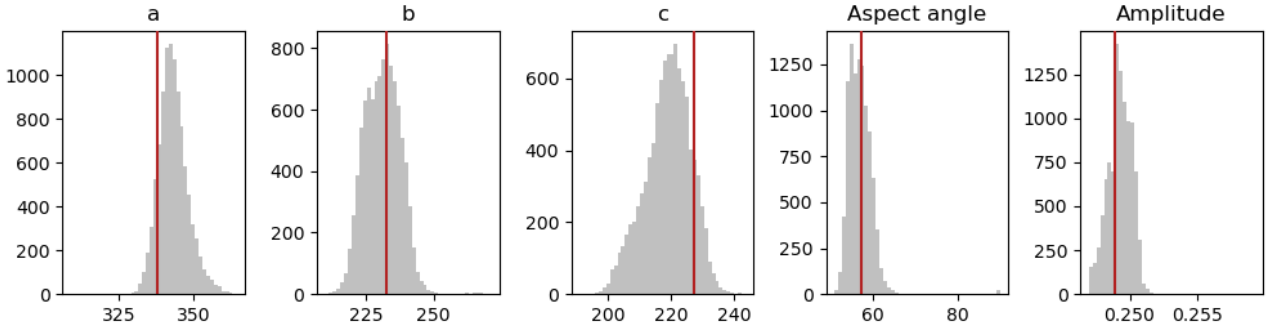


Fig. B.4: Monte Carlo distributions for the semiaxes a , b , and c (in km) of a triaxial ellipsoid compatible with the occultation observation and observed rotational light curve of 2003 VS₂, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of all the chords shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

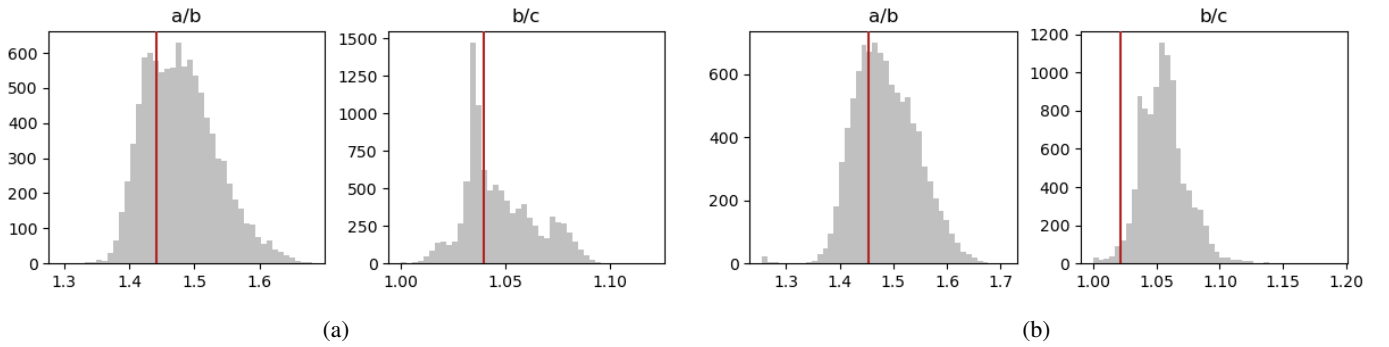


Fig. B.5: Monte Carlo distributions of the axes ratios a/b and b/c of the fitted ellipsoid for the considered case of (a) the chords unshifted with only chord 8 shifted and (b) the aligned chords, for the observed rotational light curve amplitude. The vertical red lines show the axes ratios derived from the best elliptical fit, see Table 7.

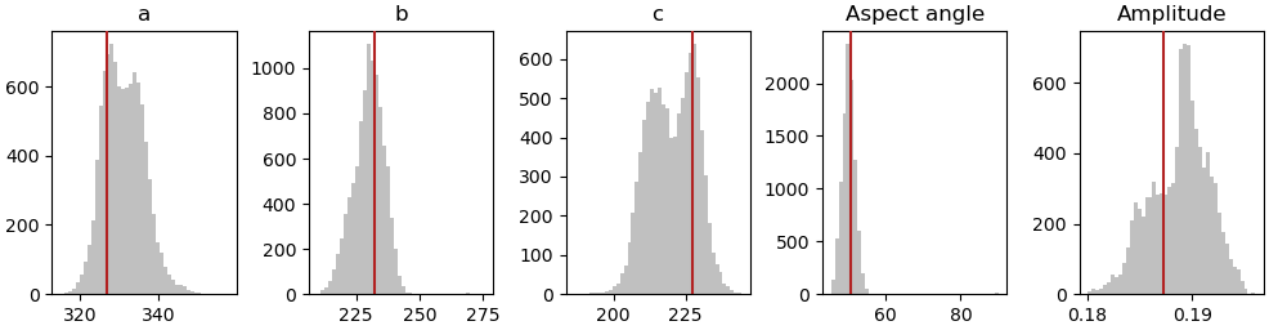


Fig. B.6: Monte Carlo distributions for the semiaxes a , b , and c (in km) of a triaxial ellipsoid compatible with the occultation observation and allowing a minimum rotational light curve amplitude of 0.18 mag, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of the original chords with only chord 8 shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

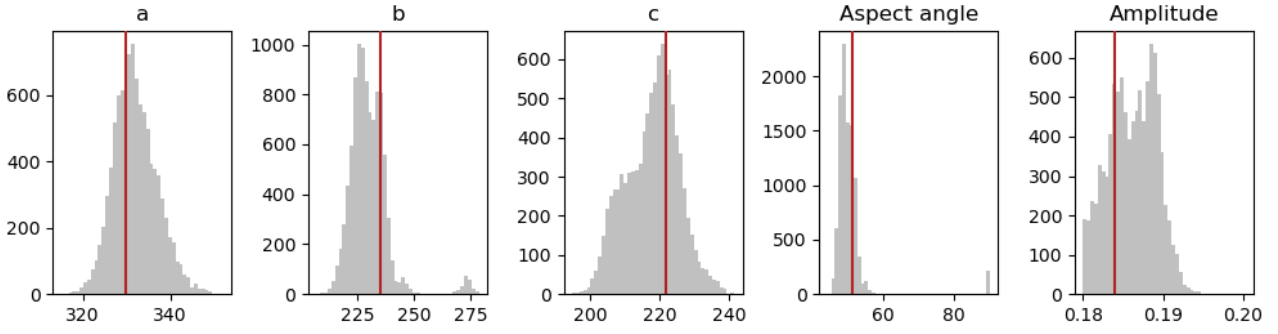


Fig. B.7: Monte Carlo distributions for the semiaxes a , b , and c (in km) of a triaxial ellipsoid compatible with the occultation observation and allowing a minimum rotational light curve amplitude of 0.18 mag, as well as distributions of the aspect angle and derived rotational light curve amplitude obtained via Eq. 3, for the case of all the chords shifted; see Sects. 4.1 and 4.3. The vertical red lines show the ellipsoidal values corresponding to the best elliptical fit, see Table 7.

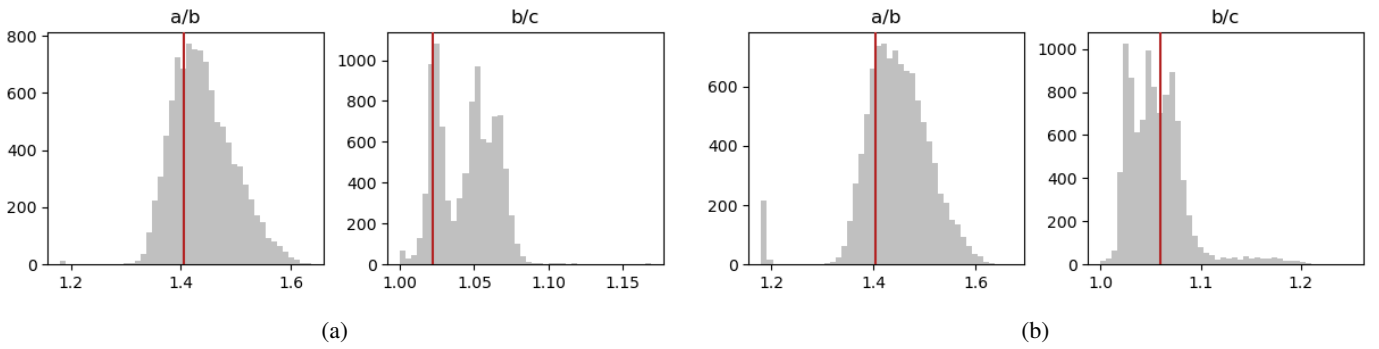


Fig. B.8: Monte Carlo distributions of the axes ratios a/b and b/c of the fitted ellipsoid for the considered case of (a) the chords unshifted with only chord 8 shifted and (b) the aligned chords, if we allow a minimum rotational light curve amplitude of 0.18 mag. The vertical red lines show the axes ratios derived from the best elliptical fit, see Table 7.