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


















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Scylla. I. A Pure-parallel, Multiwavelength Imaging Survey of the ULLYSES Fields in the LMC and SMC

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Abstract

Scylla is a deep Hubble Space Telescope (HST) survey of the stellar populations, interstellar medium, and star formation in the LMC and SMC. As a pure-parallel complement to the Ultraviolet Legacy Library of Young Stars as Essential Standards (ULLYSES) survey, Scylla obtained 342 orbits of ultraviolet (UV) through near-IR imaging of the LMC and SMC with Wide Field Camera 3. In this paper, we describe the science objectives, observing strategy, data reduction procedure, and initial results from our photometric analysis of 96 observed fields. Although our observations were constrained by ULLYSES primary exposures, we imaged all fields in at least two filters (F475W and F814W) and 64% of fields in at least three and as many as seven WFC3 filters spanning the UV to IR. Overall, we reach average 50% completeness of $m_{F225W} = 26.0$, $m_{F275W} = 26.2$, $m_{F336W} = 26.9$, $m_{F475W} = 27.8$, $m_{F814W} = 25.5$, $m_{F110W} = 24.7$, and $m_{F160W} = 24.0$ Vega mag in our photometric catalogs, which is faintward of the ancient main-sequence turnoff in all filters. The primary science goals of Scylla include characterizing the structure and properties of dust in the MCs, as well as their spatially resolved star formation and chemical enrichment histories. Our images and photometric catalogs, which represent the widest-area coverage of MCs with HST photometry to date, are available as a high-level science product at the Barbara A. Mikulski Archive for Space Telescopes.

Unified Astronomy Thesaurus concepts: [Magellanic Clouds \(990\)](#); [Hubble Space Telescope \(761\)](#); [Surveys \(1671\)](#); [Catalogs \(205\)](#); [Star formation \(1569\)](#); [Interstellar medium \(847\)](#)

Materials only available in the [online version of record](#): machine-readable tables

1. Introduction

The LMC and SMC (combined MCs) are the most massive dwarf satellites of the Milky Way (MW). Located at distances of ~ 50 – 60 kpc (V. Scowcroft et al. 2016; G. Pietrzynski et al. 2019), individual stars in the LMC and SMC can be resolved, as well as the individual structures in the interstellar medium (ISM) from which they formed (e.g., molecular clouds) and into which they feed back (e.g., H II regions, supernova remnants). Furthermore, the MCs feature a diverse range of interstellar conditions, including metallicity ($Z < 0.5 Z_{\odot}$; S. C. Russell & M. A. Dopita 1992), gas density (e.g., S. Kim et al. 1999; S. Stanimirović et al. 1999; T. Wong et al. 2011; K. E. Jameson et al. 2016; N. M. Pingel et al. 2022), and radiation field (e.g., F. Galliano et al. 2011; J. Chastenot et al. 2017, 2019). These systems therefore provide excellent

laboratories for studying the detailed properties of stellar populations and the ISM in low-metallicity environments, which are markedly different from those found in the MW.

1.1. Open Questions on the MCs

Despite the proximity of the MCs, fundamental questions regarding the nature of their ISM and stellar populations remain open owing to strong observational limitations. Both systems are so rich in neutral hydrogen (H I) that emission at 21 cm can suffer from strong optical depth effects (e.g., J. Dempsey et al. 2022). Molecular hydrogen (H₂), the most abundant molecule in the ISM, is not observable in cold, star-forming conditions in the MCs owing to its lack of excitation at low temperature. Carbon monoxide (CO), a popular H₂ tracer, cannot self-shield (unlike H₂) and therefore fails to trace the total molecular gas reservoir, especially in the low- Z MCs, where shielding by dust is minimized, and also within strong radiation fields, where molecule destruction is enhanced (A. D. Bolatto et al. 2011). As a result, we still lack a comprehensive census of the total

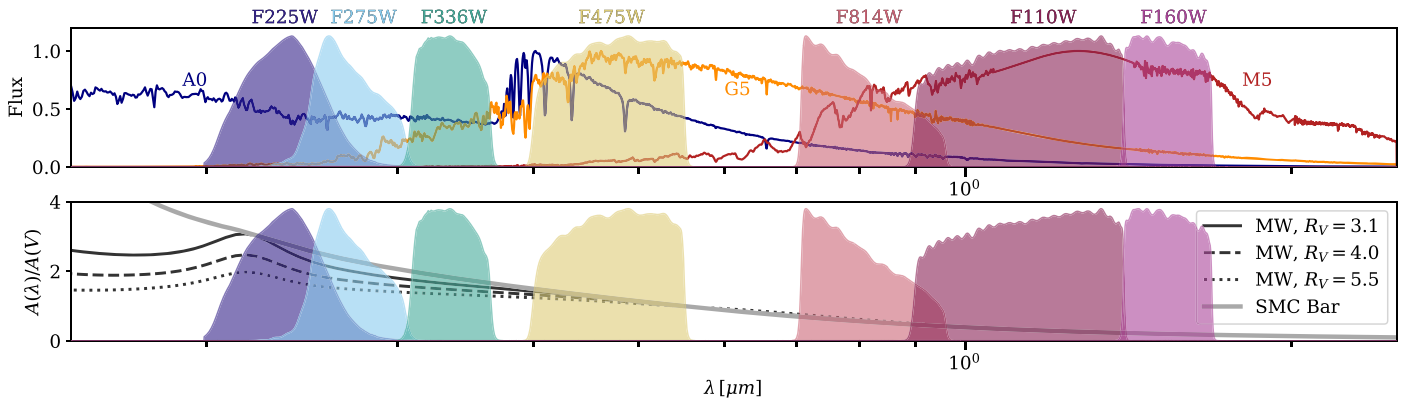


Figure 1. Top: representative stellar spectra ($F_\lambda(\lambda)$) normalized to have the same peak flux, from three stars (A0, G5, M5 types; A. J. Pickles 1998), compared with the transmission curves for six HST/WFC3 filters spanning the UV (F275W, F336W), optical (F475W, F814W), and NIR (F110W, F160W), to illustrate how distinguishing between intrinsic stellar properties (i.e., variation in temperature between A0 and M5 stars) and the effects of dust reddening requires broad wavelength coverage. Bottom: dust extinction curves of varying R_V and curve shape.

ISM mass of the MCs and how it is distributed as a function of phase (i.e., density and temperature), which severely limits our ability to diagnose the star formation efficiency and baryon cycle of these systems.

Fortunately, dust grains are distributed throughout the ISM and thus provide a convenient tracer of all phases. However, converting between dust emission in the IR—a ubiquitous tracer of dust at all redshifts—and mass is complicated. First of all, the physical properties of dust grains, including their size, shape, and composition, are known to vary *significantly* within and between galaxies (e.g., Y. C. Pei 1992; K. D. Gordon et al. 2003; G. C. Clayton et al. 2015; E. F. Schlafly et al. 2017). These properties determine the manner in which dust will absorb, scatter, and emit light as a function of wavelength (e.g., B. T. Draine 2003), which is crucial for interpreting dust observations. In addition, key ingredients of dust models are wildly uncertain. For example, the value of the dust mass absorption coefficient—required to estimate dust mass from dust emission—varies by an order of magnitude in the literature (C. J. R. Clark et al. 2019).

Even with an accurate census of the total ISM budget in the MCs, we still lack an observational view of how the ISM moves in these nearby systems. How and where stellar feedback carves holes into the surrounding ISM is fundamental for understanding mass and radiation escape into the ISM. In addition, the momentum injected by starbursts can trigger secondary generations of star formation at the intersections of colliding gas flows. Disentangling the expansion versus collapse of individual ISM structures on scales of tens of parsecs, as well as large-scale outflows and inflows on kiloparsec scales, is crucial for understanding how these systems form stars, enrich the ISM, and feed back into their circumgalactic environments. The crucial missing piece of information is the distance along the line of sight to individual ISM structures, which cannot be directly constrained by standard (emission, absorption) observations.

In addition, the physical mechanisms responsible for converting the dusty ISM into stars and planets remain uncertain (R. C. Kennicutt & N. J. Evans 2012). Standard empirical approximations for the star formation rate (SFR) and its evolution, used widely to infer how stars form, evolve, and feed back, are coarsely calibrated in time ($\Delta t \sim 100$ Myr) and derived from global scales ($\gtrsim 1$ kpc; R. C. Kennicutt & N. J. Evans 2012; A. K. Leroy et al. 2013). Not only are

these models known to break down on the scales of individual clouds (A. Schrubba et al. 2010; J. M. D. Kruijssen et al. 2014), but they are also insufficient to describe the star formation histories (SFHs) of local galaxies (K. B. W. McQuinn et al. 2010), MCs included.

For the MCs in particular, the influence of their rich history of dynamical interactions on their SFH is significant. How and where they are on their orbit through the MW halo are key open questions. Beyond the unique characteristics of the LMC/SMC system, these galaxies provide a high-resolution window into the effects of dwarf galaxy interactions on SFHs in a broader context. Resolving precisely how interactions trigger bursts of star formation is crucial for testing theories of how galaxy interactions and feedback drive inside-out versus outside-in star formation, which need detailed, radial SFHs at the oldest ages to test (e.g., G. S. Stinson et al. 2009; K. El-Badry et al. 2016).

1.2. How Scylla Will Solve MC Mysteries

A powerful technique for deciphering the intrinsic properties of stars in the MCs, as well as the nature of the intervening ISM, is to observe resolved stellar populations across a wide wavelength range. The ability to resolve individual sources is crucial to distinguish between the degenerate effects of the relative geometry between stars and dust, dust radiative transfer, and SFH on stellar observations. Wide wavelength coverage is also important—it is well known that optical photometry alone is insufficient to disentangle the effects of dust extinction and stellar temperature on observed stellar properties (K. D. Gordon et al. 2016). In Figure 1, we illustrate this point by comparing synthetic stellar spectra (top) and dust extinction curves (bottom; e.g., parameterized by R_V , the total-to-selective extinction ratio) across ultraviolet (UV) through infrared (IR) wavelengths. It is clear that in order to simultaneously and precisely distinguish between stars of different stellar types, as well as the amount and nature of dust extinction, observations spanning UV through near-IR (NIR) wavelengths are required. While UV coverage is critical for measuring dust extinction, the IR allows us to observe stars in dustier regions, improving our stellar completeness and the range of extinction we can probe (for more details on the filter selection justification, see Section 3.2 in J. J. Dalcanton et al. 2012).

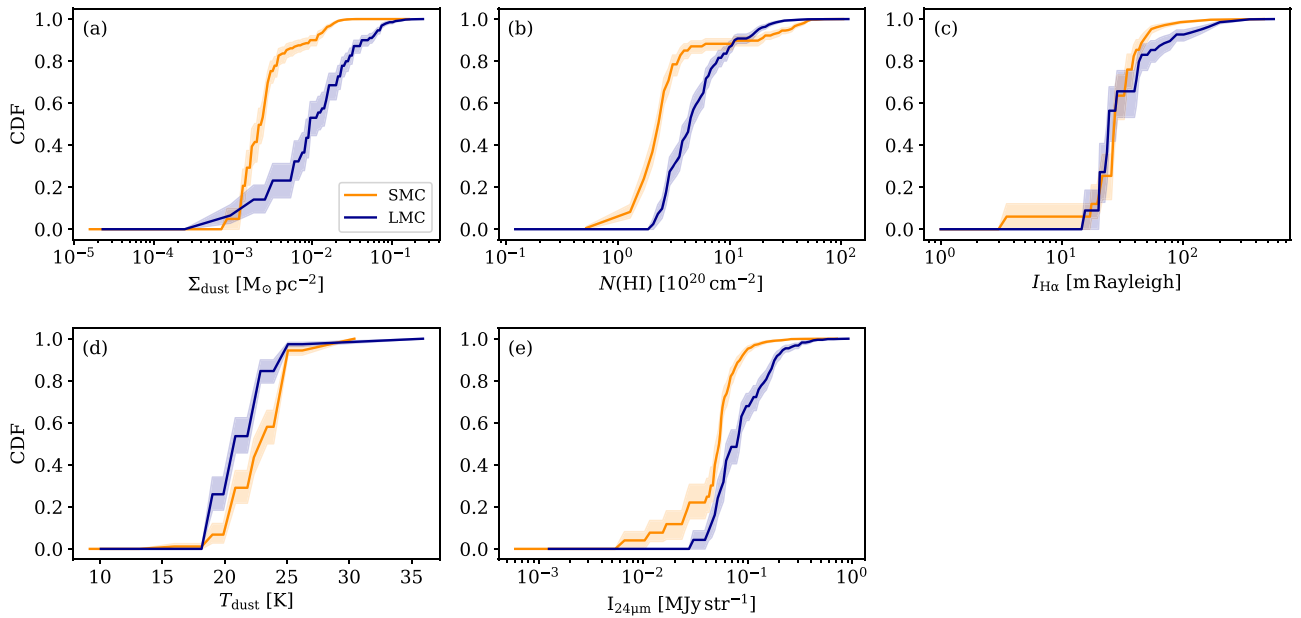


Figure 2. Cumulative distribution functions of ancillary tracers of the ISM at the position of Scylla fields in the SMC (orange) and LMC (blue). (a) Dust surface density (Σ_{dust} ; C. J. R. Clark et al. 2023); (b) H I column density ($N(\text{HI})$; S. Kim et al. 1999; N. M. Pingel et al. 2022); (c) H α emission ($I_{\text{H}\alpha}$; J. E. Gaustad et al. 2001); (d) dust temperature (T_{dust} ; C. J. R. Clark et al. 2023); (e) 24 μm intensity ($I_{24\mu\text{m}}$; K. D. Gordon et al. 2011).

Fortunately, the Hubble Space Telescope (HST) is perfectly suited to this task. At the distances of the MCs, individual stars are resolvable by HST far below the oldest main-sequence turnoff (oMSTO) in relatively short exposure times (e.g., D. R. Weisz et al. 2013; J. Roman-Duval et al. 2019). HST is also equipped with imaging detectors sensitive to UV–IR wavelengths. However, a fully resolved survey of the two systems is unfeasible given the small observing field of view of HST. This underscores the power of ground-based surveys and future wide-area facilities (e.g., the Nancy Grace Roman Space Telescope) to cover the full extents of the MCs (e.g., M. R. L. Cioni et al. 2011; D. L. Nidever et al. 2017). Existing HST observations of the LMC and SMC therefore tend to target prominent, individual regions, such as the 30 Doradus star-forming region in the LMC (E. Sabbi et al. 2016) and the Southwest Bar of the SMC (P. Yanchulova Merica-Jones et al. 2017); small numbers of deep pointings (D. R. Weisz et al. 2013; M. Cignoni et al. 2013); or regions parallel to targeted observations (METAL; J. Roman-Duval et al. 2019).

In this paper, we introduce Scylla, a pure-parallel HST imaging survey that operated alongside the Ultraviolet Legacy Library of Young Stars as Essential Standards (ULLYSES) survey (J. Roman-Duval et al. 2020). Named after the multiheaded monster from the myth of Ulysses, Scylla imaged fields parallel to spectroscopic ULLYSES targets, obtaining maximum photometric coverage from the near-UV (NUV) to the NIR with Wide Field Camera 3 (WFC3). In this paper, we describe the first 96 fields imaged over 342 orbits during Cycles 27–31 (2020–2023). In a future release, we will include all survey data (27 additional fields).

Although Scylla does not cover large, contiguous fields, our spatial coverage of the MCs is unprecedented, as the 48 fields in the LMC are distributed over the central $\sim 4.5 \times 5$ kpc, and the 48 fields in the SMC are distributed over the central $\sim 4 \times 2$ kpc. As a result, Scylla samples diverse environments in terms of metallicity, gas column densities, radiation fields, and SFHs across kiloparsec scales.

In Figure 2 we plot cumulative distribution functions of several ISM tracers: dust surface density and temperature (C. J. R. Clark et al. 2023), H I column density (S. Kim et al. 1999; N. M. Pingel et al. 2022), H α emission (J. E. Gaustad et al. 2001), and 24 μm emission (K. D. Gordon et al. 2011). These results were computed by identifying the pixels associated with each Scylla footprint in each of the sampled maps and bootstrapping the distributions with replacement to compute uncertainties. We find that Scylla observations in the LMC and SMC probe different parameter spaces in each tracer and also a wide range. For example, we probe nearly three orders of magnitude in dust surface density in the LMC and a similarly wide range of H I column densities in the LMC. The two samples are very similar in ionized hydrogen and dust temperature, and the LMC fields sample regions of enhanced star formation activity relative to the SMC (traced by 24 μm).

Armed with this data set, our science goals are to resolve how dust properties vary with interstellar environment, to constrain the multidimensional structure of the ISM of the MCs, and to probe their detailed SFHs. Scylla will probe the total dust column density (A_V) at roughly parsec-scale resolution within the MCs independently of other ISM tracers. By comparing our results with A_V derived from far-IR emission, we will probe variations in the opacity and emissivity of dust throughout the MCs. In addition, we will generate maps of the SFH and chemical enrichment history of the MCs with the highest resolution to date and resolve how the SFH evolves as a function of position and time (R. E. Cohen et al. 2024a, 2024b).

In subsequent publications, we will produce a catalog of stellar and dust parameters by fitting the multiband photometry of each star using the Bayesian Extinction and Stellar Tool (BEAST; K. D. Gordon et al. 2016), a probabilistic method for modeling spectral energy distributions (SEDs) for surveys of resolved stellar populations. Our catalogs will describe the age, initial mass, metallicity, distance, A_V , average grain size (R_V), and mixture coefficient between the MW R_V -dependent dust

Table 1
Filter Combinations

Name	F225W	F275W	F336W	F475W	F814W	F110W	F160W	N_{LMC}	N_{SMC}
A	X	X	X	X	X	X	X	3	4
B		X	X	X	X	X	X	1	6
C	X	X	X	X	X		X	2	2
D		X	X	X	X		X	3	1
E	X	X	X	X	X			2	4
F			X	X	X		X	2	3
G		X	X	X	X			2	5
H			X	X	X			13	8
I				X	X			20	15

Note. This table summarizes the 10 unique filter combinations for the Scylla survey. The number of fields with each combination in the LMC (N_{LMC}) and SMC (N_{SMC}) is indicated in the right-hand columns.

extinction and the SMC Bar dust extinction (K. D. Gordon et al. 2003) of the observed sources in the MCs.

This paper describes the observing strategy (Section 2) and data reduction procedure (Sections 3 and 4) for the Scylla survey. We present color–magnitude diagrams (CMDs) and discuss the results of analyzing the CMDs in the context of previous observations (Section 5). Finally, we present concluding remarks (Section 6).

2. Observations

Following the strategy of other large HST photometric surveys of nearby galaxies, including the Panchromatic Hubble Andromeda Treasury (PHAT; J. J. Dalcanton et al. 2012; B. F. Williams et al. 2014), the SMC Investigation of Dust and Gas Evolution (SMIDGE; P. Yanchulova Merica-Jones et al. 2017, 2021), and the Metal, Evolution, Transport and Abundance in the LMC (METAL; J. Roman-Duval et al. 2019) surveys, we observed each ULLYSES parallel field with WFC3¹¹ in filters spanning UV, optical, and NIR wavelengths.

As a pure-parallel complement to the ULLYSES survey, we were constrained by the requirements of the primary COS and STIS spectroscopic observations. Our fields are located near the primary targets ($\sim 5'$ away, based on the footprint of WFC3 relative to COS/STIS) and were positioned according to the roll angle of the telescope at the time of observation. The inability to set the roll angle means that we were unable to guarantee follow-up imaging of the same fields in the event of observing failures (e.g., guide star acquisition failures).

Based on the arrangement of buffer dump overheads imposed by the primary observations, we arranged exposures in as many of the WFC3/UVIS (F225W, F275W, F336W, F475W, F814W) and WFC3/IR (F110W and F160W) filters shown in Figure 1 as possible. We selected filters for each field based on the following hierarchy: (1) F475W and F814W, (2) F336W and F275W, (3) F110W and F160W, and (4) F225W. We prioritized observations of optical filters (F475W, F814W) to maximize the number of observed sources and to optimize the constraining power of the SFH across a large variety of stellar types. Bluer and UV observations (F336W, F275W, F225W) give us additional leverage when measuring hotter stars, extinction, and dust grain size distribution. Lastly, depending on the constraints imposed by the primary observing setup, we obtained imaging in the IR (F110W, F160W), which

allows us to measure cooler stars (e.g., asymptotic giant branch stars) and stars embedded in dustier regions. When possible, we also included one or more exposures in F225W to constrain the shape of the 2175 Å extinction feature (e.g., J. Roman-Duval et al. 2019). Overall, we used nine unique combinations of filters across the Scylla observations. These filter combinations are summarized in Table 1.

As the telescope pointing is fully determined by the primary ULLYSES observations, we were unable to dither our exposures. For WFC3/UVIS, we included multiple exposures of as close to equal length as possible in each visit to enable cosmic-ray rejection. For WFC3/IR, this is done via up-the-ramp data reduction. In many cases, the ULLYSES overheads mandated that we obtain long exposures in a small number of filters (one to two) per orbit, which hampered our ability to reliably detect bright stars (which saturate in long exposures). Therefore, where possible we inserted short (3–5 s, also known as “guard”) exposures in F475W or F814W.

In Figure 3, we summarize the positions and wavelength coverage of Scylla imaging from this release (observations through 2023 February) in the LMC and SMC. Scylla increases the total spatial coverage of UV–IR imaging from HST (e.g., D. R. Weisz et al. 2013; M. Cignoni et al. 2013; E. Sabbi et al. 2016; P. Yanchulova Merica-Jones et al. 2017; J. Roman-Duval et al. 2019) by factors of 2 and 8 in the LMC and SMC, respectively, opening unexplored parameter space in their ISM and SFHs. Assuming average distances to the LMC and SMC of 50 and 62 kpc, respectively, and a field of view of WFC3 UVIS of 160", each field covers an area of $\sim 39(48)$ pc on a side in the LMC (SMC).

In Table 2, we summarize the field names, positions, and observing parameters. Each field obtains two names: one short and one long. The short names are formatted as “L (S)MC_{*n*},” where *n* is a number corresponding to the order in which the field was observed in the survey (e.g., SMC₁₅). All values of *n* are not present in the table, as some observations overlapped with other fields, and therefore those exposures were processed together. In addition, 13 fields suffered from guide star failures or other imaging artifacts that made them unusable for photometry. Overall, 109 fields were observed through 2023 February, and the 96 without observing failures are presented here. The rest of the Scylla survey, comprising observations of 27 fields between 2023 February and August, will be presented in a future release. The long field names are formatted as “[PID]_[galaxy]–[distance from galaxy center in arcsec][cardinal direction]–[telescope rotation angle in deg (V3PA)],” where PID is the HST

¹¹ We note that although the previous surveys also used the Advanced Camera for Surveys to image regions in optical/IR filters, our parallel program status led us to choose using only WFC3 for the full survey.

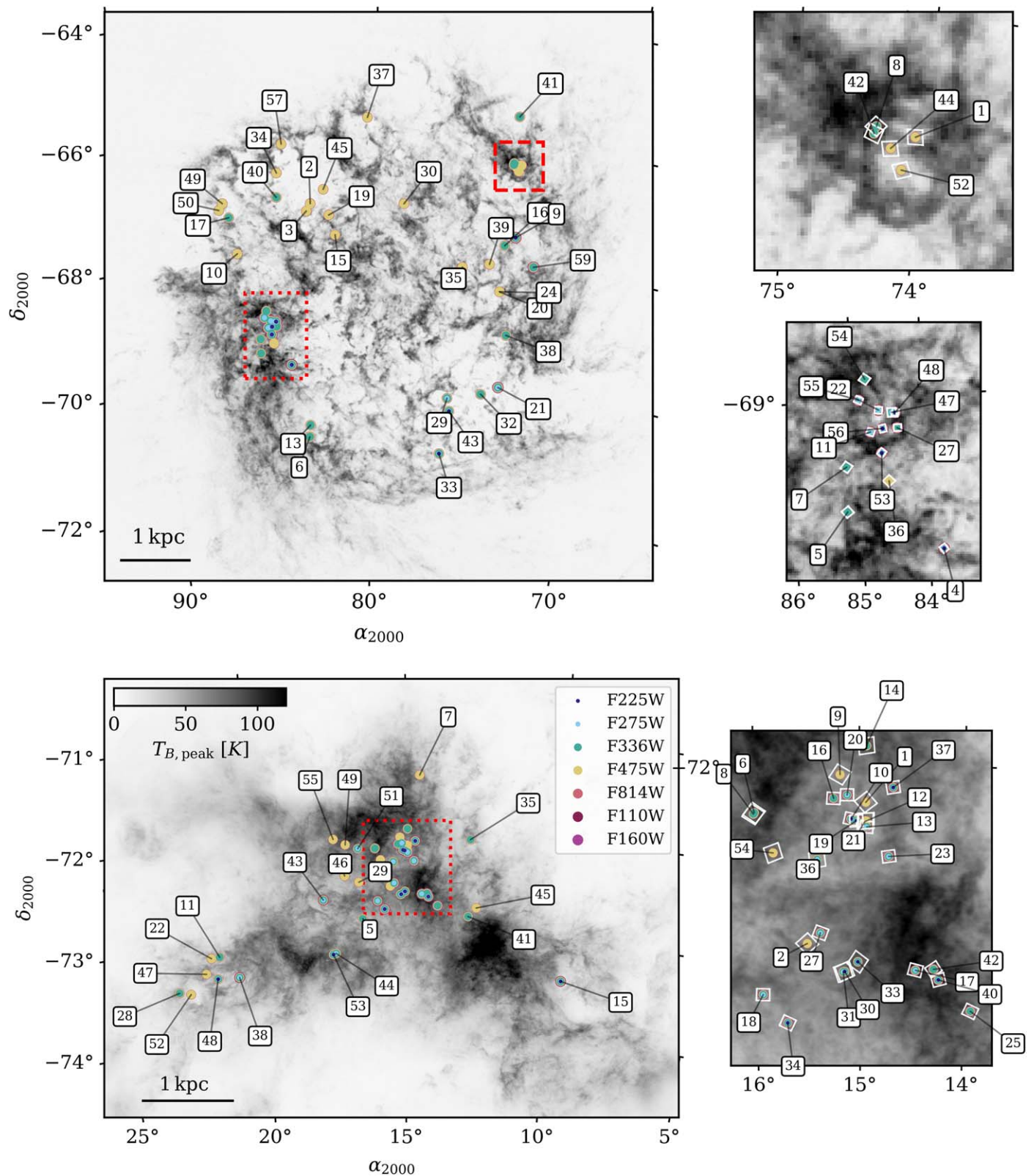


Figure 3. Maps of the peak brightness temperature of 21 cm emission in the LMC (S. Kim et al. 1999) and the SMC (N. M. Pingel et al. 2022), overlaid with observed Scylla fields. In each panel, regions where fields overlap are identified by dashed red bounding boxes and expanded at right. There are 96 fields as part of the release described in this paper, 48 in the SMC and 48 in the LMC. The number of filters used to observe each field is illustrated by the number of concentric circles overlaid at each position. Each field is tagged by its “short” name number (see Table 8). In the expanded panels, the size of each footprint is plotted in white outline.

program ID under which the observation was taken (e.g., 15891_SMC-4292_{sw}-13841). The long-form names are used for a qualitative assessment of the field location within

each galaxy, and therefore the lack of precision in the name (i.e., using an integer value for the rotation angle) is not a significant issue.

Table 2
Field Summary

Field Name (Short)	Field Name (Long)	R.A. (deg)	Decl. (deg)	N_{filters}	N_{orbits}	Combination
(1)	(2)	(3)	(4)	(5)	(6)	(7)
SMC_1	15891_SMC-3192ne-8290	14.939	-72.1032	2	2	I
SMC_2	15891_SMC-2750ne-8567	15.5059	-72.5123	2	2	I
SMC_5	15891_SMC-3514se-8584	16.4523	-72.835	3	3	H
SMC_6	15891_SMC-3956ne-9632	15.9954	-72.1305	3	2	H
SMC_7	15891_SMC-5278ne-9802	14.5464	-71.4007	2	2	I
SMC_8	15891_SMC-3955ne-9818	16.0039	-72.1341	3	2	H
SMC_9	15891_SMC-3587ne-10112	15.1806	-72.0226	2	2	I
SMC_10	15891_SMC-3149ne-12269	15.0793	-72.1515	7	4	A
SMC_11	15891_SMC-8743se-11371	21.4361	-73.1212	3	2	H
SMC_12	15891_SMC-3025ne-13499	14.9344	-72.1575	2	2	I
SMC_13	15891_SMC-2983ne-12972	14.9313	-72.1748	5	2	D
SMC_14	15891_SMC-3669ne-13972	14.9314	-71.9391	3	2	H

Note. Column (1): short field name. Column (2): long field name. Column (3): average R.A. per field. Column (4): average decl. per field. Column (5): number of filters used in the field’s observations. Column (6): number of orbits per field. Column (7): the filter combination (see Table 1).

(This table is available in its entirety in machine-readable form in the [online article](#).)

In Appendix A, we summarize the observing parameters for all exposures in the Scylla survey to date, including filters, field centers, position angles, exposure times, and post-flash exposures to reduce charge transfer efficiency (CTE) effects. The complete table is available in machine-readable form.

In Figure 4, we show example three-color images (red = F814W, green = F475W, blue = F336W) for a selection of fields with these filters available. In the images on the left (panels (a) and (c)), we see evidence of rich extended emission, likely from [O III] λ 5007, [O III] λ 4959, H β , and H γ (F475W), as well as nebular [S III] λ 9069 (F814W), as well as nebular continuum, Balmer continuum, and He II λ 3203 (F336W) emission from the ISM (as explained by J. Roman-Duval et al. 2019). While beautiful, these features confuse sources and increase photometric uncertainties, making it harder to resolve individual stars.

In SMC_21 (Figure 4(a)), we see a smattering of bright blue sources from the open cluster NGC 346, associated with the nearby star-forming region N66 (M. Rubio et al. 2018). We also see several saturated red sources with diffraction spikes, two of which are blended in the upper left quadrant. This field also contains imaging artifacts in the lower right corner, in the shape of a figure of eight.¹² These artifacts appear when light from bright objects in one quadrant is reflected on the WFC3/UVIS CCD, creating ghosts in the diagonal quadrant. In this case, the two blended sources with diffraction spikes are likely the source of these imaging artifacts. However, these artifacts were not identified as sources in our catalogs, meaning that we do not attempt to remove them from our images.

In LMC_11 (Figure 4(c)), we also see extended emission structure; however, this emission is present in all optical bands we use to construct the three-color images. Since this field is near the star-forming region 30 Doradus, this emission is likely correlated with ionized gas and/or stellar light scattered by dust. In addition, we see evidence for extinction toward some of the regions that show prominent diffuse emission.

In SMC_15 and LMC_16 (Figures 4(b) and (d)), we find no significant evidence of extended emission. However, in LMC 16, we do see a background galaxy.

3. Data Reduction

In the following section, we describe the data reduction process, including image reduction, photometry, and post-processing (quality cuts). All HST data used in this paper can be found in the Barbara A. Mikulski Archive for Space Telescopes (MAST) at doi:[10.17909/8ads-wn75](https://doi.org/10.17909/8ads-wn75).

3.1. Image Reduction

The calibrated images for the Scylla program (multicycle PIDs 15891, 16235, 16786) were downloaded from MAST during observations spanning 2020 April and 2023 February.¹³ The images were processed using CalWF3 versions 3.5.0, 3.5.1, 3.5.2, 3.6.1, and 3.6.2 (versions spanning 2018 October through 2021 April).¹⁴ For UVIS exposures, we downloaded the calibrated and CTE-corrected images (flc), and for the IR, we downloaded the calibrated images (flt), as WFC3/IR data do not require CTE correction.

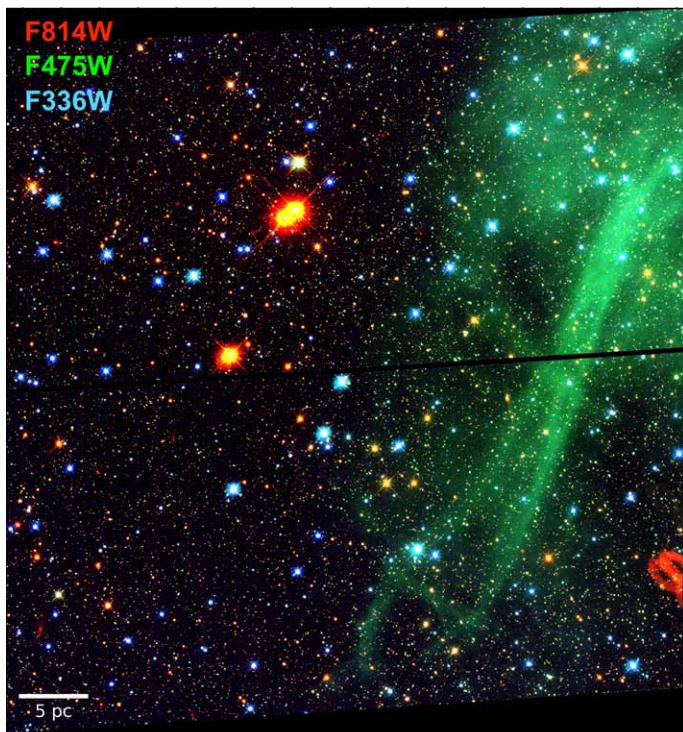
We ran all of our images through the task `astrodrizzle` from the `Drizzlepac` package¹⁵ (R. J. Avila et al. 2015), which produced combined, rectified, resampled images in each band. This process also uses a median filter to flag cosmic-ray-affected pixels on each of the input exposures by updating the data quality FITS file extensions. In addition, the quality of the combined images provides a visual check on the astrometric alignment of the individual exposures in each band. We processed our fields in batches, as they arrived between Cycles 27 and 30, and therefore not all images are aligned in the same way. At least 70% of fields were aligned to Gaia, and the remaining fields used an a priori solution based on the observed guide stars. In a future data release, we will employ the same astrometric alignment on all fields.

¹² See Instrument Science Report WFC3 2022-03 for more details (www.stsci.edu/files/live/sites/www/files/home/hst/instrumentation/wfc3/documentation/instrument-science-reports-isrs/_documents/2022/WFC3-ISR-2022-03.pdf).

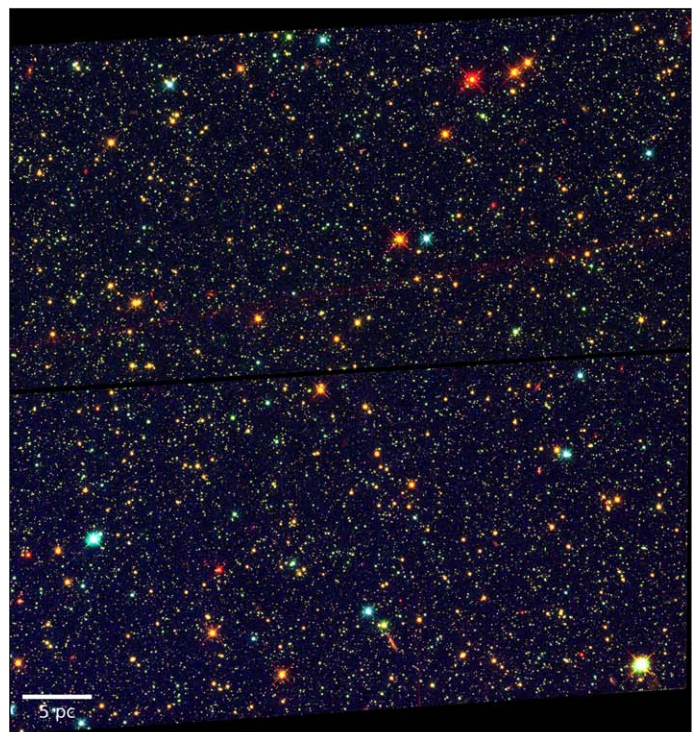
¹³ Additional Scylla observations were conducted between 2023 February and 2023 August and will be described as part of a future data release.

¹⁴ As part of a future data release, we will re-reduce all images using the same pipeline version.

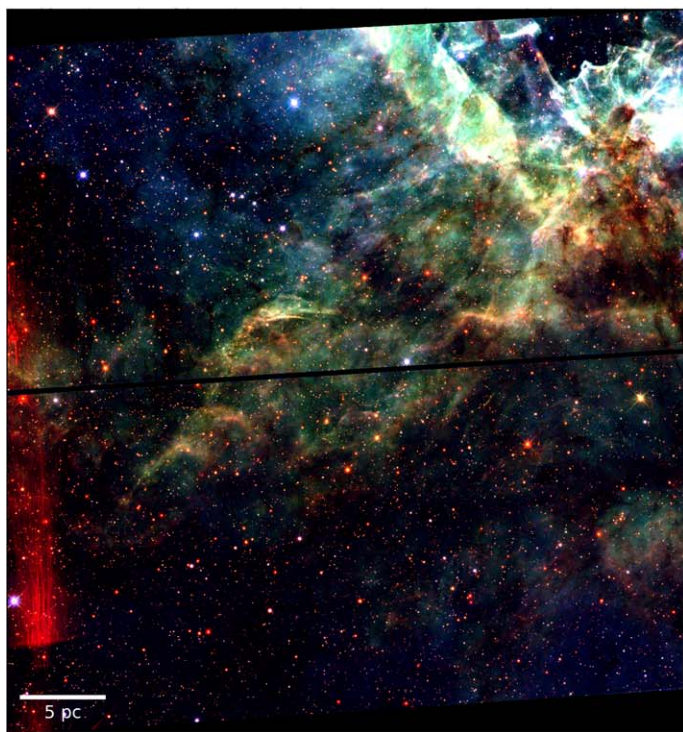
¹⁵ <https://hst-docs.stsci.edu/drizzpac>



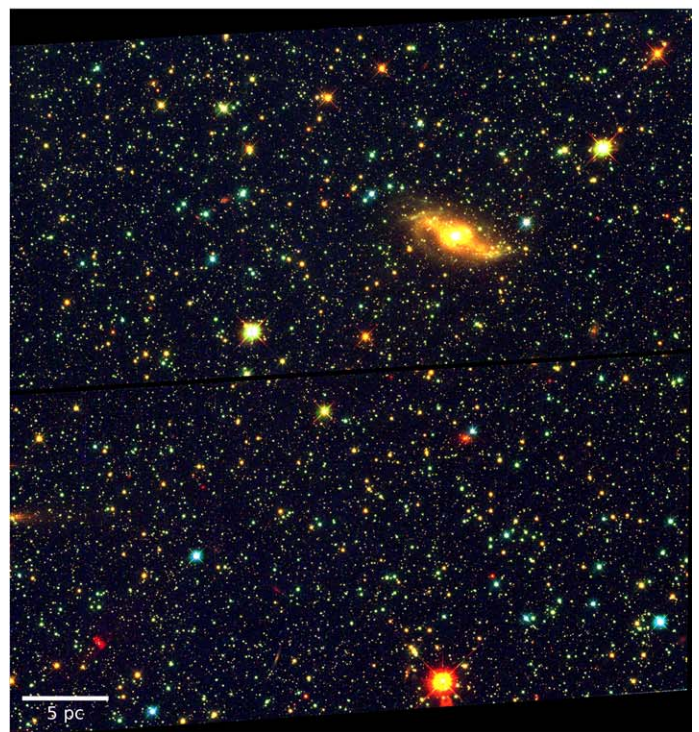
(a) SMC 21 with extended emission.



(b) SMC 15 with no emission.



(c) LMC 11 with extended emission.



(d) LMC 16 with no emission.

Figure 4. Example three-color images of fields in the SMC (top) and LMC (bottom) with (left) and without (right) extended emission. Images were created using `aplpy` by normalizing the observed flux range from the F814W (red), F475W (green), and F336W (blue) filters from WFC3. Three-color (or two-color where F336W is not available) images for all fields are available as part of our data release.

With the data quality extensions updated and the combined images available for astrometric reference, we then prepared the images for crowded-field point-spread function (PSF) photometry, as described in the following section.

3.2. Photometry

We measured the resolved photometry of all detected sources on all fit/flc exposures using the same photometry pipeline as the PHAT (B. F. Williams et al. 2014), PHATTER

(B. F. Williams et al. 2021), METAL (J. Roman-Duval et al. 2019), and SMIDGE (P. Yanchulova Merica-Jones et al. 2017) surveys. We independently analyzed the data from each field of the survey, providing a resolved star catalog for each location. Since detailed descriptions of the pipeline are available in those papers, here we provide an overview.

Our photometry pipeline uses the software package DOLPHOT (A. E. Dolphin 2002; A. Dolphin 2016). After running `astrodrizzle` on the images to flag pixels affected by cosmic rays, we applied the tools for masking and pixel area correction (`wfc3mask`) to each of the `flc/flt` images. We then separated each of the individual CCD reads from the multiextension fits files, and we applied the tool for generating a sky image for each CCD readout. All processed images covering a survey pointing were then put into the DOLPHOT task, along with a drizzled reference frame, which was the deepest single-band combined drizzled image. The DOLPHOT task then calculated the alignment of a subset of stars in each individual image to locations of stars in the reference image, applied small corrections if necessary, and returned quality statistics on the alignment. All of the Scylla data had excellent (better than 0.1-pixel standard deviation across all alignment stars) alignment statistics, which is typical for sets of exposures all taken in a single HST visit, especially when the fields are well populated.

Because we supplied all `flt/flc` exposures to DOLPHOT, it optimized sensitivity to detect stars by stacking all pixels at a given on-sky position, i.e., after applying distortion corrections to the individual `flc/flt` images, covering every location in the field to look for any location with an elevated number of counts. Then, all of these locations were fitted with the appropriate PSF to measure the brightness of the candidate point source. Each source was measured iteratively, taking into account neighboring sources with each iteration. The result was an output catalog that contains combined brightness measurements in each band, as well as measurements in each exposure for every candidate point source that was able to be fitted with the PSF. These initial catalogs are very inclusive to provide measurements of the faintest potential sources and allow thorough deblending. They are thus likely to contain contaminants and poor measurements, especially at the faint end, along with many very well measured stars. To help separate the unreliable measurements, DOLPHOT provides many quality metrics for each measurement to allow users to filter the catalog in whatever way works best for their science applications. We describe our filtering below.

3.3. Photometric Quality Cuts

There are two types of catalogs that we produced for each field: the `st` (“star”) catalog and the `vgst` (“very good star”) catalog. The difference between these two types is that `vgst` catalogs have more restrictive quality cuts compared to `st` catalogs (i.e., the `vgst` catalogs are subsets of the `st` catalogs). The two catalog types exist to facilitate different science goals—`st` for high completeness, `vgst` for high reliability. In this section, we describe the basic quality cuts we impose to create `st` catalogs and the additional quality cuts to create `vgst` catalogs.

DOLPHOT produced `phot` catalogs that combine the measurements for each source of all the individual CCD chips into one final measurement for each source in each bandpass, including the count rate, rate error, VEGA magnitude and error, χ^2 of the PSF fit, sharpness, roundness, crowding, and

signal-to-noise ratio (SNR). For the details of this step see Section 2.3 of B. F. Williams et al. (2014). The `st` catalogs are generated from `phot` and include all sources with SNR greater than 4 in at least one observed band.

To construct the `vgst` catalogs, we imposed additional quality cuts to remove spurious detections from the `st` catalogs. Specifically, we used the “sharpness,” “roundness,” and “crowding” parameters from DOLPHOT. The sharpness parameter quantifies how centrally concentrated the source flux is; high values correspond to sources with high values in the central pixels of the PSF relative to the edges (e.g., cosmic rays or hot pixels), and low values correspond to blended sources or galaxies. The roundness parameter quantifies how circular a source is, where a value of zero corresponds to a perfectly round source. The crowding parameter quantifies how neighbors affect the source photometry; large values correspond to sources whose PSF radii intersect with many others. These parameters are described in the DOLPHOT documentation,¹⁶ as well as discussed and implemented as quality cuts by other surveys (e.g., B. F. Williams et al. 2014, 2021).

First, we removed any sources not measured in all bands. This step excludes sources around the edges of each field, due to the smaller size of the WFC3 IR footprint compared to that of the UVIS camera. In future work, we plan analysis that alternately focuses on the HST UV and optical photometric bands to include sources in the UVIS camera footprint but excludes the HST IR photometric bands.

Next, we applied cuts to exclude diffraction spikes or resolved emission identified as sources by DOLPHOT. Inspecting the distribution of sources on the plane of the sky and in color–magnitude space, we found a significant number of sources that are likely spurious detections. These sources are generally bluer and fainter than the main sequence and are distributed along recognizable imaging artifacts (e.g., edges of the field and diffraction spikes). To remove these sources from the catalogs, we visually identified the contaminant sources by their locations and identified the cuts listed in Table 3 (see Appendix B for details).

Finally, we imposed cuts based on the flags generated by DOLPHOT during processing, which have values of 0 (“good”) through 8 (“extremely bad”). Per the DOLPHOT documentation, sources with flag values of 0–3 are usable. `FLAG = 1` and `FLAG = 3` indicate that the photometry aperture extends off the chip, and `FLAG = 2` indicates too many bad or saturated pixels. For some fields, a significant fraction of sources have `FLAG = 2` (~25% for SMC-6). Since sources extracted from saturated pixels may be due to bright stars of interest, we kept sources with `FLAG = 2` in the catalog. Ultimately, we kept sources with `FLAG` values of 0 and 2.

After applying these cuts to a representative sample of fields (i.e., varying depth, filter coverage, and source density), we visually confirmed that they do a reasonable job of removing spurious sources while removing few real sources, in agreement with the performance of similar quality cuts used in other surveys (e.g., B. F. Williams et al. 2014). In Figure 5, we compare the full catalog with the kept and cut sources for two fields at representative steps in the process, including sources that are not observed in all photometric bands and sources that fail our quality cuts (includes DOLPHOT and diffraction spike cuts).

¹⁶ <http://americano.dolphinim.com/dolphot/>

Table 3
Photometric Quality Cuts

Type	Parameter	Band	Value	N_{SMC}	N_{LMC}
st	Signal-to-noise ratio	*_SNR (in any one band)	>4	4,117,542	3,395,349
vgst	Sharpness ^{2, a}	F475W, F814W	≤ 0.15	968,217	972,087
	Roundness	F475W, F814W	≤ 0.6		
	Crowding	F475W, F814W	≤ 0.2		
	Quality flag ^b	*_FLAG (all bands)	0, 2		
	Flux ^c	All bands	$\neq 0$		

Notes. The *vgst* cuts include the *st* cuts.

^a The sharpness is squared to remove sources with large negative sharpness values.

^b If a source is flagged by DOLPHOT with anything other than 0 or 2 in any filter, it is cut.

^c If a source has a zero flux in any one band (meaning that no measurement was made), it is cut; however, fluxes are allowed to be positive or negative.

We note that we did not include a cut on SNR or χ^2 ; additional quality metrics quantifying the goodness of fit to the PSF used for cuts by similar surveys (e.g., B. F. Williams et al. 2014, 2021). We made this choice to ensure that even low-SNR sources (which may include rare, interesting stellar populations) are included in the *vgst* catalogs, which we will characterize with the BEAST as part of future work. Our quality cuts, as well as the total number of sources retained in each MC, are summarized in Table 3.

4. Artificial Star Tests

To quantify the uncertainty, bias, and completeness in our photometry, we performed artificial star tests (ASTs). This procedure involves placing simulated sources with known parameters within real images, repeating the same photometric extraction procedure, and comparing the derived properties of each fake source with its known intrinsic properties.

To create the AST inputs, we used the BEAST: a probabilistic method for modeling multiwavelength SEDs from large photometric surveys. With the BEAST, we generated multiwavelength SEDs in the Scylla survey bands for sources spanning a representative range of stellar and dust properties for the MCs. Using publicly available stellar atmosphere grids (F. Castelli & R. L. Kurucz 2003; T. Lanz & I. Hubeny 2003, 2007) and evolutionary tracks (PARSEC; P. Marigo et al. 2008; A. Bressan et al. 2012; Y. Chen et al. 2014, 2015; P. Marigo et al. 2017; G. Pastorelli et al. 2019, 2020), the BEAST generates the intrinsic spectrum of a star based on its age (t), mass (M_{ini}), metallicity (Z), and distance (d). This spectrum is then extinguished according to A_V , R_V , MW, and SMC/LMC dust extinction mixture coefficient (f_A ; see K. D. Gordon et al. 2016). Finally, the full extinguished stellar spectrum is converted into SEDs by integrating over the transmission curve of each HST photometric band. We summarize the seven BEAST parameters (t , M_{ini} , Z , d , A_V , R_V , and f_A), their ranges, and prior assumptions in Table 4.

For each Scylla field, we generated a set of artificial stars using the BEAST model grid. To ensure that the artificial stars cover a wide range of fluxes, we estimated the total flux range in each observed band and split that range into 40 flux bins. We then selected BEAST SEDs randomly until we had at least 50 in each bin, which resulted in ~ 2000 SEDs per field.

4.1. AST Placement by Source Density

Given the list of artificial stars, the next step is to determine where to place them within the images. Similar HST surveys have established that the noise properties of stars can vary

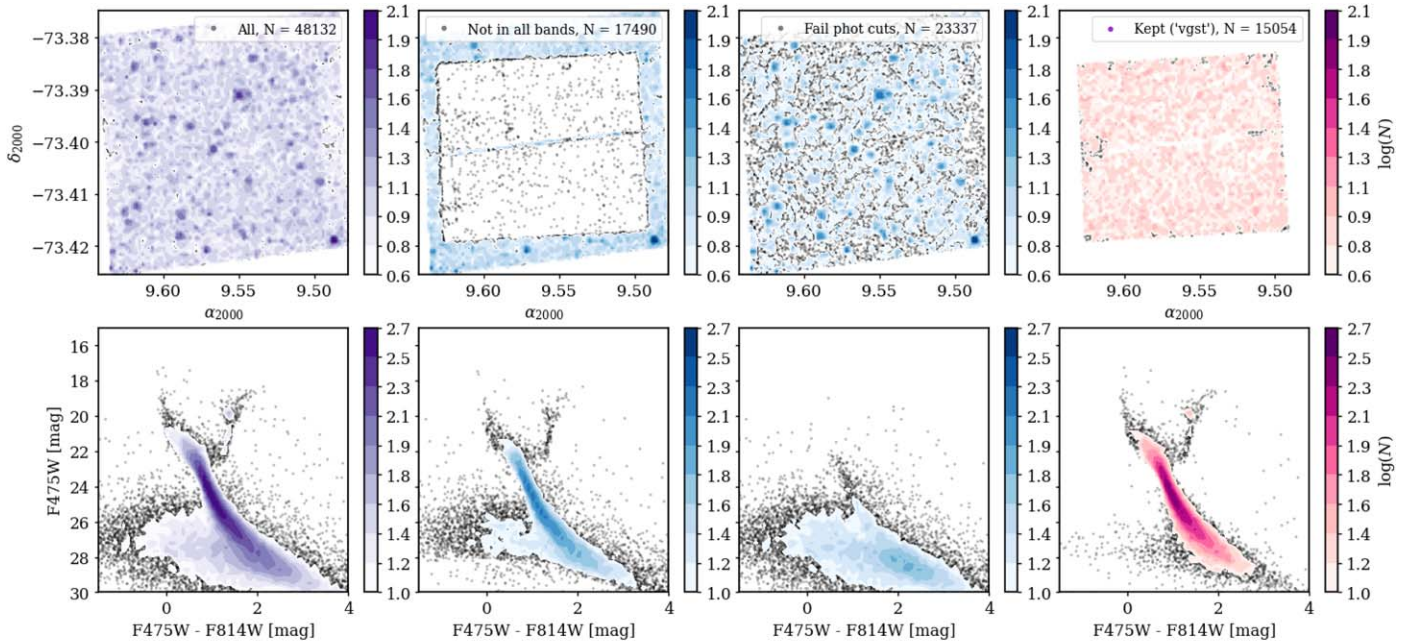
considerably with the local source density, especially for the lower-resolution NIR bands (K. D. Gordon et al. 2016). To incorporate this effect into our uncertainty modeling, we placed artificial stars according to the local source density within the field. We computed the source density across each Scylla field by counting the number of sources with photometric F475W Vega magnitudes between 26 and 15 (the range over which the catalogs are complete based on preliminary completeness tests) within pixels the size of $5''$ on each side. In addition, we eroded the field footprint boundary by $0.5''$ on all sides to avoid computing source densities at the very edge of the WFC3 chips. We find that the source densities within Scylla fields follow a roughly lognormal distribution, with most source densities ranging between 0 and 2 sources arcsec^{-2} , peaking at 0.3 and 0.6 sources for the LMC and SMC, respectively. We note that fields in the wing of the SMC (Fields 11, 22, 28, 38, 47, 48, and 52) have extremely low source densities, less than 0.1 sources arcsec^{-2} on average. By comparison, other photometric surveys like PHAT observed source densities ranging between ~ 10 and 20 in most fields of M31 with similar filters (e.g., Figure 12 in J. J. Dalcanton et al. 2012). To reflect the range of source densities observed in the MCs, we established a set of five source density bins, (0, 0.27, 0.526, 1.026, 2, >2) sources arcsec^{-2} , and placed all ~ 2000 artificial stars randomly within the pixels corresponding to each source density bin. In this step, we generated $\sim 10,000$ artificial stars per field.

4.2. AST Placement by Background Density

When there is a significant extended emission present in a field (see Figure 4), the variations in uncertainty, bias, and completeness across the field are no longer dominated by source density. We identified these ‘‘background-dominated’’ fields by analyzing the flux density in the F475W band (common to all Scylla fields). We first computed the average flux of detected sources in $5 \times 5 \text{ arcsec}^2$ pixels in each field. We then computed the interquartile range (equal to the 84th minus 16th percentiles) of flux density for all fields. We flagged a field if the interquartile range is greater than $17 \text{ mag arcsec}^{-2}$. This cutoff was selected by eye to identify the fields that have significant optical emission. Out of 96 total fields, 12 are background emission dominated (six in the LMC, six in the SMC).

For these fields, instead of placing artificial stars by source density, we generated a custom set of background density bins for each field. Specifically, we selected bin edges to generate five equal-sized (i.e., same number of sources) bins of background density.

(a) SMC_15: Quality cuts for typical Scylla field



(b) LMC_11: Quality cuts for a high-background Scylla field

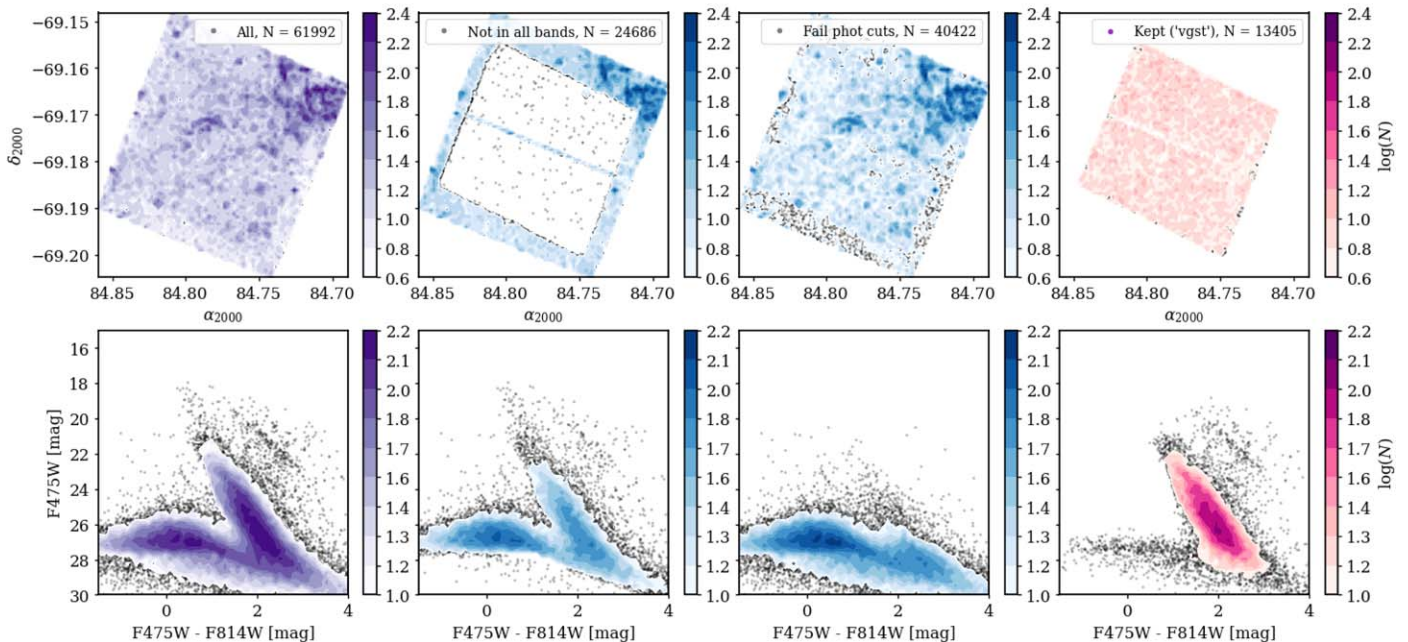


Figure 5. Illustration of the photometric quality cuts in two example fields. (a) A typical field (SMC Field 15, Figure 4(b)); (b) a field with strong background emission (LMC Field 11, Figure 4(c)). Upper rows: spatial distribution of sources; lower rows: optical CMD of the same sources. Left: full photometric catalog (.st.fits); middle left: sources removed owing to lack of detection in all bands; middle right: sources removed after failing quality cuts; right: final catalog (.vgst.fits).

In Figure 6, we plot the source density and background density maps for two representative Scylla fields, one source density dominated (top; SMC_15) and one background density dominated (bottom; LMC_11). We also compare the positions of the artificial stars for the two cases. Although the source-density-dominated field shows relatively uniform placement, the background-density-dominated field's sources are placed where the optical emission is bright.

4.3. Supplementary ASTs

In addition to the artificial stars distributed by source density (or background density), we generated supplementary artificial stars in each field to more densely sample the magnitude ranges of our observations. For each field, we generate $>10^5$ SEDs using the BEAST with similar priors to those summarized in Table 4, restricted to the observed magnitude ranges (R. E. Cohen et al.

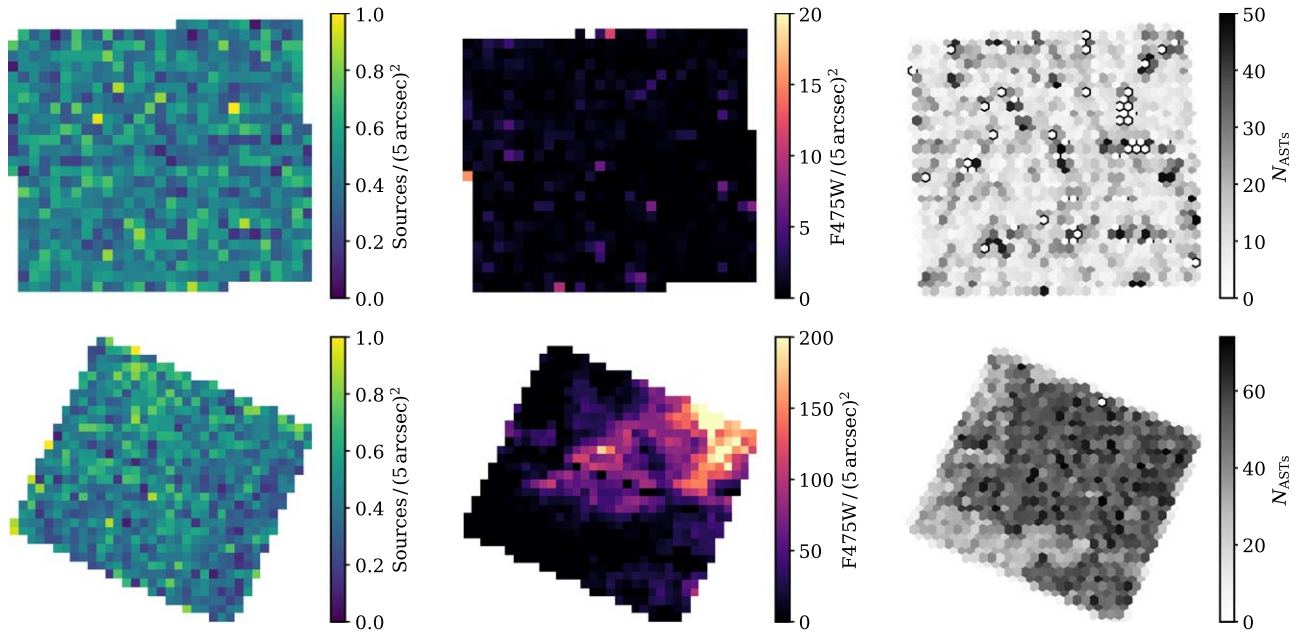


Figure 6. Summary of how ASTs are placed for two example fields, one source density dominated (SMC Field 15; top) and one background density dominated (LMC Field 11; bottom). Left: source density maps; middle: background density maps; right: positions of ASTs generated before adding supplementary sources.

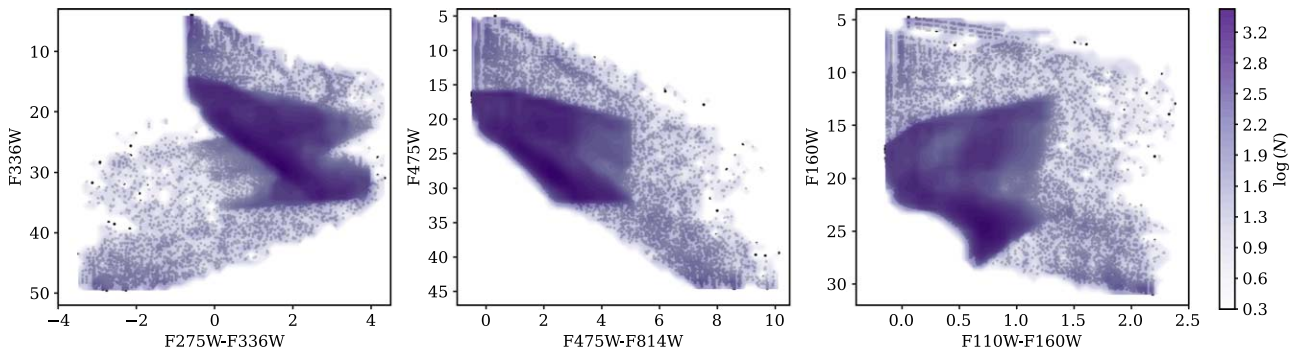


Figure 7. CMDs of the input ASTs for background-density-dominated Scylla LMC Field 11 (15891_LMC-5389ne-11134), in UV (left), optical (middle), and IR (right).

Table 4
BEAST Model Parameters Used for AST Generation

Parameter (1)	Unit (2)	Description (3)	Min. (4)	Max. (5)	Resolution (6)	Prior (7)
$\log(t)$	yr	stellar age	6.0	10.13	0.1	flat SFR
$\log(M)$	M_{\odot}	stellar mass	-0.8	2.0	variable	Kroupa IMF
$\log(Z/Z_{\odot})$...	stellar metallicity	-2.1	-0.3	0.3	flat
A_V	mag	dust column	0.01	10.0	0.05	flat
R_V	...	dust average grain size	2.0	6.0	0.5	peaked at ~ 3
f_A	...	dust mixture coefficient	0.0	1.0	0.2	peaked at 1
d [LMC]	kpc	distance	40	60	2.5	flat
d [SMC]	kpc	distance	47	77	2.5	flat

Note. Column (1): parameter name. Column (2): unit. Column (3): description. Column (4): minimum value. Column (5): maximum value. Column (6): interval. Column (7): prior model.

2024a, 2024b). These sources are spatially distributed uniformly within each field.

In Figure 7, we plot CMDs of the AST inputs for a single Scylla field (LMC_11). The initial BEAST-generated ASTs span a wide range in color–magnitude space in UV, optical, and IR bands. The supplementary ASTs are clearly identifiable

as the high-density range of sources spanning the main observable CMD ranges in each panel.

4.4. AST Results

With a final list of ASTs for each field, we repeated the full photometry procedure for each individual artificial

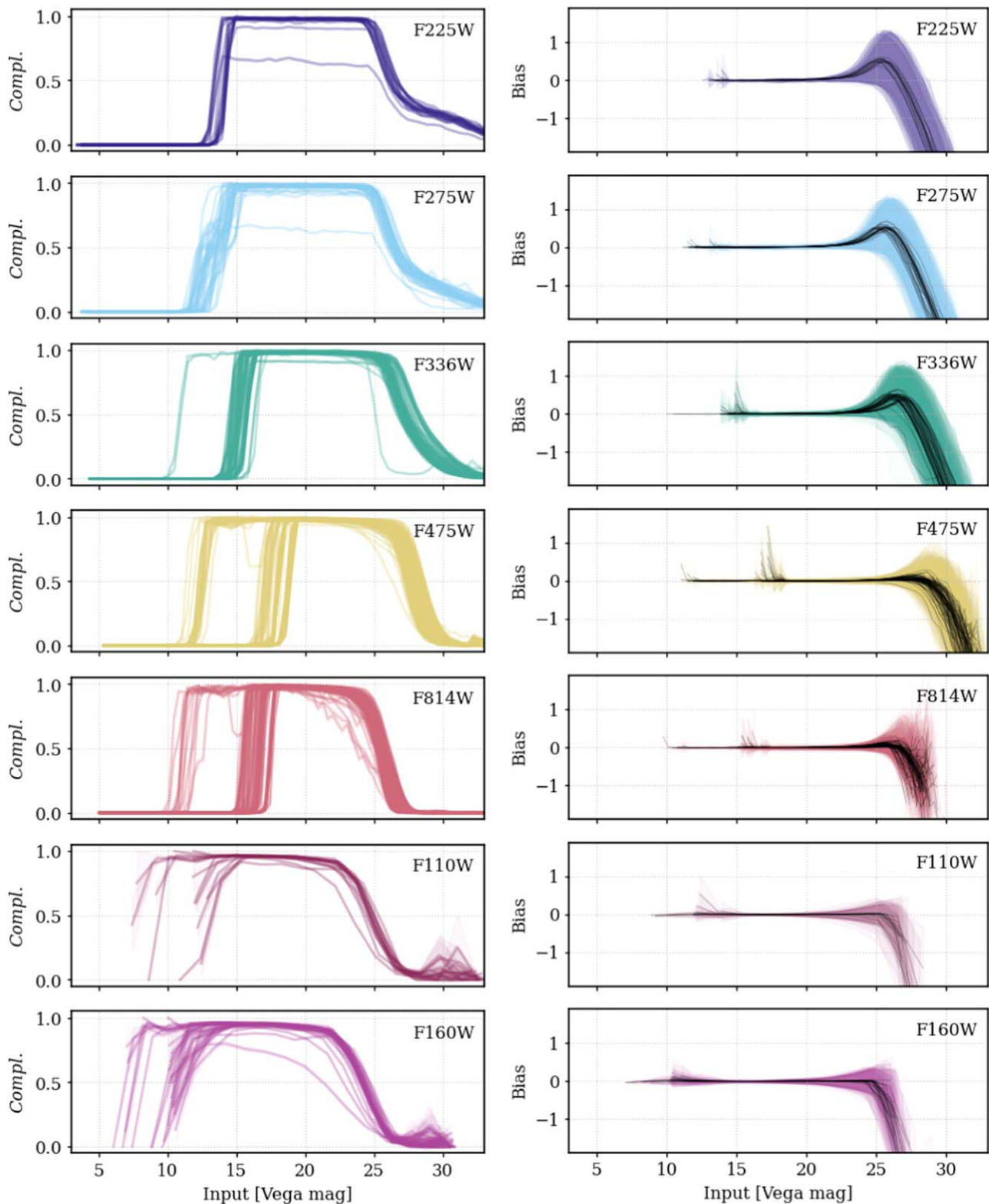


Figure 8. A summary of the completeness, bias, and uncertainty (σ) for all Scylla fields. In the left column, we plot completeness as a function of flux for the Scylla observing bands. In the right column, we plot bias (black lines), with shading corresponding to $\pm 1\sigma$.

star in each field. Once the ASTs were processed, we compared the photometric properties of the sources to their intrinsic values and computed the completeness, uncertainty (σ), and bias (μ) in each photometric band for each field. In

Figure 8, we plot these quantities as a function of Vega magnitude for all Scylla fields. In Table 8 in Appendix D we list the 50% completeness limit in each band for each field.

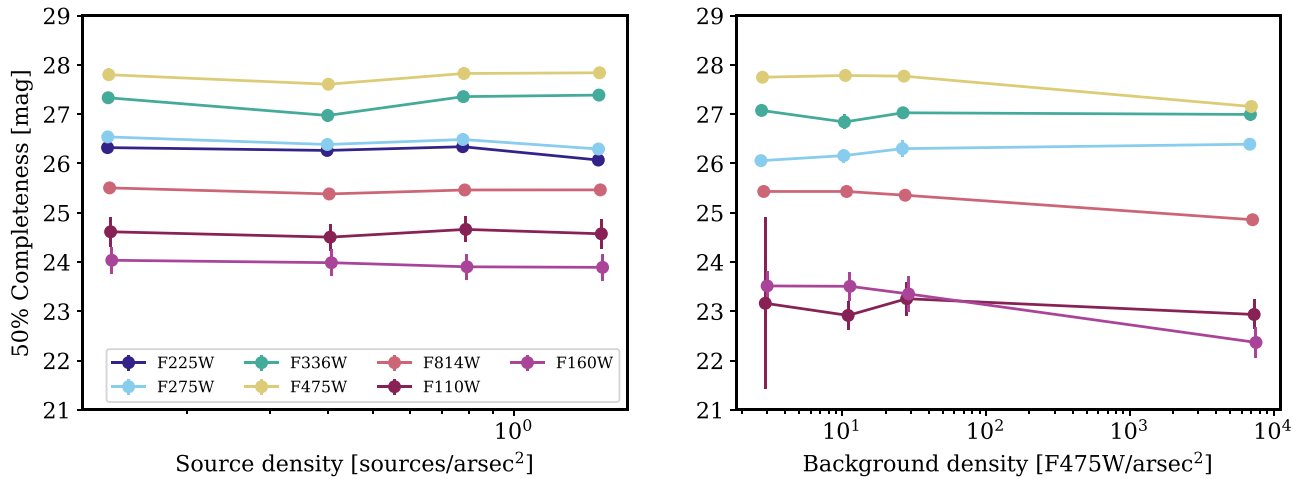


Figure 9. Left: 50% completeness as a function of source density for each band in fields that are source density dominated. Right: 50% completeness as a function of background density for each band in fields that are background density dominated.

In summary, we reach average 50% completeness of $m_{F225W} = 26.0$, $m_{F275W} = 26.2$, $m_{F336W} = 26.9$, $m_{F475W} = 27.8$, $m_{F814W} = 25.5$, $m_{F110W} = 24.7$, and $m_{F160W} = 24.0$ Vega mag in the *vgst* catalogs.

As shown in Figure 8, the Scylla survey is extremely deep in the optical bands (F475W and F814W, our first-priority filters), where we reach 50% completeness limits of 26th–28th Vega magnitudes (>3 mag below the oMSTO). For a subset of fields, the primary observation setup enabled us to include short “guard” exposures, which allowed us to extract the photometry of bright sources (which would otherwise be saturated by our standard long exposures). These fields are evident in Figure 8 as the groups of fields whose completeness extends much brighter than ~ 15 th magnitude. For the UV bands, we reach 50% completeness of <25 th magnitude. In the NIR we were limited to short, single exposures, resulting in lower completeness and higher uncertainty as a function of flux.

There is one Scylla field with completeness $\sim 60\%$ across the full magnitude range for F225W and F275W (LMC_56). In this case, during photometric processing two observations with similar spatial coverage but different observed filters were combined into the same field. As a result, the sources with no UV observations were removed during our *vgst* quality cuts, introducing artificially low completeness in those bands. In our next round of processing, we will treat each of these fields as a separate observation and correct the completeness accordingly.

To explore how our observational completeness varies within fields, we combine all ASTs and compute the 50% completeness limit in bins of source density and background density. In Figure 9, we find that the completeness remains constant within our uncertainties for all source densities probed by Scylla. The exception is the bin at ~ 0.4 sources deg⁻², where residual artifacts from incompletely subtracted diffraction spikes cause the completeness to dip slightly, especially for the shortest wavelengths. That the trend overall is flat is not surprising, given the low, uniform values of source density relative to HST’s resolution limit across the survey footprint. In contrast, we observe that the completeness starts to decline with increasing background density for the optical–NIR bands. Overall, however, our stellar recovery for high background density fields appears to be performing well. This is likely due to the fact that DOLPHOT successfully measures and subtracts the local background around sources. Our uncertainties are

highest in the IR bands, as very few background-dominated fields also had IR exposures. For the shorter-wavelength filters, the completeness is similar to the source-density-dominated cases.

4.5. Comparison of Overlapping Fields

In total, 15 Scylla fields have some overlap in terms of the positions of detected sources in their *vgst* catalogs. To verify that the photometry is consistent, for all filters that are the same between the overlapping fields, we compare the output fluxes in Figure 10. For comparison, we include the average uncertainty and bias in each filter based on the ASTs (Figure 8) between the overlapping fields. We find that the average difference in the derived fluxes (ΔF) falls within the uncertainties from the ASTs in all cases, and the spread in ΔF is contained within $\pm 3\sigma$ for all bands.

4.6. Data Release

Our catalogs are made publicly available as *st* fits files, with flags for *vgst*, at MAST as a High Level Science Product (HLSP; <https://archive.stsci.edu/hlsp/scylla>), with a permanent DOI of 10.17909/mk54-kg51. In this data release, we also include the drizzled *.fits* images, as well as the *.pngs* of the three-color images for all fields (two-color where only F475W and F814W are available). ASTs will be made available as part of the HLSP at a later date.

5. Results

The result of the Scylla survey is a deep, diverse sampling of the photometric properties of stars in the LMC and SMC. With the addition of the per-field ASTs, we also have a detailed understanding of our completeness, uncertainty, and bias in each observation. In this section, we highlight these products and their potential for impact on our understanding of the ISM and stellar populations of the MCs.

5.1. Color–Magnitude Diagrams

To summarize our results, we create stacked CMDs for the optical, UV, and NIR bands in Scylla LMC and SMC fields (Figure 11). We observe that both galaxies include similar

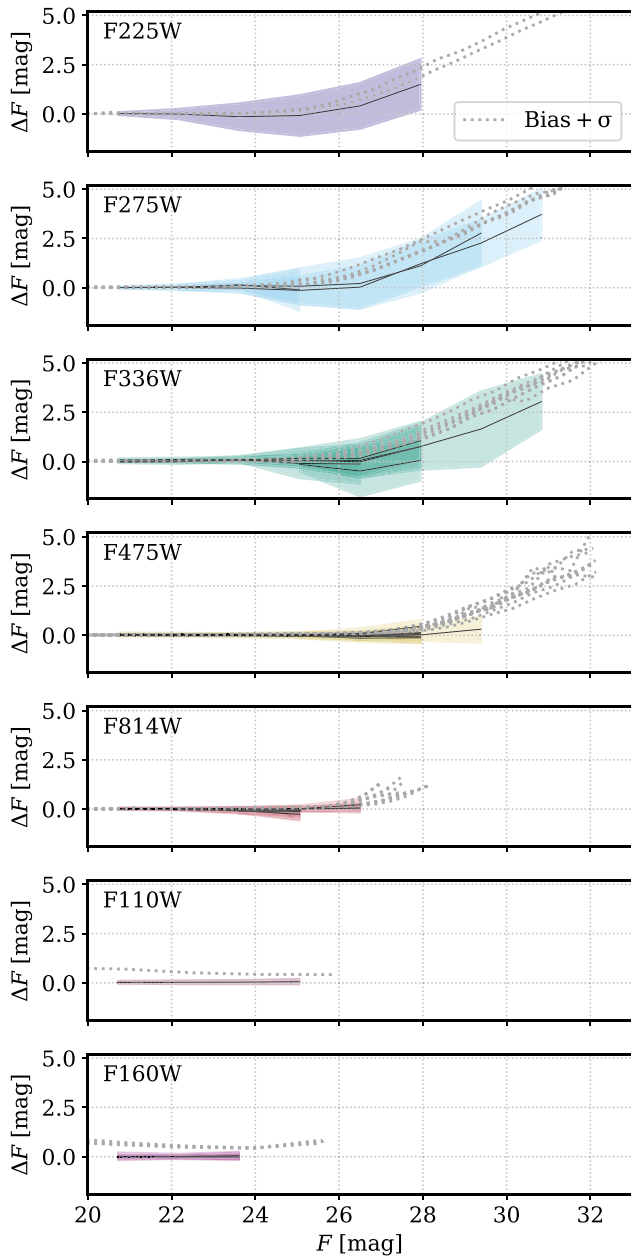


Figure 10. The difference in observed magnitude as a function of flux for fields that overlap spatially in Scylla. The black lines indicate the median difference in bins of decreasing flux (increasing magnitude), and the color-shaded regions denote $\pm 1\sigma$ within each band. We overplot the results of the ASTs (bias + uncertainty) for each field in each panel and observe that the differences are all within measured uncertainties.

stellar populations, as seen by the range of color and magnitude spanned by each stacked CMD.

In both MCs, our observations sample the unevolved main sequence down to $\gtrsim 2$ mag faintward of the ancient main-sequence turnoff in optical (F475W and F814W) filters, while cool lower main-sequence stars preferentially evade detection in the bluest filters (F275W, F336W). At the bright end, saturation prevents detection of massive main-sequence stars younger than a few hundred megayears. However, in all bandpasses the breadth of the observed evolutionary sequences is due in part to age and metallicity distributions that vary across the sight lines sampled by our fields.

A key difference between the LMC and SMC in CMD space is the shape of the red clump and the main sequence in the optical bands. The red clump consists of K giants, and although its morphology in the CMD depends on the age and metallicity distribution of the underlying population, it can function as a standard candle (see L. Girardi 2016 for a review).

In the presence of dust, however, these stars experience a shift toward fainter magnitudes and redder colors. Varying amounts of dust cause this effect to a varying degree, and the result is that the red clump appears extended in color and magnitude along a reddening vector (shown in gray in the upper right corner of each panel of Figure 11). We observe from the optical and IR CMDs that the LMC red clump is highly elongated along this reddening vector whereas the SMC red clump is much more compact along the color axis. The effect appears reversed for the UV CMD, but we note that this is due to the higher density of sources in the SMC UV CMD since in the SMC there are more Scylla fields that were observed with HST’s UV filters than in the LMC (see Table 1). Additionally, in the optical bands the LMC main sequence appears significantly wider than the SMC main sequence. This suggests that the LMC contains more dust, which is consistent with previous analysis of IR emission in these systems (K. D. Gordon et al. 2014; J. Chastenot et al. 2017; D. Utomo et al. 2019; C. J. R. Clark et al. 2023). CMD-based analysis of LMC and SMC dust extinction properties similarly suggests that the amount of dust in the two galaxies causes noticeably different effects, specifically in the appearance of CMD features such as the red clump and the red giant branch (RGB; G. De Marchi et al. 2016; P. Yanchulova Merica-Jones et al. 2017).

To further investigate the variations between fields, we plot the individual optical CMDs (as F475W and F814W are common to all observations) for all Scylla fields in the LMC (Figure 12) and SMC (Figure 13). We observe that there is a significant variation in CMD shapes between the 96 fields, between and within each MC. Due to the lack of short exposures in most fields, our saturation limits (shown in shaded gray) illustrate our lack of sensitivity to the youngest and oldest bright sources. Even within each MC, differences in extinction among and within different sight lines are readily apparent (e.g., compare LMC_2, LMC_53, and LMC_54 in Figure 12). Furthermore, recent star formation is evident in some cases (e.g., LMC_44 and LMC_52 near the N11 star-forming complex) from the presence of pre-main-sequence stars redward of and parallel to the main sequence (D. A. Gouliermis 2012).

In Figures 12 and 13 we observe that some fields retain a significant number of sources blueward of the main sequence with roughly constant F475W magnitude (e.g., LMC_1, LMC_44, SMC_1, SMC_22). These fields are all cases with two- or three-filter coverage, meaning that the DOLPHOT source detection had access to the fewest number of images relative to the rest of the fields. In addition, by inspection of these cases, our quality cuts to remove diffraction spikes did not perform as well. We note that these sources can be removed by imposing an SNR cut on the photometry but were explicitly included in the *vgst* catalogs for completeness considerations (see Section 3).

5.2. Spectral Energy Distributions

In addition to CMDs, we can also analyze how the observed SEDs of sources from different CMD regions compare. In

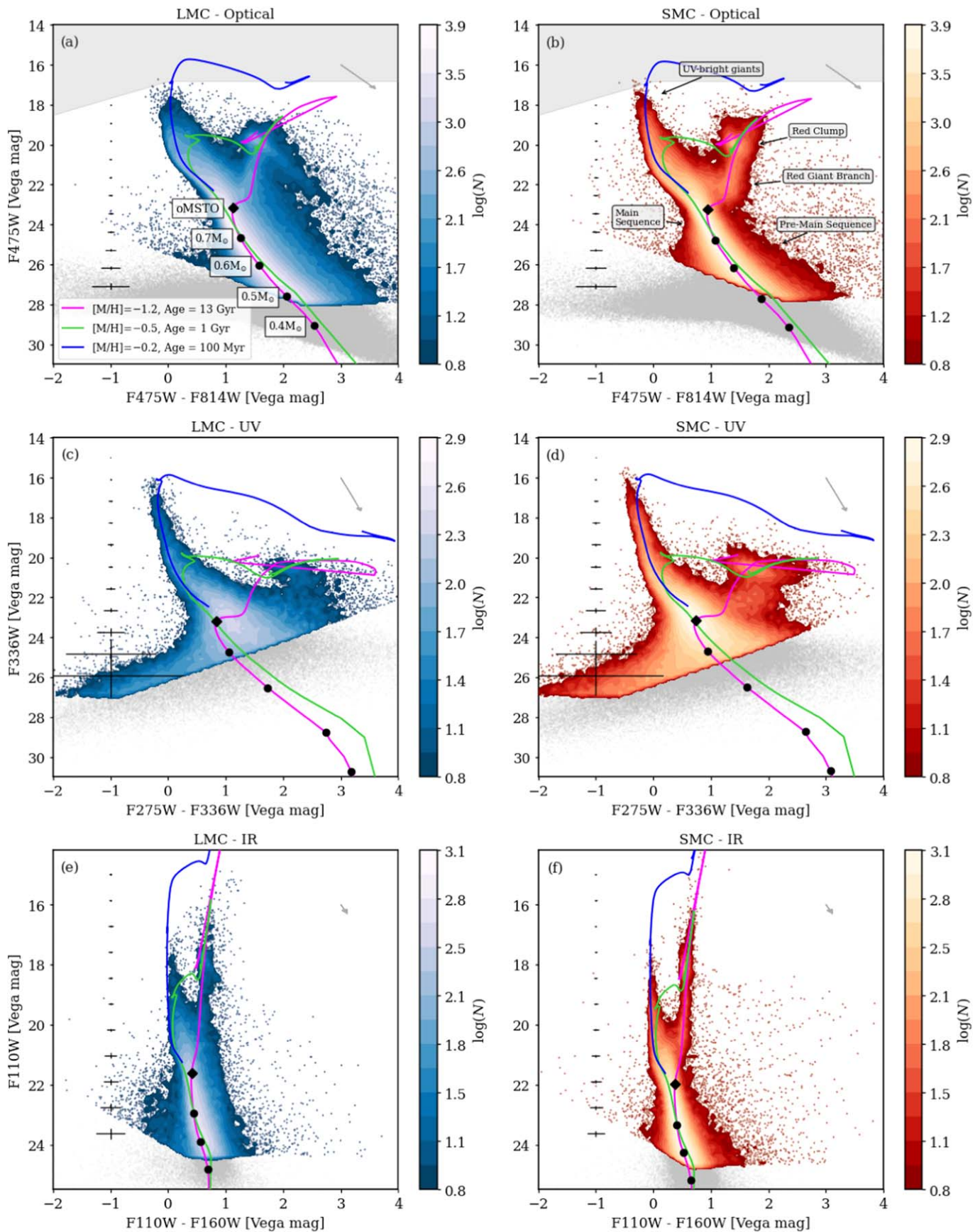


Figure 11. Stacked CMDs in the optical (top), UV (middle), and IR (bottom) bands for the LMC (left) and SMC (right) for the *vgst* catalogs. Sources brighter than the average 50% completeness limits are shown in color contours, and fainter sources are gray. Average photometric uncertainties as a function of flux, computed from the ASTs, are shown as error bars. Saturation limits in each band are shaded gray in the optical panels (all others are much brighter). Reddening vectors (gray, top right) and PARSEC isochrones are overlaid, indicating the main-sequence stellar mass range typical of our photometric catalogs.



Figure 12. Optical CMDs for all LMC fields from the *vgst* catalogs. Saturation limits in each band are shaded gray in each panel. Reddening vectors (gray, upper right corner) and PARSEC isochrones are overlaid, indicating the main-sequence stellar mass range typical of our photometric catalogs (same as Figure 11). We indicate the number of filters obtained for each field in the lower right corners.



Figure 13. Optical CMDs for all SMC fields from the *vgst* catalogs. Saturation limits in each band are shaded gray in each panel. Reddening vectors (gray, upper right corner) and PARSEC isochrones are overlaid, indicating the main-sequence stellar mass range typical of our photometric catalogs (same as Figure 11). We indicate the number of filters obtained for each field in the lower right corners.

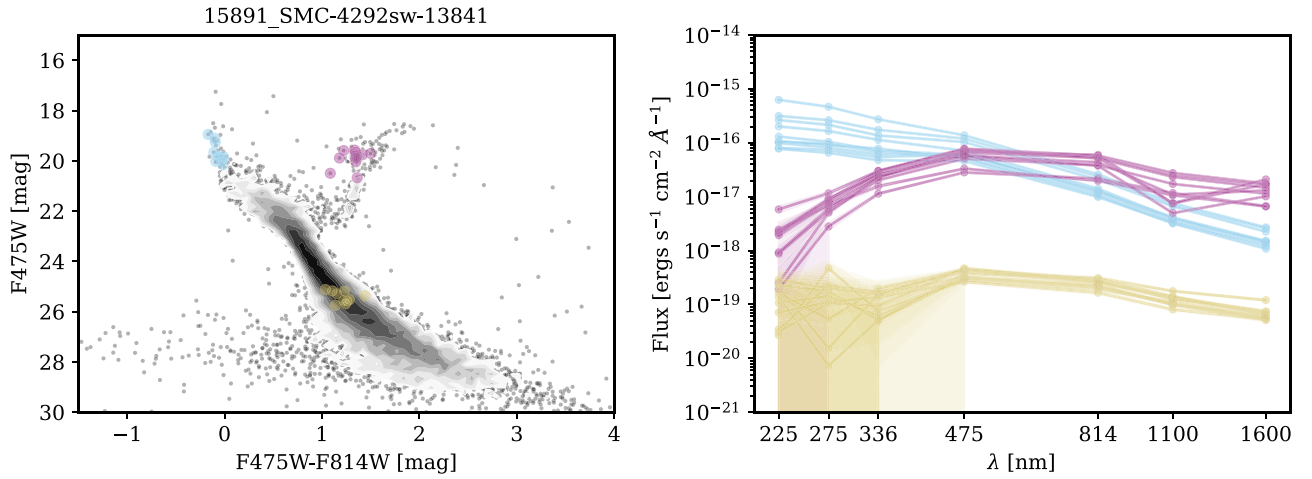


Figure 14. Example SEDs of sources from different regions on a CMD. Left: an optical CMD of SMC Field 15 (15891_SMC-4292sw-13841) with 10 randomly selected sources in each of the three regions (young main sequence: blue; RGB: pink; dim main sequence: yellow). Right: SED of each source, colored by region on the CMD.

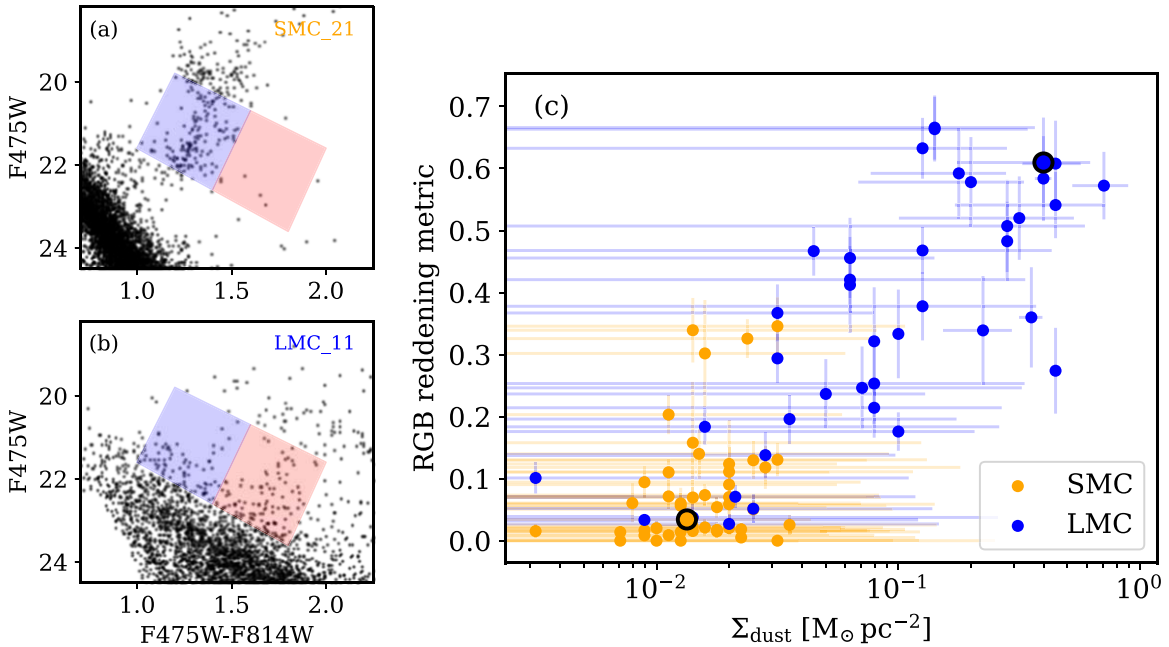


Figure 15. Demonstration of simple reddening analysis of Scylla optical CMDs. In panels (a) and (b), we zoom in to the RGB of two example fields, SMC_21 and LMC_11. Two regions of the CMD are highlighted, and the reddening metric is computed by dividing the number of sources in the red box by the total number of sources in both boxes. In panel (c), we compare this RGB reddening metric with the dust mass surface density (C. J. R. Clark et al. 2023), finding a positive correlation. The two example fields from panels (a) and (b) are highlighted with black outlines.

Figure 14, we show the SEDs of 10 different sources from three separate regions on an optical CMD in SMC_15. This field was observed in seven bands, so the SEDs span from the NUV to the NIR. We can see that for the brightest sources on the main sequence (blue) the SEDs are brightest in the NUV as expected, as opposed to RGB sources (pink), which are brightest in the redder optical bands (F475W and F814W). For dimmer sources along the main sequence, we observe significant variability in their NUV detections within large uncertainties.

5.3. Tracing Dust Reddening with a CMD

In addition to insight into the stellar populations of the MCs, the Scylla CMDs also trace the properties of dust reddening along the line of sight. In Figure 15, we demonstrate a proof of concept. In panels (a) and (b) we zoom in to the RGB region of

one example Scylla field each from the SMC and LMC. We define two regions in CMD space based on the structure of the RGB analyzed by the SMIDGE survey (P. Yanchulova Merica-Jones et al. 2017), one red and one blue. We then define an RGB reddening metric by counting the number of sources in the red box and dividing by the total number of sources in both boxes. This simple metric aims to capture, in broad strokes, the elongation of the RGB due to the effect of dust reddening.

To test this simple reddening metric, in Figure 15(c) we plot it against the dust mass surface density (Σ_{dust}) derived from IR SED modeling of the MCs (C. J. R. Clark et al. 2023). Despite considerable scatter, we find that the two tracers are positively correlated. This suggests that we are able to recover properties of MC dust using the Scylla CMDs. In future work, we will refine this analysis to disentangle the effects of reddening

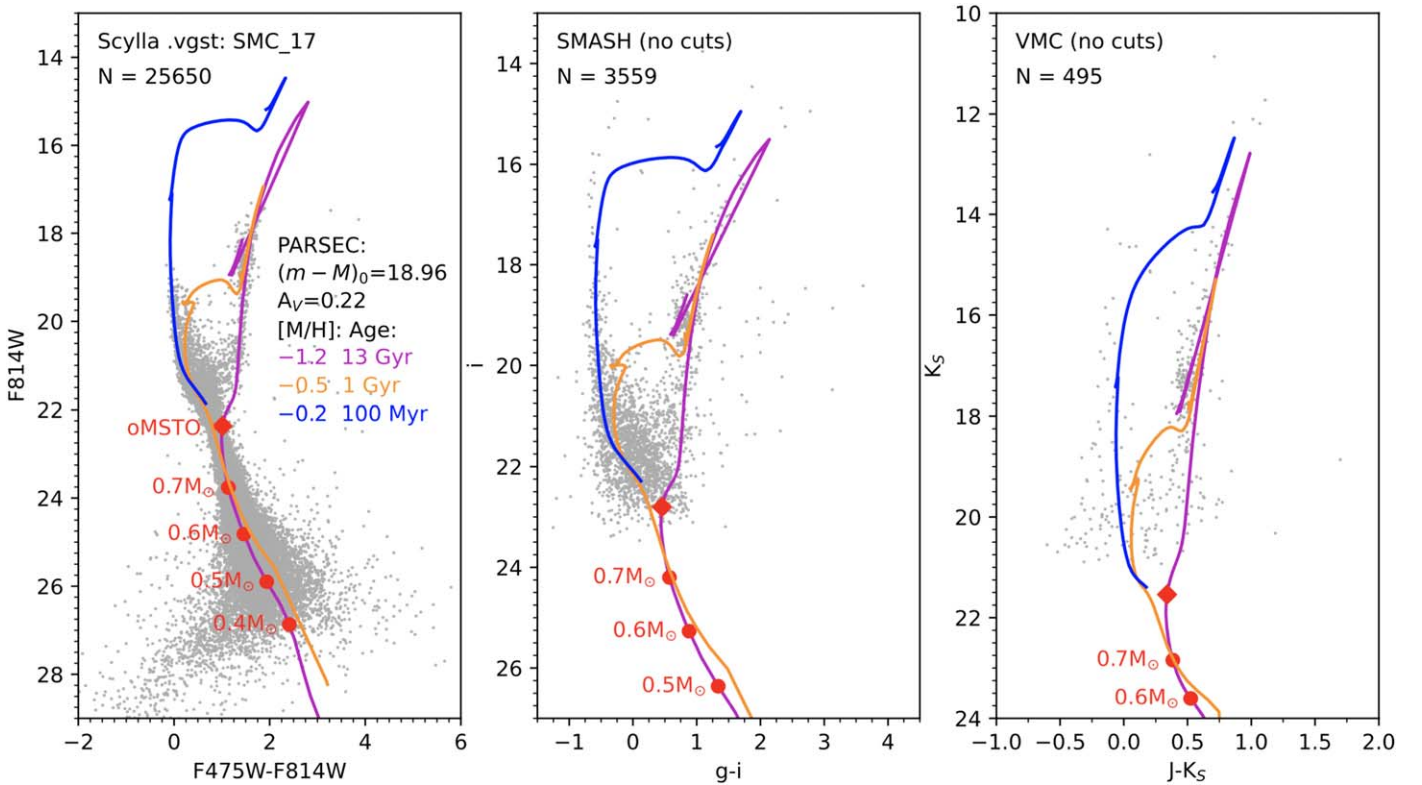


Figure 16. A comparison of optical CMDs from Scylla and ground-based surveys within Scylla (*vgst*; left), SMASH (middle), and VMC (right). For the ground-based surveys, we extract all sources within a single Scylla field (SMC_17) and apply no quality cuts. In each panel, we overlay PARSEC isochrones, indicating the main-sequence stellar mass range typical of our photometric catalogs.

on the structure of the RGB from the effects of line-of-sight distance.

5.4. Comparison with Other MC Surveys

As discussed in Section 1, Scylla’s overall spatial coverage is larger than any other HST survey of comparable depth to date. To date, Scylla has observed 48 fields each in the LMC and SMC. In comparison, the METAL survey covered 33 fields (J. Roman-Duval et al. 2019) dispersed throughout the LMC, the HTTP survey (E. Sabbi et al. 2016) focused on 30 Doradus (equivalent to ~ 36 fields), and the SMIDGE survey covered the equivalent of 18 fields in the SMC SW Bar.

Similarly to METAL, Scylla uses HST’s WFC3/UVIS (F225W, F275W, F336W, F475W, F814W) and WFC3/IR (F110W and F160W) filters. HTTP also uses the Advanced Camera for Surveys (ACS)/WFC in the F555W, F658N, and F775W filters (without F225W), and SMIDGE also uses the ACS/WFC in the F550M and F658N filters.

The Scylla survey’s significant photometric depth can be compared to that of other HST MC surveys. For example, the 50% completeness limit of METAL is on average ~ 1.5 mag deeper in F225W, F275W, F336W, and F475W, while it is about the same in F814W and F160W and ~ 0.5 mag shallower in F110W. The 50% completeness limit of SMIDGE, which only targets the SMC, is 0.8–1.8 mag deeper in all filters except F475W, where it is only 0.4 mag deeper than Scylla’s completeness. In 30 Doradus, HTTP reaches an average 50% completeness of ~ 2 mag shallower in F275W and F336W and 0.5 mag shallower in F110W, while it has the same completeness as Scylla in F160W. As for the optical bands, the two surveys observed in different filters, where the HTTP

50% completeness in F555W is 25.8 mag and the Scylla 50% completeness in F475W is 27.8 mag.

While there are rich MC observations from ground-based surveys, the spatial resolution and depth of HST are unsurpassed. Due to stellar crowding, ground-based surveys show significantly shallower CMDs in the central regions of the galaxies compared to the outer parts. For example, in Figure 16 we compare the CMD of a typical Scylla field (SMC_17) with the results from the VISTA Survey of the MCs (VISTA; M. R. L. Cioni et al. 2011) and the Survey of the Magellanic Stellar History (SMASH; D. L. Nidever et al. 2017, 2021). Within the same footprint as SMC_17, we extracted SMASH photometry¹⁷ in similar filters (g, i) to F475W and F814W and VMC photometry¹⁸ in oft-used J, K_S bandpasses.

Based on Figure 16, it is clear that Scylla reaches several magnitudes faintward of the oMSTO, deeper than SMASH or VMC in crowding conditions typical of the inner SMC. We overlay stellar isochrones from PARSEC onto the CMDs to demonstrate that the photometric depth attained by Scylla corresponds to a stellar mass sensitivity of $\sim 0.5 M_{\odot}$. Despite our improved sensitivity, Scylla is limited to 96 HST imaging footprints (to date) across both MCs sampling only $\sim 0.08 \text{ kpc}^2$ in the LMC and $\sim 0.1 \text{ kpc}^2$ in the SMC, whereas SMASH and VMC provide contiguous spatial coverage of both systems, as well as enhanced sensitivity to bright sources that are saturated by our typically long exposures.

¹⁷ <https://datalab.noirlab.edu/smash/smash.php>

¹⁸ <http://horus.roe.ac.uk/vsa/>

6. Conclusions

In this work, we present an overview of the Scylla HST survey. Scylla was designed to investigate the stellar populations, ISM, and star formation in the MCs. We describe the science goals, observing strategy, data reduction procedures, and initial results from the photometric analysis of 96 observed fields.

Although constrained by the ULLYSES primary observations, we imaged all fields in at least two filters (F475W and F814W), with 64% of fields imaged in three or more filters spanning the UV to IR. This comprehensive approach allowed us to achieve an average 50% completeness of $m_{F225W} = 26.0$, $m_{F275W} = 26.2$, $m_{F336W} = 26.9$, $m_{F475W} = 27.8$, $m_{F814W} = 25.5$, $m_{F110W} = 24.7$, and $m_{F160W} = 24.0$ Vega mag in our `vgst` (quality-restricted) catalogs, reaching faintward of the ancient main-sequence turnoff in all filters.

Our photometric analysis revealed significant variations in CMDs between and within each MC, highlighting the diverse stellar populations and dust extinction properties across the observed fields. In addition, we demonstrated the feasibility of extracting dust reddening information from the shape of the RGB in the CMDs.

Additionally, we compared our survey statistics with previous HST and ground-based surveys of the MCs. We found that Scylla achieves significantly better sensitivity than ground-based surveys in the crowded central regions of the MCs, despite being limited to a smaller observing footprint.

The Scylla survey's deep, multiwavelength imaging and photometric catalogs provide a valuable resource for characterizing the structure and properties of dust, as well as the spatially resolved star formation and chemical enrichment histories of the MCs. These data products are available as an HLSP at MAST via doi:[10.17909/mk54-kg51](https://doi.org/10.17909/mk54-kg51). This release will enable further investigations into the complex interplay between stellar populations and the ISM in these nearby galaxies.

In future work, we will further refine our analysis to disentangle the effects of reddening and line-of-sight distance on the structure of the MCs. We will also utilize the BEAST to

generate a comprehensive catalog of stellar and dust parameters, providing detailed insights into the diverse environments within the MCs.

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Facilities: HST (WFC3/IR), HST (WFC3/UVIS).

Software: DOLPHOT, BEAST (K. D. Gordon et al. 2016), TOPCAT (M. B. Taylor 2005), astropy (Astropy Collaboration 2013, 2018), numpy (C. R. Harris et al. 2020), scipy (P. Virtanen et al. 2020), matplotlib (J. D. Hunter 2007), glue (C. Beaumont et al. 2015; T. Robitaille et al. 2018).

Appendix A Exposure Information

In this appendix we include a table with exposure-level information for each target field. Table 5 includes the field name, file root (i.e., MAST identifier), position, position angle, detector, filter, exposure time, and level of post-flash for each exposure.

Table 5
Observing Parameters

Name	Field	Root	R.A.	Decl.	PA	Detector	Filter	T_{exp}	Post-flash
(1)	(2)	(3)	(deg)	(deg)	(deg)	(7)	(8)	(s)	(s)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
SMC_15	15891_SMC-4292sw-13841	ie9m68jaq	9.5636155	-73.400765	138.409	IR-FIX	F160W	499.2	...
SMC_15	15891_SMC-4292sw-13841	ie9m68jdaq	9.5604039	-73.403033	138.409	UVIS	F225W	382.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m68jfq	9.5604039	-73.403033	138.409	UVIS	F225W	436.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m68jjq	9.5604039	-73.403033	138.409	UVIS	F275W	433.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m69jmq	9.5636155	-73.400765	138.409	IR-FIX	F110W	349.2	...
SMC_15	15891_SMC-4292sw-13841	ie9m69joq	9.5604039	-73.403033	138.409	UVIS	F475W	710.0	2.3
SMC_15	15891_SMC-4292sw-13841	ie9m69jq	9.5604039	-73.403033	138.409	UVIS	F275W	400.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m69juq	9.5604039	-73.403033	138.409	UVIS	F475W	469.0	4.3
SMC_15	15891_SMC-4292sw-13841	ie9m69jq	9.5604039	-73.403033	138.409	UVIS	F475W	3.0	7.7
SMC_15	15891_SMC-4292sw-13841	ie9m70jyq	9.5604039	-73.403033	138.409	UVIS	F336W	359.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m70k1q	9.5604039	-73.403033	138.409	UVIS	F814W	685.0	2.0
SMC_15	15891_SMC-4292sw-13841	ie9m70k3q	9.5604039	-73.403033	138.409	UVIS	F336W	380.0	7.3
SMC_15	15891_SMC-4292sw-13841	ie9m70k7q	9.5604039	-73.403033	138.409	UVIS	F814W	455.0	3.9
SMC_15	15891_SMC-4292sw-13841	ie9m70k9q	9.5604039	-73.403033	138.409	UVIS	F814W	3.0	7.7
SMC_45	15891_SMC-641nw-12753	ie9m08ccq	12.617763	-72.723821	127.527	UVIS	F475W	1330.0	0.0
SMC_45	15891_SMC-641nw-12753	ie9m08ctq	12.617763	-72.723821	127.527	UVIS	F814W	734.0	1.6
SMC_45	15891_SMC-641nw-12753	ie9m09d0q	12.617763	-72.723821	127.527	UVIS	F475W	1720.0	0.0
SMC_45	15891_SMC-641nw-12753	ie9m09d3q	12.617763	-72.723821	127.527	UVIS	F814W	715.0	2.0
SMC_35	16235_SMC-2773nw-32334	iehs77b9q	12.884764	-72.03468	323.343	UVIS	F814W	1040.0	0.0
SMC_35	16235_SMC-2773nw-32334	iehs77bcq	12.884764	-72.03468	323.343	UVIS	F814W	1040.0	0.0
SMC_35	16235_SMC-2773nw-32334	iehs78bhq	12.884764	-72.03468	323.343	UVIS	F475W	1183.0	0.0
SMC_35	16235_SMC-2773nw-32334	iehs78bkq	12.884764	-72.03468	323.343	UVIS	F475W	1182.0	0.0
SMC_35	16235_SMC-2773nw-32334	iehs79bpq	12.884764	-72.03468	323.343	UVIS	F336W	1178.0	7.1
SMC_35	16235_SMC-2773nw-32334	iehs79bsq	12.884764	-72.03468	323.343	UVIS	F336W	1177.0	7.1

Note. Observing parameters of all Scylla exposures. Column (1): name. Column (2): long-form name. Column (3): root (aka the exposure identifier in MAST). Column (4): R.A. Column (5): decl. Column (6): position angle. Column (7): detector name. Column (8): filter. Column (9): exposure time. Column (10): post-flash level for the exposure.

(This table is available in its entirety in machine-readable form in the [online article](#).)

Appendix B

Identifying “vgst” Cuts

To determine photometric quality cuts for the Scylla survey, we manually inspected a test field, SMC_6 (selected randomly owing to its wide filter coverage). Using TOPCAT (M. B. Taylor 2005), we selected sources that were spatially coincident with diffraction spikes. From a total of 93,628 detected sources, 43% (40,516)

were flagged. In Figure 17, we display the spatial distribution of the sources and the distributions of sharpness, roundness, and crowding for the visually selected contaminant sample and the rest of the field in bands F475W and F814W (filters common to all Scylla fields).

Based on this result, we selected the *vgst* quality cuts summarized in Table 3.

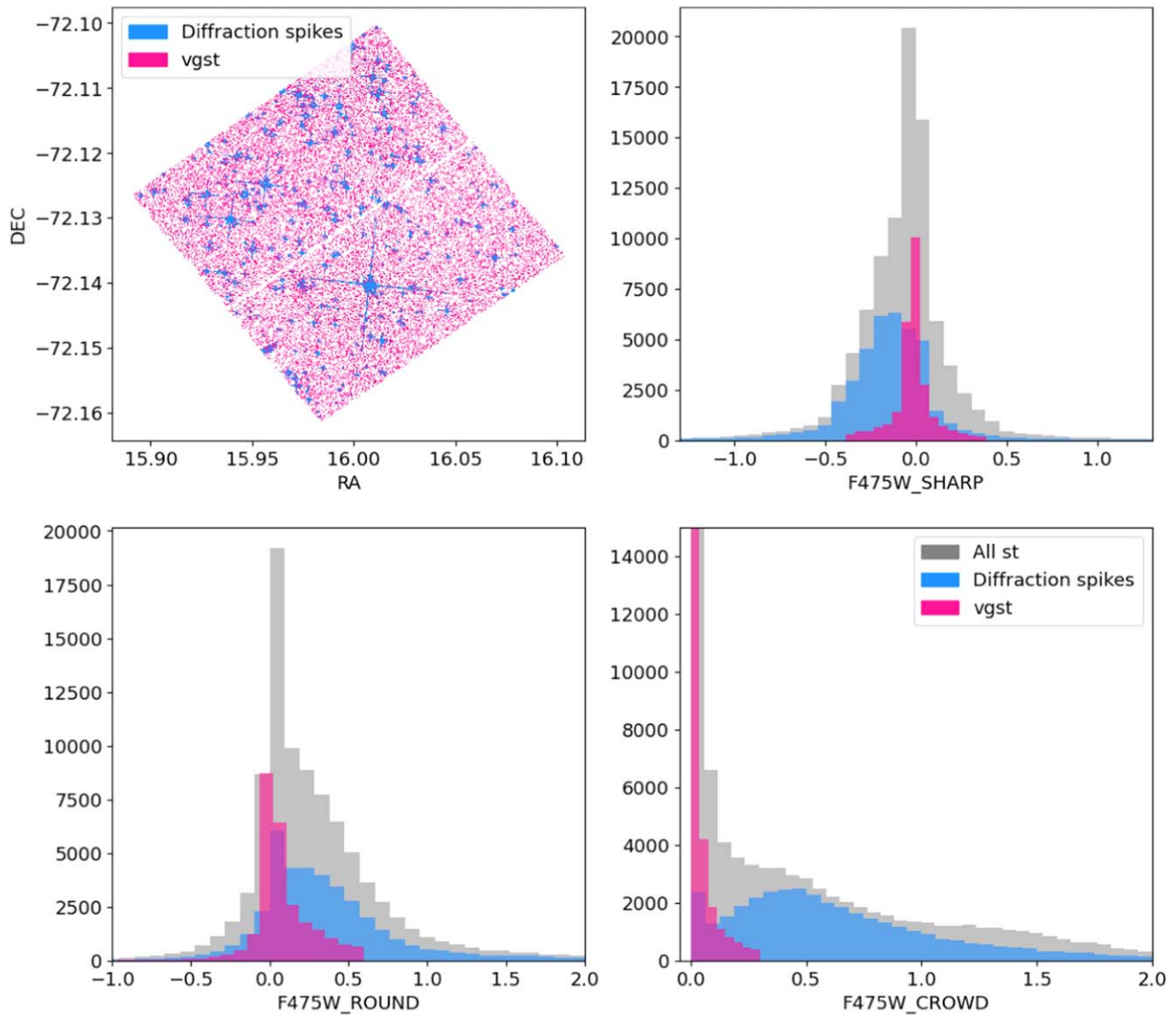


Figure 17. A sample field illustrating the selection process for the *vgst* quality cuts. Top left: first, we visually identify contaminant sources as indicated by diffraction spikes in the imaging footprint (blue). We then evaluate the sharpness, roundness, and crowding parameters of the full sample of stars and the diffraction spike sources, and we select ranges for these photometric quality parameters, which optimize the removal of diffraction spikes.

Appendix C WFC3 UVIS Variations

Following photometry and quality cuts, we observed residual magnitude variations in our *vgst* catalog sources as a function of position relative to the WFC3 UVIS camera chip gap. One potential explanation for these variations could be the effects of CTE loss. Space-based CCD detectors can suffer from CTE degradation when exposed to cosmic rays, impacting precision in photometric observations.¹⁹ WFC3 UVIS is a two-chip CCD, meaning that there is a chip gap that runs along the center of the detector, while readout amplifiers are located at the corners of the detector. As charge is transferred to the readout amplifiers and passes through pixels that may have been damaged by cosmic rays, it gets temporarily trapped in the

lattice. It has been documented that charge obtained farther from the amplifiers will result in an underestimation of flux, and since regions near the chip gap are the farthest, charge from this region of the lattice is most affected. Without corrections for CTE loss, we expect sources near the chip gap to appear dimmer than the sources near the readout across all filters.

This issue is usually remedied by including CTE corrections in the photometric reduction pipeline, increasing background levels by post-flashing the images, or obtaining dithered observations of the same field. However, since Scylla is a pure-parallel program, we were not able to obtain dithered exposures for any of our fields. However, post-flash exposures were obtained for most fields (see Table 5). Despite these precautions, we still find spatial variations in the distribution of observed magnitudes.

As an example, in Figure 18 we show the distributions of CTE-corrected magnitudes for the 16% of sources closest to the chip gap (purple) and farthest from the chip gap (green). This field, SMC_41, was selected as an example owing to its homogeneous distribution of sources and lack of extended

¹⁹ CTE degradation for the WFC3 UVIS detector has been monitored since 2009 and is documented in several Instrument Science Reports (www.stsci.edu/hst/instrumentation/wfc3/documentation/instrument-science-reports-isrs?keyword=2017-09%26itemsPerPage=15), such as WFC3 2017-09 and WFC3 2018-15.

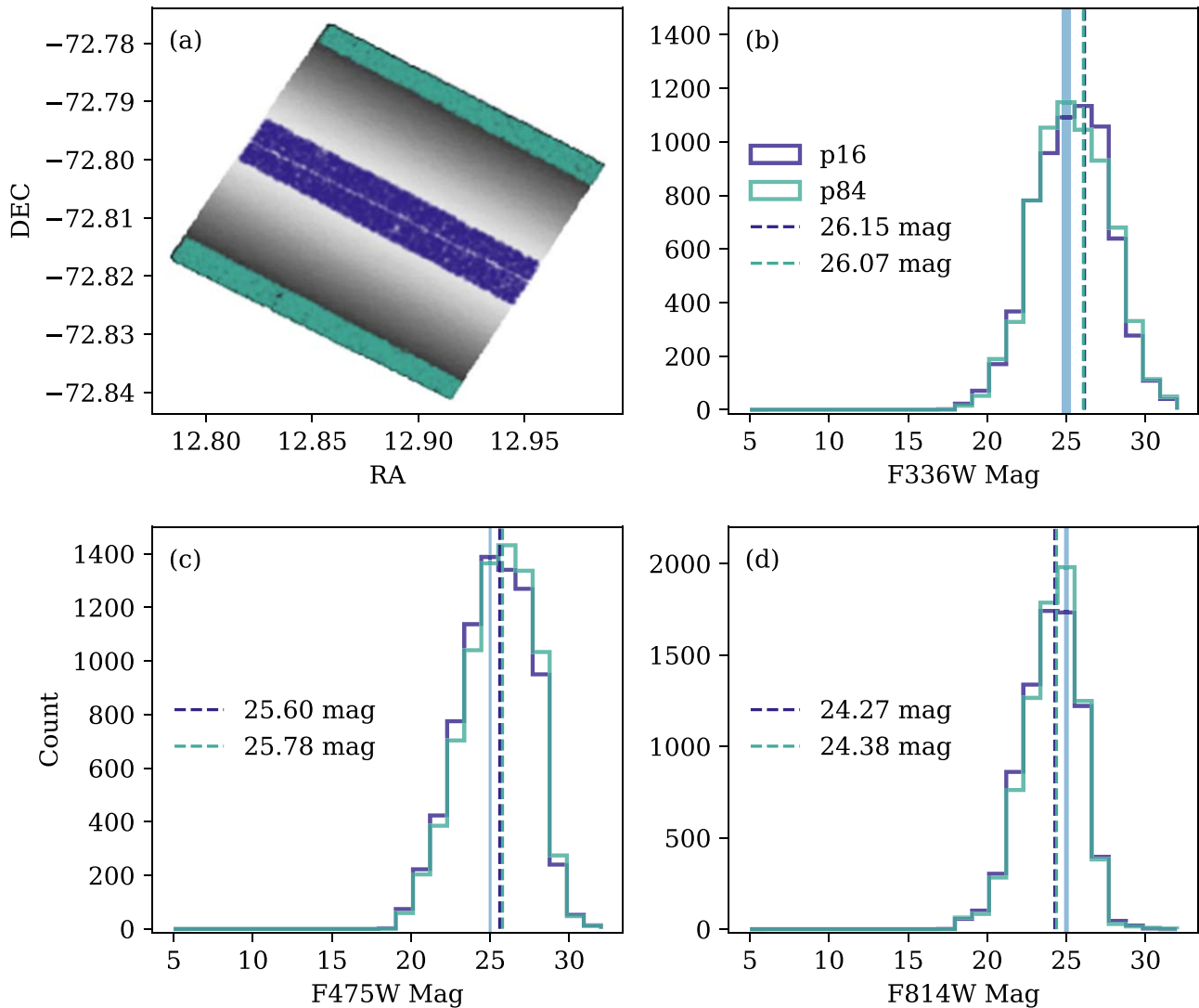


Figure 18. Observed CTE-corrected magnitude variations as a function of chip gap distance. (a) On-sky spatial distribution of all *vgst* sources in SMC Field 41 (16235_SMC-286sw-34349). Sources in the lower 16th/upper 84th percentile of distances from the chip gap are highlighted in purple/green, respectively. The magnitude distributions of these sources in WFC3 UVIS filters F336W, F475W, and F814W are shown in panels (b), (c), and (d), respectively. The median values of each distribution are shown with a dashed line, and the 3σ magnitude uncertainties at 25 mag derived from ASTs in this particular field are shaded in blue (0.3, 0.09, and 0.15, respectively).

emission. For filters F475W and F814W, sources near the chip gap are systematically brighter than sources near the readouts. This is the opposite for F336W, where sources near the chip gap are dimmer than sources near the readouts.

We test whether the two populations in each filter are statistically different by performing a Kolmogorov–Smirnov test, which measures the probability that two populations were drawn from the same distribution. For the two populations in F475W, we obtain a p -value of 1.13×10^{-4} , which rejects the null hypothesis at a 95% confidence level (p -value < 0.05) that the two populations are drawn from the same distribution, i.e., the two populations *are not drawn* from the same distribution. We report similar findings for the two populations in F814W, which have a p -value of 6.66×10^{-4} . Conversely, for the two populations in F336W, we obtain a p -value of 0.347, indicating that the two populations *are drawn* from the same distribution.

We find statistically significant differences in the observed magnitudes as a function of chip gap distance in filters F475W and F814W. Based on the observed magnitude distributions, we suspect that the CTE correction algorithm is overcorrecting the *flc* images in filters F475W and F814W. However, these differences are not significantly larger than photometric uncertainties in our filters ($\sigma = 0.1$, 0.03, and 0.05 mag around the 25th magnitude for F336W, F475W, and F814W, respectively). In Figure 18, we plot the median magnitude of each distribution (dashed lines) to compare with the 25th magnitude 3σ uncertainty derived from our ASTs in the field (shaded). From this, we see that the differences in the median magnitudes are comparable to our photometric uncertainties, indicating that there may be lingering CTE effects present in our photometry. However, these effects will vary in strength depending on the conditions of each observation and should not impact our overall science.

Appendix D

Summary of AST Results

In Tables 6 and 7, we summarize the AST results for the source-density-dominated and background-density-dominated fields, respectively. For bins of increasing source or background density, we compute the completeness, bias, and uncertainty as a

function of flux in each observed band. We include uncertainties on these parameters, which were computed by bootstrapping the binned sample with replacement over 1000 trials.

In Table 8, we summarize the 50% completeness limits as a function of observing band for each field.

Table 6
Summary of ASTs for Source-density-dominated Fields

Source Density (sources arcsec ⁻²) (1)	Filter (2)	Magnitude (Vega mag) (3)	Completeness (4)	Bias (Vega mag) (5)	Uncertainty (Vega mag) (6)
0.0–0.27	F225W	18.25	1.0 ± 0.006	0.001 ± 0.002	0.003 ± 0.026
0.0–0.27	F225W	18.75	0.834 ± 0.06	0.004 ± 0.001	0.008 ± 0.01
0.0–0.27	F225W	19.25	0.989 ± 0.032	0.001 ± 0.001	0.007 ± 0.006
0.0–0.27	F225W	19.75	1.0 ± 0.013	0.009 ± 0.001	0.008 ± 0.001
0.0–0.27	F225W	20.25	0.846 ± 0.069	0.021 ± 0.003	0.018 ± 0.001
0.0–0.27	F225W	20.75	1.0 ± 0.014	0.027 ± 0.004	0.028 ± 0.002
0.0–0.27	F225W	21.25	0.983 ± 0.009	0.031 ± 0.008	0.043 ± 0.006
0.0–0.27	F225W	21.75	0.995 ± 0.013	0.047 ± 0.004	0.046 ± 0.007
0.0–0.27	F225W	22.25	1.0 ± 0.01	0.066 ± 0.003	0.068 ± 0.003
0.0–0.27	F225W	22.75	0.983 ± 0.015	0.07 ± 0.01	0.132 ± 0.013
0.0–0.27	F225W	23.25	1.0 ± 0.012	0.117 ± 0.01	0.132 ± 0.013
0.0–0.27	F225W	23.75	1.0 ± 0.009	0.22 ± 0.013	0.171 ± 0.025
0.0–0.27	F225W	24.25	0.997 ± 0.008	0.301 ± 0.027	0.312 ± 0.042
0.0–0.27	F225W	24.75	0.969 ± 0.013	0.41 ± 0.036	0.429 ± 0.065
0.0–0.27	F225W	25.25	0.915 ± 0.027	0.616 ± 0.048	0.668 ± 0.069
0.0–0.27	F225W	25.75	0.73 ± 0.029	0.455 ± 0.155	0.926 ± 0.059
0.0–0.27	F225W	26.25	0.487 ± 0.048	0.567 ± 0.066	0.82 ± 0.206
0.0–0.27	F225W	26.75	0.639 ± 0.065	0.055 ± 0.121	0.946 ± 0.145
0.0–0.27	F225W	27.25	0.353 ± 0.046	−0.772 ± 0.294	1.24 ± 0.134
0.0–0.27	F225W	27.75	0.256 ± 0.043	−1.304 ± 0.459	0.833 ± 0.288
0.0–0.27	F225W	28.25	0.223 ± 0.056	−0.87 ± 0.169	1.139 ± 0.387
0.0–0.27	F275W	18.25	1.0 ± 0.013	−0.002 ± 0.001	0.005 ± 0.007
0.0–0.27	F275W	18.75	0.978 ± 0.017	0.001 ± 0.001	0.006 ± 0.002
0.0–0.27	F275W	19.25	1.0 ± 0.007	0.003 ± 0.001	0.007 ± 0.011
0.0–0.27	F275W	19.75	0.999 ± 0.011	0.007 ± 0.001	0.015 ± 0.003
0.0–0.27	F275W	20.25	0.97 ± 0.024	0.009 ± 0.001	0.015 ± 0.003
0.0–0.27	F275W	20.75	0.968 ± 0.006	0.012 ± 0.001	0.022 ± 0.002
0.0–0.27	F275W	21.25	0.98 ± 0.007	0.045 ± 0.014	0.169 ± 0.051
0.0–0.27	F275W	21.75	0.978 ± 0.017	0.031 ± 0.002	0.041 ± 0.002
0.0–0.27	F275W	22.25	0.976 ± 0.005	0.061 ± 0.008	0.17 ± 0.046
0.0–0.27	F275W	22.75	0.98 ± 0.008	0.075 ± 0.003	0.084 ± 0.01
0.0–0.27	F275W	23.25	0.991 ± 0.006	0.114 ± 0.004	0.131 ± 0.006
0.0–0.27	F275W	23.75	0.986 ± 0.007	0.182 ± 0.005	0.16 ± 0.01
0.0–0.27	F275W	24.25	0.962 ± 0.011	0.274 ± 0.018	0.313 ± 0.027
0.0–0.27	F275W	24.75	0.947 ± 0.016	0.334 ± 0.044	0.449 ± 0.031
0.0–0.27	F275W	25.25	0.887 ± 0.015	0.645 ± 0.046	0.756 ± 0.036
0.0–0.27	F275W	25.75	0.807 ± 0.02	0.617 ± 0.068	0.859 ± 0.109
0.0–0.27	F275W	26.25	0.602 ± 0.038	0.529 ± 0.044	1.121 ± 0.054
0.0–0.27	F275W	26.75	0.463 ± 0.012	−0.143 ± 0.155	0.9 ± 0.073
0.0–0.27	F275W	27.25	0.351 ± 0.044	0.403 ± 0.176	1.408 ± 0.123
0.0–0.27	F275W	27.75	0.275 ± 0.034	−0.309 ± 0.246	1.128 ± 0.064
0.0–0.27	F275W	28.25	0.248 ± 0.03	−0.791 ± 0.173	1.063 ± 0.098
0.0–0.27	F336W	18.25	0.964 ± 0.014	−0.002 ± 0.001	0.003 ± 0.003
0.0–0.27	F336W	18.75	1.0 ± 0.004	−0.001 ± 0.001	0.004 ± 0.004
0.0–0.27	F336W	19.25	0.972 ± 0.014	0.0 ± 0.001	0.024 ± 0.014
0.0–0.27	F336W	19.75	0.982 ± 0.004	0.001 ± 0.001	0.007 ± 0.013

(This table is available in its entirety in machine-readable form in the [online article](#).)

Table 7
Summary of ASTs for Background-density-dominated Fields

Background Density (sources arcsec ⁻²) (1)	Filter (2)	Magnitude (Vega mag) (3)	Completeness (4)	Bias (Vega mag) (5)	Uncertainty (Vega mag) (6)
0.006–5.385	F275W	18.25	1.0 ± 0.003	–0.003 ± 0.001	0.003 ± 0.002
0.006–5.385	F275W	18.75	1.0 ± 0.004	0.0 ± 0.002	0.006 ± 0.002
0.006–5.385	F275W	19.25	1.0 ± 0.008	0.0 ± 0.002	0.005 ± 0.005
0.006–5.385	F275W	19.75	1.0 ± 0.006	0.007 ± 0.003	0.004 ± 0.004
0.006–5.385	F275W	20.25	1.0 ± 0.021	0.01 ± 0.003	0.0 ± 0.008
0.006–5.385	F275W	20.75	1.0 ± 0.002	0.017 ± 0.008	0.029 ± 0.004
0.006–5.385	F275W	21.25	1.0 ± 0.003	0.033 ± 0.006	0.034 ± 0.009
0.006–5.385	F275W	21.75	1.0 ± 0.001	0.038 ± 0.015	0.037 ± 0.011
0.006–5.385	F275W	22.25	0.992 ± 0.025	0.019 ± 0.015	0.013 ± 0.013
0.006–5.385	F275W	22.75	0.938 ± 0.035	0.104 ± 0.011	0.078 ± 0.009
0.006–5.385	F275W	23.25	1.0 ± 0.001	0.062 ± 0.032	0.097 ± 0.035
0.006–5.385	F275W	23.75	1.0 ± 0.002	0.3 ± 0.055	0.285 ± 0.13
0.006–5.385	F275W	24.25	1.0 ± 0.005	0.438 ± 0.059	0.03 ± 0.172
0.006–5.385	F275W	24.75	0.813 ± 0.142	0.754 ± 0.306	0.414 ± 0.109
0.006–5.385	F275W	25.25	1.0 ± 0.047	0.569 ± 0.253	1.065 ± 0.133
0.006–5.385	F275W	25.75	1.0 ± 0.049	0.914 ± 0.26	1.653 ± 0.324
0.006–5.385	F275W	26.25	0.611 ± 0.068	0.182 ± 0.169	0.755 ± 0.172
0.006–5.385	F275W	26.75	0.604 ± 0.128	0.514 ± 0.311	0.703 ± 0.297
0.006–5.385	F275W	27.25	0.244 ± 0.087	–1.037 ± 0.385	0.315 ± 0.455
0.006–5.385	F275W	27.75	0.141 ± 0.094	–0.903 ± 0.22	0.651 ± 0.454
0.006–5.385	F275W	28.25	0.236 ± 0.117	–1.414 ± 0.485	0.638 ± 0.324
0.006–5.385	F336W	18.25	1.0 ± 0.013	–0.004 ± 0.001	0.002 ± 0.002
0.006–5.385	F336W	18.75	1.0 ± 0.014	0.001 ± 0.001	0.002 ± 0.001
0.006–5.385	F336W	19.25	0.994 ± 0.044	0.0 ± 0.001	0.005 ± 0.001
0.006–5.385	F336W	19.75	0.893 ± 0.06	0.006 ± 0.001	0.007 ± 0.01
0.006–5.385	F336W	20.25	1.0 ± 0.001	0.006 ± 0.001	0.006 ± 0.001
0.006–5.385	F336W	20.75	1.0 ± 0.006	0.007 ± 0.003	0.011 ± 0.004
0.006–5.385	F336W	21.25	1.0 ± 0.018	0.012 ± 0.005	0.024 ± 0.008
0.006–5.385	F336W	21.75	1.0 ± 0.021	0.076 ± 0.029	0.029 ± 0.007
0.006–5.385	F336W	22.25	1.0 ± 0.008	0.047 ± 0.006	0.022 ± 0.032
0.006–5.385	F336W	22.75	1.0 ± 0.034	0.038 ± 0.003	0.023 ± 0.016
0.006–5.385	F336W	23.25	0.952 ± 0.067	0.108 ± 0.015	0.051 ± 0.005
0.006–5.385	F336W	23.75	1.0 ± 0.013	0.086 ± 0.011	0.067 ± 0.011
0.006–5.385	F336W	24.25	1.0 ± 0.004	0.118 ± 0.025	0.084 ± 0.022

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Table 8
Completeness

Field Name	Avg R.A. (deg)	Avg Decl. (deg)	F225W (mag)	F275W (mag)	F336W (mag)	F475W (mag)	F814W (mag)	F110W (mag)	F160W (mag)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
SMC_15	9.562	-73.4013	25.7	25.7	26.5	27.7	25.4	24.9	24.2
SMC_45	12.6171	-72.721	27.6	25.3
SMC_35	12.8842	-72.0375	27.8	28.2	25.7
SMC_41	12.8849	-72.809	27.5	27.7	25.2
SMC_25	13.9209	-72.7082	26.9	27.3	25.0	...	23.9
SMC_40	14.238	-72.6194	26.0	26.2	26.7	27.2	25.1	...	23.9
SMC_42	14.2842	-72.5876	27.6	27.4	25.2	...	24.0
SMC_17	14.4594	-72.5908	...	26.4	26.7	27.8	25.5	24.9	24.3
SMC_7	14.5464	-71.4007	28.5	25.8
SMC_37	14.6806	-72.0602	26.2	26.3	26.8	27.3	25.2	24.6	24.0
SMC_23	14.7218	-72.2607	...	25.7	26.5	27.0	25.0	...	23.8
SMC_13	14.9313	-72.1748	...	25.8	26.1	27.1	25.1	...	23.9
SMC_14	14.9314	-71.9391	26.5	28.1	25.6
SMC_12	14.9344	-72.1575	27.7	25.3
SMC_1	14.939	-72.1032	27.9	25.4
SMC_19	14.944	-72.1573	27.6	25.4
SMC_33	15.0188	-72.5662	26.4	26.3	27.0	26.9	25.0
SMC_21	15.0284	-72.1569	25.9	26.3	26.9	27.0	25.0	24.5	23.8
SMC_10	15.0793	-72.1515	25.6	26.3	26.6	27.1	25.1	24.5	23.8
SMC_20	15.1156	-72.082	...	26.1	26.7	27.3	25.3	25.0	24.4
SMC_30	15.1462	-72.5944	26.4	26.3	27.0	27.1	25.0
SMC_31	15.1603	-72.5957	...	26.3	26.7	27.0	25.0
SMC_9	15.1806	-72.0226	28.1	25.7
SMC_16	15.2425	-72.0912	26.7	27.7	25.4	...	24.1
SMC_27	15.379	-72.4818	...	26.6	27.0	26.8	25.7	24.0	23.5
SMC_36	15.403	-72.269	...	25.9	25.1	27.1	23.5
SMC_2	15.5059	-72.5123	28.4	25.8
SMC_34	15.706	-72.7421	26.1	26.2	26.8	27.2	25.1	...	23.9
SMC_54	15.8228	-72.2482	28.0	25.5
SMC_18	15.9413	-72.6589	...	26.6	27.3	27.4	25.3	24.7	24.1
SMC_6	15.9954	-72.1305	26.7	27.7	25.4
SMC_8	16.0039	-72.1341	26.9	27.9	25.6
SMC_5	16.4523	-72.835	27.6	28.2	25.7
SMC_29	16.5543	-72.4672	28.0	25.6
SMC_51	16.5598	-72.1312	...	25.9	26.8	27.5	25.2
SMC_49	16.9656	-72.089	27.9	25.6
SMC_46	17.0188	-72.4006	27.6	25.3
SMC_55	17.3518	-72.0303	28.1	25.7
SMC_44	17.4098	-73.1726	...	26.2	27.2	27.3	25.0
SMC_53	17.5022	-73.1831	26.0	26.4	27.2	27.4	25.2
SMC_43	17.7638	-72.633	...	26.8	26.8	27.1	25.1	24.5	23.8
SMC_38	20.82	-73.3444	...	26.8	26.8	27.3	25.2	24.6	24.0
SMC_11	21.4361	-73.1212	26.7	28.2	25.7
SMC_48	21.5707	-73.3374	26.0	26.3	27.2	27.6	25.5
SMC_22	21.723	-73.127	28.6	26.0
SMC_47	21.9543	-73.2765	28.2	25.7
SMC_52	22.5766	-73.4592	28.3	25.9
SMC_28	22.9725	-73.4318	27.3	27.6	25.3
LMC_59	73.0619	-68.0239	27.7	28.1	25.6	...	24.4
LMC_38	73.9133	-69.1883	27.7	28.0	25.5
LMC_9	73.9739	-67.58	25.6	25.8	26.9	27.5	25.2	24.7	24.1
LMC_21	74.0085	-70.0449	...	25.7	26.9	27.8	25.5	...	24.4
LMC_1	74.0669	-66.3823	28.1	25.6
LMC_52	74.1457	-66.4882	27.2	24.9
LMC_44	74.2426	-66.4259	27.4	25.1
LMC_41	74.3489	-65.5843	27.9	28.4	26.0
LMC_8	74.3669	-66.3639	26.9	28.1	25.5
LMC_42	74.3789	-66.3868	27.7	28.1	25.7
LMC_16	74.4019	-67.735	27.4	28.3	25.8
LMC_24	74.4057	-68.4953	28.1	25.5
LMC_20	74.4149	-68.4953	28.1	25.6
LMC_32	74.7905	-70.1869	27.6	28.1	25.5











Table 8
(Continued)




Field Name	Avg R.A. (deg)	Avg Decl. (deg)	F225W (mag)	F275W (mag)	F336W (mag)	F475W (mag)	F814W (mag)	F110W (mag)	F160W (mag)
(1)	(2)	(3)	(4)	(5)	(6)	(7)	(8)	(9)	(10)
LMC_39	74.9718	-68.067	28.0	25.5
LMC_35	76.1349	-68.1583	27.8	25.4
LMC_43	76.2446	-70.5113	25.8	26.1	26.8	27.1	25.0
LMC_29	76.3845	-70.3046	...	26.1	26.9	27.3	25.2
LMC_33	76.5819	-71.2035	26.2	26.3	27.0	27.2	25.1
LMC_30	78.7752	-67.1761	28.1	25.6
LMC_37	80.3377	-65.7595	28.3	25.7
LMC_15	81.675	-67.7037	28.2	25.7
LMC_19	81.9348	-67.3754	28.3	25.7
LMC_45	82.1444	-66.9497	28.2	25.7
LMC_2	82.6984	-67.1787	28.0	25.7
LMC_3	82.8737	-67.3005	28.2	25.8
LMC_13	83.0047	-70.7874	26.5	27.6	25.3
LMC_6	83.06	-70.9753	26.5	27.8	25.5
LMC_4	83.7896	-69.8	25.6	25.9	26.5	27.1	24.9	24.5	24.0
LMC_57	83.793	-66.1738	27.7	25.4
LMC_34	84.0445	-66.6476	28.0	25.5
LMC_40	84.1042	-67.0506	27.7	28.1	25.7
LMC_27	84.3877	-69.1643	26.8	27.1	24.6	...	23.1
LMC_48	84.4053	-69.0844	26.0	26.3	26.8	26.9	24.9	...	23.8
LMC_47	84.4514	-69.0828	...	26.3	26.8	27.1	25.0
LMC_36	84.55	-69.4356	27.7	25.3
LMC_56	84.5993	-69.1631	25.3	25.1	26.8	26.9	24.8	24.2	23.7
LMC_53	84.6333	-69.2891	26.0	26.0	26.8	26.9	24.9	...	23.9
LMC_22	84.6425	-69.0671	...	26.0	26.3	26.0	23.7	23.1	22.6
LMC_11	84.7783	-69.1768	...	26.1	26.7	27.3	23.9	...	29.4
LMC_54	84.8145	-68.9016	27.6	27.8	25.4
LMC_55	84.9219	-69.0085	...	26.1	26.8	27.3	25.1	...	24.1
LMC_7	85.1607	-69.3493	26.9	27.8	25.5
LMC_5	85.1885	-69.5809	26.7	27.7	25.3
LMC_10	85.9116	-67.9299	28.3	25.8
LMC_17	86.1452	-67.325	27.5	28.3	25.8
LMC_49	86.3622	-67.0819	27.7	25.4
LMC_50	86.5299	-67.1893	28.1	25.7

Note. Summary of 50% completeness limits for each field. Column (1): name. Column (2): average catalog RA. Column (3): average catalog decl. Columns (4)–(10): 50% completeness limits in F225W, F275W, F336W, F475W, F814W, F110W, and F160W.

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