



HAL
open science

MINDS: Mid-infrared atomic and molecular hydrogen lines in the inner disk around a low-mass star

Riccardo Franceschi, Thomas Henning, Benoît Tabone, Giulia Perotti, Alessio Caratti O Garatti, Giulio Bettoni, Ewine F. van Dishoeck, Inga Kamp, Olivier Absil, Manuel Güdel, et al.

► To cite this version:

Riccardo Franceschi, Thomas Henning, Benoît Tabone, Giulia Perotti, Alessio Caratti O Garatti, et al.. MINDS: Mid-infrared atomic and molecular hydrogen lines in the inner disk around a low-mass star. *Astronomy & Astrophysics - A&A*, 2024, 687, <10.1051/0004-6361/202348034>. <insu-04843528>

HAL Id: insu-04843528

<https://insu.hal.science/insu-04843528v1>

Submitted on 17 Dec 2024

HAL is a multi-disciplinary open access archive for the deposit and dissemination of scientific research documents, whether they are published or not. The documents may come from teaching and research institutions in France or abroad, or from public or private research centers.

L'archive ouverte pluridisciplinaire HAL, est destinée au dépôt et à la diffusion de documents scientifiques de niveau recherche, publiés ou non, émanant des établissements d'enseignement et de recherche français ou étrangers, des laboratoires publics ou privés.



Distributed under a Creative Commons CC BY 4.0 - Attribution - International License

MINDS: Mid-infrared atomic and molecular hydrogen lines in the inner disk around a low-mass star

Riccardo Franceschi¹, Thomas Henning¹, Benoît Tabone², Giulia Perotti¹, Alessio Caratti o Garatti^{3,4}, Giulio Bettoni⁵, Ewine F. van Dishoeck^{6,5}, Inga Kamp⁷, Olivier Absil⁸, Manuel Güdel^{9,10}, Göran Olofsson¹¹, L. B. F. M. Waters^{12,13}, Aditya M. Arabhavi¹⁴, Valentin Christiaens⁷, Danny Gasman¹⁵, Sierra L. Grant⁵, Hyerin Jang¹², Donna Rodgers-Lee⁴, Matthias Samland¹, Kamber Schwarz¹, Milou Temmink⁶, David Barrado¹⁶, Anthony Boccaletti¹⁷, Vincent Geers¹⁸, Pierre-Olivier Lagage¹⁹, Eric Pantin¹⁹, Tom P. Ray⁴, Silvia Scheithauer¹, Bart Vandenbussche¹⁵, and Gillian Wright¹⁸

(Affiliations can be found after the references)

Received 21 September 2023 / Accepted 25 March 2024

ABSTRACT

Context. Understanding the physical conditions of circumstellar material around young stars is crucial to star and planet formation studies. In particular, very low-mass stars ($M_{\star} < 0.2 M_{\odot}$) are interesting sources to characterize as they are known to host a diverse population of rocky planets. Molecular and atomic hydrogen lines can probe the properties of the circumstellar gas.

Aims. This work aims to measure the mass accretion rate, the accretion luminosity, and more generally the physical conditions of the warm emitting gas in the inner disk of the very low-mass star 2MASS-J16053215-1933159. We investigate the source mid-infrared spectrum for atomic and molecular hydrogen line emission.

Methods. We present the full *James Webb* Space Telescope (JWST) Mid-Infrared Instrument (MIRI) Medium Resolution Spectrometer (MRS) spectrum of the protoplanetary disk around the very low-mass star 2MASS-J16053215-1933159 from the MINDS GTO program, previously shown to be abundant in hydrocarbon molecules. We analyzed the atomic and molecular hydrogen lines in this source by fitting one or multiple Gaussian profiles. We then built a rotational diagram for the H_2 lines to constrain the rotational temperature and column density of the gas. Finally, we compared the observed atomic line fluxes to predictions from two standard emission models.

Results. We identify five molecular hydrogen pure rotational lines and 16 atomic hydrogen recombination lines in the 5–20 μm spectral range. The spectrum indicates optically thin emission for both species. We use the molecular hydrogen lines to constrain the mass and temperature of the warm emitting gas. We derive a total gas mass of only $2.3 \times 10^{-5} M_{\text{Jup}}$ and a temperature of 635 K for the warm H_2 gas component located in the very inner disk ($r < 0.033$ au), which only accounts for a small fraction of the upper limit for the disk mass from continuum observations ($0.2 M_{\text{Jup}}$). The HI (7-6) recombination line is used to measure the mass accretion rate ($4.0 \times 10^{-10} M_{\odot} \text{ yr}^{-1}$) and luminosity ($3.1 \times 10^{-3} L_{\odot}$) onto the central source. This line falls close to the HI (11-8) line, however at the spectral resolution of JWST MIRI we managed to measure both separately. Previous studies based on *Spitzer* have measured the combined flux of both lines to measure accretion rates. HI recombination lines can also be used to derive the physical properties of the gas using atomic recombination models. The model predictions of the atomic line relative intensities constrain the atomic hydrogen density to about 10^9 – 10^{10} cm^{-3} and temperatures up to 5000 K.

Conclusions. The JWST-MIRI MRS observations for the very low-mass star 2MASS-J16053215-1933159 reveal a large number of emission lines, many originating from atomic and molecular hydrogen because we are able to look into the disk warm molecular layer. Their analysis constrains the physical properties of the emitting gas and showcases the potential of JWST to deepen our understanding of the physical and chemical structure of protoplanetary disks.

Key words. accretion, accretion disks – line: identification – radiation mechanisms: general – techniques: spectroscopic – protoplanetary disks – infrared: planetary systems

1. Introduction

While molecular hydrogen is the most abundant molecule in the universe, it is intrinsically hard to observe since it is a symmetric and homonuclear diatomic molecule (Thi et al. 2001; Carmona et al. 2011; Trapman et al. 2017). Dipole ro-vibrational transitions are forbidden and only weak electric-quadrupole transitions are allowed. Molecular hydrogen emission can only be observed in high-temperature or high-luminosity environments. Its emission has been detected in several classical T Tauri disks (Beck et al. 2008; Manara et al. 2021; Gangi et al. 2023), coming from a warm layer above the optically thick dust emission

(Carmona et al. 2011). Atomic hydrogen lines are also often detected and originate from accretion flows (Fang et al. 2009; Hartmann et al. 2016). Given the challenges in observing hydrogen emission, protoplanetary disks are usually characterized through other observational signatures. Dust emission is more easily observed (Andrews et al. 2018), and extensive modeling efforts can be found in the literature (see, e.g., Testi et al. 2014; Miotello et al. 2023 for reviews). The disk gas structure remains less characterized, but with the development of observational facilities, it is now possible to study the gas through the emission of molecular tracers. For instance, the Atacama Large Millimeter/submillimeter Array (ALMA) observations of optically thin

and optically thick CO isotopologues have been used to constrain the disk mass and temperature (Miotello et al. 2016, 2018; Zhang et al. 2021; Calahan et al. 2021; Schwarz et al. 2016, 2021; Pascucci et al. 2023). Yoshida et al. (2022) measured the pressure broadening of CO line emission and derived the gas surface density in the inner disk of TW Hya. Law et al. (2022) and Paneque-Carreño et al. (2023) used CO emission to study the disk vertical gas structure, while Lodato et al. (2023) measured the disk mass from the perturbations on the CO gas orbit induced by the disk self-gravity in the disks around IM Lup and GM Aur.

The HD molecule serves as a more direct indicator of H₂ since they are chemically very similar. The disk mass derived from HD measurements for TW Hya is higher than the CO-based estimate by about two orders of magnitudes, likely due to the depletion of CO from the gas phase (Bergin et al. 2013; McClure et al. 2016; Schwarz et al. 2016; Kama et al. 2020). However, HD emission primarily originates from warm regions within the disk (around $T \approx 30\text{--}50$ K; Bergin et al. 2013; Trapman et al. 2017), providing only lower limits to the gas mass. Additionally, HD has been detected only in TW Hya (Bergin et al. 2013), GM Aur, and DM Tau (McClure et al. 2016), and no new HD observations have been conducted since the conclusion of the *Herschel* mission.

Other emission lines are now commonly observed as well, each tracing different disk properties (Öberg et al. 2021). HCO⁺ and N₂H⁺ trace the gas ionization in the molecular layer, a fundamental parameter to disk chemistry (Cleeves et al. 2015; Teague et al. 2015; Aikawa et al. 2021) and can also be used to measure gas depletion (Smirnov-Pinchukov et al. 2020; Anderson et al. 2021; Trapman et al. 2022).

The InfraRed Spectrograph (IRS) on the *Spitzer* Space Telescope observed emission lines that can be used to characterize the gas surrounding young stellar objects (YSOs). For instance, a higher [Ne II] luminosity in the inner disk originates either from stellar jets or from photoevaporative disks driven by stellar high-energy photons (van Boekel et al. 2009; Lahuis et al. 2007; Pascucci et al. 2007; Güdel et al. 2010; Baldovin-Saavedra et al. 2011, 2012). Pascucci et al. (2006) estimated the upper limit of gas masses in the inner disk of a sample of young solar-like stars from the non-detection of mid-infrared H₂ lines. They found that none of their systems have a mass greater than 0.05 M_{Jup} within the first tens of au of the inner disk. Atomic hydrogen recombination lines are now commonly detected in the infrared spectra of YSOs and are used as an accretion tracer, as high HI luminosity is correlated with high accretion luminosities in the observed disks (e.g., Antonucci et al. 2011; Alcalá et al. 2014; Antonucci et al. 2016; Beuther et al. 2023). In particular, Rigliaco et al. (2015) detected the strongest of the Humphreys Series lines in the mid-infrared, the HI (7-6) line (12.37 μm), and derived an empirical relation between the HI (7-6) line luminosity and the accretion luminosity.

This paper presents observations of the disk surrounding the very low-mass star 2MASS-J16053215-1933159 (referred to as J160532) using the *James Webb* Space Telescope (JWST) Mid-Infrared Instrument (MIRI) Medium Resolution Spectrometer (MRS). These observations were obtained as part of the MIRI mid-Infrared Disk Survey (MINDS). A comprehensive analysis of the MIRI data is detailed in Tabone et al. (2023), with a primary focus on hydrocarbon chemistry. The study reveals strong emission lines of C₂H₂, C₄H₂, and C₆H₆ while detecting no H₂O. This suggests that the inner disk of J160532 exhibits an abundance of hydrocarbons, indicating more active carbon chemistry in comparison to disks around higher-mass stars and a notably high C/O ratio.

The MIRI spectrum of this source shows numerous atomic and molecular hydrogen emission lines that we investigate in this paper. We used the HI recombination lines to measure the source accretion rate and the atomic gas properties and the H₂ emission to constrain the gas rotational temperature and column density.

2. Observations

The target source, J160532, is an M dwarf star, the most common class of stars found in our Galaxy, and frequently hosting exoplanets (Dressing & Charbonneau 2015; Sabotta et al. 2021). Constraining the physical properties of disks around these stars provides important information about planet formation models. The star is located in the Upper Scorpius star-forming region, with a distance of 152 ± 1 pc (Gaia Collaboration 2023) and an age of 2.6 ± 1.6 Myr (Miret-Roig et al. 2022). The source is an M4.75 spectral type star, with mass $M_{\star} = 0.14 M_{\odot}$ and luminosity $L_{\star} = 0.04 L_{\odot}$ (Pascucci et al. 2013; Luhman et al. 2018; Carpenter et al. 2014). The spectral energy distribution (SED) of the source indicates the presence of a protoplanetary disk and the non-detection of millimeter continuum emission indicates a dust mass less than 0.75 M_{Earth} , corresponding to a total gas and dust mass of less than 0.2 M_{Jup} for a standard gas-to-dust ratio of 100 (Barenfeld et al. 2016).

The data analyzed in this paper are introduced in Tabone et al. (2023), where details about the observations, data reduction, and analysis are presented. Below we briefly summarize the main aspects. The data were observed by the JWST MIRI instrument (Rieke et al. 2015; Wright et al. 2015; Labiano et al. 2021; Wright et al. 2023) in Medium Resolution Spectroscopy (MRS) mode (Wells et al. 2015; Argyriou et al. 2023) on 2022 August 1, for a total exposure time of 2.22 h. These data were observed as part of the Cycle 1 Guaranteed Time Observation (GTO) program 1282, the MINDS program (PI: Thomas Henning). The details of the data reduction of the uncalibrated data can be found in Tabone et al. (2023). The spectral resolving power R ranges from 3500 to 1500 from the shortest to the longest wavelengths, equivalent to a velocity resolution of about 90 km s⁻¹ to 200 km s⁻¹. The spectrum has been corrected for the source small radial velocity -1.98 km s⁻¹ (Jönsson et al. 2020). The extinction for this source is particularly low, and no correction was needed for these data. To better analyze the spectral features, we performed a low-order fit of the continuum and of the two broad bumps at 7.7 and 13.7 μm , corresponding to optically thick C₂H₂ emission (see Tabone et al. 2023 for a detailed description of the procedure). The resulting spectrum is shown in Fig. 1.

3. Analysis and results

3.1. Molecular hydrogen

As reported in Tabone et al. (2023), the spectrum shows HI, H₂, and molecular (C₂H₂, C₄H₂, CO₂ among many others) features. A study of these molecular components of the spectrum is also included. Here we focused on the HI and H₂ emission lines, reported in Fig. 1. Thanks to the high resolving power and sensitivity of the MIRI-MRS instrument, we detected five pure rotational $\nu = 0 - 0$ transitions of H₂, while the typically fainter rotational transitions with higher vibrational states remain undetected (Fig. 2). The emitting region is unresolved, with no evidence of extended emission, meaning that it is coming from the inner, warm region of the disk. These lines are expected to be optically thin due to the small transition probabilities, and they

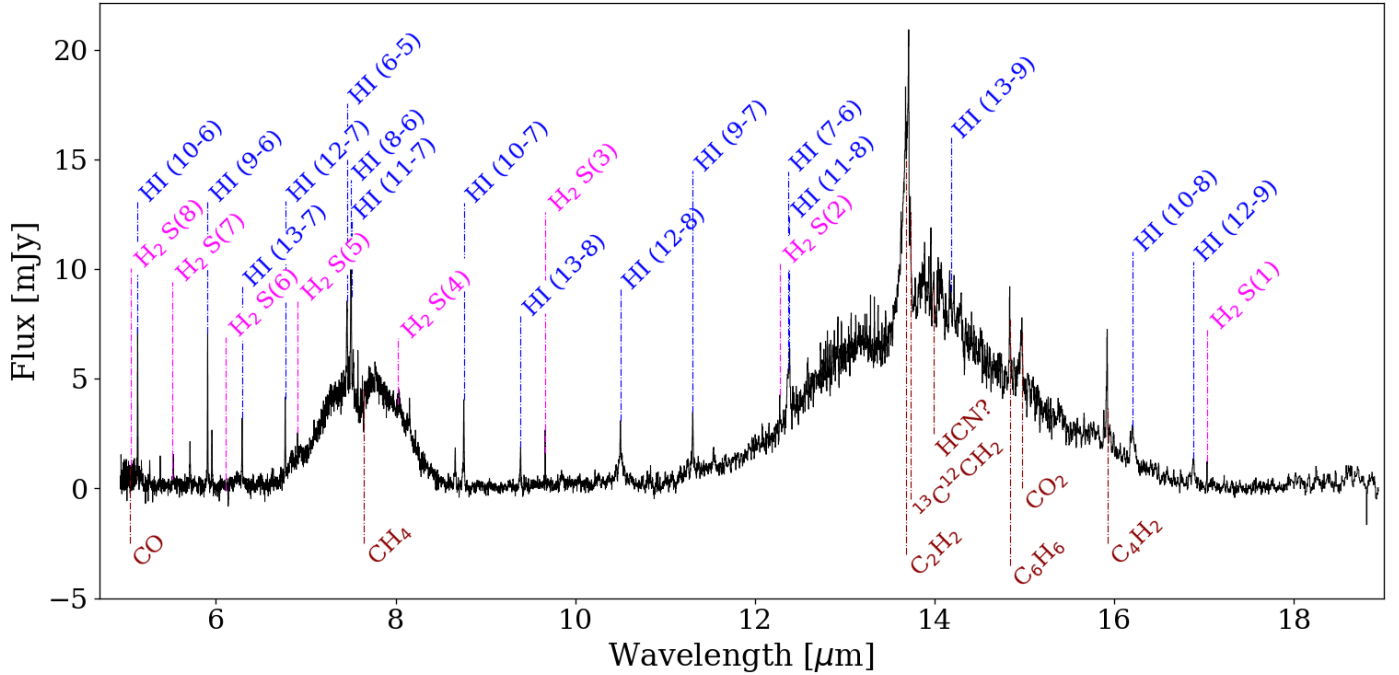


Fig. 1. JWST MIRI-MRS continuum-subtracted spectrum of J160532. We identify several molecular (pink) and atomic hydrogen (blue) transitions, highlighted in the figure. A study of other molecular species (dark red) can be found in Tabone et al. (2023).

can be used to constrain the temperature and column density of the emitting gas.

To accurately measure the line fluxes, we performed a linear fit of the continuum in a wavelength window of $0.03 \mu\text{m}$ centered on each line. The error in the flux density is the standard deviation on a nearby spectral segment that does not contain spectral features. The H_2 lines were fitted using one or multiple Gaussian profiles when the line showed evidence of blending with other spectral features to remove their contribution. This was done using the Levenberg-Marquardt least-squares minimization. The H_2 line flux was then calculated by integrating the main Gaussian component. In our calculations, the width of the Gaussian profiles was left as a free parameter. However, when working at the spectral resolution scale, one could also fix the width of the Gaussian profiles at the local spectral resolution element. We found that both methods give very similar results, within our error estimates, and we opted to leave the Gaussian width as a free parameter. The $\nu = 0 - 0 S(6)$, $S(7)$, and $S(8)$ lines are in a spectral region rich in other spectral features, and their profiles could not be extracted. In particular, the $S(8)$ line falls between 4.9 and $5.1 \mu\text{m}$, where several $\text{CO } \nu = 1 - 0$ P-branch lines are found, indicative of high-temperature gas ($T > 1000 \text{ K}$; Tabone et al. 2023). The $S(1)$ and $S(4)$ are partially blended to C_2H_2 emission lines, identified in a previous analysis (Tabone et al. 2023). However, the C_2H_2 emission contributes only a small amount to the total, and they produce results consistent with those from the other H_2 rotational lines. The line profiles and the best-fit profiles are shown in Fig. 2, and the line parameters are reported in Table 1.

We then used the measured fluxes to build a rotational diagram (Goldsmith & Langer 1999) to estimate the emitting gas temperature and column density, assuming that the H_2 emission is in local thermodynamical equilibrium (LTE). This assumption is justified by the low critical density of H_2 emission, which is the typical minimal density required to populate the energy levels by collisions bringing them to LTE. The critical densities

Table 1. Measured parameters of the identified pure rotational molecular hydrogen lines.

Transition	Line center [μm]	Integrated flux $10^{-13} [\text{erg s}^{-1} \text{cm}^{-2}]$
$S(1)$	17.035	0.9 ± 0.2
$S(2)$	12.279	2.4 ± 0.4
$S(3)$	9.665	2.8 ± 0.2
$S(4)$	8.025	3.1 ± 0.9
$S(5)$	6.910	2.2 ± 0.7

for H_2 rotational levels are orders of magnitude lower than for CO , the next most abundant molecule, and lower than the typical densities in the warm molecular layers. The difference in critical densities is due to the lower Einstein coefficient probability of rotational H_2 transitions compared to CO rotational transitions. The low critical density for H_2 rotational transitions, therefore, ensures LTE emission.

Assuming a Boltzmann distribution, for optically thin emission the observed fluxes F_J must satisfy the relations:

$$\frac{4\pi F_J}{h\nu_J \Omega g_J A_J} = \frac{N_{\text{tot}}}{Q_{\text{rot}}} e^{-E_J/k_B T_{\text{rot}}} \quad (1)$$

where h is the Planck constant, k_B the Boltzmann constant, ν_J is the line frequency, $\Omega = \pi(r_{\text{em}}/d)^2$ is the emitting area solid angle, A_J the Einstein coefficient, $g_J = (2I + 1) \times (2J + 1)$ is the statistical weight of the transition, with I the total nuclear spin, N_{tot} is the molecule column density, $Q(T_{\text{rot}})$ is the partition function, E_J is the transition energy, and T_{rot} is the rotation temperature. The ortho-to-para ratio is assumed to be 3 in LTE. Assuming optically thin emission at LTE with temperature T_{rot} , we estimated the column density and temperature of the gas:

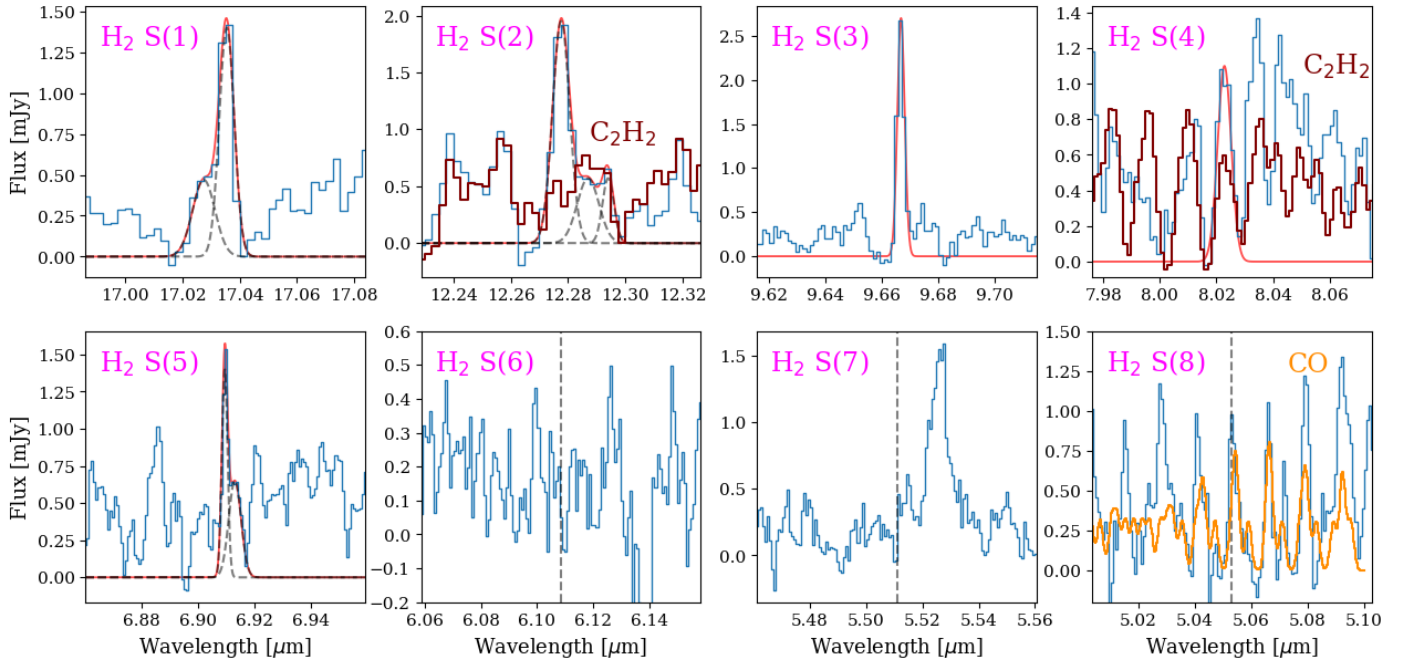


Fig. 2. Molecular hydrogen pure rotational lines within the spectral range. The red profiles show line fitting using one or multiple Gaussian functions (dashed lines) in the case of line blending with other spectral features. The rest wavelength of the undetected lines is indicated by a slashed vertical line. In dark red we show the C₂H₂ rescaled model from Tabone et al. (2023) to show the features location. In orange, we show an eye-fit of a CO slab model. The S(2) line is blended with a C₂H₂ line. The S(6) and S(7) transitions fall in a crowded spectral region, and could not be identified. The S(8) transition is found in a region populated by several CO $v = 1 - 0$ P-branch lines, and cannot be identified either.

$$\ln\left(\frac{4\pi F_J}{h\nu_J g_J A_J}\right) = -\frac{E_J}{k_B T_{\text{rot}}} - \ln\left(\frac{Q_{\text{rot}}}{N_{\text{tot}} \Omega}\right). \quad (2)$$

This expression allows the construction of a rotational diagram. In this graphical representation, the variables are $x = E_J/k_B$ and $y = \log(N_J/g_J)$, where $N_J = 4\pi F_J/h\nu_J A_J$. The second term on the right-hand side is treated as a constant. When plotted, this diagram is expected to produce a straight line with a slope of $-1/T_{\text{rot}}$ and an intercept that provides the mass of the emitting gas, under the assumption of a specific emission area.

The resulting rotation diagram is shown in Fig. 3. This figure has been produced using the tool `pdrtpy`¹, introduced in Berné et al. (2022); Pound & Wolfire (2023). The emitting area can be estimated from an optically thick line. We assumed the H₂ emission to have the same emitting radius r_{em} as the C₂H₂ emission, 0.033 au (Tabone et al. 2023). Indeed, the emitting radius of the H₂ emission cannot be directly estimated due to its optically thin nature, which leads to a degeneracy between column density and temperature of unresolved H₂ emission. The C₂H₂ emission, on the other hand, is optically thick and independent from column density, allowing us to estimate its emitting radius. Since the emission of C₂H₂ and H₂ are excited by similar physical conditions, it is reasonable to assume that the C₂H₂ and H₂ emission originate from the same disk region, and have the same emitting radius. Using this estimate for the emitting radius, we found a temperature of 635 ± 94 K and a column density of 190 ± 110 g cm⁻², corresponding to a total mass of warm H₂ gas of $(2.3 \pm 1.3) \times 10^{-5} M_{\text{Jup}}$.

We tested the consistency of these results by running a slab model for the H₂ emission. In this model, the line emission is

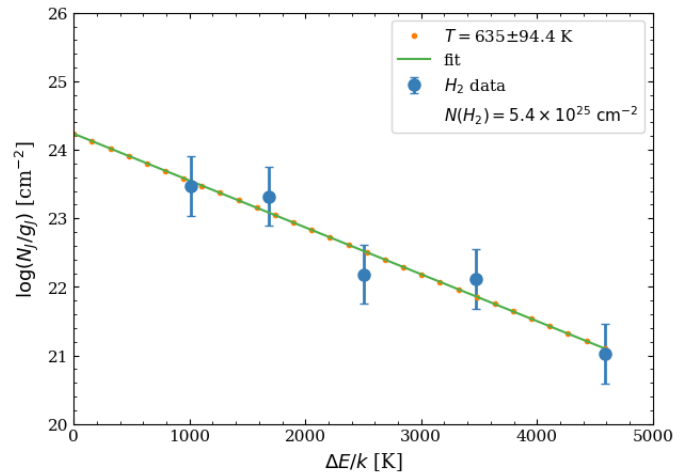


Fig. 3. Rotation diagram of H₂ rotational transitions. A linear fit was performed on the measured fluxes, and the diagram shows a single component for the emitting gas with a temperature of about 635 K and mass $2.5 \times 10^{-5} M_{\text{Jup}}$.

assumed to originate from a gas layer of temperature T , column density of N , and emitting radius r , similarly as in Tabone et al. (2023). Using our temperature and mass estimates and our assumed emitting radius we obtained results consistent with the observed line fluxes. Moreover, the predicted fluxes for the S(6), S(7), and S(8) lines fall below the noise level of our observations, consistent with their non-detection in the data.

The linear trend observed in the rotational diagram is an empirical argument justifying the assumption of LTE emission. In instances where a single excitation temperature is unable to describe the data, either a temperature gradient or non-LTE processes should be taken into account. In the case of non-LTE

¹ <https://pdrtpy.readthedocs.io/en/latest/>

effects, they often lead to a sub-thermal population of lines, implying that our mass estimate would serve as a lower limit, and the actual column density could be higher. However, taking into consideration the previous discussion on the low critical density for rotational H₂ emission, lower than the typical densities in warm molecular layers, the effect of a temperature gradient in the emitting gas would be a more likely explanation for a non-linear trend in the rotational diagram.

The observed H₂ emission traces only the innermost region of the disk and its mass represents only a small fraction of the total disk mass. The disk midplane, even at such small radii, could have temperatures lower than our estimate. Therefore, even in the inner disk, warm H₂ is a fraction of the total H₂ reservoir, and our mass estimate for the warm H₂ component is a lower limit to the total H₂ mass within the emitting radius. From the non-detection of millimeter continuum emission, [Barenfeld et al. \(2016\)](#) give an upper limit to the disk mass of 0.2 M_{Jup} . By making a few assumptions, we can give a rough estimate of the disk mass based on our measured warm H₂ mass. We assumed a small disk size $r_0 = 10$ au, and a power-law distribution for the disk surface density distribution:

$$\Sigma(r) = \Sigma_0 \left(\frac{r}{r_0} \right)^{-1}. \quad (3)$$

By knowing the gas surface density at $r = 0.033$ au from our H₂ emission analysis, we can use this equation to get a total disk mass estimate of 0.05 M_{Jup} , which is in agreement with the upper limit of 0.2 M_{Jup} from [Barenfeld et al. \(2016\)](#).

3.2. Atomic hydrogen

Atomic hydrogen recombination lines trace stellar accretion and ejection processes, as demonstrated by their broad profiles consistent with almost free-falling material (e.g., [Alcalá et al. 2014](#)). These lines are now commonly used to measure mass accretion rates in YSOs ([Antoniucci et al. 2011, 2016, 2014](#); [Biazzo et al. 2012](#); [GRAVITY Collaboration 2023](#)), although they can also originate from winds and jets ([Beck et al. 2010](#); [Bally 2016](#); [Ercolano & Pascucci 2017](#)). While it is necessary to account for both accretion and wind processes to reproduce all the spectral features, a tight empirical correlation has been measured between the integrated flux of HI recombination lines and the accretion luminosities L_{acc} . This correlation has been extensively studied for optical and near-infrared lines, such as the H α , Br γ , Pa β ([Antoniucci et al. 2014](#); [Alcalá et al. 2014, 2017](#)), and Pf β (at 4.6 μm) lines ([Salyk et al. 2013](#)). However, for the embedded phases of star formation or sources significantly affected by extinction, we must rely on longer wavelengths. Only a few space-based studies are available in the literature, based on data obtained with the Infrared Space Observatory and *Spitzer* space telescope (e.g., [van Dishoeck et al. 1998](#); [Gibb et al. 2000](#); [Molinari et al. 2008](#); [An et al. 2011](#)). In particular, [Rigliaco et al. \(2015\)](#) explores the relations between the HI (7-6) line luminosity at 12.37 μm and the accretion luminosity using *Spitzer* data on a sample of classical T Tauri stars. With the MIRI instrument, it is now possible to study in detail the mid-infrared atomic hydrogen recombination lines.

We identified 16 atomic hydrogen recombination lines in the spectrum of J160532, with the principal quantum number of the lower energy levels between 5 and 8 (Pfund, Humphreys, and higher series). To estimate the line fluxes, we followed the same procedure used to analyze H₂ in Sect. 3.1, and the lines are shown

Table 2. Measured parameters of the identified atomic hydrogen recombination lines.

Transition	Line center [μm]	Integrated flux 10^{-16} [erg s ⁻¹ cm ⁻²]
6–5	7.460	25.9 ± 5.8
7–6	12.372	9.4 ± 1.0
8–6	7.503	20.9 ± 1.0
9–6	5.908	37.7 ± 1.3
10–6	5.129	44.6 ± 2.2
9–7	11.309	7.1 ± 1.0
10–7	8.760	15.7 ± 1.0
11–7	7.508	32.0 ± 5.3
12–7	6.772	19.4 ± 5.3
13–7	6.292	14.9 ± 1.9
14–7	5.957	17.1 ± 1.9
10–8	16.209	4.5 ± 1.3
11–8	12.387	9.7 ± 1.1
12–8	10.504	7.1 ± 0.9
13–8	9.392	7.7 ± 1.0
14–8	8.665	5.3 ± 1.0

in Fig. 4. The main Gaussian component of the line fits has a width ranging from about 100–200 km s⁻¹, compatible with the local spectral resolution. Our results are summarized in Table 2. The line fluxes are shown in an excitation diagram in Fig. 5. This diagram highlights that the intensities of the HI lines deviate from what would be expected in optically thin emission conditions. Moreover, lines with the same upper energy level do not overlap in the plot, contrary to what would be expected in the scenario of optically thin emission.

We measured the mass accretion rate from the HI (7-6) line at 12.37 μm using the method introduced by [Rigliaco et al. \(2015\)](#), who derives the following empirical relation between the HI (7-6) integrated flux and the accretion luminosity in classical T Tauri stars:

$$\log L_{\text{HI (7-6)}}/L_{\odot} = (0.48 \pm 0.09) \times \log L_{\text{acc}}/L_{\odot} - (4.68 \pm 0.10). \quad (4)$$

The standard relation between the mass accretion rate \dot{M}_{acc} and L_{acc} was introduced by [Gullbring et al. \(1998\)](#):

$$\dot{M}_{\text{acc}} = \frac{L_{\text{acc}} R_{\star}}{GM_{\star}} \left(1 - \frac{R_{\star}}{R_{\text{in}}} \right)^{-1}, \quad (5)$$

where $M_{\star} = 0.14 M_{\odot}$ is the stellar mass, R_{in} the radius at which the stellar magnetosphere disperses the accreting gas, and $R_{\star} = 0.45 R_{\odot}$ is the stellar radius. R_{in} is usually assumed to be $5 R_{\star}$ ([Herczeg & Hillenbrand 2008](#); [Alcalá et al. 2014](#); [Rigliaco et al. 2015](#)), while R_{\star} is derived from the stellar effective temperature $T_{\text{eff}} = 3850$ K and the stellar luminosity L_{\star} using the Stefan–Boltzmann relation $L_{\star} = 4\pi R_{\star}^2 \sigma T_{\text{eff}}^4$, where σ is the Stefan–Boltzmann constant and T_{eff} is estimated from the spectral type M4.75 ([Fang et al. 2017](#)).

Notably, the HI (7-6) line at 12.372 μm is located near another atomic hydrogen recombination line, the HI (11-8) transition at 12.387 μm . This is evident from the observed line profile (see Fig. 6), which displays two dominant Gaussian components, each approximately 150 km s⁻¹ wide. The HI (7-6) line

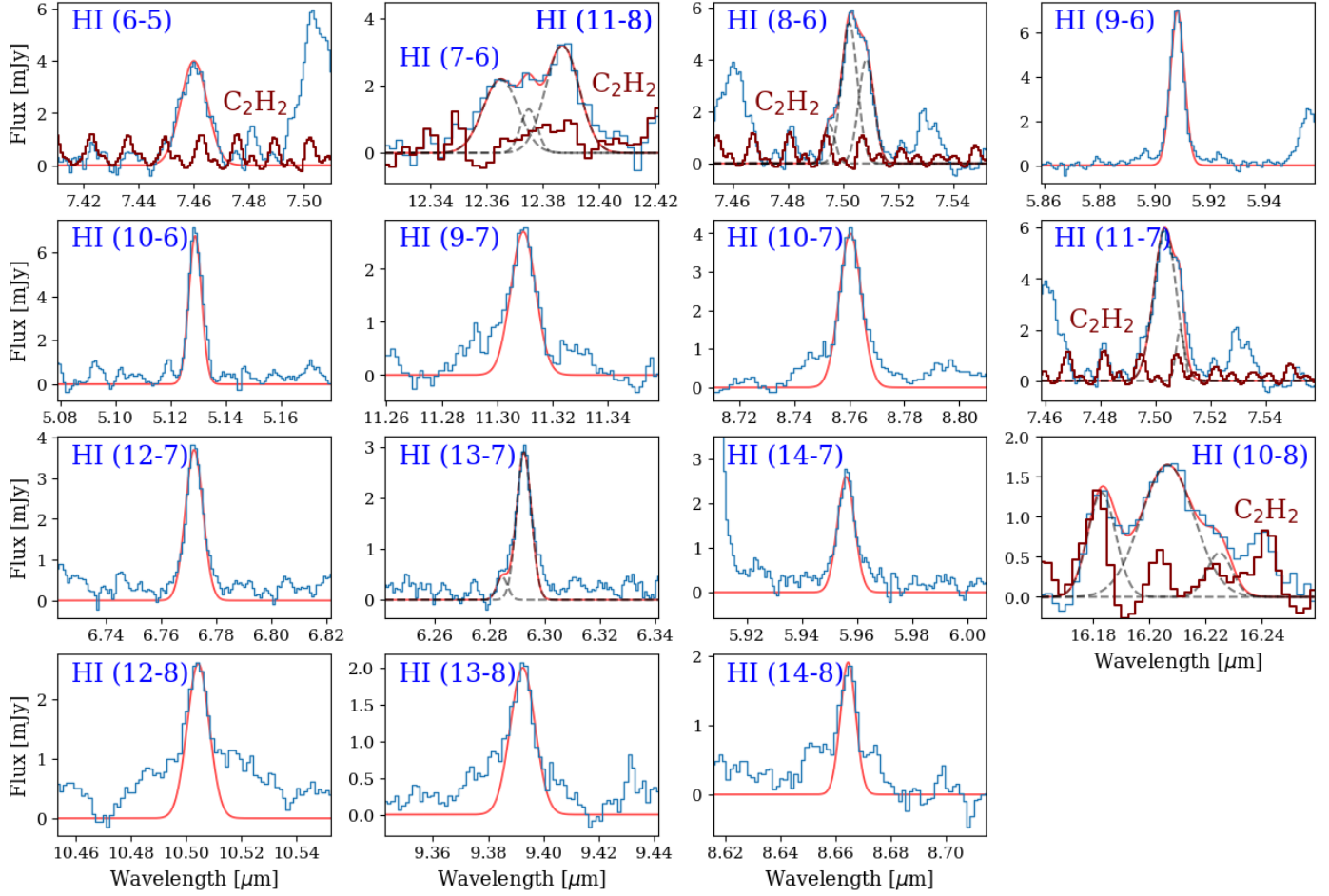


Fig. 4. Detected atomic hydrogen recombination lines. The red profiles show line fitting using one or multiple Gaussian functions (dashed lines) in the case of line blending with other spectral features. The dark red profile is the C_2H_2 model from Tabone et al. (2023).

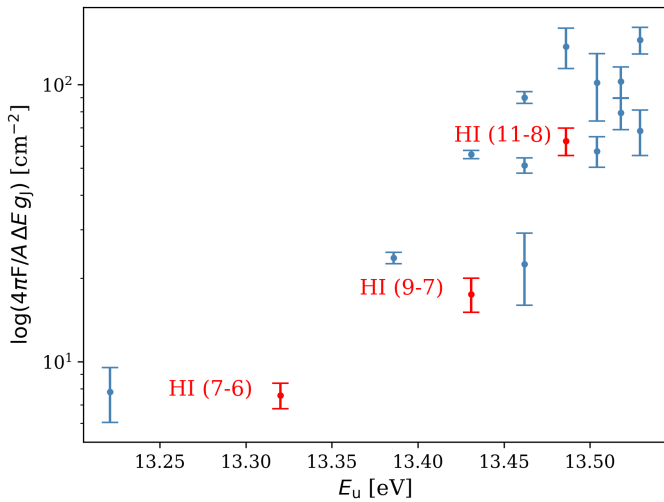


Fig. 5. Excitation diagram of the HI lines. These lines are produced by shocked material for which the Boltzmann distribution does not apply, therefore the y-axis is in terms of integrated line flux rather than column density. In red, we highlight the lines used for the comparison to atomic line emission models.

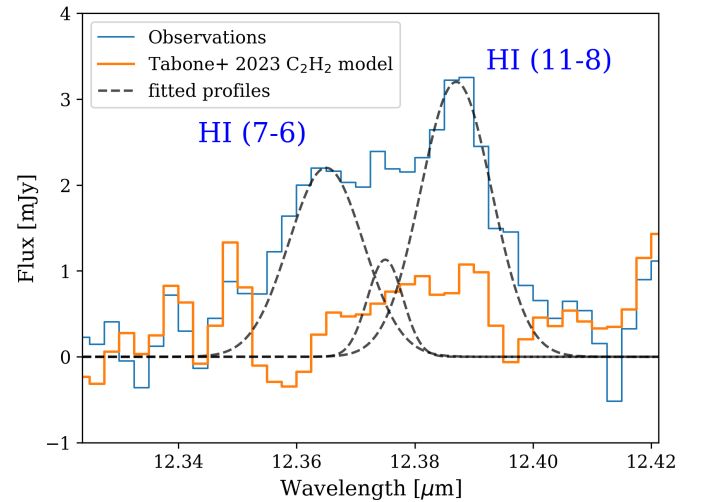


Fig. 6. Emission lines used to measure the accretion rates. The black dashed profile on the left is the HI (7-6) line, the one on the right is the HI (11-8) line, while the central component is C_2H_2 emission, according to the predictions of Tabone et al. (2023, the orange profile).

is shifted by about -4 km s^{-1} compared to its rest frequency, while the HI (11-8) by about 64 km s^{-1} , which is compatible with the line rest frequency given the local resolution of

about 110 km s^{-1} . Additionally a third, weaker component is present, with a width similar to the local resolution element, around $\sim 60 \text{ km s}^{-1}$. Comparatively, the widths of the two main

components align with those of the other atomic hydrogen lines. We identified the component at the shorter wavelength as the HI (7-6) line, while the component at the longer wavelength corresponds to the HI (11-8) line. The third component could be the Q -branch of a weak P , Q , and R -branch pattern of optically thick C_2H_2 emission modeled by Tabone et al. (2023), centered at $12.379 \mu\text{m}$ (Fig. 6). The two identified components have peak flux densities of 2.3 mJy and 3.5 mJy, and luminosities of $4.9 \times 10^{-7} L_\odot$ and $6.8 \times 10^{-7} L_\odot$ respectively. From the HI (7-6) luminosity we get an accretion luminosity of $4.0 \pm 2.5 \times 10^{-4} L_\odot$ using Eq. (4), corresponding to a mass accretion rate of $1.0 \times 10^{-10} M_\odot \text{ yr}^{-1}$. However, this relation was calibrated by Rigliaco et al. (2015) using *Spitzer* data, which lacked the spectral resolution necessary to distinguish between the HI(7-6) and HI(11-8) profiles. Since both lines are excited at similar energies, the excitation of one line implies the excitation of the other. Therefore, we argue that the fluxes used by Rigliaco et al. (2015) to measure the accretion luminosity are not solely the HI (7-6) line but rather the combined fluxes of both the HI (7-6) and HI (11-8) lines. The contribution to the total flux given by the C_2H_2 component is negligible and does not affect the measured accretion luminosity. By considering the total line flux, we obtained an accretion luminosity of $(3.1 \pm 1.9) \times 10^{-3} L_\odot$, corresponding to a mass accretion rate of $(4.0 \pm 2.5) \times 10^{-10} M_\odot \text{ yr}^{-1}$. This result is in good agreement with the mass accretion rate estimate of $4.17 \times 10^{-10} M_\odot \text{ yr}^{-1}$ reported by Fang et al. (2023) based on the $H\alpha$ emission. A follow-up study utilizing the higher spectral resolution data from the JWST on a sample of sources could validate the hypothesis that the accretion rate measured using Eq. (4) uses the combined HI (7-6) and HI(11-8) fluxes.

Finally, Antonucci et al. (2016) study the Balmer and Paschen decrements in a sample of 36 low-mass Class II sources. While they do not find any correlation between the line decrements and the source properties, they find a tentative correlation between the decrement shape and the mass accretion rate. If we assume their results to be valid also for other HI series (a reasonable assumption since they originate from the same physical process), the decrement shape in our source (Fig. 7) falls in their Type 3, bumpy profile category. The disks in this category have mass accretion rates between 10^{-10} and $10^{-9} M_\odot \text{ yr}^{-1}$, which is in accord with our measured value of $(4.0 \pm 2.5) \times 10^{-10} M_\odot \text{ yr}^{-1}$.

3.3. Atomic line emission models

To constrain the physical properties of the gas from which the HI lines originate, we compared the observed line fluxes to predictions from two standard emission models. The first model employed was the classic Case B recombination line model (Baker & Menzel 1938; Storey & Hummer 1995), which assumes that Lyman lines are optically thick and all other transitions are optically thin, and that level populations are determined through radiative cascade from the continuum. The second model utilized is the more recent local line excitation approach developed by Kwan & Fischer (2011), which considers the local physical conditions of the gas to consistently evaluate the line emissivities. The input physical parameters of this model are the gas temperature, the hydrogen density, the ionization rate (which was not specified in the Case B models), and the gas velocity gradient transverse to the radial direction. These parameters are chosen to reflect the typical physical condition in accreting YSOs.

The Case B recombination model is commonly used in the literature to derive the physical properties of accreting gas onto T Tauri stars. It assumes that the gas is optically thick to Lyman

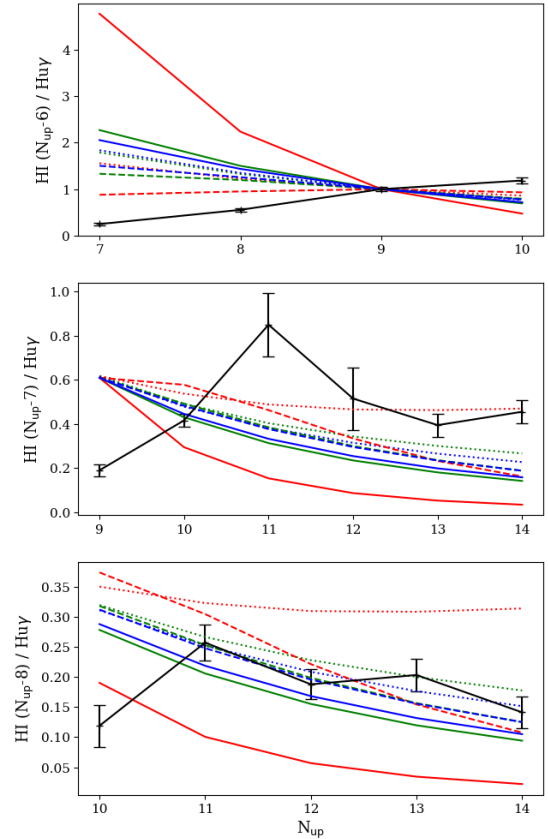


Fig. 7. Comparison of the three detected line series to the Case B model predictions. The solid black lines show the observed line fluxes for transition with lower energy levels (from top to bottom) 6, 7, and 8 relative to the reference line HI (9-6). The lines with different colors are the predicted values by the Case B model for different densities and temperatures of the emitting gas. The model densities are 10^8 cm^{-3} (dotted line), 10^{10} cm^{-3} (slashed line), and 10^{12} cm^{-3} . The temperatures are 500 K (red), 5000 K (green), and 10000 K (blue).

series photons and optically thin to the photons linked to other transitions. Storey & Hummer (1995) computed the line fluxes for hydrogenic atoms as a function of temperature and electron density for the Case B recombination model, which are available online². Using this model, we compare in Fig. 8 the HI (9-7) and HI (7-6) lines for a temperature range of 500–30 000 K and electron density n_e between 10^8 cm^{-3} and 10^{12} cm^{-3} to our measured value of $\sim 0.76 \pm 0.35$. According to this model, the gas temperature varies widely and it is poorly constrained, ranging from 50 to 12 500 K, electron density between 10^8 cm^{-3} and 10^{11} cm^{-3} . The highest temperatures are observed at column densities of $\sim 10^{10} \text{ cm}^{-3}$.

While we detected numerous HI lines, our analysis centers on the HI (9-7), (7-6), and (11-8) transitions. The physical properties of the emitting gas can be probed using any line of choice. These bright lines are favored due to observational limitations, enabling us to effectively constrain gas properties using only a subset of observed lines. Although a broader selection of lines would provide tighter constraints, our options are constrained by available atomic line emission models and our spectral window, limiting us to the HI (9-7), (7-6), and (11-8) transitions that have been previously modeled (Storey & Hummer 1995;

² <http://cdsarc.u-strasbg.fr/viz-bin/Cat?VI/64>

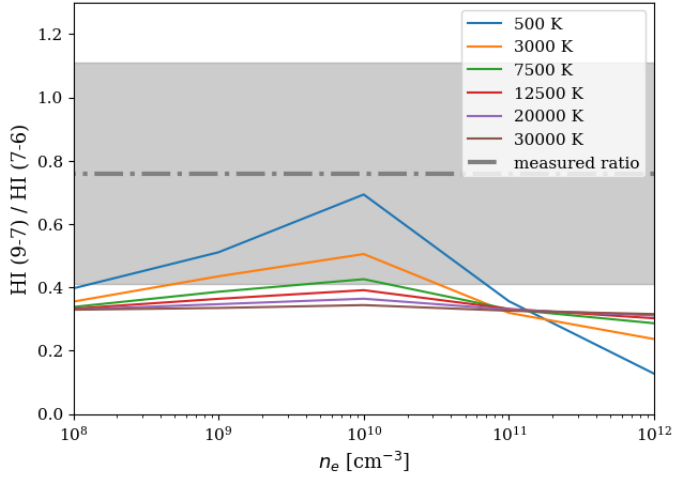


Fig. 8. HI (9-7) and HI (7-6) line ratio as a function of electron density from the Case B recombination model compared to our measured ratio (gray dotted line). The shaded area is the uncertainty of the measured ratio.

Kwan & Fischer 2011). Our study highlights how the JWST-MIRI instrument expands the range of observable lines, leading to a more comprehensive understanding of protoplanetary disk properties through improved modeling efforts with this new data.

In Fig. 7 we show the line ratios as a function of the upper level quantum number N_{up} , for different line series, with lower energy levels 6, 7, and 8. In computing the ratio, we used as reference the HI (9-6) (H γ) line, the cleanest line in our sample. However, choosing a different line does not significantly affect the shape of the line ratio distribution. In Fig. 7 we also show the predicted values for the Case B model for a range of temperatures from 500 to 10 000 K and densities from 10^8 to 10^{12} cm^{-3} . The Case B line ratios poorly reproduce the observations, especially for the lower lines in the series. This plot highlights the effect of the optical depth on the line fluxes. In the case of optically thin emission, we expect the ratio of lines with the same upper level to be the ratio of the Einstein coefficients of the transitions. However, our observed lines contradict this expectation. For instance, the ratio of the HI (9-7) and HI (9-6) lines is about 5.3. This value is much higher than the Einstein coefficient ratio of about 0.9. This discrepancy suggests that these lines cannot be considered optically thin, as is assumed in the Case B recombination theory. This model therefore is not an appropriate description of these data.

The HI (9-7) and HI (7-6) line ratio was also computed following a different approach by Kwan & Fischer (2011), hereafter denoted KF model, where they measure the optical depth of each transition instead of assuming optically thin emission. They also included in their calculation the gas ionization rate and the velocity gradient, not accounted for in the Case B recombination model and tuned these parameters to reflect the typical physical condition of accreting YSOs. Using the data available online³ we compared again the model ratio of the HI (9-7) and HI (7-6) line to our measurement, for temperatures between 54 000 K and 20 000 K and atomic hydrogen densities between 10^8 cm^{-3} and 10^{12} cm^{-3} (see Fig. 9). The KF model predicts that all the temperatures are compatible with our measured line ratios, and constrains the hydrogen density to 7×10^9 – 4×10^{10} cm^{-3} .

³ <https://iopscience.iop.org/article/10.1088/0004-637X/778/2/148>

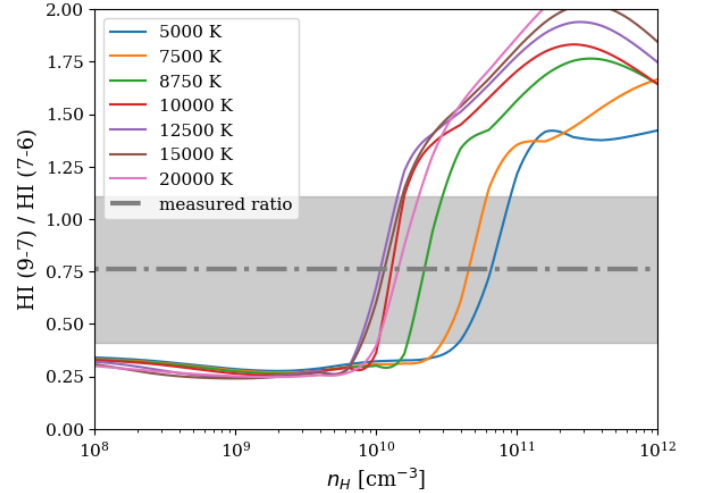


Fig. 9. HI (9-7) and HI (7-6) line ratios as a function of atomic hydrogen density according to the KF atomic emission model (Kwan & Fischer 2011) compared to our measured ratio, the gray dotted line. The shaded area is the uncertainty of the measured ratio.

Compared to the Case B model, a proper treatment of the lines optical depth gives us much tighter constraints. Since the ionization degree of n_e/n_H typically ranges between 0.1 and 1.0, this density range is in agreement with the constraints from the Case B model. Case B models with temperatures up to 7500 K and atomic hydrogen densities of about 10^{10} – 10^{11} cm^{-3} (with $n_e/n_H = 0.1$) are in agreement with the low temperature (5000 K), high density (a few 10^{10} cm^{-3}) KF models, while densities below 10^{10} cm^{-3} cannot be reached given our measured line ratio. However, we could not give an upper limit to the atomic hydrogen densities according to the KF models, as the line ratio was not calculated for temperatures lower than 5000 K. When looking at the model prediction of the other HI line fluxes of the Case B model, the fit to the data is quite poor. The observed line fluxes require temperatures lower than 500 K. These low temperatures are unlikely for shocked accreting HI gas. This indicates that these lines may not be optically thin, as assumed in the Case B model (van den Ancker et al. 2000; Edwards et al. 2013; Rigliaco et al. 2015), and this model cannot be used to model the emission of high-density regions such as accretion flows in YSOs.

4. Discussion

H₂ emission in classical T Tauri stars can originate from either the inner region of a protoplanetary disk, excited by UV radiation, or shocked gas in stellar outflows (Beck et al. 2010; Carmona et al. 2011). We analyzed the profiles of the detected H₂ lines to constrain the origin and formation mechanism of these lines. As discussed in Sect. 3.1, the H₂ emission comes from a spatially unresolved component with no evidence of outflows. The emission is observed at a velocity consistent with the rest velocity of the star. However, the data are also compatible with a stellar wind or outflow coming from the very inner disk with a projected velocity $\lesssim 90$ km s^{-1} . To further constrain the origin of these lines, we measured the full width at zero intensity (FWZI) of the profiles, giving us the maximum projected velocity of the emitting H₂, and compared this value to the Keplerian velocity at the emitting radius 0.033 au. We measured a mean FWZI of 160 ± 110 km s^{-1} and given that the Keplerian velocity at the

emitting radius is 60 km s^{-1} , this measurement is compatible with Keplerian motions of the gas within the assumed emitting radius. However, it is important to note that we lack information about the disk geometry, and the projected Keplerian velocity might be below 60 km s^{-1} , depending on the disk inclination.

We performed the same analysis of the FWZI of the HI line profiles to investigate their origin. As discussed in Sect. 3.2, we assumed that HI emission traces accretion flows of free-falling gas on the stellar surface. Following the assumption that the gas is accreting from a radius $R_{\text{in}} = 5 R_{\star}$ (see Sect. 3.2), we got a free-fall velocity of 315 km s^{-1} . The average FWZI for HI lines is $370 \pm 140 \text{ km s}^{-1}$, supporting the hypothesis that this emission comes from free-falling gas accreting on the central object.

The accretion rate we derived using the atomic hydrogen lines is consistent with those measured in other YSOs. Alcalá et al. (2017) measured the accretion rate of YSOs in the Lupus region. Their measured accretion rates range from $\sim 5 \times 10^{-12} M_{\odot} \text{ yr}^{-1}$ to $\sim 6 \times 10^{-8} M_{\odot} \text{ yr}^{-1}$, and our measured mass accretion rate of $\sim 4 \times 10^{-10} M_{\odot} \text{ yr}^{-1}$ correspond to an average accretion for the YSOs in their sample. Alcalá et al. (2017) also derive an analytical relation between $L_{\text{acc}} - L_{\star}$ and one between $\dot{M}_{\text{acc}} - M_{\star}$. Using their relation for L_{acc} :

$$\log L_{\text{acc}} = (1.26 \pm 0.14) \cdot \log L_{\star} - (1.60 \pm 0.13), \quad (6)$$

we found $L_{\text{acc}} = (4.4 \pm 2.4) \times 10^{-4} L_{\odot}$. This measurement is about one order of magnitude lower than our measured value $(3.1 \pm 1.9) \times 10^{-3} L_{\odot}$. There is, however, a significant scatter in the data used to derive Eq. (6), and our measurement falls within this scatter. Using their relation for \dot{M}_{acc} for low-mass stars ($< 0.2 M_{\odot}$):

$$\log \dot{M}_{\text{acc}} = (4.58 \pm 0.68) \cdot \log M_{\star} - (6.11 \pm 0.61), \quad (7)$$

we derived $\dot{M}_{\text{acc}} = (1.0 \pm 1.8) \times 10^{-10} M_{\odot} \text{ yr}^{-1}$, in agreement with our measured value of $(4.0 \pm 2.5) \times 10^{-10}$.

Numerous spectroscopic surveys in several star formation regions have observed a correlation between the accretion rate \dot{M}_{acc} and the stellar mass M_{\star} (Alcalá et al. 2014; Venuti et al. 2014, 2019; Manara et al. 2016, 2017; Hartmann et al. 2016), with reported slopes ranging from approximately 1.6 to 2. Manara et al. (2023) compare these data and presented their findings in their Fig. 4. Their analysis demonstrates that the observations align well with the relationship $\dot{M}_{\text{acc}} \propto M_{\star}^2$. The accretion rate for this particular source adheres to the $\dot{M}_{\text{acc}} \propto M_{\star}^2$ relationship, classifying it as a standard accretor compared to other YSOs. It is important to highlight that our determined accretion rate is not exceptionally high, despite the prominent presence of HI lines in the spectrum. In fact, it falls within the lower range when compared to similar sources with measured accretion rates (e.g., Rigliaco et al. 2015; Manara et al. 2016; Alcalá et al. 2017). This clear detection of HI lines is likely attributed to the remarkably low dust extinction, enabling us to penetrate deeply into the disk and obtain an unobstructed view of the accretion layer.

5. Conclusions

In this paper, we present the full JWST-MIRI spectrum of the disk around the very low-mass star 2MASS-J16053215-1933159. These observations reveal numerous atomic and molecular hydrogen lines in the mid-infrared spectrum of this source. We utilize HI recombination lines to measure the accretion

rate of the source and H_2 emission to constrain the rotational temperature and column density of the disk gas.

With the high resolving power and sensitivity of the MIRI-MRS instrument, we detect several pure rotational $\nu = 0 - 0$ H_2 transitions. The observed lines are optically thin due to their small transition probabilities, allowing us to constrain the temperature and column density of the gas. To accurately measure the line fluxes, we performed a linear fit of the continuum and fit Gaussian profiles to the H_2 line profiles, while accounting for blending with other spectral features. Using the measured fluxes, we constructed a rotational diagram to estimate the temperature and column density of the emitting gas. Under the assumption of optically thin emission in local thermodynamic equilibrium, our calculations yield a total mass of warm H_2 gas at $(2.3 \pm 1.3) \times 10^{-5} M_{\text{Jup}}$ with a corresponding temperature of $635 \pm 94 \text{ K}$. This represents only a minor portion of the entire disk mass, which is constrained to be below $0.2 M_{\text{Jup}}$ due to the upper limit given by the non-detection of millimeter continuum emission.

We investigate the HI (7-6) line profile, a tracer for mass accretion. Through analyzing its profiles, we identify two main Gaussian components corresponding to the HI (7-6) and HI (11-8) atomic hydrogen recombination lines. These lines are in close proximity in wavelength, and the flux measurement used by Rigliaco et al. (2015) to assess accretion luminosity likely included contributions from both lines, rather than solely the HI (7-6) flux. When considering the combined flux, our estimated accretion luminosity $4.17 \times 10^{-10} M_{\odot} \text{ yr}^{-1}$ is consistent with the measurement based on $\text{H}\alpha$ emission by Fang et al. (2023). This suggests that the previous calibration of accretion luminosities may include the contribution from both lines. To further validate our hypothesis, future observations with higher spectral resolution than *Spitzer* data, such as those possible with the JWST, would be valuable. We compare our measured accretion rate with other YSOs, finding that this source has typical accretion properties.

To derive the physical properties of the accreting gas onto the star, we compare the HI (9-7) to HI (7-6) line ratio with the Case B recombination model (Baker & Menzel 1938; Storey & Hummer 1995) and the Kwan & Fischer (KF) model (Kwan & Fischer 2011). The Case B model provides limited constraints on the density and temperature, primarily because the lines are not optically thin, as assumed by this model. The KF model predicts that our measured line ratios are compatible with a range of temperatures from 5000 to 20000 K, and atomic hydrogen densities between 10^9 cm^{-3} and 10^{11} cm^{-3} . The atomic hydrogen density is well constrained by the KF model, while the temperature is not. However, the Case B model fails to reproduce the observed line fluxes and points to temperatures lower than 500 K, not physical for accretion flows.

In summary, our observation and analysis of one of the first JWST MIRI MRS data of a faint disk reveal new important results on the circumstellar gas properties. We observed several H_2 rotational lines, which constrain the warm emitting gas temperature and mass. Moreover, the measurement of HI recombination lines can be used to estimate the mass accretion rate onto the central source, while the comparison with theoretical models gives us information on the physical condition of the emitting gas. These results, while interesting on their own, demonstrate how further investigations of protoplanetary disks with JWST will dramatically enhance our understanding of these complex astronomical environments.

Acknowledgements. The following National and International Funding Agencies funded and supported the MIRI development: NASA; ESA; Belgian

Science Policy Office (BELSPO); Centre Nationale d'Études Spatiales (CNES); Danish National Space Centre; Deutsches Zentrum für Luftund Raumfahrt (DLR); Enterprise Ireland; Ministerio De Economía y Competividad; Netherlands Research School for Astronomy (NOVA); Netherlands Organisation for Scientific Research (NWO); Science and Technology Facilities Council; Swiss Space Office; Swedish National Space Agency; and UK Space Agency. R.F., T.H., and K.S. acknowledge support from the European Research Council under the Horizon 2020 Framework Program via the ERC Advanced Grant Origins 83 24 28. B.T. is a Laureate of the Paris Region fellowship program, which is supported by the Ile-de-France Region and has received funding under the Horizon 2020 innovation framework program and Marie Skłodowska-Curie grant agreement No. 945298. A.C.G. has been supported by PRIN-INAF MAIN-STREAM 2017 "Protoplanetary disks seen through the eyes of new-generation instruments" and from PRIN-INAF 2019 "Spectroscopically tracing the disk dispersal evolution (STRADE)" G.B. thanks the Deutsche Forschungsgemeinschaft (DFG) - grant 138 325594231, FOR 2634/2. E.v.D. acknowledges support from the ERC grant 101019751 MOLDISK and the Danish National Research Foundation through the Center of Excellence "InterCat" (DNRF150). I.K., A.M.A., and E.v.D. acknowledge support from grant TOP-1 614.001.751 from the Dutch Research Council (NWO). I.K. acknowledge funding from H2020-MSCA-ITN-2019, grant no. 860470 (CHAMELEON). O.A. and V.C. acknowledge funding from the Belgian F.R.S.-FNRS. O.A., V.C., and D.G. thank the Belgian Federal Science Policy Office (BELSPO) for the provision of financial support in the framework of the PRODEX Programme of the European Space Agency (ESA). D.R.L. acknowledges support from Science Foundation Ireland (grant number 21/PATH-S/9339). M.T. acknowledges support from the ERC grant 101019751 MOLDISK.

References

- Aikawa, Y., Cataldi, G., Yamato, Y., et al. 2021, *ApJS*, 257, 13
 Alcalá, J. M., Natta, A., Manara, C. F., et al. 2014, *A&A*, 561, A2
 Alcalá, J. M., Manara, C. F., Natta, A., et al. 2017, *A&A*, 600, A20
 An, D., Ramírez, S. V., Sellgren, K., et al. 2011, *ApJ*, 736, 133
 Anderson, D. E., Blake, G. A., Cleeves, L. I., et al. 2021, *ApJ*, 909, 55
 Andrews, S. M., Huang, J., Pérez, L. M., et al. 2018, *ApJ*, 869, L41
 Antoniucci, S., García López, R., Nisini, B., et al. 2011, *A&A*, 534, A32
 Antoniucci, S., García López, R., Nisini, B., et al. 2014, *A&A*, 572, A62
 Antoniucci, S., Podio, L., Nisini, B., et al. 2016, *A&A*, 593, L13
 Argyriou, I., Glasse, A., Law, D. R., et al. 2023, *A&A*, 675, A111
 Baker, J. G., & Menzel, D. H. 1938, *ApJ*, 88, 52
 Baldovin-Saavedra, C., Audard, M., Güdel, M., et al. 2011, *A&A*, 528, A22
 Baldovin-Saavedra, C., Audard, M., Carmona, A., et al. 2012, *A&A*, 543, A30
 Bally, J. 2016, *ARA&A*, 54, 491
 Barenfeld, S. A., Carpenter, J. M., Ricci, L., & Isella, A. 2016, *ApJ*, 827, 142
 Beck, T. L., McGregor, P. J., Takami, M., & Pyo, T.-S. 2008, *ApJ*, 676, 472
 Beck, T. L., Bary, J. S., & McGregor, P. J. 2010, *ApJ*, 722, 1360
 Bergin, E. A., Cleeves, L. I., Gorti, U., et al. 2013, *Nature*, 493, 644
 Berné, O., Habart, É., and Peeters, E., et al. 2022, *PASP*, 134, 054301
 Beuther, H., van Dishoeck, E. F., Tychoniec, L., et al. 2023, *A&A*, 673, A121
 Biazzo, K., Alcalá, J. M., Covino, E., et al. 2012, *A&A*, 547, A104
 Calahan, J. K., Bergin, E. A., Zhang, K., et al. 2021, *ApJS*, 257, 17
 Carmona, A., van der Plas, G., van den Ancker, M. E., et al. 2011, *A&A*, 533, A39
 Carpenter, J. M., Ricci, L., & Isella, A. 2014, *ApJ*, 787, 42
 Cleeves, L. I., Bergin, E. A., Qi, C., Adams, F. C., & Öberg, K. I. 2015, *ApJ*, 799, 204
 Dressing, C. D., & Charbonneau, D. 2015, *ApJ*, 807, 45
 Edwards, S., Kwan, J., Fischer, W., et al. 2013, *ApJ*, 778, 148
 Ercolano, B., & Pascucci, I. 2017, *R. Soc. Open Sci.*, 4, 170114
 Fang, M., van Boekel, R., Wang, W., et al. 2009, *A&A*, 504, 461
 Fang, M., Kim, J. S., Pascucci, I., et al. 2017, *AJ*, 153, 188
 Fang, M., Pascucci, I., Edwards, S., et al. 2023, *ApJ*, 945, 112
 Gaia Collaboration (Vallenari, A., et al.) 2023, *A&A*, 674, A1
 Gangi, M., Nisini, B., Manara, C. F., et al. 2023, *A&A*, 675, A153
 Gibb, E. L., Whittet, D. C. B., Schutte, W. A., et al. 2000, *ApJ*, 536, 347
 Goldsmith, P. F., & Langer, W. D. 1999, *ApJ*, 517, 209
 GRAVITY Collaboration (Wojtczak, J. A., et al.) 2023, *A&A*, 669, A59
 Güdel, M., Lahuis, F., Briggs, K. R., et al. 2010, *A&A*, 519, A113
 Gullbring, E., Hartmann, L., Briceño, C., & Calvet, N. 1998, *ApJ*, 492, 323
 Hartmann, L., Herczeg, G., & Calvet, N. 2016, *ARA&A*, 54, 135
 Herczeg, G. J., & Hillenbrand, L. A. 2008, *ApJ*, 681, 594
 Jönsson, H., Holtzman, J. A., Allende Prieto, C., et al. 2020, *AJ*, 160, 120
 Kama, M., Trapman, L., Fedele, D., et al. 2020, *A&A*, 634, A88
 Kwan, J., & Fischer, W. 2011, *MNRAS*, 411, 2383
 Labiano, A., Argyriou, I., Álvarez-Márquez, J., et al. 2021, *A&A*, 656, A57
 Lahuis, F., van Dishoeck, E. F., Blake, G. A., et al. 2007, *ApJ*, 665, 492
 Law, C. J., Crystian, S., Teague, R., et al. 2022, *ApJ*, 932, 114
 Lodato, G., Rampinelli, L., Viscardi, E., et al. 2023, *MNRAS*, 518, 4481
 Luhman, K. L., Herrmann, K. A., Mamajek, E. E., Esplin, T. L., & Pecaut, M. J. 2018, *AJ*, 156, 76
 Manara, C. F., Rosotti, G., Testi, L., et al. 2016, *A&A*, 591, L3
 Manara, C. F., Testi, L., Herczeg, G. J., et al. 2017, *A&A*, 604, A127
 Manara, C. F., Frasca, A., Venuti, L., et al. 2021, *A&A*, 650, A196
 Manara, C. F., Ansdell, M., Rosotti, G. P., et al. 2023, *ASP Conf. Ser.*, 534, 539
 McClure, M. K., Bergin, E. A., Cleeves, L. I., et al. 2016, *ApJ*, 831, 167
 Miotello, A., van Dishoeck, E. F., Kama, M., & Bruderer, S. 2016, *A&A*, 594, A85
 Miotello, A., Facchini, S., van Dishoeck, E. F., & Bruderer, S. 2018, *A&A*, 619, A113
 Miotello, A., Kamp, I., Birnstiel, T., Cleeves, L. C., & Kataoka, A. 2023, in *ASP Conf. Ser.*, 534, Protostars and Planets VII, ed. S. Inutsuka, Y. Aikawa, T. Muto, K. Tomida, & M. Tamura, 501
 Miret-Roig, N., Galli, P. A. B., Olivares, J., et al. 2022, *A&A*, 667, A163
 Molinari, S., Faustini, F., Testi, L., et al. 2008, *A&A*, 487, 1119
 Öberg, K. I., Guzmán, V. V., Walsh, C., et al. 2021, *ApJS*, 257, 1
 Panque-Carreño, T., Miotello, A., van Dishoeck, E. F., et al. 2023, *A&A*, 669, A126
 Pascucci, I., Gorti, U., Hollenbach, D., et al. 2006, *ApJ*, 651, 1177
 Pascucci, I., Hollenbach, D., Najita, J., et al. 2007, *ApJ*, 663, 383
 Pascucci, I., Herczeg, G., Carr, J. S., & Bruderer, S. 2013, *ApJ*, 779, 178
 Pascucci, I., Skinner, B. N., Deng, D., et al. 2023, *ApJ*, 953, 183
 Pound, M. W. and Wolfire, M. G. 2023, *AJ*, 165, 25
 Rieke, G. H., Wright, G. S., Böker, T., et al. 2015, *PASP*, 127, 584
 Rigliaco, E., Pascucci, I., Duchene, G., et al. 2015, *ApJ*, 801, 31
 Sabotta, S., Schlecker, M., Chaturvedi, P., et al. 2021, *A&A*, 653, A114
 Salyk, C., Herczeg, G. J., Brown, J. M., et al. 2013, *ApJ*, 769, 21
 Schwarz, K. R., Bergin, E. A., Cleeves, L. I., et al. 2016, *ApJ*, 823, 91
 Schwarz, K. R., Calahan, J. K., Zhang, K., et al. 2021, *ApJS*, 257, 20
 Smirnov-Pinchukov, G. V., Semenov, D. A., Akimkin, V. V., & Henning, T. 2020, *A&A*, 644, A4
 Storey, P. J., & Hummer, D. G. 1995, *MNRAS*, 272, 41
 Tabone, B., Bettoni, G., van Dishoeck, E. F., et al. 2023, *Nat. Astron.*, 7, 805
 Teague, R., Semenov, D., Guilloteau, S., et al. 2015, *A&A*, 574, A137
 Testi, L., Birnstiel, T., Ricci, L., et al. 2014, in *Protostars and Planets VI*, eds. H. Beuther, R. S. Klessen, C. P. Dullemond, & T. Henning (Tucson: University of Arizona Press), 339
 Thi, W. F., van Dishoeck, E. F., Blake, G. A., et al. 2001, *ApJ*, 561, 1074
 Trapman, L., Miotello, A., Kama, M., van Dishoeck, E. F., & Bruderer, S. 2017, *A&A*, 605, A69
 Trapman, L., Zhang, K., van't Hoff, M. L. R., Hogerheijde, M. R., & Bergin, E. A. 2022, *ApJ*, 926, L2
 van Boekel, R., Güdel, M., Henning, T., Lahuis, F., & Pantin, E. 2009, *A&A*, 497, 137
 van den Ancker, M. E., Wesselius, P. R., & Tielens, A. G. G. M. 2000, *A&A*, 355, 194
 van Dishoeck, E. F., Wright, C. M., Cernicharo, J., et al. 1998, *ApJ*, 502, L173
 Venuti, L., Bouvier, J., Flaccomio, E., et al. 2014, *A&A*, 570, A82
 Venuti, L., Stelzer, B., Alcalá, J. M., et al. 2019, *A&A*, 632, A46
 Wells, M., Pel, J. W., Glasse, A., et al. 2015, *PASP*, 127, 646
 Wright, G. S., Wright, D., Goodson, G. B., et al. 2015, *PASP*, 127, 595
 Wright, G. S., Rieke, G. H., Glasse, A., et al. 2023, *PASP*, 135, 048003
 Yoshida, T. C., Nomura, H., Tsukagoshi, T., Furuya, K., & Ueda, T. 2022, *ApJ*, 937, L14
 Zhang, K., Booth, A. S., Law, C. J., et al. 2021, *ApJS*, 257, 5

¹ Max Planck Institut für Astronomie (MPIA), Königstuhl 17, 69117 Heidelberg, Germany
 e-mail: franceschi@m pia.de

² Université Paris-Saclay, CNRS, Institut d'Astrophysique Spatiale, 91405 Orsay, France

³ INAF – Osservatorio Astronomico di Capodimonte, Salita Moiarriello 16, 80131 Napoli, Italy

⁴ Dublin Institute for Advanced Studies, 31 Fitzwilliam Place, D02 XF86 Dublin, Ireland

⁵ Max-Planck Institut für Extraterrestrische Physik (MPE), Giessenbachstr. 1, 85748 Garching, Germany

- ⁶ Leiden Observatory, Leiden University, 2300 RA Leiden, The Netherlands
- ⁷ Kapteyn Astronomical Institute, University of Groningen, The Netherlands
- ⁸ STAR Institute, Université de Liège, Allée du Six Août 19c, 4000 Liège, Belgium
- ⁹ Dept. of Astrophysics, University of Vienna, Türkenschanzstr. 17, 1180 Vienna, Austria
- ¹⁰ ETH Zürich, Institute for Particle Physics and Astrophysics, Wolfgang-Pauli-Str. 27, 8093 Zürich, Switzerland
- ¹¹ Department of Astronomy, Stockholm University, AlbaNova University Center, 10691 Stockholm, Sweden
- ¹² Department of Astrophysics/IMAPP, Radboud University, PO Box 9010, 6500 GL Nijmegen, The Netherlands
- ¹³ SRON Netherlands Institute for Space Research, Niels Bohrweg 4, 2333 CA Leiden, The Netherlands
- ¹⁴ Kapteyn Astronomical Institute, Rijksuniversiteit Groningen, Postbus 800, 9700AV Groningen, The Netherlands
- ¹⁵ Institute of Astronomy, KU Leuven, Celestijnenlaan 200D, 3001 Leuven, Belgium
- ¹⁶ Centro de Astrobiología (CAB), CSIC-INTA, ESAC Campus, Camino Bajo del Castillo s/n, 28692 Villanueva de la Cañada, Madrid, Spain
- ¹⁷ LESIA, Observatoire de Paris, Université PSL, CNRS, Sorbonne Université, Université de Paris, 5 place Jules Janssen, 92195 Meudon, France
- ¹⁸ UK Astronomy Technology Centre, Royal Observatory Edinburgh, Blackford Hill, Edinburgh EH9 3HJ, UK
- ¹⁹ Université Paris-Saclay, Université Paris Cité, CEA, CNRS, AIM, 91191 Gif-sur-Yvette, France