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Discovery of a Rare Group of Dwarf Galaxies in the Local Universe

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Abstract

We report the discovery of a rare isolated group of five dwarf galaxies located at $z = 0.0086$ ($D = 36$ Mpc). All member galaxies are star forming, blue, and gas rich, with $g - r$ indices ranging from 0.2 to 0.6 mag, and two of them show signs of ongoing mutual interaction. The most massive member of the group has a stellar mass that is half of the Small Magellanic Cloud stellar mass, and the median stellar mass of the group members is $7.87 \times 10^7 M_{\odot}$. The derived total dynamical mass of the group is $M_{\text{dyn}} = 6.02 \times 10^{10} M_{\odot}$, whereas its total baryonic mass (stellar + H I) is $2.6 \times 10^9 M_{\odot}$, which gives us the dynamical to baryonic mass ratio of 23. Interestingly, all galaxies found in the group are aligned along a straight line in the plane of the sky. The observed spatial extent of the member galaxies is 154 kpc, and their relative line-of-sight velocity span is within 75 km s^{-1} . Using the spatially resolved optical spectra provided by the Dark Energy Spectroscopic Instrument early data release, we find that three group members share a common rotational direction. With these unique properties of the group and its member galaxies, we discuss the possible importance of such a system in the formation and evolution of dwarf galaxy groups and in testing the theory of large-scale structure formation.

Unified Astronomy Thesaurus concepts: [Interacting galaxies \(802\)](#); [Dwarf galaxies \(416\)](#); [Galaxy groups \(597\)](#)

1. Introduction

The Lambda cold dark matter (Λ CDM) model is successful in explaining the large-scale structure of the Universe, but it encounters difficulties with explaining many aspects of dwarf galaxies (S. D. M. White & M. J. Rees 1978; C. S. Frenk & S. D. M. White 2012; J. S. Bullock & M. Boylan-Kolchin 2017). According to the Λ CDM theory, even low-mass halos possess detectable substructures, which implies that dwarf galaxies ought to have subhalos and form individual groups (J. Diemand et al. 2008; C. Wheeler et al. 2015). Consequently, simulations predict mergers involving only dwarf galaxies (A. Deason et al. 2014).

However, statistical analysis of both observational and simulation data estimated that fewer than 5% of dwarf galaxies have close companions, and there is less than a 0.004% chance that a dwarf galaxy would be part of a quad group (with the number of members equal to 4) of dwarf galaxies (S. Stierwalt et al. 2017; G. Besla et al. 2018). Moreover, very few of these close companion dwarf galaxies could undergo a merger.

The interaction between the Large Magellanic Cloud (LMC) and the Small Magellanic Cloud (SMC) is a unique example that has been extensively studied in both observations and theoretical models (M. E. Putman et al. 1998; G. Besla et al. 2010, 2012; K. Glatt et al. 2010; N. Kallivayalil et al. 2013, 2018; E. D'Onghia & A. J. Fox 2016). In addition, recent deep observation has revealed several new satellites around the LMC–SMC system (A. Drlica-Wagner et al. 2015; S. E. Koposov et al. 2015, 2018). The discovery of the LMC group has revitalized the importance of LMC–SMC on the evolution of the Milky Way (MW) satellite system and the MW itself; particularly, it is hypothesized that a fraction of the MW satellites is accreted as part of the LMC–SMC

group, and that helps to explain the observed overdensity of globular clusters and satellite dwarfs in the direction of the LMC (S.-J. Yoon & Y.-W. Lee 2002; L. V. Sales et al. 2017; D. Erkal & V. A. Belokurov 2020). Understanding the origin of the LMC–SMC group in the vicinity of the MW has been an active area of research (M. Cautun et al. 2019; T. A. Evans et al. 2020; E. Vasiliev 2023).


In this Letter, we report the discovery of a system of interacting dwarf galaxies within a group of at least five confirmed members. The most massive galaxy in this group has a stellar mass of $M_{*} = 2.7 \times 10^8 M_{\odot}$, which is half of the SMC stellar mass. This Letter is organized as follows. In Section 2, we introduce the system and provide the details of data analysis. In Section 3, we discuss our findings, and we summarize them in Section 4.

2. Interacting Dwarfs in a Group Environment

2.1. Identification

Our primary focus is to conduct an in-depth analysis of the merging systems of dwarf galaxies in various environments. To achieve this, we have conducted a comprehensive search for such objects by visually inspecting color images of dwarf galaxies in the local volume ($z < 0.02$) provided by the Sloan Digital Sky Survey (SDSS) and Legacy survey (H. Aihara et al. 2011; A. Dey et al. 2019). A comprehensive list of such objects is published in a catalog of merging dwarf galaxies (S. Paudel et al. 2018). In this work, we present a particular case of dwarf galaxy interaction, identifying their unique environment. In particular, we find that dwarf galaxies interact inside a group of dwarf galaxies.

We have identified a unique group of five star-forming dwarf galaxies located in a near-isolated environment at a distance of $D = 36$ Mpc.⁶ We find no massive galaxies ($M_{*} > 10^{10} M_{\odot}$)

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⁶ Based on Hubble flow with a mean redshift of the group $z = 0.0086$.

within a 700 kpc sky-projected radius and line-of-sight radial velocity range of $\pm 700 \text{ km s}^{-1}$ as we have searched in the NASA Extragalactic Database.⁷ These values typically correspond to twice the projected viral radius and velocity dispersion of a group with a central galaxy of stellar mass $\sim 10^{10} M_{\odot}$, respectively (D. Makarov & I. Karachentsev 2011).

In Figure 1, we show the color image cutout of the system, which we have obtained from the Legacy survey viewer.⁸ We see that member dwarf galaxies are arranged in almost a vertical straight line spanning from north to south, where the most massive galaxy is a face-on spiral, D2, which has a B -band absolute luminosity of -16.71 mag (similar to that of the SMC; $M_B = -16.79 \text{ mag}$). As the color image reveals, all member galaxies are blue and star forming. The spatial span of the group (largest sky-projected separation between member galaxies) is 154 kpc, and the velocity dispersion is 33 km s^{-1} . A principal component analysis was employed to fit a planar structure, yielding a semimajor axis of 57 kpc and a semiminor axis of 4 kpc, which results in an axis ratio of 0.06.

In the vicinity of the most massive galaxy of the group, D2, we identify a pair of dwarf galaxies (D3 and D4) located at a sky-projection distance of 10.3 kpc southward with a relative line-of-sight velocity of 30 km s^{-1} . A careful visual inspection of the optical images⁹ of the pair reveals their disturbed morphology, showing a sign of tidal interaction where the tidal tails are elongated in the opposite direction of the interaction. The relative line-of-sight velocity between the pair is 17 km s^{-1} . The B -band absolute luminosities of D3 and D4 are -14.74 and -13.98 mag , respectively, and their stellar mass ratio is 3.

2.2. Data Analysis

This work benefits from the substantial multiwavelength data available in public archives, which allowed us to perform the required detailed measurements of the physical properties of the group and its member galaxies. Particularly, Dark Energy Spectroscopic Instrument (DESI¹⁰) early data release (EDR; DESI Collaboration et al. 2024) spectroscopic data enabled us to confirm its group member with radial velocity measurement. We used the Legacy survey imaging data to calculate optical g -, r -, and z -band magnitudes and Strasbourg Astronomical Data Center¹¹ catalog services to fetch other physical properties of the member galaxies.

We downloaded g -, r -, and z -band images from the Legacy survey database, and the observation was made through the Beijing–Arizona Sky Survey (H. Zou et al. 2017). We performed aperture photometry of the member galaxy using the similar method we had implemented in our main catalog paper (S. Paudel et al. 2018). Finally, the B -band magnitudes were derived from the g and r magnitudes using the filter conversion equation, $B = g + 0.227(g - r) - 0.337$.¹² The B -band magnitude of member galaxies ranges from -16.71 to -13.98 mag , the former being similar to the SMC and the latter being similar to the Fornax dwarf spheroidal galaxy. As

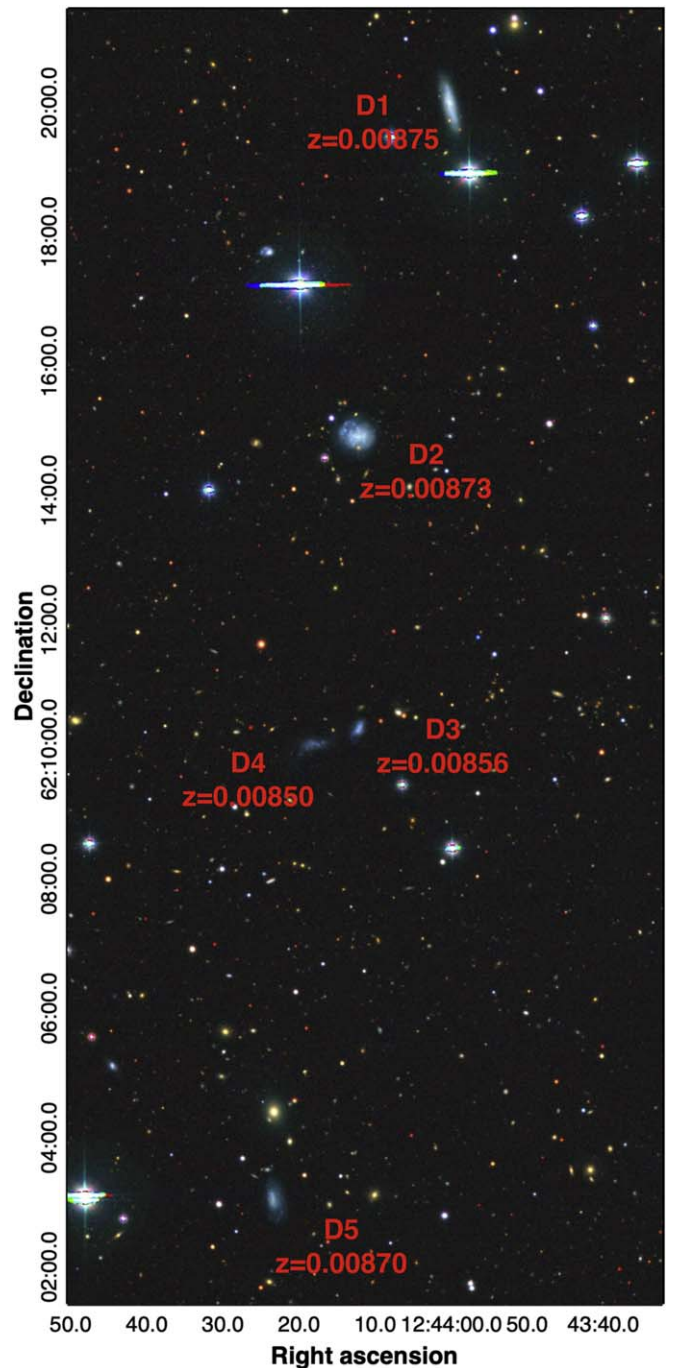


Figure 1. The $g - r - z$ combined color image cutout of the system obtained from the Legacy survey viewer page. The five member dwarf galaxies are designated, and we also provide their redshift information. All five members are arranged in almost a vertical straight line spanning from north to south. The most massive galaxy is a face-on spiral galaxy, D2, which has a B -band absolute luminosity of -16.71 mag .

expected for star-forming dwarf galaxies, the member dwarf galaxies have $g - r$ indices ranging from 0.2 to 0.5 mag.

We obtained infrared band W1 and W2 magnitudes from the Wide-field Infrared Survey Explorer (WISE; E. L. Wright et al. 2010) all-sky survey catalog (R. M. Cutri et al. 2021) and derived the stellar mass using an empirical formula from M. Eskew et al. (2012). The brightest central galaxy has a stellar mass of $2.75 \times 10^8 M_{\odot}$, which is nearly half of the SMC stellar mass (A. W. McConnachie 2012). The least massive

⁷ <https://ned.ipac.caltech.edu/>

⁸ <https://www.legacysurvey.org/viewer>

⁹ To gain signal in the low-surface brightness region tidal tails, we coadded the Legacy g -, r -, and z -band images.

¹⁰ <https://data.desi.lbl.gov/>

¹¹ <https://cds.unistra.fr>

¹² <https://www.sdss3.org/dr8/algorithms/sdssUBVRITransform.php>

member galaxy, D4, has a stellar mass of $1.47 \times 10^7 M_{\odot}$, similar to that of Fornax dwarf spheroidal, and the median of stellar masses of the five member galaxies is $7.87 \times 10^7 M_{\odot}$.

Using the Five-hundred-meter Aperture Spherical Radio Telescope (FAST¹³), the FAST All Sky HI survey (FASHI; C.-P. Zhang et al. 2024) has covered the sky around our group. In our galaxy group region, three sources are detected in 21 cm emission line observation with a radial velocity similar to the member galaxies: D1, D2, and D3. For D5, there is no 21 cm radio emission detection in FASHI, and for D3 and D4, the detection is combined, likely due to the separation between them being smaller than FASHI 2'.9 beam size. We derive total neutral hydrogen (HI) mass from 21 cm emission fluxes using a formula:

$M_{\text{HI}}(M_{\odot}) = 2.36 \times 10^5 D^2 \int F dV$, where $\int F dV$ is integrated flux in Jy km s^{-1} , and D is distance to galaxy in Mpc. For the undetected galaxy, D5, we estimated the limit of detection using the formula from G. Gavazzi et al. (2008):

$M_{\text{HI,lim}}(M_{\odot}) = \text{SNR} \times \text{rms} \times W_{50} \times D^2 \times 2.36 \times 10^5$, with the minimum signal-to-noise ratio of 1, $W_{50} = 50 \text{ km s}^{-1}$ of our dwarf galaxies, the median noise rms of the survey as $\text{rms} = 1.5 \text{ mJy}$, and the distance $D = 36 \text{ Mpc}$.

Fortunately, the sky region is also covered by DESI EDR. We identify all member galaxies targeted by the DESI spectroscopic observation, including multiple observations for D1, D2, and D5. Indeed, this helps us to confirm the membership of these galaxies with radial velocity measurement. We find that the overall range of radial velocity of the member galaxies ranges from 2550 to 2624 km s^{-1} with a standard deviation $\sigma_v = 33 \text{ km s}^{-1}$. The optical spectra of these galaxies strongly show H_{α} emission, as shown in Figure 2. We estimated the oxygen abundance, $12 + \log(\text{O}/\text{H})$, using a combination of line ratios $\text{H}_{\alpha}/[\text{N II}]$ and $[\text{O III}]/\text{H}_{\beta}$ (R. A. Marino et al. 2013). We list the photometric and physical parameters of member galaxies in Table 1.

The total dynamical mass of the system is calculated as the total projected mass of the group (J. Heisler et al. 1985) using the following equation:

$$M = \frac{32}{\pi} \frac{1}{G(N-3/2)} \sum_i^N R_{p,i} \Delta v_i^2, \quad (1)$$

where v_i and $R_{p,i}$ are, respectively, radial velocity and projected distance of the i th galaxy relative to the system's center. N is the total number of the members of galaxies. Using this equation, we obtained the total dynamical mass of the system $M_{\text{dyn}} = 6.02 \times 10^{10} M_{\odot}$, whereas the total baryonic mass (stellar + HI) of the group is $2.6 \times 10^9 M_{\odot}$, which gives us the dynamical to baryonic mass ratio of 23.

The 3D velocity dispersion (σ_{3D}) is calculated using line-of-sight velocities of all the group members and corrected for potential transverse motion using the following equation: $\sigma_{3D} = \sqrt{3} \times \sqrt{\langle v^2 \rangle - \langle v \rangle^2}$, where v is line-of-sight velocities. We obtained the value of $\sigma_{3D} = 51 \text{ km s}^{-1}$. Finally, we calculated the minimum mass ratio required for the group to be a gravitationally bound structure. For this, we used the escape velocity of the group equal to σ_{3D} and found that the mass required for the bound group is $2.99 \times 10^{10} M_{\odot}$, which is half of the total dynamical mass and approximately 10 times the total Baryonic mass.

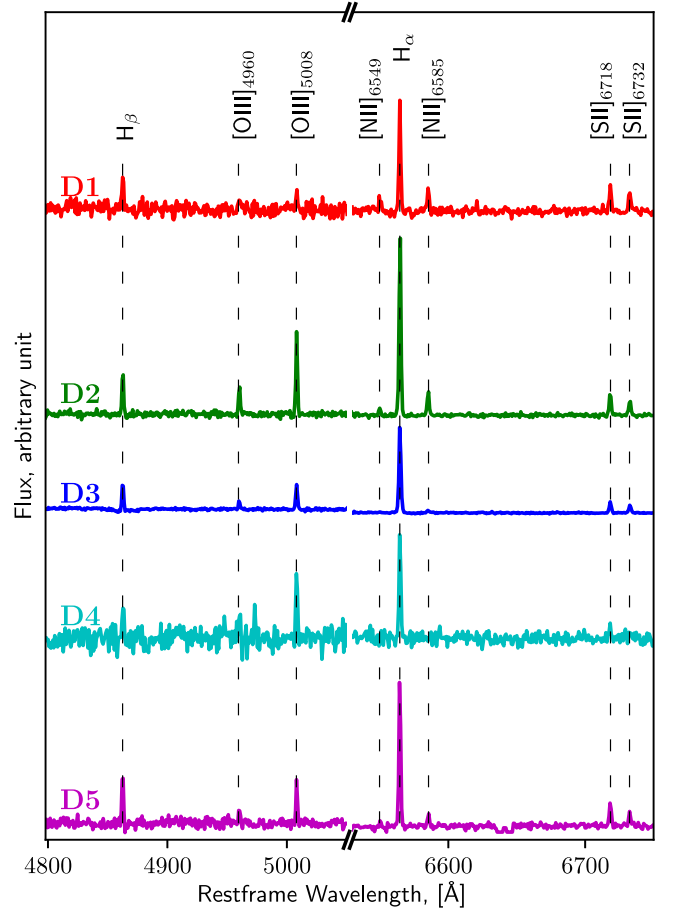


Figure 2. The DESI EDR optical spectra of the central region of member galaxies. The observed spectra are shifted to the rest-frame wavelength and smoothed with a three-pixel Gaussian kernel. The major emission lines are marked with vertical dash lines.

3. Discussion

In this work, we presented a unique group of five dwarf galaxies in a relatively isolated environment. All member galaxies of the group are star forming and blue. Despite searching the SDSS and DESI EDR spectroscopic databases within a 500 kpc radius and a relative line-of-sight velocity range of $\pm 500 \text{ km s}^{-1}$, no additional members were identified. The SDSS has targeted D1, D2, and D5, while the DESI has targeted all five members. Indeed, this may give a biased view, potentially overlooking non-star-forming dwarf galaxies. Notably, measuring the redshift of star-forming dwarf galaxies is considerably easier and more effective due to the presence of emission lines compared to non-star-forming dwarf galaxies characterized by absorption-dominated spectra. It is important to note that this analysis is based solely on the five confirmed star-forming dwarf galaxy members. There is a possibility that additional member dwarfs could be discovered through a dedicated, in-depth spectroscopic survey of the area.

3.1. Comparison to Previous Work

S. Stierwalt et al. (2017); hereafter S17) studied seven isolated compact groups of galaxies of redshift range from 0.02 to 0.05 ($D \simeq 80\text{--}200 \text{ Mpc}$). Compared to the S17 groups, our group represents a much closer system at $D = 36 \text{ Mpc}$. In the S17 sample, only one group has five members, which is

¹³ <https://fast.bao.ac.cn>

Table 1
Physical Properties of the Member Galaxies

R. A.	Decl.	M_B	v	R	g	r	z	$g-r$	W1	W2	M_*	Z	M_{HI}	
hh:mm:ss	dd:mm:ss	(mag)	(km s^{-1})	kpc	(mag)	(mag)	(mag)	(mag)	(mag)	(mag)	$\log(M_\odot)$	(dex)	$\log(M_\odot)$	
D1	12:43:59	+62:20:00	-16.42	2618	49	15.97	15.45	15.25	0.52	14.76	14.58	8.29	8.36	8.57
D2	12:44:12	+62:14:51	-16.71	2624	00	15.71	15.28	15.18	0.43	14.56	14.47	8.44	8.33	9.02
D3	12:44:12	+62:10:19	-14.74	2567	43	17.74	17.50	17.50	0.24	16.29	16.07	7.65	8.23	8.62 ^a
D4	12:44:20	+62:09:58	-13.98	2550	44	18.49	18.22	18.30	0.27	17.31	16.99	7.17	8.00	...
D5	12:44:23	+62:03:06	-15.26	2610	99	17.16	16.75	16.78	0.41	16.12	16.10	7.87	8.03	< 6.98 ^b

Notes. R is a sky-projected distance of the member galaxies from the most massive member, D2. A typical error on the magnitudes is 0.01 mag in all-optical band photometry. The stellar masses are derived from infrared band W1 and W1 magnitudes from the WISE all-sky survey. Z represents the emission line metallicity, $12 + \log(\text{O}/\text{H})$, which we have derived from the DESI EDR spectra of the central region of galaxies. The total neutral hydrogen masses are derived from the cataloged values of the 21 cm emission line observation of the FASHI survey.

^a Since the FASHI survey uses single-dish observations with a beam size of $2''.9$, the interacting pairs D3 and D4 are unresolved, and the listed value of HI mass is combined.

^b D5 is not detected in the FASHI survey; we provide the upper limit of detection based on the survey depth; see the text for details.

also the most massive group with a total dynamical mass of $1.12 \times 10^{12} M_\odot$, and there are two quads and three triplets. Among the three triplets, the group 1440 + 14 is located near a relatively bright ($M_B < -19$ mag) galaxy (at a distance of 230 kpc with a radial velocity separation of 90 km s^{-1}). We, therefore, drop this group from our comparison sample. In Figure 3, we show a phase-space diagram and compare our group to the S17 groups, where the solid and dash lines represent the escape velocity curve for a halo mass of 10^{10} and $10^{11} M_\odot$, respectively. It is interesting to note that the member galaxies of our group fall well below the escape velocity curve for a halo mass of $10^{10} M_\odot$, confirming the bound structure. In fact, a few members of the most massive group of S17 are outside of the escape velocity curve of the halo mass of $10^{11} M_\odot$ (red and green dots), but reported dynamical masses of these groups are on the order of $10^{12} M_\odot$. One main difference between the S17 groups and our group is that the S17 groups are compact, with the largest group having a size of 80.5 kpc, whereas our group has 154 kpc. This is probably due to the selection criteria of S17: they deliberately selected a sample of compact groups. On the other hand, S17 groups have a significantly larger velocity of dispersion compared to our group. The largest S17 group (which also has five members) has a velocity of dispersion of $> 130 \text{ km s}^{-1}$, whereas our group has a velocity dispersion of 33 km s^{-1} .

In Figure 4, we show the baryonic Tully–Fisher relation (BTFR) plotted as total baryonic mass against σ_{3D} of dwarf galaxy groups. For comparison, we plotted a well-studied sample of galaxies of F. Lelli et al. (2016; hereafter L16). It is evident that the majority of S17 groups fall on the right side of the BTFR defined by L16 galaxies, indicating an excess of baryon mass in these systems. We also highlighted the triplet groups of S17 (with blue dots), which have nearly similar total baryon mass. On the other hand, our group stands out, being a significant outlier on the left side of the BTFR; in fact, it nearly falls along the line of the cosmic average baryon fraction, indicating that there are no missing baryons in this group, and all baryons from the big bang nucleosynthesis are locked inside its member dwarf galaxies.

3.2. Uniqueness, Chance, or Real?

The ongoing interaction between the LMC and SMC is indeed happening in a group environment in the vicinity of the

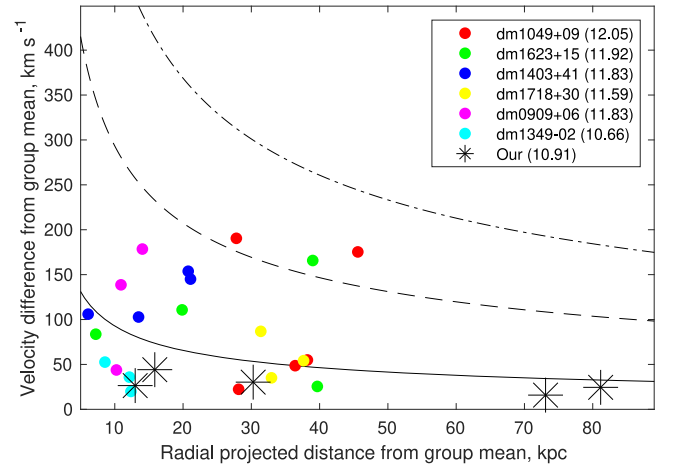


Figure 3. The position of the member galaxies (black asterisk) in a phase-space diagram. For comparison, we use the S17 groups, where different colors represent different groups, and their respective group dynamical masses ($\log(M_{\text{dyn}}/M_\odot)$) are listed in the legend. The solid, dashed, and dashed-dotted lines represent the escape velocity curve for a halo mass of 10^{10} , 10^{11} , and $10^{11.5} M_\odot$, respectively.

MW (M. E. Putman et al. 1998). However, satellite pair interaction in a group environment is rare (A. S. G. Robotham et al. 2012; S. Paudel & C. Sengupta 2017; S. Paudel et al. 2020). In our merging dwarf galaxy catalog, only 5% of dwarf–dwarf galaxy interactions happen in group environments. In our group of this study, we also find that the dwarf galaxies, D3 and D4, of stellar masses similar to Fornax dwarf spheroidal are interacting in the vicinity of the group’s most massive host. The LMC–SMC pair has been studied in great detail and is regarded as an essential laboratory for studying the evolution of dwarf galaxies and the formation of anisotropic distribution of satellite galaxies around their host. Recent analyses have shown that the LMC–SMC pair is not only interacting but also constitutes a group of dwarf galaxies, which includes at least five other MW satellites (D. Zhang et al. 2019; E. Patel et al. 2020).

One of the unique features of our group is the near-straight-line distribution of the member galaxies in sky projection. To confirm whether this is a chance projection or real is beyond the observational limit of this work. In particular, we do not know the distances of individual galaxies from which we could

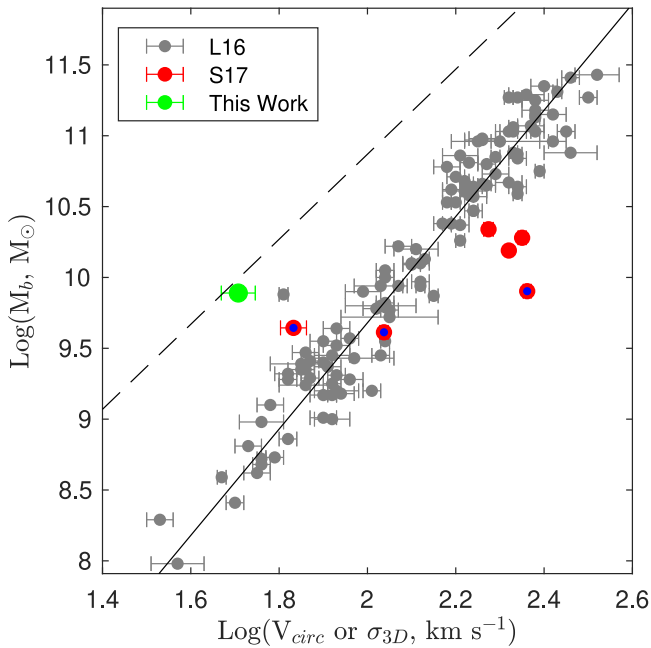


Figure 4. The BTFR plotted as total baryonic mass against σ_{3D} of dwarf galaxy groups. For comparison, we plot the total baryonic mass against the rotation velocity of galaxies from L16. The S17 sample is shown by red circles, where the blue dots represent the triplet group. The best-fit value of L16 data shown by a solid line and the dashed line is derived from the cosmic baryon fraction using the equation $M_{200}f_b \propto M_{200}^3 f_V^3$, with a value of $f_b = 0.17$ and $f_V = 1$. Our group (green circles) stands out, being a significant outlier on the left side of the BTFR, and nearly falls along the line of the cosmic average baryon fraction.

make a 3D picture of the group. In fact, the satellites around our MW galaxy are distributed in an extremely thin plane with an axis ratio of 0.29 (M. S. Pawłowski et al. 2012, 2014), and there are a dozen similar structures that have been identified in various galaxy groups (N. I. Libeskind et al. 2019). Extending this phenomenon, we may have observed similar structures on a much smaller scale in contrast to previously discovered planar satellite systems around the massive host with a stellar mass of $\sim 10^{11} M_\odot$.

In addition to their planar distribution, we also find that the member galaxies in this group have similar rotational directions. In Figure 5, we show a zoom-in view of the member galaxies, where the relative line-of-sight velocity directions with respect to the galaxy centers are marked with blue and red symbols for redshifted and blueshifted velocities, respectively. Thanks to the DESI, the EDR provides multiple spectroscopic observations for D1, D2, and D5 from different parts of galaxies, which allow us to tentatively estimate their direction of rotation. D1 and D5 are fairly edge-on galaxies, and with the radial velocity measurement from different ends of galaxies, we can clearly determine their rotational direction. D2, on the other hand, has a nearly face-on orientation. While we can still estimate its rotational direction using spatially resolved spectroscopy, the data statistics are, unfortunately, quite limited. We noticed that D1, D2, and D5 have a common rotational direction in that the north parts of these galaxies are redshifted, and the lower south parts are blueshifted.

4. Summary

In this work, we have studied a rare group of dwarf galaxies consisting of at least five confirmed members, two of which are

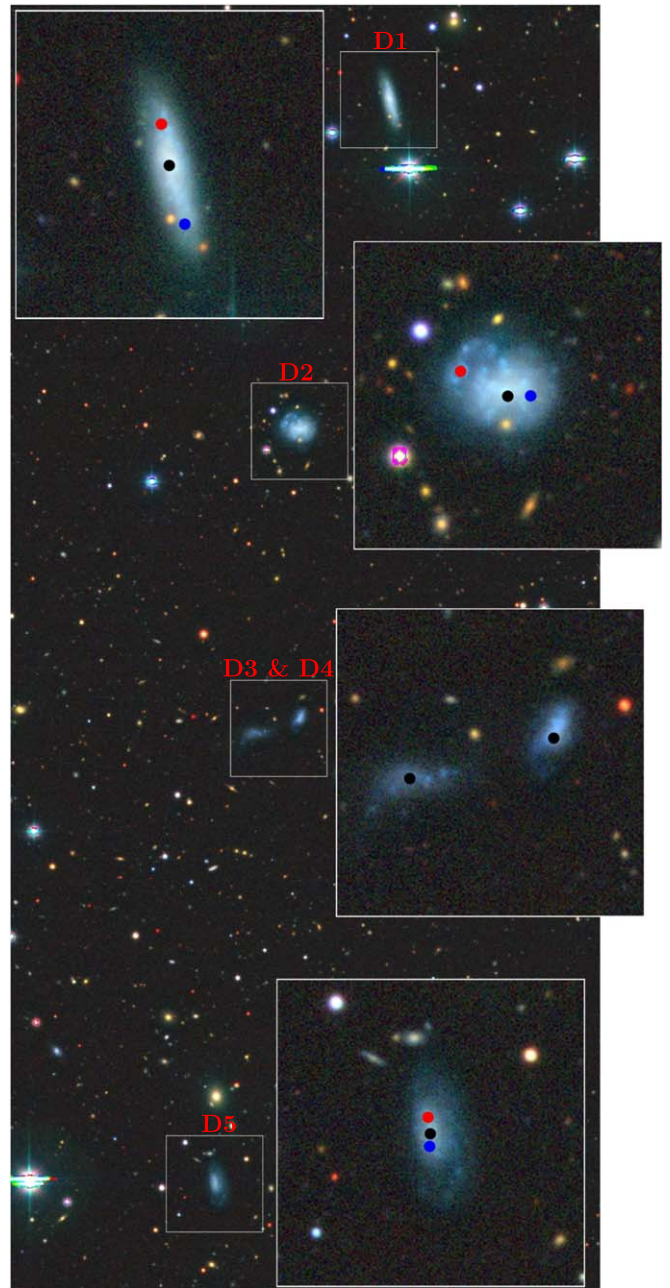


Figure 5. A zoomed-in view of the member galaxies. The $1' \times 1'$ square box is zoomed in to a better view, where we marked the central positions with the black dots. The red and blue dots represent the redshifted and blueshifted parts of the galaxy with respect to their center, respectively.

interacting. The member galaxies are star forming and blue. The brightest member galaxy has a stellar mass of $2.75 \times 10^8 M_\odot$, which is nearly half of SMC's stellar mass. The median stellar mass of group members is $7.87 \times 10^7 M_\odot$. We find that all galaxies in the group are arranged nearly in a straight line, forming an exceptionally thin planar structure. The observed spatial span of the member galaxy (distance from one end to another) is 154 kpc, and the maximum difference in line-of-sight velocity between them is 75 km s^{-1} . The derived total dynamical of the system is $M_{\text{dyn}} = 6.02 \times 10^{10} M_\odot$, whereas the total baryonic mass (stellar + H I) of the group is $2.6 \times 10^9 M_\odot$, which gives us the dynamical to baryonic mass ratio of 23. Using σ_{3D} , we derived the minimum mass ratio

required for the group to be a gravitationally bound structure, which comes out to be $2.99 \times 10^{10} M_{\odot}$. We also found that our group is a significant outlier from the standard BTFR defined by galaxies, indicating that all baryons of the group are locked inside the member galaxies.

We have identified a distinct group of dwarf galaxies, with all five members aligned along a straight line in the celestial plane and three sharing a common rotational direction. The significance of the observed features and whether they are coincidental or genuine remains uncertain. It has been argued that observations of coherent structures like a plane of satellites or rotating satellite systems are against our well-known Λ CDM cosmology, which assumes the hierarchical buildup of large-scale structures (N. I. Libeskind et al. 2015; M. S. Pawlowski et al. 2017; M. S. Pawlowski 2021). Further investigation of this system, both observationally and theoretically, could shed light on the prevalence of such configurations in the Λ CDM Universe and potentially reveal the accuracy of our current understanding of large-scale structure formation. Moreover, subsequent high-resolution 21 cm radio observations may provide detailed insights into the internal kinematics of each dwarf galaxy, potentially confirming their individual rotational characteristics.

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